

IXPE Observations of the Blazar Mrk 501 in 2022: A Multiwavelength View

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ABSTRACT

The blazar Markarian 501 (Mrk 501) was observed on three occasions over a 4-month period between 2022 March and 2022 July with the *Imaging X-ray Polarimetry Explorer* (IXPE). In this paper, we report for the first time on the third IXPE observation, performed between 2022 July 9 and 12, during which IXPE detected a linear polarization degree of $\Pi_X = 6 \pm 2$ per cent at a polarization angle, measured east of north, of $\Psi_X = 143^\circ \pm 11^\circ$ within the 2–8 keV X-ray band. The X-ray polarization angle and degree during this observation are consistent with those obtained during the first two observations. The chromaticity of the polarization across radio, optical, and X-ray bands is likewise consistent with the result from the simultaneous campaigns during the first two observations. Furthermore, we present two types of models to explain the observed spectral energy distributions (SEDs) and energy-resolved polarization: a synchrotron self-Compton model with an anisotropic magnetic field probability distribution in the emitting volume, as well as an energy-stratified shock model. Our results support both the shock scenario as well as support that small levels of magnetic field anisotropy can explain the observed polarization.

Key words: Galaxies: active – Galaxies: jets – BL Lacertae objects: individual – Mrk501

1 INTRODUCTION

Blazars are a class of active galactic nuclei (AGN) characterized by their flux and spectral variability, high polarization, and collimated relativistic jets of plasma. The observed emission from such objects arises from the collimated plasma outflow (jet), which emits over a wide range of the electromagnetic spectrum and is oriented at small angles ($\leq 10^\circ$) with respect to the line of sight of the observer. Multi-wavelength blazar emissions are characterized by broad, double-peaked spectral energy distributions (SEDs). Synchrotron emission from relativistic electrons in the jet results in a low-energy component extending from radio to the optical, UV, or X-ray bands, at which point the high-energy component starts to dominate the emission. The high-energy peak consists either of photons Compton-scattered from the same electrons producing the synchrotron emission, or of synchrotron emission from secondary particles created by high-energy protons (Sol & Zech 2022).

Mrk 501 is a nearby ($z = 0.034$) high synchrotron peaked (HSP) blazar, with its low-energy emission component peaking in the X-ray band, and its high-energy component peaking in the GeV - TeV energy range (Pian et al. 1998; Aharonian et al. 1999; Krawczynski et al. 2000). The recently-launched *Imaging X-ray Polarimetry Explorer* (IXPE) satellite has allowed us to measure, for the first time, the polarization degree and direction of the synchrotron X-rays in the 2 - 8 keV energy band (Weisskopf et al. 2022). These measurements help constrain the uniformity and direction of the magnetic fields experienced by the high-energy electrons emitting the X-rays.

IXPE observed Mrk 501 three times in 2022. The first two observations (March 8 - 10, 2022, 100 ksec, and March 26 - 28, 2022, 86 ksec) have been reported in a previous paper (Liodakis et al. 2022). Those two observations revealed X-ray polarization degrees of $\Pi_X = 10 \pm 2$ per cent and $\Pi_X = 11 \pm 2$ per cent, at electric vector position angles of $\Psi_X = 134^\circ \pm 5^\circ$ and $\Psi_X = 115^\circ \pm 4^\circ$, respectively. Liodakis et al. (2022) interpreted the strong increase of the linear polarization degree, from $\Pi_O \approx 4$ per cent in the optical band to $\Pi_X \approx 10$ per cent in the X-ray band, as evidence for electron acceleration by shocks. In such models, X-rays are emitted by higher-energy electrons closer to the acceleration region, while the optical emission comes from a larger region with a less uniform magnetic field that thereby yields a lower polarization degree for the emission. The energy-stratified electron populations can thus explain the observed multiwavelength polarization characteristics of Mrk 501. On the other hand, it is possible that models with particle acceleration from magnetic reconnection might lead to faster and/or greater polarization swings in X-rays, especially as compared to energy-stratified shocks, with lower net polarization degrees in the X-ray band owing to averaging over different polarization directions (Zhang et al. 2021; Di Gesu et al. 2022a), a result contrary to the observational findings. However, the results depend on the particular model, and reconnection cannot yet be ruled out. Finally, as mentioned above, the X-rays from Mrk 501 were found to be polarized in a direction approximately parallel to the jet axis angle of $\Psi_{jet} = 120^\circ \pm 12^\circ$ (Weaver et al. 2022). The polarization direction, too, can be explained by shocks from an energy-stratified model, with the shocks perpendicular to the jet direction and the magnetic field oriented parallel to the shock front.

In this paper, we report on a third IXPE campaign performed from July 9 - 12, 2022, with contemporaneous radio and optical polarization observations and contemporaneous X-ray spectral coverage. Similar to the first two observations, this third pointing found the source in a low-flux, or quiescent, state. The rest of the paper is organized in the following manner. We present the data sets from

the three campaigns, including the new data from this campaign, and their analysis methods in Sect. 2. Discussion of the light curves, the fitting of X-ray energy and polarization spectra, and the multi-wavelength polarization results can be found in Sect. 3. We describe our two models of the multi-wavelength spectral and polarization results in Sect. 4, and report on the results from fitting these two SED models in Sect. 5, before closing with a discussion of all results in Sect. 6.

2 DATA SETS AND ANALYSIS METHODS

We combine the IXPE observations with campaigns covering various bands of the electromagnetic spectrum. During the first two IXPE observations of Mrk 501, several other observatories provided contemporaneous coverage in radio, infrared, optical, ultraviolet, and X-rays. For the observations, data, and analysis of those two March observations, which are similar to those added here from the third epoch, we refer the reader to (Liodakis et al. 2022). We describe below the bands of coverage for the third (July) observation, along with the processing and analysis methods used.

2.1 IXPE X-ray Observations

The new IXPE Mrk 501 observations were taken between 2022-07-09 23:35 UTC and 2022-07-12 00:40 UTC. In all, 97 ksec of spectropolarimetric data were acquired. A circular region of radius 1.5 arcminutes and an annulus region with radii 2.5 and 4.5 arcminutes, centered on the source, were used for extracting the source and background counts, respectively.

In order to obtain for this paper the overall polarization degree and angle measurements for each of the three observations, the *ixpeobssim* software suite (Baldini et al. 2022) was used to combine and visualise the data from the source region files of the three IXPE detectors and yield model-free polarization results. The background counts were subtracted from the source counts before the spectropolarimetric fits were done in *Sherpa* (Freeman et al. 2001).

2.2 Additional X-ray Observations

We also use data from the XRT onboard the *Neil Gehrels Swift* observatory (Gehrels et al. 2004), reduced using the *xrtpipeline* (version 0.13.7). The *Swift*-XRT observations presented in this paper cover a time range from the beginning of 2022 until the end of August 2022, and were performed in “window timing mode”.

We furthermore include two *XMM-Newton* (Jansen et al. 2001) observations taken on 22 and 24 March 2022. For the data extraction, we have used the standard methods with XMMSAS (version 20.0.0). In our analysis, we only consider the EPIC-pn data, which was taken in “Small Window Mode”. We extracted the source from a circular region centered on the source coordinates (RA = 253.4675867°, Dec. = 39.7600296°), with a radius of 35″. As pile-up was detected, we excluded an inner circular region with a radius of 5″ from our source region. For the background region, we extracted a circular region centered on (RA = 253.4846759°, Dec. = 39.8063927°), with a radius of 60″, as far from the source as possible.

The hard X-ray range presented in this work is covered by four

observations with the Nuclear Spectroscopic Telescope Array (*NuSTAR*; Harrison et al. 2013). We extracted the data of both Focal Plane Modules A and B (FPMA and FPMB) with nupipeline (version 0.4.9), HEASoft (version 6.31.1), and CALDB (version 20230504). The source region was a circular region centered on (RA = 253.4657281°, Dec. = 39.7601030°), with a radius of 80''. The background region was a circular region centered on (RA = 253.5406350°, Dec. = 39.8665981°), with a radius of 160''.

The X-ray energy spectral fits of the *Swift*-XRT observations (year-long) and the joint *XMM*, *Swift*, and *NuSTAR* observations (March 2022) were performed with an absorbed log-parabolic model. The neutral hydrogen column density was set to the Galactic value of $N_H = 1.69 \times 10^{20} \text{ cm}^{-2}$ (HI4PI Collaboration et al. 2016). We used Vern cross-sections (Verner et al. 1996) and the *wilm* abundances (Wilms et al. 2000). The *Swift*-XRT observations were fit from 0.5 to 10 keV. The *NuSTAR* FPMA and FPMB data were fit with the spectral parameters α , β , and flux tied together, while accounting for the normalisation difference via a constant left free in the fit. The best-fit results of the simultaneous observations and of the *Swift* data alone are given in Tables 1 and 2 (with Tables 3 and 4 being continuations of the latter) in the Appendix, respectively.

2.3 Radio Observations

For the radio observations, we used the Owens Valley Radio Observatory (15 GHz) 40m telescope and the Metsähovi Radio Observatory (22 and 37 GHz) 13.7m diameter telescope, as well as radio polarization measurements from the IRAM-30m, taken at 86 GHz and 230 GHz on 11 July (MJD 59771.755). The OVRO observation and data reduction procedures are described in detail in Richards et al. (2011). For the Metsähovi observations, the telescope detection limit at 37 GHz is ~ 0.2 Jy under optimal conditions. We consider all data points with a signal-to-noise ratio $< 4\sigma$ as non-detections. The errors in the flux density are estimated by accounting for both the contribution of the measurement root-mean-square (RMS) and the uncertainty in the absolute flux calibration. A detailed description of the data reduction and analysis can be found in Teräsranta et al. (1998).

We obtained millimeter-wave polarimetric observations at the Pico Veleta Observatory (Sierra Nevada, Granada, Spain) using the IRAM 30 m telescope. The observations were taken at 3.5 mm (86.24 GHz) and 1.3 mm (230 GHz) as part of the POLAMI (Polarimetric Monitoring of AGN at Millimeter Wavelengths¹) program (Agudo et al. 2018b,a; Thum et al. 2018). In this observing setup, we simultaneously recorded the four (I , Q , U , V) Stokes parameters, using the XPOL procedure (Thum et al. 2008). A more detailed description of the data reduction, calibration, and managing and flagging procedures can be found in Agudo et al. (2018b).

Additional mm-radio observations at 225.5 GHz were obtained using the Submillimeter Array (SMA) Ho et al. (2004) as part of the SMA Monitoring of AGN with POLarization (SMAPOL) program on the 2022 July 10 (MJD 59770.46). The SMA observations are taken using the SMA polarimeter Marrone & Rao (2008) and SWARM correlator Primiani et al. (2016) in full polarization mode. The polarization information for the source is extracted from the Stokes I , Q , U visibilities and the data are then calibrated using the

MIR².

2.4 Infrared, Optical, and UV Observations

In the optical band, the Burst Optical Observer and Transient Exploring System (BOOTES) supplied g , r , i photometry. A description of the observations can be found in Castro-Tirado et al. (2012) and Hu et al. (2023). The Perkins Telescope, Calar Alto, St. Petersburg, NOT, T60, BOOTES, and KANATA observatories provided data in units of magnitudes, which were converted to fluxes for the SEDs via the appropriate Johnsons-Cousins filter (or, for BOOTES, the SDSS filter) conversions. We also included optical and UV observations from *Swift*-UVOT.

Optical (R-band) polarization measurements were taken with the Nordic Optical Telescope (NOT) and RoboPol polarimeter at the Skinakas observatory. The polarization results are corrected for the unpolarized host-galaxy emission (Nilsson et al. 2007; Hovatta et al. 2016). Further details on the data acquisition and reduction for these observatories can be found in Liodakis et al. (2022); Di Gesu et al. (2022b); Middei et al. (2023).

2.5 Gamma-ray Observations

The *IXPE* observations were performed during a period in which the γ -ray flux of Mrk 501 was near its average. Hence, we use the time-averaged *Fermi*-LAT 4FGL spectrum as a proxy for a simultaneous GeV energy spectrum (Abdollahi et al. 2020).

3 OBSERVATIONAL RESULTS

3.1 X-ray Light Curves

Figure 1 displays the *Swift*-XRT X-ray light curve (0.5 – 10 keV) from January to August 2022. We obtained the flux values, shown in the upper panel, by fitting a log-parabola model to each spectrum. We display the best-fit spectral parameters, the photon index α at 1 keV and the spectral curvature β , in the two lower panels. The flux of Mrk 501 varied by a factor of up to three during the monitoring coverage of eight months. During some of the times when the source shows a rising flux, the spectrum becomes significantly harder (e.g., at MJD ~ 59640 , ~ 5965 , ~ 59755), but not every time (e.g., MJD ~ 59705). The curvature of the X-ray spectrum slightly varies over time, too, but with less of an apparent connection to the flux state. As can be seen from the X-ray monitoring with *Swift*-XRT, the first and third *IXPE* observations caught the source at a similar, rather low flux state. The second *IXPE* observation occurred during a fast flux increase. Figure 2 shows the light curves, as well as the 4–8 keV/2–4 keV hardness ratios, for each of the three *IXPE* three observations of Mrk 501. The higher average flux during the second observation accompanies a corresponding average increase in the hardness ratios (and thus harder energy spectra), the same correlation as exhibited in earlier campaigns (e.g., Krawczynski et al. 2000).

¹ <http://polami.iaa.es/>

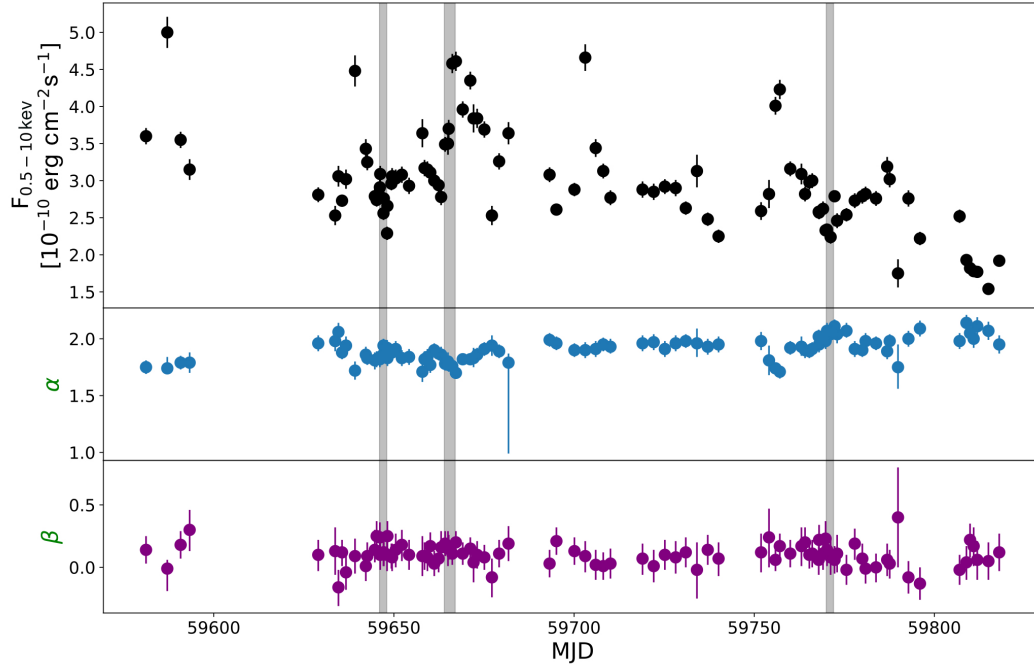


Figure 1. Fluxes (upper panel) and spectral parameters α (middle panel) and β (lower panel) retrieved via modelling the *Swift*-XRT Mrk 501 observations from January to August 2022 with a log-parabola model. The grey vertical bands show the times of the three *IXPE* observations campaigns. Detailed results from fitting the *XMM-Newton*, *NuSTAR*, and *Swift*-XRT data are found in Tables 2 - 4 in the Appendix.

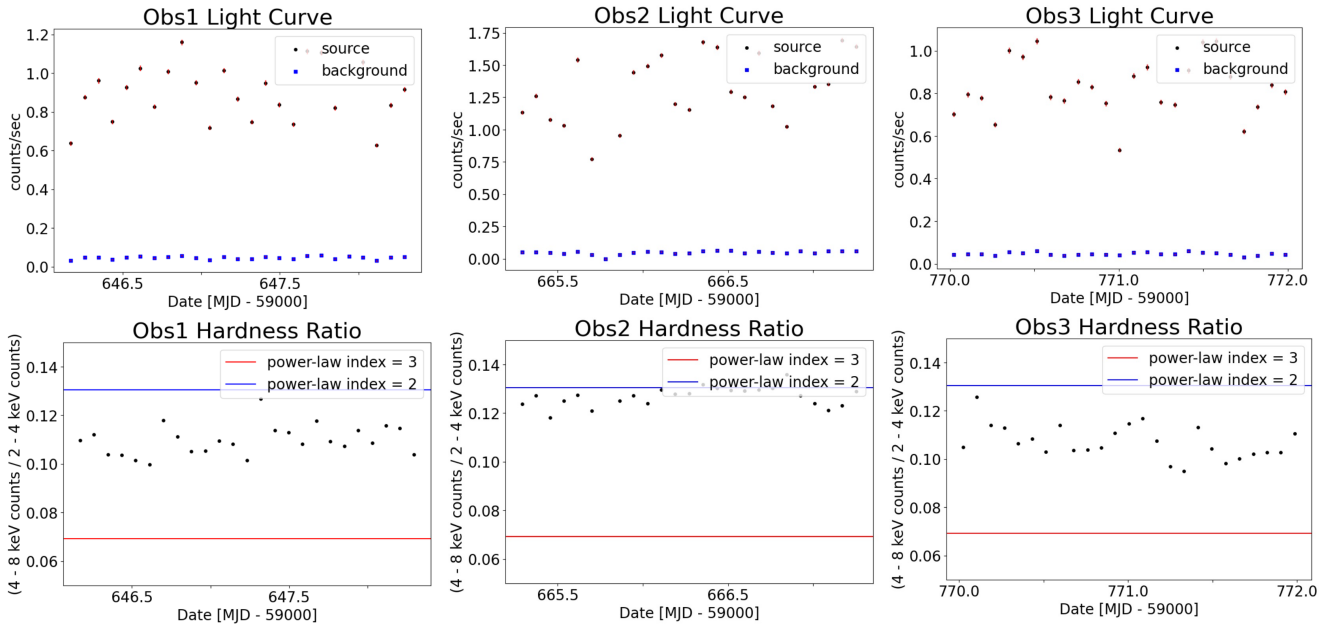


Figure 2. Light curves and 4–8 keV/2–4 keV hardness ratios for each of the three *IXPE* campaigns. The first observation has 7647 seconds per bin for both the light curve and hardness ratios; the second, 7060 seconds per bin; the third, 7069 seconds per bin. A more coarse-grained time-binning of the counts in the second observation will reveal a similar rise in flux to that shown in the *SWIFT*-XRT light curve of Figure 1. The (higher up) blue line gives the ratio of 4–8 keV/2–4 keV counts that would be obtained by *IXPE*, as calculated by *WEBPIMMS*, if the spectrum consisted solely of a power-law component with index 2. Likewise, the red line gives the ratio, if the spectrum consisted solely of a power-law component with index 3 (i.e., a softer spectrum).

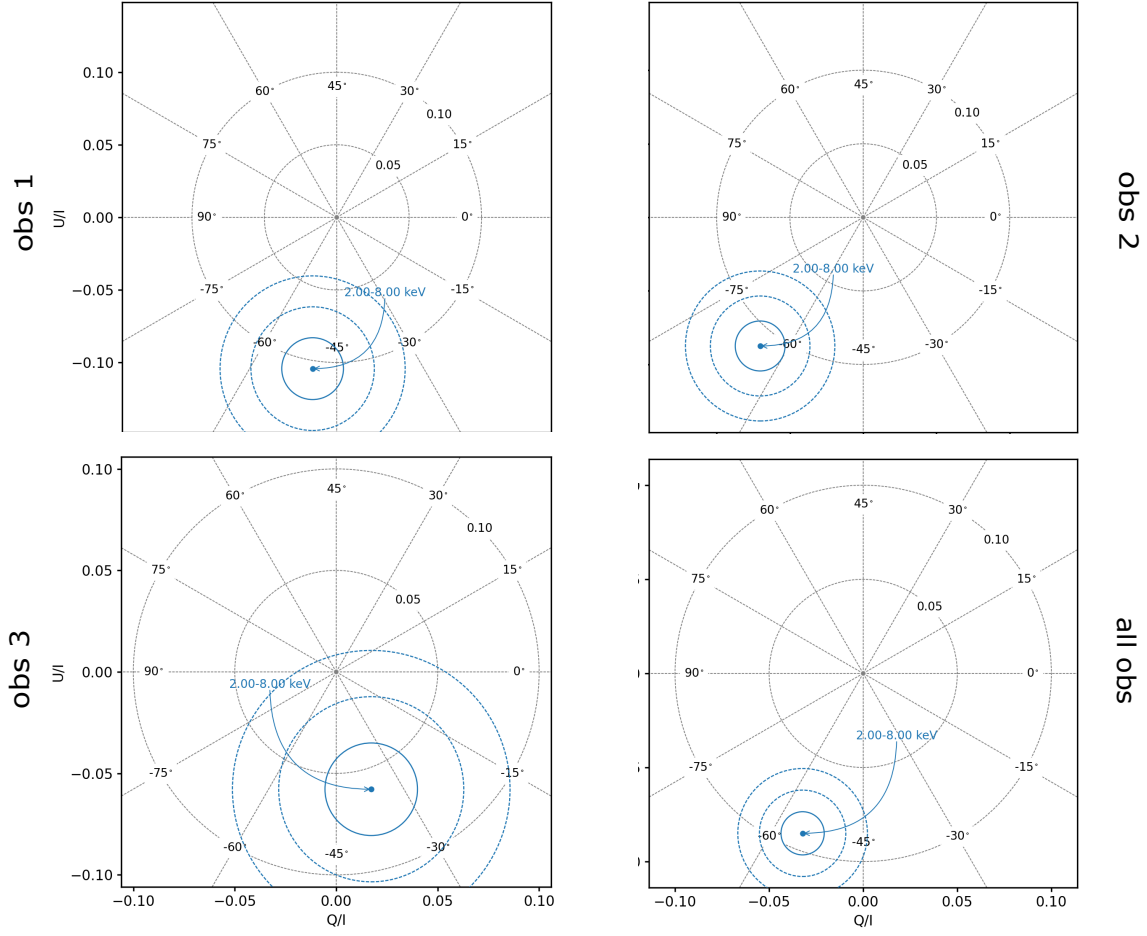


Figure 3. Detected Stokes Q and U for the three *IXPE* observations of Mrk 501, as well as for all three observations taken together. The circles around the result display, from innermost to outermost, the 1, 2, and 3 σ uncertainty regions. Note that a polarization angle of 135° is equivalent to −45°.

Table 1. Summary of the X-ray polarization results from the model-free analysis enabled by the *PCUBE* algorithm of the *ixpeobssim* software (Baldini et al. 2022). Results are given for the three *IXPE* observations, as well as for combining the data from all three observations, just as is shown in Figure 3. Errors are given at the 1 σ confidence level. The lower significance of the third observation is mainly a result of the polarization degree being measured to be of a lower magnitude than in the first two observations. Note that a polarization angle of 135° is also equivalent to −45°.

Parameter	Obs 1	Obs 2	Obs 3	All Obs
Stokes I	5133 ± 24	7712 ± 30	4332 ± 22	17177 ± 44
Stokes Q/I	-0.017 ± 0.021	-0.070 ± 0.017	0.017 ± 0.023	-0.032 ± 0.011
Stokes U/I	-0.104 ± 0.021	-0.087 ± 0.017	-0.058 ± 0.023	-0.085 ± 0.011
PD	0.106 ± 0.021	0.112 ± 0.017	0.060 ± 0.023	0.091 ± 0.011
PA	-50 ± 6	-64 ± 4	-37 ± 11	-55 ± 4
Significance [σ]	4.43	6.17	1.88	7.53

3.2 Spectropolarimetric Results

Figure 3 and Table 1 present the polarimetric results from each of the three *IXPE* observations, as well as the combined results from

² <https://lweb.cfa.harvard.edu/~cqj/mircook.html> software package (Sault et al. 1995)

all three observations, produced with the *PCUBE* algorithm from *ixpeobssim* (Baldini et al. 2022). Polarization was detected at the $> 4\sigma$ -level in observations 1 and 2 and detected at close to the 2σ -level in the third observation. For the third observation, a linear polarization degree of 6.0 ± 2.3 per cent and a polarization angle of $143^\circ \pm 11^\circ$ were found. Overall, we do not find significant evidence for the X-ray polarization changing between the three observations. Although the third observation, when taken by itself, seems to indicate a polarization angle more discrepant from the jet angle, the result obtained from the combination of all three observations remains quite consistent with the previous finding of alignment between the jet ($\psi_{\text{jet}} = 120^\circ \pm 12^\circ$) and X-ray polarization ($\psi_X = 125^\circ \pm 4^\circ$) angles³, with the combined polarization detection being at an even higher significance of 7.5σ .

We then carried out joint spectropolarimetric fits for the third *IXPE* observation and the three *SWIFT*-XRT observations that occurred on the same UTC calendar days. The data for all three Stokes parameters were background-subtracted before fitting. The data were fit to two separate models: an absorbed power-law spectrum, as well as an absorbed log-parabola model, both commonly used for describing blazar X-ray SEDs. A constant polarization model (one polarization degree and one polarization angle applicable throughout the 2–8 keV range) was applied to each of the models. The results are given in Table 2, and the spectral fits are plotted in Figure 4. The fitting of the two models produces very similar (and reasonable) χ^2 goodness-of-fit statistics, even though the log-parabolic model has one more fitting parameter. Moreover, both models give polarization results for the third *IXPE* observation consistent with each other and with the spectral, model-free analysis: a polarization degree of 7 ± 2 per cent and a polarization angle of $135^\circ \pm 8^\circ$.

For the duration of the third *IXPE* observation, we have obtained contemporaneous polarization measurements in the radio and optical bands, as described in Sect. 2. The radio emission exhibits a polarization degree of $\Pi_R = 1.51 \pm 0.33$ per cent along a radio polarization angle of $\psi_R = 144^\circ \pm 6^\circ$ at 86 GHz. In the mm-radio range, the IRAM-30m observations at 230 GHz yielded only an upper limit of $\Pi_R < 4.6$ per cent (with a 99 per cent confidence). At 225.5 GHz, using SMA, the source shows $\Pi_R = 1.27 \pm 0.28$ per cent along $\psi_R = 130 \pm 5^\circ$. Thus, the radio measurements of the polarization angle are consistent with those from the X-rays. The host-galaxy-corrected (Nilsson et al. 2007; Hovatta et al. 2016) optical (R-band) polarization degree from NOT and RoboPol during the same time period corresponds to $\Pi_O = 2.43 \pm 0.53$ per cent along $\psi_O = 113^\circ \pm 8^\circ$.

Figure 5 summarizes the multiwavelength polarization results obtained during the time of the third *IXPE* observation. The X-ray polarization degree is about a factor of three higher than that detected in the optical band and a factor of four higher than that detected in the radio band. The multi-wavelength polarization results from this campaign are consistent with those from the previous two, featuring a strongly chromatic polarization degree that is roughly aligned with the jet axis on the plane of the sky at $\psi_{\text{jet}} = 120^\circ \pm 12^\circ$.

4 LEPTONIC EMISSION MODELS FOR THE SPECTRAL ENERGY DISTRIBUTIONS

In leptonic emission models, the low-energy component is produced by synchrotron emission from electrons and possibly positrons,

while the high-energy component is produced by electrons inverse-Compton-scattering synchrotron photons or other photons (e.g., accretion disk photons). We focus here on synchrotron self-Compton (Jones et al. 1974) (SSC) models which have successfully been used to fit blazar energy spectra (see Boettcher (2010) for a review). The first model is a one-zone model with a non-uniform magnetic field. The increase of the polarization from the optical to the X-ray bands implies that the electrons emitting the optical (X-ray) emission see a less (more) uniform magnetic field distribution. The electrons emitting in the optical have lower energies, lose their energy on longer time scales, and can therefore travel through regions with greater variation of magnetic field direction. In the shock-in-jet model, we model the evolution of the magnetic field in the downstream medium, and use the magnetic field evolution to predict the polarization degrees.

4.1 Synchrotron Self-Compton Model with Anisotropic Magnetic Field

One of the two models that we use to describe the data uses the SSC code described in (Krawczynski et al. 2004). The SSC code assumes a spherical emission volume of radius R traveling with velocity βc (c being the speed of light) along the jet axis. The observer views the emission volume at an observer-frame angle θ_{obs} away from the jet axis. According to the standard equations of special relativity, the relativistic Doppler factor is given by the equation

$$\delta = [\Gamma(1 - \beta \cos \theta_{\text{obs}})]^{-1}. \quad (1)$$

The cosine of the angle to the jet axis $\mu_{\text{obs}} \equiv \cos \theta_{\text{obs}}$ transforms into the plasma frame according to

$$\mu_j = \frac{\mu_{\text{obs}} - \beta}{1 - \mu_{\text{obs}}\beta}. \quad (2)$$

An isotropic distribution of electrons emits synchrotron radiation in magnetic fields having strength B . The synchrotron emission scattering off the electrons gives rise to the high-energy emission component. The electron energy spectrum is described by a broken power law with index p_1 ($dN_e/d\gamma \propto \gamma^{-p_1}$) from electron Lorentz factor γ_{min} to γ_b and with index p_2 from γ_b to γ_{max} .

For this work, we have added the ability to predict the polarization of the synchrotron emission. To do so, we assume that the magnetic field in the emission region has the same strength everywhere in the spherical emission volume, but that the magnetic field direction varies locally between each minuscule sub-volume of the sphere. The myriad of small sub-volumes thus has each exhibiting its own magnetic field direction, not necessarily related to the direction of the field in an adjacent sub-volume. The condition imposed on these directions, however, is that the overall distribution of magnetic field directions (when looking across all sub-volumes) exhibits an anisotropy that can be described with an axisymmetric probability distribution $p_B(\mu)$ depending only on the cosine μ of the angle between the magnetic field direction within the sub-volume and the jet axis (with all angles in the jet frame) for a given n . n , then, is used to provide a simple parameterization:

$$p_B(\mu, n) = \frac{n+1}{2} (1 - |\mu|)^n. \quad (3)$$

This prescription thus yields an isotropic distribution of magnetic fields for $n = 0$ and a distribution of magnetic fields all perpendicular to the jet axis for large n , and is normalized to give $\int_{-1}^1 d\mu p_B(\mu, n) = 1$. Figure 6 shows the two sample probability distributions used later to fit the observed polarization degrees.

³ https://www.bu.edu/blazars/VLBA_GLAST/1652.html

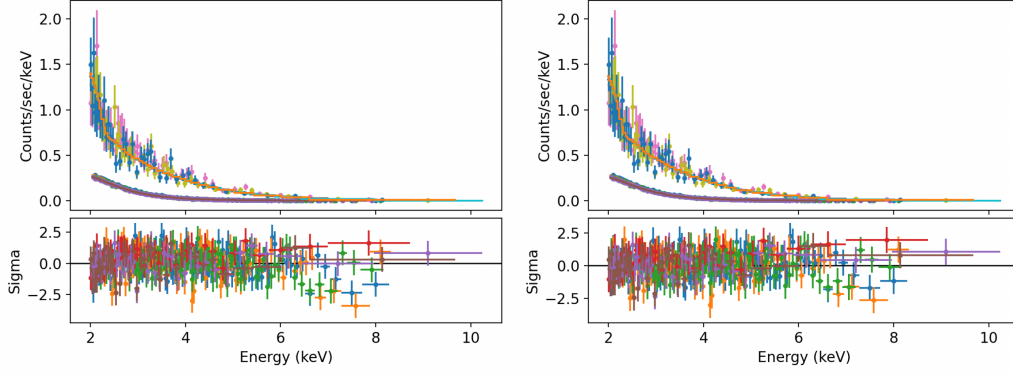


Figure 4. *SWIFT*-XRT and *IXPE* data for Stokes I, fitted with an absorbed power-law model (left) and an absorbed log-parabola model (right). The data have been re-binned to have a minimum signal-to-noise ratio of 5 in each bin. The lower flux blue-green-purple data points, with error bars, are those of the three *IXPE* detectors from the third *IXPE* observation of Mrk 501. The higher flux blue-pink-yellow data points, with error bars, are those of the three *SWIFT* observations concurrent with the third *IXPE* observation. The solid lines are the spectral models, obtained with the *Sherpa* fitting software, that best explain each data set. The residuals are given in the subplots below the main figures, with the y-axis being in units of the error of a given data point.

Table 2. Results from the joint spectropolarimetric fits of all of the *IXPE* and *Swift*-XRT data during the time of the third *IXPE* observation. The hydrogen absorption n_H is fixed to $0.0155 \times 10^{22} \text{ cm}^{-2}$. The pivot energy of the log-parabola model was set to 2 keV. The instrumental cross-calibration factors are normalized such that the *IXPE* DU1 factor is always unity. Errors are given at the 1σ confidence level.

Model Parameter	Absorbed Power-Law	Absorbed Log-Parabola
<i>IXPE</i> DU1	1 (frozen)	1 (frozen)
<i>IXPE</i> DU2	0.950 ± 0.009	0.950 ± 0.009
<i>IXPE</i> DU3	0.892 ± 0.008	0.892 ± 0.008
<i>SWIFT</i>	$1.15^{+0.03}_{-0.02}$	1.15 ± 0.03
power-law index	$2.41^{+0.01}_{-0.02}$	n/a
power-law norm [at 1 keV]	0.067 ± 0.001	n/a
log-parabola alpha	n/a	$2.20^{+0.04}_{-0.05}$
log-parabola beta	n/a	0.5 ± 0.1
log-parabola norm	n/a	0.0124 ± 0.0001
polarization degree (%)	7 ± 2	7 ± 2
polarization angle (deg)	135 ± 8	135 ± 8
χ^2 for spectral parameters	399	375
d.o.f. for spectral parameters	425	424
χ^2 for polarization parameters	186	185
d.o.f. for polarization parameters	178	178

We calculate the emitted synchrotron Stokes I and Stokes Q parameters as a function of the emitted frequency by integrating over the contributions from all electron Lorentz factors and all magnetic field orientations (see Laing 2002, for a similar calculation). In the jet frame, the emitted synchrotron power per angular frequency ω is given by

$$I = \int_{\gamma_{\min}}^{\gamma_{\max}} d\gamma \frac{dN_e}{d\gamma} \int_{-1}^1 d\mu \int_0^{2\pi} d\phi p_B(\mu, n) \frac{\sqrt{3} B q^3 \sin \alpha}{4\pi^2 m c^2} F(x). \quad (4)$$

The first integral runs over the Lorentz factors γ of the electrons, and the second and third integrals run over the magnetic field directions (i.e., over μ , which is the cosine of the angle between the magnetic field direction and the jet axis, as well as over the azimuthal angle ϕ). q is the charge of the electron, B is the magnetic field strength, α is the angle between the magnetic field direction and the line of sight, and m is the electron mass. The variable x depends on all of the integration variables and is given by $x = \omega/\omega_c$, with the synchrotron critical frequency $\omega_c = 3\gamma^2 q B \sin \alpha / 2mc$. The function $F(x)$ (Rybicki &

Lightman 1986) is an integral of the modified Bessel function of order 5/3:

$$F(x) = x \int_x^\infty K_{5/3}(\xi) d\xi.$$

The corresponding power in Stokes Q parallel to the jet axis projected on the plane of the sky is given by

$$Q = \int_{\gamma_{\min}}^{\gamma_{\max}} d\gamma \frac{dN_e}{d\gamma} \int_{-1}^1 d\mu \int_0^{2\pi} d\phi p_B(\mu, n) \times \frac{\sqrt{3} B q^3 \sin \alpha}{4\pi^2 m c^2} G(x) \cos 2\chi. \quad (5)$$

The terms on the right hand side of Equation (5) up to the $\cos 2\chi$ term give the intensity of the linearly polarized synchrotron emission. The factor α , in particular, accounts for the dependency of the intensity on the angle between the magnetic field direction and the direction of the line of sight. The function $G(x)$ (Rybicki & Lightman 1986) is a function of the modified Bessel function of order 2/3:

$$G(x) = x K_{2/3}(x).$$

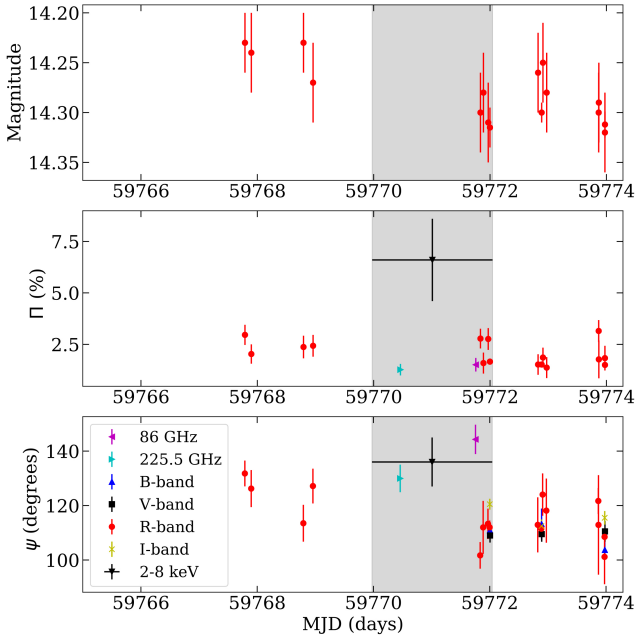


Figure 5. Contemporaneous radio, optical, and X-ray polarization observations around the time of the third *IXPE* observation. The top panel shows the R-band magnitudes, the middle panel the measured polarization degrees, and the bottom panel the measured polarization angles. The grey-shaded area marks the duration of the third *IXPE* observation.

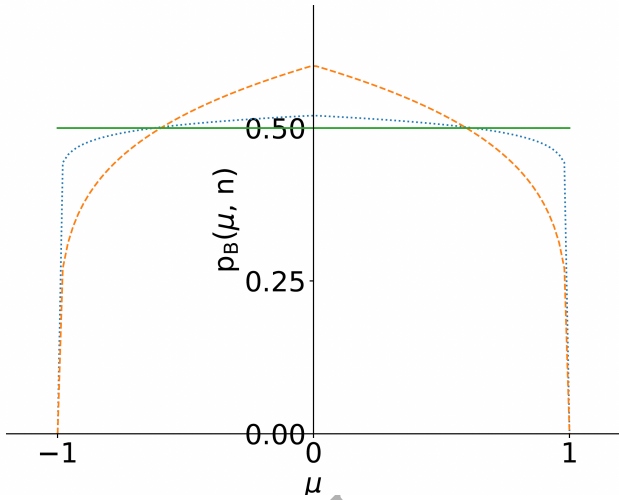


Figure 6. We use a simple phenomenological prescription to study the impact of anisotropies in the magnetic field distribution. We assume that the overall distribution of magnetic field directions in all minuscule sub-volumes of the sphere can be described by a probability distribution $p_B(\mu, n)$ depending on the cosine μ of the angle between the magnetic field direction within that sub-volume and the jet axis, with the parameter n characterizing the amount of anisotropy in the overall distribution. We show here the probability distributions for $n = 0.041$ (dotted blue line) and $n = 0.205$ (dashed orange line), which can explain the optical and X-ray polarization degree results, respectively, reported in Section 3. The two distributions demonstrate that small deviations from a magnetic field distribution having all directions equally likely are sufficient to explain the varied observed polarization degrees. The flat (isotropic) distribution of $p_B = 1/2$ is shown by the green solid line.

χ is the angle between the electric field polarization direction (perpendicular to the magnetic field) and the jet axis projected onto the plane of the sky. Thus, the factor $\cos 2\chi$ extracts the Q -component of the intensity (Rybicki & Lightman 1986) where, in our treatment, $Q > 0$ denotes the linear polarization parallel to the jet axis, and $Q < 0$ denotes linear polarization perpendicular to the jet axis. The azimuthal symmetry of the distribution of magnetic fields, p_B , implies that p_B is ultimately anisotropic only with respect to the angle of declination from the jet axis (with no anisotropies with respect to the azimuthal angle around the jet axis). Thus, after summing the contributions from all azimuthal angles, the net polarization can only be parallel to or perpendicular to the jet axis, and Stokes U (denoting the plasma-frame intensity of the emission linearly polarized along the diagonals between the jet axis and the directions perpendicular to it) must thus vanish. The polarization degree can then be given just by $|Q|/I$: the polarization angle is, by symmetry, either 0° ($Q > 0$) or 90° ($Q < 0$) with respect to the jet axis.

Our code integrates over all directions of the 4π sphere. For each direction, we calculate μ and χ with the help of simple vector calculus equations. Each result contributes with a weight proportional to $p_B(\mu, n)$ (equation 3). After calculating I and Q , we account for the effect of synchrotron self-absorption on the emitted intensity, but not on the polarization (Longair 2010) and account for the effect of relativistic boosting of the specific intensities for the transformation from the jet frame to the stationary frame, and calculate the observable fluxes while accounting for the source’s luminosity distance (Rybicki & Lightman 1986).

We use the standard synchrotron self-Compton kernel (Blumenthal & Gould 1970) to calculate the self-Compton emission, and the Cosmic Infrared Background (CIB) model of Franceschini et al. (2008) to calculate the γ -ray fluxes reaching the observer, accounting for the extinction owing to $\gamma + \gamma \rightarrow e^+e^-$ processes. The calculation of the CIB extinction assumes a Hubble constant of $67.45 \text{ Mpc s}^{-1} \text{ km}^{-1}$ (Workman 2022).

4.2 Shock-in-Jet Model

The second model is a refined version of the model described in Tavecchio et al. (2018). A detailed description will be presented in a forthcoming paper (Tavecchio et al., in preparation). Relativistic electrons are injected at the front of a mildly relativistic shock (as observed in the flow frame) of a cylindrical jet with radius R . The upstream flow is assumed to carry a weak field approximately parallel to the jet axis with intensity $B_{\parallel}$. As demonstrated by PIC simulations, just downstream of the shock, an intense field parallel to the shock (and hence normal to the jet axis) is generated by streaming high-energy particles. We assume that the intensity of the self-generated field decays along the downstream flow according to the relation $B_{\perp}(d) = B_{\perp,0}[1 + d/\lambda]^{-m}$, where d is the distance from the shock. Here, λ plays the role of an effective decay length. The turbulent-like structure of the self-generated magnetic field is modeled with a cell structure, with each cell representing a coherence domain. In each cell we specify the total magnetic field as the sum of a constant parallel (poloidal) field $B_{\parallel} = B_z$ and the orthogonal field $B_{\perp}(d)$ (evaluated at the distance d of the cell), with the components B_x and B_y randomly selected in each cell with the condition $B_x^2 + B_y^2 = B_{\perp}^2(d)$. Building upon the original model developed in Tavecchio et al. (2018), we now add the self-consistent calculation of the electron energy distribution by using the full continuity equation, including radiative losses, treated as in Chiaberge & Ghisellini (1999). We assume that electrons injected at the shock

follow a truncated power-law energy distribution with index n and minimum and maximum cut-off Lorentz factors of $\gamma_{\min}$ and $\gamma_{\max}$, respectively. Emitting (and cooling) particles are advected downstream with velocity $v_{\text{adv}} = c/3$ up to a distance d_{max} , where we assume the emission is quenched by jet expansion and consequent adiabatic losses. We calculate the SED and the polarization in the quiescent state by summing the Stokes parameters of the emission from all cells in the observer frame (Lyutikov et al. 2005).

Due to the different cooling lengths characterizing electrons with different energies, radiation at different frequencies is produced in different volumes of the downstream region (Marscher & Gear 1985). The short cooling distance of the particles at the highest energies, radiating in the X-ray band at the peak of the synchrotron component, allows the emission to occur only in a thin region after the shock, where the dominant field is the self-generated, orthogonal magnetic field. When the jet is observed at a typical angle $\sim 1/\Gamma$, where Γ is the downstream bulk Lorentz factor, the emitted synchrotron radiation will be characterized by a high polarization degree. On the other hand, radiation at lower frequencies is produced by electrons with relatively long cooling times. This radiation is thus emitted in a larger volume, where both the orthogonal and the parallel fields are important, determining the effective dilution of the emerging polarization.

5 SPECTRAL AND POLARIMETRIC ENERGY DISTRIBUTION MODELING RESULTS

5.1 Synchrotron Self-Compton Model with Anisotropic Magnetic Field

Figure 7 shows the multiwavelength SED and polarization energy spectra, together with the results from the SSC model with an anisotropic magnetic field distribution. The top panel shows the spectral energy distribution, while the two bottom panels show the polarization degree and direction. The model uses $\gamma_{\min}mc^2 = 10$ MeV, $\gamma_b mc^2 = 2.5$ GeV, $\gamma_{\max}mc^2 = 398$ GeV, $p_1 = 1.44$, $p_2 = 2.44$, $B = 0.033$ G, $R = 7.4 \times 10^{15}$ cm, $\beta = 0.9995$, $\Gamma = 33$, and $\theta_{\text{obs}} = 1^\circ.6$. The last two parameters give a relativistic Doppler factor of $\delta = 35.69$. The radius corresponds to an observer frame light travel time $t_R = R/(c\delta)$ of 1.92 hours. The synchrotron emission at the synchrotron peak is emitted by 2.5 GeV electrons. These electrons cool from synchrotron radiation and from self-Compton emission. The observer-frame synchrotron cooling time is around 1100 hours. The combined synchrotron and self-Compton cooling time would then be somewhat shorter.

We can reproduce the polarization degrees observed in the optical and X-ray bands with anisotropy parameters of $n = 0.041$ and $n = 0.205$, respectively. The corresponding probability distributions of the magnetic field are shown in Fig. 6. These results show that the magnetic field seen by the X-ray-emitting electrons is significantly more uniform (less isotropic) than the one seen by the electrons emitting at optical wavelengths.

We note that the inferred amount of magnetic field anisotropy is degenerate with the angle at which the observer views the emitting plasma. The closer to the jet axis that the observer views the emitting plasma, the more anisotropic the magnetic field needs to be in order to fit the observed polarization degree. The required magnetic field anisotropy is thus minimized for a jet-frame viewing angle θ_j of 90° . We note, furthermore, that different $\theta_{\text{obs}} - \Gamma$ combinations can give the same Doppler factor δ (equation 1) and jet-frame viewing angle θ_j , leading to identical predicted SEDs for different combinations

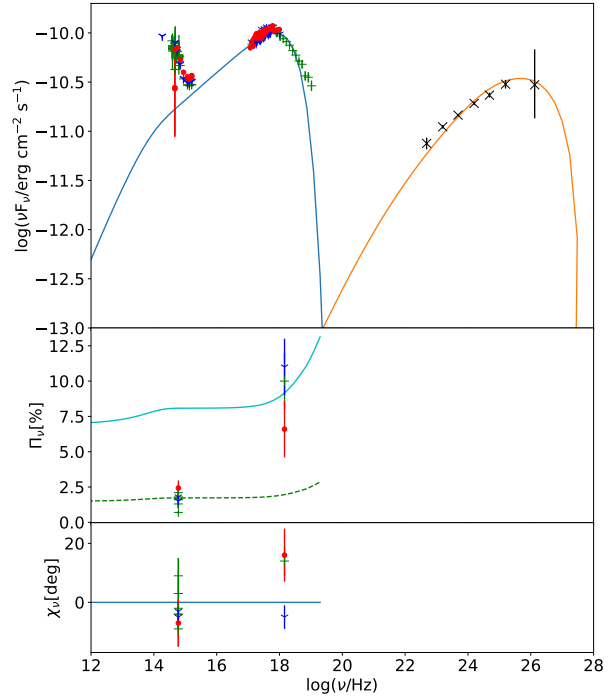


Figure 7. SED (upper panel), polarization degree (middle), and EVPA (lower) of Mrk 501. In the SED panel, the solid lines show the flux predictions from the SSC model with an anisotropic magnetic field. For the polarization degree we show the modeling results for magnetic field anisotropies described by $n = 0.041$ (dashed green line, to match optical polarization results) and $n = 0.205$ (solid cyan line, for X-ray polarization results) (see Figure 6). The modeled polarization angle $\chi_v = 0^\circ$ corresponds to a polarization parallel to the jet, as indeed is generally observed to be true on average for Mrk 501. In all three panels, the observational results are described in section 2, taken during the times of the three *IXPE* observations. The green crosses are from the time of the first observation, the blue three-pointed points from the time of the second observation, and the red dots from the time of the third observation. The black diagonal crosses are the data from the archival *Fermi*-LAT 4FGL spectrum.

of those parameters. We thus fix $\theta_{\text{obs}} = 1^\circ.6$ and $\Gamma = 33$, giving $\theta_j = 90^\circ$. The values of n inferred in our models thus represent the minimum amount of magnetic field anisotropy required to describe the observed polarization degrees.

5.2 Shock-In-Jet Model

Figure 8 displays the flux and polarization energy spectra, together with the results from the shock-in-jet model. The top panel shows the spectral energy distribution, while the two bottom panels show the polarization degree and direction. The model assumes the following parameters: $\gamma_{\min} = 100$, $\gamma_{\max} = 8 \times 10^5$, $n = 2.2$, $B_{\perp,0} = 0.27$ G, $B_z = 0.034$ G, $R = 4.3 \times 10^{15}$ cm, $\lambda = 5 \times 10^{13}$ cm, $m = 0.5$, $\Gamma = 22$, and $\theta_{\text{obs}} = 1^\circ.3$. The last two parameters give a relativistic Doppler factor of $\delta = 35$. For this figure, even though we show the results down to 10^{10} Hz, we note that this model does not include the effect of synchrotron self-absorption or Faraday rotation on the predicted properties.

A critical parameter for the predicted polarization degree is the angle θ_j between the jet axis and the observer, as measured in the jet plasma frame. A viewing angle of $\theta_j = 90^\circ$ maximizes the polarization in the X-ray band, since the projected self-generated field

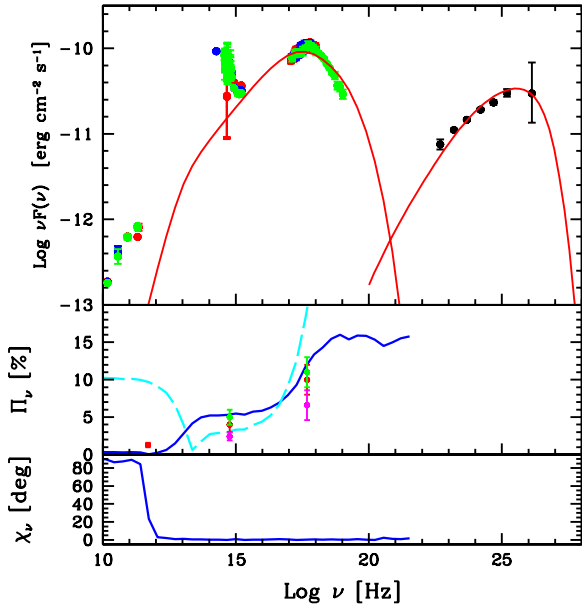


Figure 8. Comparison of the observed (data points) and modeled (lines) SED (top), polarization degree (middle), and EVPA (bottom) of Mrk 501 for the shock-in-jet model. The solid blue line shows our fiducial model, and the dashed cyan line shows the results for a different viewing angle. The modeled polarization angle swings from $\chi_v \approx 90^\circ$ below 10^{12} Hz to $\chi_v \approx 0^\circ$ above 10^{12} Hz, since the radio-emitting electrons experience the magnetic field parallel to the jet, and higher energy electrons experience the shock magnetic field perpendicular to the jet.

is fully orthogonal to the jet axis. For smaller viewing angles, both projected (randomly-distributed) B_x and B_y components contribute, resulting in a lower polarization. In the limiting case $\theta_j = 0^\circ$, the effective polarization is zero, since an observer located exactly along the jet axis sees an equal contribution of the two components of the magnetic field. To show the relatively strong dependence of the observed polarization on the observing angle, we report Π_v (dashed cyan line in the middle panel) for a different choice of the bulk Lorentz factor $\Gamma = 30$ and viewing angle $\theta_{\text{obs}} = 1.6^\circ$ giving the same Doppler factor $\delta = 35$. For the latter choice, the X-ray emission exhibits a higher polarization degree and a more pronounced wavelength dependence.

6 CONCLUSIONS

We have presented an analysis of the third observation of the X-ray polarization of the blazar Mrk 501. We do not find statistically significant evidence for a change in the soft X-ray polarization degree or angle between the three *IXPE* observations in 2022. Multi-wavelength measurements confirm the earlier result (Liodakis et al. 2022) that the polarization degree increases by at least a factor of 2 from the optical band to the X-ray band. Additionally, the polarization direction remains consistent with an approximate alignment along the radio jet.

These observations can be explained by a scenario of particle acceleration at a shock, oriented perpendicular to the radio jet axis, which generates a magnetic field parallel to the shock front that de-

cays with distance from the front. These observations can also be explained by modeling the spectropolarimetric data with synchrotron-Compton models. For the latter, we emphasize the degeneracy of the anisotropy of the magnetic field with the jet-plasma-frame angle at which this anisotropic magnetic field is seen. We have demonstrated that relatively small levels of magnetic field anisotropy can explain the observed polarization results. The consistency of the polarization degree and direction between observations spaced four months apart supports the shock scenario, if shocks in blazar jets are generally oriented perpendicular to the axis. It is not yet clear, however, whether similar models would also apply to those blazars whose X-ray emission is not primarily from synchrotron radiation (such as low synchrotron peaked (LSP) blazars).

During all of the observational data sets obtained thus far, Mrk 501 has not been seen to exhibit a rotation or variation of the X-ray polarization direction (Chen et al. 2024), including the rapid swings at a rate of $\sim 80^\circ/\text{day}$ – $90^\circ/\text{day}$ that have been reported for the blazar Mrk 421 (Di Gesu et al. 2023). The latter object is otherwise similar to Mrk 501 in terms of luminosity and a synchrotron SED that peaks at X-ray energies. The polarization swings in Mrk 421 were interpreted by Di Gesu et al. (2023) as evidence for a shock accelerating particles as it moves along a helical magnetic field threading the jet. The phenomenology of such polarization direction swings therefore remains to be determined. Future observations of the X-ray and multiwavelength polarization of Mrk 421, Mrk 501, and other blazars will be pivotal for sampling and determining how particle acceleration in the jets relates to the magnetic field structure.

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DATA AVAILABILITY

The data underlying this article will be shared on reasonable request to the corresponding author.

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Table 1. Best fit results for simultaneous *XMM-Newton*/*Swift* and *NuSTAR* data. Flux is unabsorbed value.

Date	MJD	ObsID		α	β	$F_{0.5-80\text{ keV}}$ [$10^{-10}\text{ erg cm}^{-2}\text{ s}^{-1}$]	Fit statistic $\chi^2/\text{d.o.f.}$
		NuSTAR	XMM/Swift				
2022-03-09	59647.0	60701032002	00011184180 (S)	1.79 ± 0.04	0.340 ± 0.023	6.82 ± 0.35	984.66/942
2022-03-22	59660.0	60502009002	0843230301 (X)	1.887 ± 0.005	0.189 ± 0.006	12.73 ± 0.24	2179.38/1914
2022-03-24	59662.0	60502009004	0843230401 (X)	1.925 ± 0.004	0.214 ± 0.005	10.10 ± 0.15	2323.12/2002
2022-03-27	59665.0	60702062004	00011184187 (S)	1.76 ± 0.04	0.262 ± 0.019	9.85 ± 0.57	1132.89/1131

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Table 2. Best fit results for the *Swift*-XRT observations from January to August 2022. Flux is unabsorbed value.

Date	MJD	ObsID	α	β	$F_{0.5-10\text{keV}}$ [10^{-10} erg cm $^{-2}$ s $^{-1}$]	Fit statistic $\chi^2/\text{d.o.f.}$
2022-1-2	59581.2	00011184166	1.75 ± 0.06	0.14 ± 0.11	3.60 ± 0.11	222.16/204
2022-1-8	59587.1	00011184168	$1.74^{+0.05}_{-0.10}$	$-0.01^{+0.18}_{-0.07}$	5.00 ± 0.21	153.05/184
2022-1-11	59590.8	00011184169	1.79 ± 0.06	0.18 ± 0.11	3.55 ± 0.11	184.32/202
2022-1-14	59593.3	00011184170	1.79 ± 0.09	$0.30^{+0.17}_{-0.16}$	3.15 ± 0.14	106.74/117
2022-2-19	59629.0	00096028003	1.96 ± 0.07	$0.10^{+0.13}_{-0.12}$	2.81 ± 0.10	169.70/164
2022-2-23	59633.7	00096029001	1.98 ± 0.09	$0.13^{+0.19}_{-0.18}$	2.53 ± 0.13	119.65/91
2022-2-24	59634.6	00096029002	2.06 ± 0.08	$-0.16^{+0.15}_{-0.14}$	3.06 ± 0.14	119.63/121
2022-2-25	59635.6	00096029003	1.88 ± 0.06	$0.12^{+0.11}_{-0.10}$	2.73 ± 0.08	205.56/200
2022-2-26	59636.7	00096029004	$1.94^{+0.07}_{-0.08}$	$-0.04^{+0.14}_{-0.13}$	3.02 ± 0.13	142.59/147
2022-3-1	59639.2	00011184176	1.72 ± 0.08	0.09 ± 0.14	4.48 ± 0.21	146.19/144
2022-3-4	59642.2	00011184177	1.86 ± 0.07	0.01 ± 0.12	3.43 ± 0.13	165.42/174
2022-3-4	59642.6	00096029005	1.83 ± 0.06	0.09 ± 0.11	3.25 ± 0.11	189.71/192
2022-3-6	59644.7	00096029007	$1.81^{+0.08}_{-0.07}$	$0.14^{+0.10}_{-0.16}$	2.79 ± 0.11	167.68/178
2022-3-7	59645.2	00011184178	1.82 ± 0.07	0.25 ± 0.12	2.74 ± 0.09	224.31/184
2022-3-8	59646.1	00096029008	1.82 ± 0.07	0.24 ± 0.12	2.91 ± 0.10	144.82/176
2022-3-8	59646.2	00096029009	1.84 ± 0.07	$0.10^{+0.12}_{-0.11}$	3.09 ± 0.11	206.11/183
2022-3-9	59647.1	00011184179	1.94 ± 0.06	$0.12^{+0.12}_{-0.11}$	2.56 ± 0.09	221.15/178
2022-3-9	59647.2	00011184180	1.87 ± 0.06	0.12 ± 0.11	2.76 ± 0.09	183.76/187
2022-3-10	59648.1	00011184181	1.92 ± 0.07	0.10 ± 0.13	2.29 ± 0.09	148.06/161
2022-3-10	59648.2	00011184182	1.83 ± 0.07	0.25 ± 0.12	2.66 ± 0.09	164.23/180
2022-3-11	59649.3	00011184183	1.88 ± 0.06	0.09 ± 0.10	2.96 ± 0.09	208.45/199
2022-3-11	59649.5	00096029010	1.88 ± 0.06	0.08 ± 0.11	3.06 ± 0.11	147.89/180
2022-3-12	59650.5	00096029011	1.91 ± 0.07	$0.14^{+0.14}_{-0.13}$	3.03 ± 0.12	156.14/147
2022-3-14	59652.2	00011184184	1.83 ± 0.07	$0.18^{+0.13}_{-0.12}$	3.08 ± 0.11	178.92/176
2022-3-16	59654.2	00096028004	1.84 ± 0.07	0.10 ± 0.12	2.93 ± 0.11	199.00/176
2022-3-19	59657.9	00096029013	$1.71^{+0.09}_{-0.10}$	0.09 ± 0.16	3.64 ± 0.19	125.43/106
2022-3-20	59658.5	00096029014	1.82 ± 0.07	0.09 ± 0.11	3.17 ± 0.11	182.60/183
2022-3-21	59659.3	00096029015	$1.84^{+0.06}_{-0.07}$	0.08 ± 0.11	3.15 ± 0.11	216.16/185
2022-3-22	59660.1	00096029016	1.77 ± 0.07	$0.17^{+0.12}_{-0.11}$	3.11 ± 0.11	150.25/176
2022-3-23	59661.2	00096029017	$1.90^{+0.05}_{-0.06}$	0.03 ± 0.10	3.00 ± 0.09	195.53/213
2022-3-24	59662.4	00096029018	1.87 ± 0.06	0.07 ± 0.11	2.94 ± 0.10	144.02/195
2022-3-25	59663.1	00011184185	1.86 ± 0.07	0.16 ± 0.13	2.78 ± 0.11	128.39/159
2022-3-26	59664.2	00011184186	1.78 ± 0.06	0.19 ± 0.10	3.49 ± 0.10	261.58/212
2022-3-27	59665.0	00096029019	1.80 ± 0.08	0.15 ± 0.14	3.50 ± 0.15	134.75/138
2022-3-27	59665.2	00011184187	1.77 ± 0.06	0.16 ± 0.10	3.70 ± 0.12	194.72/206
2022-3-28	59666.2	00011184188	1.76 ± 0.05	0.11 ± 0.09	4.58 ± 0.13	240.58/231
2022-3-29	59667.2	00011184189	1.70 ± 0.05	0.20 ± 0.09	4.61 ± 0.13	282.44/241
2022-3-31	59669.1	00011184190	1.82 ± 0.05	0.11 ± 0.09	3.96 ± 0.11	249.07/234
2022-4-2	59671.2	00011184191	1.82 ± 0.05	$0.15^{+0.10}_{-0.09}$	4.35 ± 0.12	223.48/224
2022-4-3	59672.1	00011184192	$1.83^{+0.08}_{-0.09}$	$0.04^{+0.16}_{-0.15}$	3.84 ± 0.19	87.04/113
2022-4-4	59673.1	00011184193	1.86 ± 0.06	0.10 ± 0.11	3.84 ± 0.13	212.22/201
2022-4-6	59675.1	00011184194	1.91 ± 0.06	0.08 ± 0.10	3.69 ± 0.11	188.06/203

Table 3. Continuing Table 2.

Date	MJD	ObsID	α	β	$F_{0.5-10\text{keV}}$ [$10^{-10}\text{ erg cm}^{-2}\text{ s}^{-1}$]	Fit statistic $\chi^2/\text{d.o.f.}$
2022-4-8	59677.2	00011184195	1.94 ± 0.09	$-0.08^{+0.16}_{-0.15}$	2.53 ± 0.13	92.47/99
2022-4-10	59679.2	00011184196	1.89 ± 0.06	0.11 ± 0.11	3.26 ± 0.11	170.09/187
2022-4-12	59681.8	00011184197	1.79 ± 0.08	0.19 ± 0.14	3.64 ± 0.15	155.98/144
2022-4-24	59693.2	00011184199	1.99 ± 0.06	0.03 ± 0.11	3.08 ± 0.10	157.03/182
2022-4-26	59695.1	00011184200	1.96 ± 0.06	0.21 ± 0.11	2.61 ± 0.08	210.47/195
2022-5-1	59700.1	00011184202	1.90 ± 0.06	0.13 ± 0.11	2.88 ± 0.09	235.15/193
2022-5-4	59703.1	00011184203	1.90 ± 0.07	$0.09^{+0.13}_{-0.12}$	4.66 ± 0.18	168.14/162
2022-5-7	59706.0	00011184204	1.91 ± 0.07	0.02 ± 0.12	3.44 ± 0.12	150.23/173
2022-5-9	59708.1	00011184205	1.95 ± 0.06	0.01 ± 0.11	3.13 ± 0.10	183.52/193
2022-5-11	59710.1	00011184206	$1.93^{+0.06}_{-0.07}$	0.03 ± 0.12	2.77 ± 0.10	181.39/175
2022-5-20	59719.0	00011184207	1.96 ± 0.07	$0.07^{+0.13}_{-0.12}$	2.88 ± 0.11	121.46/162
2022-5-23	59722.1	00011184208	1.97 ± 0.07	0.01 ± 0.13	2.85 ± 0.11	140.77/148
2022-5-26	59725.2	00011184209	1.91 ± 0.07	0.10 ± 0.12	2.92 ± 0.10	136.26/170
2022-5-29	59728.2	00011184210	1.96 ± 0.07	0.08 ± 0.13	2.90 ± 0.11	149.94/167
2022-5-31	59731.0	00011184211	1.98 ± 0.06	0.12 ± 0.12	2.63 ± 0.09	170.36/174
2022-6-4	59734.1	00011184212	$1.96^{+0.12}_{-0.13}$	$-0.02^{+0.23}_{-0.22}$	3.13 ± 0.22	54.54/57
2022-6-7	59737.1	00011184213	1.93 ± 0.07	0.14 ± 0.12	2.48 ± 0.09	151.22/170
2022-6-10	59740.1	00011184214	1.95 ± 0.07	$0.07^{+0.14}_{-0.13}$	2.25 ± 0.09	159.42/151
2022-6-21	59751.9	00011184215	1.98 ± 0.08	$0.12^{+0.16}_{-0.15}$	2.59 ± 0.12	109.53/109
2022-6-24	59754.1	00011184216	1.81 ± 0.13	$0.24^{+0.24}_{-0.23}$	2.82 ± 0.19	75.31/56
2022-6-25	59755.9	00011184217	1.74 ± 0.05	0.06 ± 0.09	4.01 ± 0.12	207.56/233
2022-6-27	59757.1	00011184218	1.71 ± 0.06	0.17 ± 0.10	4.23 ± 0.13	237.63/218
2022-6-30	59760.0	00011184220	1.92 ± 0.06	0.11 ± 0.11	3.16 ± 0.10	170.09/186
2022-7-3	59763.1	00096558001	$1.93^{+0.08}_{-0.09}$	$0.17^{+0.16}_{-0.15}$	3.09 ± 0.14	122.49/111
2022-7-4	59764.1	00096558002	1.90 ± 0.07	$0.20^{+0.13}_{-0.12}$	2.82 ± 0.10	123.15/170
2022-7-5	59765.3	00096558003	1.89 ± 0.06	0.10 ± 0.11	2.98 ± 0.10	206.52/198
2022-7-6	59766.1	00096558004	1.91 ± 0.06	0.11 ± 0.11	3.00 ± 0.10	211.27/195
2022-7-7	59767.9	00096558005	2.02 ± 0.06	0.06 ± 0.12	2.58 ± 0.09	189.73/165
2022-7-8	59768.0	00096558006	1.95 ± 0.07	$0.22^{+0.13}_{-0.12}$	2.57 ± 0.09	168.07/170
2022-7-9	59769.1	00011184222	1.99 ± 0.07	0.11 ± 0.13	2.62 ± 0.10	175.35/165
2022-7-9	59769.8	00096558007	1.98 ± 0.07	0.23 ± 0.14	2.33 ± 0.09	123.78/146
2022-7-10	59770.2	00011184223	2.07 ± 0.07	0.14 ± 0.13	2.34 ± 0.08	133.44/161
2022-7-11	59771.2	00011184224	2.04 ± 0.07	0.08 ± 0.14	2.24 ± 0.09	148.66/151
2022-7-12	59772.3	00096558008	2.11 ± 0.06	0.06 ± 0.11	2.79 ± 0.08	201.07/191
2022-7-13	59773.0	00096558009	2.04 ± 0.08	0.11 ± 0.15	2.46 ± 0.10	112.22/133
2022-7-15	59775.6	00011184225	$2.07^{+0.06}_{-0.07}$	-0.02 ± 0.12	2.54 ± 0.09	168.80/170
2022-7-17	59777.9	00011184226	1.91 ± 0.07	$0.19^{+0.13}_{-0.12}$	2.73 ± 0.10	188.04/175
2022-7-19	59780.0	00011184227	$1.90^{+0.06}_{-0.07}$	0.07 ± 0.12	2.79 ± 0.10	160.98/180
2022-7-20	59780.9	00011184228	$1.98^{+0.06}_{-0.07}$	$-0.01^{+0.12}_{-0.11}$	2.82 ± 0.10	177.72/179
2022-7-23	59783.8	00011184229	1.96 ± 0.06	$0.00^{+0.12}_{-0.11}$	2.76 ± 0.10	163.93/176
2022-7-26	59786.9	00011184230	1.89 ± 0.07	0.06 ± 0.13	3.19 ± 0.13	151.36/160
2022-7-27	59787.6	00011184231	1.98 ± 0.06	$0.03^{+0.12}_{-0.11}$	3.02 ± 0.11	213.93/177
2022-7-29	59789.9	00011184232	$1.75^{+0.19}_{-0.20}$	0.4 ± 0.4	1.75 ± 0.19	17.22/23

Table 4. Continuing Table 2.

Date	MJD	ObsID	α	β	$F_{0.5-10\text{keV}}$ [$10^{-10}\text{ erg cm}^{-2}\text{ s}^{-1}$]	Fit statistic $\chi^2/\text{d.o.f.}$
2022-8-1	59792.8	00011184233	2.00 ± 0.07	-0.08 ± 0.13	2.76 ± 0.11	178.10/157
2022-8-5	59796.0	00011184234	2.09 ± 0.07	-0.13 ± 0.13	2.22 ± 0.09	122.86/155
2022-8-16	59807.0	00011184236	1.98 ± 0.07	-0.02 ± 0.12	2.52 ± 0.09	179.60/172
2022-8-17	59808.9	00011184237	2.14 ± 0.07	$0.04^{+0.14}_{-0.13}$	1.93 ± 0.07	123.65/146
2022-8-18	59809.9	00011184238	2.05 ± 0.07	$0.22^{+0.14}_{-0.13}$	1.82 ± 0.06	179.37/158
2022-8-19	59810.9	00011184239	2.00 ± 0.08	$0.17^{+0.16}_{-0.15}$	1.78 ± 0.08	98.93/114
2022-8-20	59811.9	00011184240	2.11 ± 0.08	$0.06^{+0.16}_{-0.15}$	1.77 ± 0.08	113.33/119
2022-8-23	59815.0	00011184241	2.07 ± 0.08	0.05 ± 0.15	1.54 ± 0.07	103.73/126
2022-8-26	59818.0	00011184242	1.95 ± 0.08	0.12 ± 0.15	1.92 ± 0.08	109.47/126