

The discovery of three hot Jupiters, NGTS-23b, 24b, and 25b, and updated parameters for HATS-54b from the Next Generation Transit Survey

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ABSTRACT

We report the discovery of three new hot Jupiters with the Next Generation Transit Survey (NGTS) as well as updated parameters for HATS-54b, which was independently discovered by NGTS. NGTS-23b, NGTS-24b, and NGTS-25b have orbital periods of 4.076, 3.468, and 2.823 d and orbit G-, F-, and K-type stars, respectively. NGTS-24 and HATS-54 appear close to transitioning off the main-sequence (if they are not already doing so), and therefore are interesting targets given the observed lack of hot Jupiters around sub-giant stars. By considering the host star luminosities and the planets' small orbital separations (0.037–0.050 au), we find that all four hot Jupiters are above the minimum irradiance threshold for inflation mechanisms to be effective. NGTS-23b has a mass of $0.61 M_J$ and radius of $1.27 R_J$ and is likely inflated. With a radius of $1.21 R_J$ and mass of $0.52 M_J$, NGTS-24b has a radius larger than expected from non-inflated models but its radius is smaller than the predicted radius from current Bayesian inflationary models. Finally, NGTS-25b is intermediate between the inflated and non-inflated cases, having a mass of $0.64 M_J$ and a radius of $1.02 R_J$. The physical processes driving radius inflation remain poorly understood, and by building the sample of hot Jupiters we can aim to identify the additional controlling parameters, such as metallicity and stellar age.

Key words: planets and satellites: detection – stars: individual: NGTS-23 (CTOI-77287067), NGTS-24 (TOI-4270), NGTS-25 and HATS-54 – planetary systems.

1 INTRODUCTION

Although hot Jupiters have a low occurrence rate of <1 per cent (Mayor et al. 2011; Hsu et al. 2019; Zhou et al. 2019), they still represent one of the most common categories of planets found due to biases inherent in exoplanet discovery techniques. Understanding the structure, composition, and evolution of these planets is of fundamental importance in understanding the processes of formation, evolution, and migration acting on the wider exoplanet population. Increasing the sample of known hot Jupiters continues to help constrain and parameterize planetary models. For example, there are models that attempt to explain the evolution of planetary radii whilst accounting for the effects of stellar irradiation and heavy

element cores (Fortney, Marley & Barnes 2007; Baraffe, Chabrier & Barman 2008) but they struggle to explain the observed radii of highly-irradiated gas giants.

There are a number of proposed mechanisms to explain inflation in hot Jupiters such as kinetic heating (Guillot & Showman 2002), Ohmic heating through magnetohydrodynamic effects (Batygin & Stevenson 2010; Perna, Menou & Rauscher 2010; Ginzburg & Sari 2016), double diffusive convection (Chabrier & Baraffe 2007), and tidal dissipation (Bodenheimer, Lin & Mardling 2001; Bodenheimer, Laughlin & Lin 2003; Jermyn, Tout & Ogilvie 2017). In order to distinguish between such models, a robust sample of well-characterized hot Jupiters is required to enable statistical studies to constrain these models (such as that by Sestovic, Demory & Queloz 2018). Of particular interest are planets irradiated by fluxes greater than approximately $2 \times 10^5 \text{ W m}^{-2}$, which appear to be preferentially inflated (Demory & Seager 2011; Miller & Fortney

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Table 1. Summary of photometric observations.

Target	Instrument	Nights observed	N_{images}	Exptime (s)	Filter	Comments
NGTS-23	NGTS	2017/08/16–2018/03/18	187,478	10	NGTS	
	TESS	2018/10/19–2018/11/14	1027	1800	TESS	Sector 4
	TESS	2018/11/15–2018/12/11	1176	1800	TESS	Sector 5
	TESS	2020/10/22–2020/11/16	3445	600	TESS	Sector 31
	TESS	2020/11/20–2020/12/16	3589	600	TESS	Sector 32
	EulerCam	2018/12/27	171	90	VG	
	SAAO 1.0 m	2018/12/15	300	60	V	
	SAAO 1.9 m	2018/12/19	295	60	V	
NGTS-24	NGTS	2017/12/10–2018/06/26	235,369	10	NGTS	
	TESS	2019/02/28–2019/03/25	1084	1800	TESS	Sector 9
	TESS	2019/03/26–2019/04/22	1103	1800	TESS	Sector 10
	TESS	2021/03/07–2021/04/01	3467	600	TESS	Sector 36
NGTS-25	NGTS	2018/03/24–2018/11/01	208,001	10	NGTS	
	TESS	2020/07/05–2020/07/30	2898	600	TESS	Sector 27
	SAAO 1.0 m	2019/10/03	940	20	V	
HATS-54 (NGTS-22)	NGTS	2015/12/18–2016/09/01	124,203	10	NGTS	
	TESS	2019/04/23–2019/05/20	1145	1800	TESS	Sector 11
	TESS	2021/04/02–2021/04/28	3461	600	TESS	Sector 37
	EulerCam	2019/03/25	130	~80	VG	
	EulerCam	2019/04/27	136	75	VG	
	SAAO 1.0 m	2018/07/09	700	20	I	
	SAAO 1.0 m	2019/04/30	360	60	V	

2011). However, it is still unclear what impact additional parameters, such as metallicity, have on whether the planet will appear inflated.

We present the discovery of three new hot Jupiters from the Next Generation Transit Survey (NGTS; Wheatley et al. 2018). We also refine the parameters for HATS-54b, which was independently discovered by NGTS. In Section 2, the initial NGTS photometry is described as well as the follow-up observations (both photometry and spectroscopy). Section 3 describes how both the stellar and planetary parameters were derived. The results will be discussed in Section 4 including investigations into the orbital decay timescales (Section 4.1) and the potential inflation (Section 4.2) of these planets. Our final conclusions are given in Section 5.

2 OBSERVATIONS

The planets presented in this work were confirmed using both photometric and spectroscopic observations. These observations are outlined in this section.

2.1 Photometry

All planets presented in this paper were detected using photometry from NGTS. To confirm the planetary nature of the stellar companions, and to further constrain the system parameters, additional photometry was obtained using the Transiting Exoplanet Survey Satellite (TESS; Ricker et al. 2014), EulerCam (Lendl et al. 2012) and the 1-/1.9-m telescopes at the South African Astronomical Observatory (SAAO; Coppejans et al. 2013). A summary of the photometric observations is given in Table 1, and the phase-folded transit light curves can be seen in Fig. 1. Light curves not used in the global fitting can be found in the appendix (Fig. A1).

2.1.1 NGTS

NGTS is a wide-field photometric survey located at ESO's Paranal Observatory in Chile. NGTS has a fully robotic array of twelve 20cm

Newtonian telescopes operating using the unique NGTS filter band (520–890 nm). Each telescope is equipped with a $2K \times 2K$ e2V deep-depleted Andor iKon-L CCD camera. It is optimized to detect and characterize transiting exoplanets around K- and early M-type stars and has achieved sufficient precision to discover exoplanets as small as $3R_{\oplus}$ (Wheatley et al. 2018; West et al. 2019; Smith et al. 2021).

NGTS-23 (CTOI-77287067; $V = 14.010$), NGTS-24 (TOI-4270; $V = 13.192$), and NGTS-25 ($V = 14.266$) were observed with NGTS during the 2018 campaign for a total of 162 nights, 158 nights and 147 nights, respectively. HATS-54 (TOI-1920, NGTS-22; $V = 13.914$) was observed during the 2016 campaign for a total of 119 nights and was detected by NGTS independently of its discovery by Espinoza et al. (2019). All the objects were observed with a 10s exposure time and over 120 000 images were collected for each object. The data were reduced and aperture photometry performed via the CASUTOOLS¹ package, before nightly trends were removed using an adapted version of the SYSREM algorithm (Tamuz, Mazeh & Zucker 2005).

The data was then searched for transit-like events by ORION, a custom implementation of the box-fitting least-squares (BLS) fitting algorithm (Kovács, Zucker & Mazeh 2002; Collier Cameron et al. 2006; Wheatley et al. 2018). ORION identified nine partial and one full transit for HATS-54b (NGTS-22b); ten partial and one full transit for NGTS-23b; nine partial and seven full transits for NGTS-24b; and eleven partial and eight full transits for NGTS-25b. The initial fits to the NGTS data provided by ORION revealed that the depths, widths and shapes of the transits for each object were compatible with transiting hot Jupiters.

Further checks were performed to ensure that the nature of the signal was planetary and not stellar or a false positive. These included comparing the transit depth of consecutive odd and even transits, which can show a depth difference that can demonstrate an

¹<http://casu.ast.cam.ac.uk/surveys-projects/software-release>

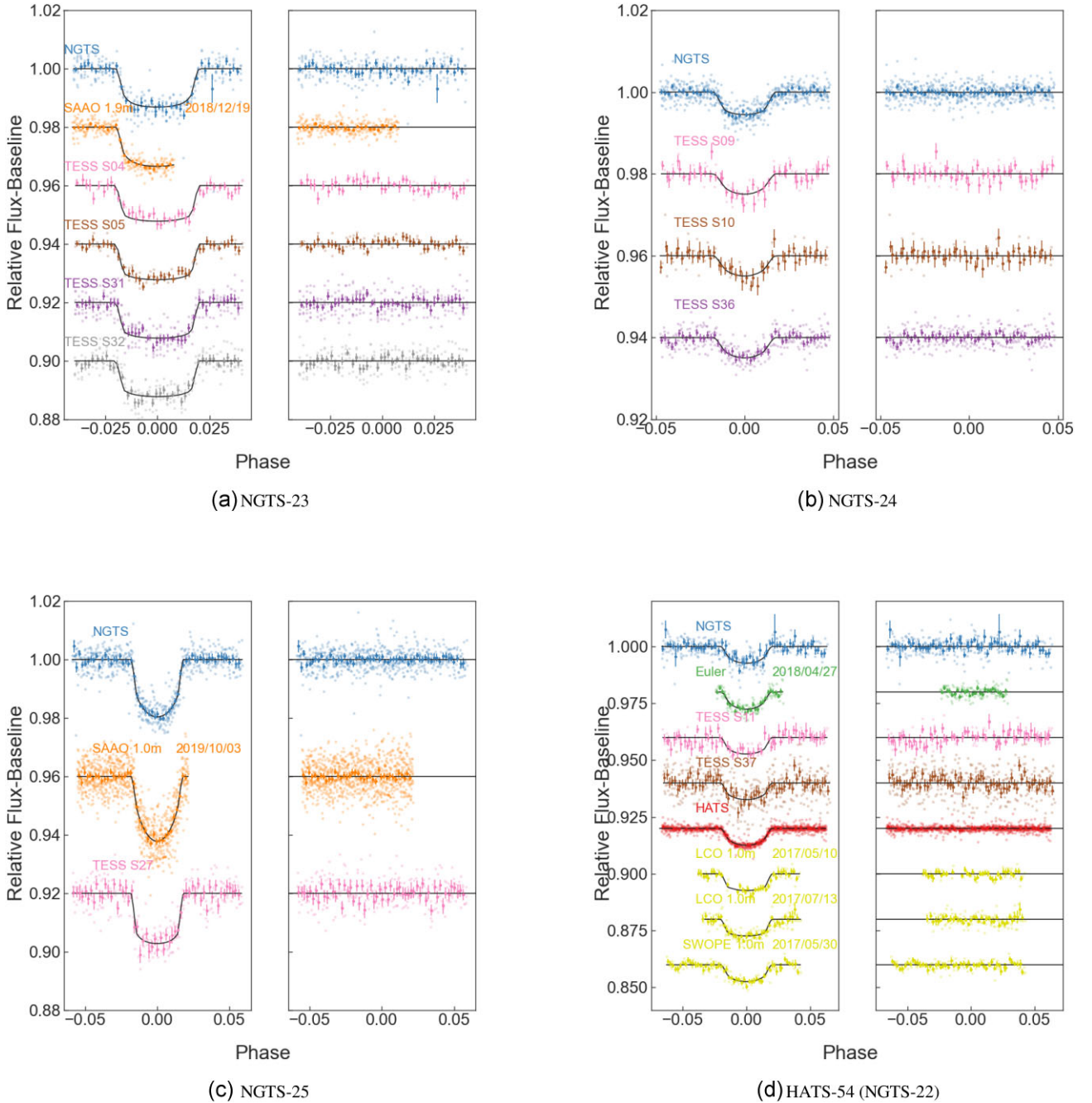


Figure 1. NGTS discovery photometry (blue points) and follow up photometry for each planet. The lighter points show the original photometry except for NGTS, which has been binned to 10-min intervals. The darker datapoints have been binned to 0.002 phase. Photometry from HATS is shown in red while the HATS follow-up photometry that was also used in our combined fits are shown in yellow. The solid black line shows the light-curve model for the planet. Residuals are shown to the right of the light-curves. Light curves offset for clarity.

eclipsing binary system has been mis-folded on half the true period. Also, secondary eclipses around 0.5 phase in the phase-folded light curves were searched for as this could also indicate that the orbiting companion is stellar in nature. All candidates passed this vetting and thus follow-up observations were taken.

2.1.2 SAAO

Between 2018 July and 2019 October, we took follow-up photometry for each target (except for NGTS-24) using either the 1.0-m or 1.9-m

telescope at SAAO (Copejans et al. 2013). The 1.0-m telescope was equipped with the full-frame transfer CCD Sutherland High-speed Optical Camera (SHOC), ‘SHOC’ n’awe’, that has a $2.85 \text{ arcmin} \times 2.85 \text{ arcmin}$ field of view. The 1.9-m telescope was equipped with the ‘SHOC’ n’disbelief’ optical camera, that had a slightly narrower field of view of $1.3 \text{ arcmin} \times 1.3 \text{ arcmin}$. We were able to simultaneously observe at least one comparison star of similar brightness for each object. All observations were taken using either the I band or V band filters.

Two light curves were obtained for NGTS-23 one each on the 1.0-m and 1.9-m telescopes, on the 2018 December 15 and 2018 December 19, respectively. Both observations used an exposure time of 60-s with the V band filter. Both observations ended early due to high humidity however, we obtained 300 images and 295 images, respectively.

A single V-band light curve was taken for NGTS-25b on the 2019 October 3 using the 1.0-m telescope. 940 images were obtained using an exposure time of 20-s. However, the middle of these observations were affected by bad seeing.

Two observations of HATS-54b were conducted using the 1.0-m telescope on the nights of the 2018 July 9 and the 2019 April 30. The 2018 observation obtained 700 images with an exposure time of 20 s using the I band filter. This observation was ended early due to high humidity. The 2019 observation obtained 360 images with an exposure time of 60s using the V band filter.

Each light curve was bias and flat-field corrected using the local Python-based SAAO SHOC pipeline, which uses IRAF (Tody 1986) photometry tasks (PyRAF- Science Software Branch at STScI 2012) and facilitates the extraction of raw and differential light curves. We used the Starlink package AUTOPHOTOM (Currie et al. 2014) to perform aperture photometry on both our target and comparison stars, and chose apertures that gave the maximum signal-to-noise ratio. Background apertures were adjusted to account for changes in the stellar point-spread function over the night as the atmospheric conditions varied. Finally, the measured fluxes of the comparison stars for each object were used for differential photometry of our targets.

2.1.3 EulerCam

We also observed NGTS-23 and HATS-54 with EulerCam (Lendl et al. 2012) on the 1.2-m *Euler Telescope* at La Silla Observatory. All observations presented in this work were done through a *Geneva V* filter (Rufener & Nicolet 1988). NGTS-23 was observed on the 2018 December 27, with 171 images being acquired with an exposure time of 90 s and a 0.1-mm defocus.

HATS-54 was observed twice with EulerCam. The first observation was performed on the 2019 March 5, with 130 images being acquired with an exposure time of approximately 80-s. No defocus was applied for this observation. The second observation was carried out on the 2019 April 27. 136 images were acquired with an exposure time of 75-s and a defocus of 0.5 mm. Only a partial transit of NGTS-23 was observed, whereas the observations of HATS-54b covered one full transit and one partial transit.

For both targets, the data were reduced using the standard procedure of bias subtraction and flat field correction. The aperture photometry as well as x- and y-positions, full width half-maximum (FWHM), airmass, and sky background of the target star were extracted using the PyRAF-based routines. The comparison stars and the photometric aperture radius were carefully chosen in order to reduce the RMS in the out-of-transit light-curve parts.

2.1.4 TESS

TESS is a space-based NASA survey telescope that searches for transiting planets around bright stars (Ricker et al. 2014). It has a wide field of view, with four $24 \times 24^\circ$ cameras, each equipped with four $2k \times 2k$ CCDs. Its typical observing baseline of ≈ 27 d makes it well-suited to detecting short-period transiting exoplanets. We searched for our candidates in the TESS full frame images (FFIs)

using the TESSCut² tool and found that all four stars were observed by TESS, with observations between the 2018 October 19 and the 2021 April 28. NGTS-23 was observed in Sectors 4, 5, 31, and 32; NGTS-24 was observed in Sectors 9, 10, and 36; HATS-54 was observed in Sectors 11 and 37 and NGTS-25 was observed in Sector 27. HATS-54b falls onto a CCD in TESS Sector 38 but we do not use it as it appears to be in the overscan region of the camera. A summary of this information can be found in Table 1.

Observations before 2020 July were within the primary mission of TESS, so the TESS FFIs for these sectors were observed at 30-min cadence. For observations that are in the first extended mission of TESS (which is set to finish in 2022 September) the TESS FFIs were observed at 10-min cadence. This means that all of the stars were observed at both 10-min and 30-min cadence with TESS apart from NGTS-25, which was only observed with 10-min cadence. For each of our candidates we used an automatically determined optimal aperture to exclude neighbouring stars, ranging from three to eight pixels. However, the large pixel scale of TESS (21 arcsec) meant that there was still some slight dilution of the FFI light curves.

For each candidate, we found that the best BLS period from TESS was consistent with the ORION value for the NGTS photometry. TESS detected six full transits for NGTS-25, eighteen full transits for NGTS-24, two partial and twenty full transits for NGTS-23, and seventeen full transits for HATS-54.

For all four objects, the transits observed with TESS, as well as the other photometric follow-up, was consistent with planetary companions. Therefore, all objects were subsequently followed up with spectroscopy to determine the mass of the companions.

2.2 Spectroscopy

In addition to the photometric follow-up, spectroscopy was also used to obtain radial velocity measurements for all of the systems. This allowed for the mass, and therefore the density, of the respective transiting companion to be determined. A summary of the spectroscopic observations is given in Table 2 and phase-folded radial-velocity curves can be seen in Fig. 2.

The three different échelle spectrographs used are all located at La Silla Observatory in Chile. The High Accuracy Radial velocity Planet Searcher (HARPS; Mayor et al. 2003) has a resolving power of $R = 115\,000$ and is mounted on the ESO 3.6-m telescope. The $R = 60\,000$ CORALIE spectrograph (Queloz et al. 2001a) is installed on the Swiss 1.2-m Leonhard Euler telescope while the Fibre-fed Extended Range Optical Spectrograph (FEROS; Kaufer & Pasquini 1998) is mounted on the MPG/ESO 2.2-m telescope and has $R = 48\,000$. All four systems were observed by a minimum of two spectrographs and have at least 11 datapoints (see Table 2).

The HARPS data was reduced via the offline data reduction system (DRS) before being cross-correlated with a binary mask to extract radial velocities (Baranne et al. 1996). For NGTS-24 and NGTS-23 a G2 mask was used whereas a G9 mask and a K0 mask were used for HATS-54 and NGTS-25, respectively. This procedure was also employed for the CORALIE data, for which the spectra were reduced using the standard CORALIE DRS. For the data collected with FEROS, the CERES reduction pipeline (Brahm, Jordán & Espinoza 2017) performed a radial velocity extraction by the same method.

All four targets had semi-amplitudes consistent with a Jupiter-mass companion ($K \sim 100 \text{ m s}^{-1}$) and had variations in phase with the derived period from the photometric data. We investigated whether

²<https://mast.stsci.edu/tesscut/>

Table 2. Summary of the spectroscopic observations. $\text{SNR}_{\text{combined}}$ is the SNR of the co-added spectra that were used in the analysis outlined in Section 3.

Target	Instrument	Nights observed	N_{spectra}	$\text{SNR}_{\text{combined}}$	Programme
NGTS-23	HARPS	2019/09/12–2020/01/20	7	29.13	0104.C-0588
	FEROS	2018/12/24–2020/01/04	8	–	0102.A-9011 & 0103.A-9004 & 0103.C-0719 & 0104.A-9012
NGTS-24	CORALIE	2018/10/27–2018/11/06	3	–	N/A
	FEROS	2019/12/30–2020/01/04	6	–	0104.A-9012
	CORALIE	2019/04/22–2019/12/26	5	–	N/A
NGTS-25	HARPS	2021/06/12–2021/09/13	12	20.92	105.20G9
	CORALIE	2019/08/09–2021/08/22	2	–	N/A
HATS-54 (NGTS-22)	HARPS	2019/05/10–2019/06/08	5	26.45	0103.C-0719
	FEROS	2019/04/18–2019/04/26	8	–	0103.A-9004
	CORALIE	2019/02/23–2019/04/28	6	–	N/A

there were any correlations between the line bisector (BIS) of the CCF to ensure that the signals were of planetary origin and did not arise from stellar activity (e.g. stellar spots) or a blended eclipsing binary (Queloz et al. 2001b). No evidence was found to contradict the existence of a planetary companion.

3 ANALYSIS

The measured planetary parameters are dependent on the characterization of the host star. Therefore, accurately constraining the properties of the host star is vital to determine accurate planetary parameters.

3.1 Stellar analysis

Most of our stellar parameters were obtained by fitting the spectral energy distributions (SEDs). However, to obtain suitable priors to constrain the SED fits, we first analysed spectra for each object. For the spectral analysis we used the open source framework *ISPEC* (Blanco-Cuaresma et al. 2014; Blanco-Cuaresma 2019). *ISPEC* allows users to choose from a list of models for spectral synthesis. We chose to employ the *SPECTRUM* (Gray & Corbally 1994) radiative transfer code and the solar abundance model from Asplund et al. (2009). Additionally, we use the Gaia-ESO Survey (GES) line list (version 6.0; Heiter et al. 2021), which covers most of the wavelength range of our spectra (420–920 nm), and the MARCES.GES model atmosphere (Gustafsson et al. 2008).

We used our high-resolution HARPS data for the analysis of NGTS-23, NGTS-25, and HATS-54. For each object, we shift the spectra to the laboratory frame of reference before co-adding them to produce a high signal-to-noise ratio (SNR) combined spectrum. The SNRs of these spectra are reported in Table 2. For NGTS-24, we attempted to follow a similar procedure using the CORALIE/FEROS spectra. However, the SNR of the co-added CORALIE spectra were too low to obtain meaningful results. For FEROS, the Th-Ar lamp used for calibration appears to have contaminated the stellar spectrum. Therefore, we are unable to use the FEROS spectra for spectral analysis but they still yield reliable results for the radial velocity measurements.

We determined the stellar effective temperature, (T_{eff}) and the surface gravity ($\log g$), via fits to the $\text{H}\alpha$, Na I D, and Mg I b lines whereas Fe I and Fe II lines were used to determine the metallicity ($[Fe/H]$), alpha enhancement ($[\alpha/Fe]$), microscopic velocity (v_{mic}),

macroscopic velocity (v_{mac}), and the rotational line-of-sight broadening ($v \sin i_*$).

3.1.1 Gaia DR3

For all four objects, we obtained additional photometric and astrometric data from *Gaia* DR3 (Gaia Collaboration 2016, 2022).

Our host stars show no evidence of being unresolved binaries, with all four targets having an astrometric excess noise of approximately zero as well as a low astrometric goodness of fit in the along-scan direction (*GOF_{AL}*) in *Gaia* DR3 (Evans 2018) and RUWE (re-normalized unit weight error) of approximately one. They also pass the photometric filtering processes outlined by Arenou et al. (2018) to identify blended stars.

The Gaia positions, proper motions, and absolute radial velocities of each star were used to determine the Galactic velocity components (U_{LSR} , V_{LSR} , W_{LSR}), assuming a Local Standard of Rest of $U_{\text{LSR}} = (11.1, 12.14, 7.25 \text{ km s}^{-1})$; Schönrich, Binney & Dehnen (2010). By using the selection criteria for kinematically thin-disc ($v_{\text{tot}} < 50 \text{ km s}^{-1}$) and kinematically thick-disc ($70 < v_{\text{tot}} < 180 \text{ km s}^{-1}$) objects (Gaia Collaboration 2018), we conclude that three of the stars belong to the thin disc population while NGTS-23 is in-between the two populations (see Table 3). This appears consistent with our determination that NGTS-23 is the only metal-poor star out of our systems (see Table 4).

3.1.2 SED fitting

We performed fits to the SEDs of each object to obtain the stellar parameters for each system. For this, we employed the Python tool *ARIADNE* (Vines & Jenkins 2022). *ARIADNE* fits catalogue photometry to different atmospheric model grids such as Phoenix V2 (Husser et al. 2013), BT-Settl, BT-Cond, BT-NextGen (Hauschildt, Allard & Baron 1999; Allard, Homeier & Freytag 2012), Castelli & Kurucz (2004), and Kurucz (1993).

The SEDs were modelled by interpolating the model grids in $T_{\text{eff}} - \log g - [Fe/H]$ space. Any underestimated uncertainties in the photometry were accounted for by including an excess noise term. Gaussian priors for T_{eff} , $\log g$, and $[Fe/H]$ were taken from our stellar analysis with *ISPEC* (see Table 4), whilst Gaussian priors for the stellar radius and distance were obtained from the *Gaia* DR2/DR3 values, although we used an inflated value for the radius error. For NGTS-24b, we used values from *Gaia* but with inflated

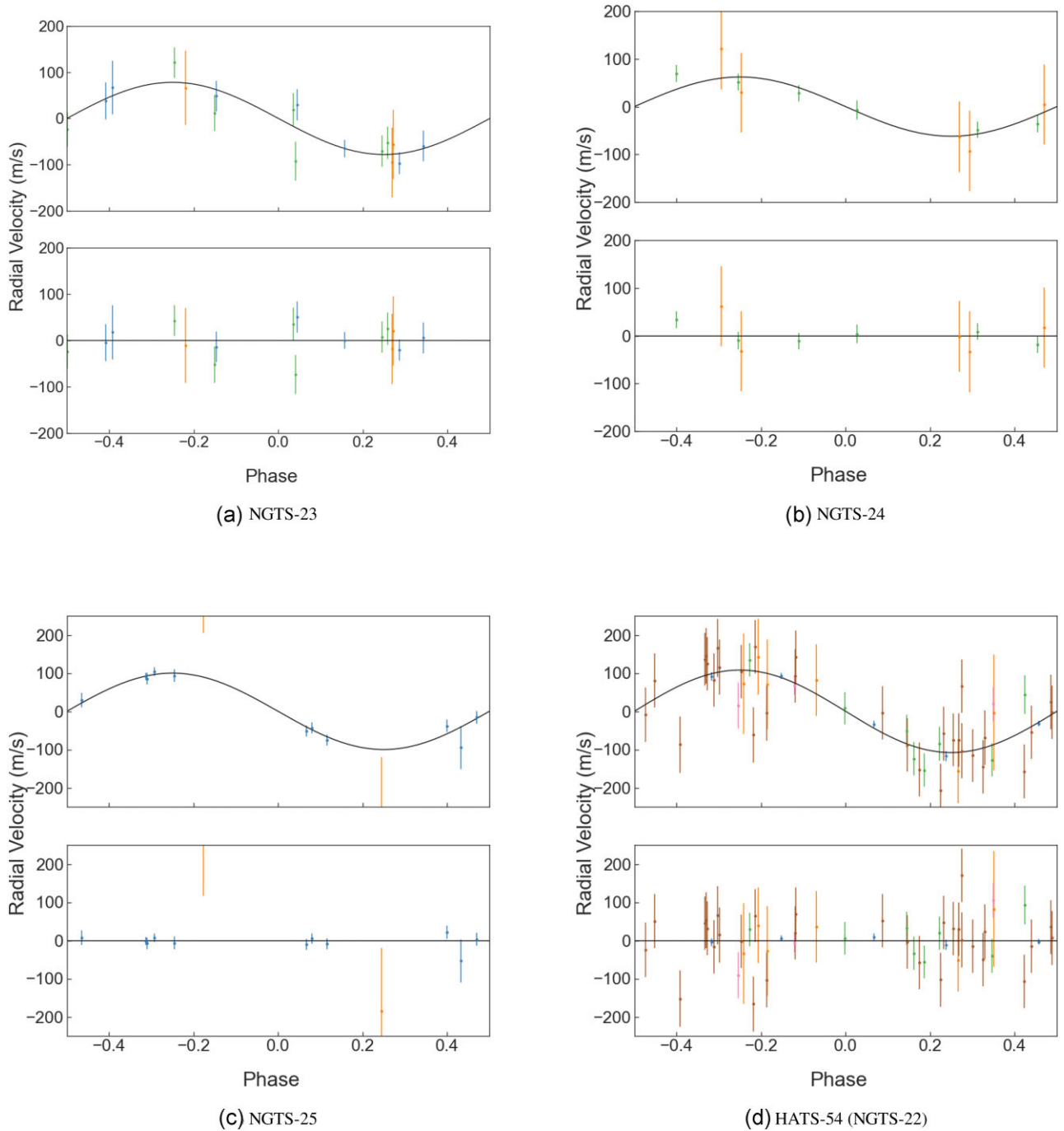


Figure 2. Phase-folded radial velocity measurements for NGTS-23, NGTS-24, NGTS-25, and HATS-54 (NGTS-22). These include data obtained from HARPS (blue points), FEROS (green points), and CORALIE (orange points). HARPS spectroscopy used by the HATS team is shown in peak while the FEROS spectroscopy used by HATS is shown in brown. The radial velocity model is shown by the solid black line. The residuals are shown below the radial velocity measurements.

Table 3. The galactic velocity components and total space velocities for the four systems presented in this paper.

Target	$U_{\text{LSR}}, V_{\text{LSR}}, W_{\text{LSR}} (\text{km s}^{-1})$	$v_{\text{tot}} (\text{km s}^{-1})$
NGTS-23	−54.39, 19.95, −1.52	57.95
NGTS-24	5.11, 25.74, 2.54	26.37
NGTS-25	−14.40, 27.52, 11.98	33.29
HATS-54 (NGTS-22)	31.40, −38.39, −4.24	49.78

errors as were not able to obtain results using iSPEC. Inflated errors were chosen so as not to unnecessarily bias the results from ARIADNE. In addition, A_V was limited to the maximum line-of-sight value taken from galactic dust maps (Schlegel, Finkbeiner & Davis 1998; Schlafly & Finkbeiner 2011). The priors for the excess noise parameters were set to a normal distribution with a mean of zero and a variance equal to five times the associated uncertainty. The mass value was estimated from MIST isochrones (Choi et al. 2016)

Table 4. Stellar properties of the four systems.

Property	NGTS-23	NGTS-24	NGTS-25	HATS-54 (NGTS-22)	Source
				Updated	Discovery
2MASS ID	J04414374-4002404	J11141533-3754365	J20294038-3901565	J13223237-4441196	2MASS
GAIA Source ID	4816462091938129408	5397086969654628864	6682513094425008640	6087996849371141248	<i>Gaia</i> DR2
TIC ID	77287067	151920165	290742484	363851359	TIC8
RA	04 ^h 41 ^m 43.6 ^s	11 ^h 14 ^m 15.3 ^s	20 ^h 29 ^m 40.3 ^s	13 ^h 22 ^m 32.4 ^s	<i>Gaia</i> DR3
DEC	−40°02′41.0 arcsec	−37°54′36.5 arcsec	−39°01′55.5 arcsec	−44°41′20.0 arcsec	<i>Gaia</i> DR3
μ_{RA} (mas y ^{−1})	0.86 ± 0.01	−0.69 ± 0.01	5.17 ± 0.02	−3.50 ± 0.02	<i>Gaia</i> DR3
μ_{DEC} (mas y ^{−1})	12.27 ± 0.02	0.57 ± 0.01	6.03 ± 0.02	−7.96 ± 0.02	<i>Gaia</i> DR3
Parallax (mas)	0.95 ± 0.01	1.39 ± 0.01	1.90 ± 0.02	1.36 ± 0.02	<i>Gaia</i> DR3
V (mag)	14.010 ± 0.062	13.192 ± 0.063	14.266 ± 0.048	13.914 ± 0.027	APASS
B (mag)	14.578 ± 0.024	13.948 ± 0.053	15.146 ± 0.052	14.710 ± 0.070	APASS
g (mag)	14.291 ± 0.028	13.561 ± 0.014	14.656 ± 0.048	14.258 ± 0.045	APASS
r (mag)	13.940 ± 0.049	13.007 ± 0.011	14.000 ± 0.049	13.694 ± 0.041	APASS
i (mag)	13.845 ± 0.043	12.799 ± 0.051	13.844 ± 0.092	13.497 ± 0.074	APASS
G (mag)	13.9406 ± 0.0002	13.0410 ± 0.0002	14.0742 ± 0.0003	13.7621 ± 0.0003	<i>Gaia</i> DR3
J (mag)	13.056 ± 0.023	11.879 ± 0.019	12.818 ± 0.019	12.611 ± 0.021	2MASS
H (mag)	12.797 ± 0.024	11.577 ± 0.022	12.417 ± 0.020	12.273 ± 0.024	2MASS
K (mag)	12.730 ± 0.035	11.510 ± 0.019	12.324 ± 0.022	12.170 ± 0.017	2MASS
W1 (mag)	12.710 ± 0.024	11.455 ± 0.023	12.287 ± 0.023	12.157 ± 0.023	WISE
W2 (mag)	12.744 ± 0.023	11.505 ± 0.020	12.328 ± 0.024	12.165 ± 0.023	WISE
T (mag)	13.586 ± 0.006	12.570 ± 0.006	13.546 ± 0.006	13.276 ± 0.006	TIC8
Spectral Type	F9V	G2IV	K0V	G6V	ARIADNE
T_{eff}	6057 ± 64	5820 ⁺⁹³ _{−8}	5321 ± 58	5621 ± 70	5702 ± 26
log g	4.04 ± 0.15	4.09 ^{+0.05} _{−0.04}	4.48 ± 0.14	4.20 ± 0.20	4.240 ± 0.023
[Fe/H]	−0.26 ± 0.11	0.27 ^{+0.18} _{−0.03}	0.12 ± 0.05	0.37 ± 0.07	0.396 ± 0.031
[α /Fe]	0.07 ± 0.17	—	−0.08 ± 0.10	−0.15 ± 0.10	—
$v \sin i$ (km s ^{−1})	1.46 ± 1.08	—	1.96 ± 0.68	<2.50	3.83 ± 0.42
v_{mic} (km s ^{−1})	1.24 ± 0.20	—	0.89 ± 0.21	0.88 ± 0.23	0.948 ± 0.034
v_{mac} (km s ^{−1})	5.71 ± 0.84	—	3.24 ± 0.78	4.49 ± 0.68	3.61 ± 0.11
R_* (R _⊙)	1.24 ± 0.03	1.64 ^{+0.03} _{−0.09}	0.86 ^{+0.02} _{−0.05}	1.23 ± 0.02	1.317 ± 0.036
M_* (M _⊙)	1.01 ^{+0.06} _{−0.07}	1.26 ^{+0.05} _{−0.21}	0.91 ^{+0.03} _{−0.04}	1.05 ± 0.05	1.097 ± 0.022
ρ_* (g cm ^{−3})	0.755 ^{+0.027} _{−0.034}	0.406 ± 0.059	2.219 ^{+0.085} _{−0.110}	0.795 ± 0.052	0.678 ± 0.054
L_* (L _⊙)	1.920 ^{+0.093} _{−0.106}	2.792 ^{+0.215} _{−0.289}	0.511 ^{+0.045} _{−0.052}	1.302 ^{+0.049} _{−0.055}	1.631 ^{+0.114} _{−0.076}
d (pc)	991 ⁺¹⁹ _{−23}	725 ⁺³³ _{−28}	517 ⁺¹⁸ _{−22}	720 ⁺¹¹ _{−12}	769 ± 21
A_v (mag)	0.019 ^{+0.029} _{−0.016}	0.249 ^{+0.035} _{−0.067}	0.077 ^{+0.061} _{−0.054}	0.114 ^{+0.042} _{−0.038}	0.279 ± 0.018
Age (Gyrs)	6.39 ^{+1.93} _{−1.60}	4.77 ^{+3.56} _{−0.91}	3.49 ^{+3.49} _{−3.37}	7.96 ^{+2.97} _{−2.26}	6.60 ± 0.76

Notes. *denotes that, for NGTS-24b, ARIADNE was used instead of iSPEC for reasons outlined in Section 3.1. The second last column contains the stellar parameters for HATS-54 from the discovery paper (Espinoza et al. 2019). The final column gives the data source, for which the following references are given: 2MASS (Skrutskie et al. 2006); APASS (Henden & Munari 2014); Gaia (Gaia Collaboration 2016, 2018, 2022); WISE (Wright et al. 2010); TIC8 (Stassun et al. 2019).

using the previously derived values. The derived stellar parameters are listed in Table 4.

3.1.3 Stellar rotation

As well as modelling the stellar parameters, we also wished to constrain the stellar rotation period. Coupled to a knowledge of the stellar radius and $v \sin i_*$, determination of the stellar rotation period can then allow constraints to be placed on the inclination angle of the stellar rotation axis. This can enable misaligned star-planet systems to be identified (Watson et al. 2010).

The first method used to constrain the rotation period was to spectroscopically measure the chromospheric activity of the star. Using the formulae described in Lovis et al. (2011), the log R'_{HK} was measured from the co-added HARPS spectra used in Section 3.1. Therefore, the only system we were not able to obtain a log R'_{HK}

value for was NGTS-24b. From these co-added spectra, we found log R'_{HK} values of −5.43, −4.61, and −4.74 for NGTS-23, NGTS-25, and HATS-54, respectively. Using the relation from Noyes et al. (1984), the most probable stellar rotation periods are predicted to be 27 d for NGTS-23, 24 d for NGTS-25, and 29 d for HATS-54.

Our second approach used the $v \sin i_*$ of the star measured from iSPEC. By assuming that the stellar inclination angle, i_* , is 90° (i.e. $\sin i_* = 1$), an upper limit can be placed on the stellar rotation period when the stellar radius is known. NGTS-23 and NGTS-25 were found to have a $v \sin i_*$ of 1.46 ± 1.08 km s^{−1} and 1.96 ± 0.68 km s^{−1}, respectively, which correspond to upper limits of the rotation periods of 43 ± 32 d and 22 ± 8 d. We were only able to obtain an upper-limit to the $v \sin i_*$ for HATS-54, finding it to be <2.50 km s^{−1}. This corresponds to an upper limit of 25 d. However, if the value of $v \sin i_*$ is lower then this upper limit will increase, which means that HATS-54 is likely a slow rotator. These are consistent with the

previous estimates from the $\log R'_{\text{HK}}$ and appear consistent with the age and spectral types of the stars.

Finally, we searched the Generalised Lomb–Scargle periodograms (Lomb 1976; Scargle 1982; Zechmeister & Kürster 2009) of the NGTS photometry after the transits were removed (see appendix D1 for plots). NGTS-23 has significant peaks at approximately 15 and 33 d. The 33-d peak is more consistent with the previous estimate of 27 d from the $\log R'_{\text{HK}}$. The most significant peak for NGTS-24 is at approximately 30 d. However, it is possible that this peak is indicative of the moon contaminating the photometry. However, given that the next most significant peak is about 47 d, it appears likely that NGTS-24 is slowly rotating. NGTS-25 has peaks in the periodogram at approximately 14 and 28 d, which could be caused by moonlight in the photometry. There is also a significant peak between 40 and 50 d but this is not consistent with the upper limit from the $v \sin i_*$ and the estimate of 24 d from the $\log R'_{\text{HK}}$ so it is unlikely to correspond with the rotation period. Finally, HATS-54 has numerous significant peaks. There is a significant peak at roughly 27 d, which is broadly consistent with the previous estimates. We conclude, like the other stars presented in this paper, that HATS-54 is slowly rotating.

As a final check, we explored alternative methods, including autocorrelation and Gaussian Process (GP) regression in addition to LS periodogram (Gillen et al. 2020), via the rotation Toolkit (RoTo),³ but do not detect clear rotation patterns beyond the aforementioned suggestive signals.

3.1.4 Stellar properties

Using the parameters obtained for the stellar radii, masses, effective temperatures, and luminosities, we can compare our host stars to the general population of stars that host planets, where the planets have defined masses and radii. This is a sample of approximately 950 stars, and a stellar mass-radius diagram and a Hertzsprung–Russell diagram can be seen in Figs 3 and 4, respectively.

From these figures, it appears that NGTS-25 is on the main-sequence and lies in the general population of stars. NGTS-23 and HATS-54 both occupy a sparse region of mass-radius parameter space. Finally, NGTS-24 is in an even less populated region of mass-radius space and appears to be transitioning off the main-sequence if it has not already done so. This appears a reasonable conclusion given its age and spectral type (see Table 4). Therefore, it is an interesting target given the rarity of hot Jupiters around sub-giant stars (Villaver & Livio 2009; Hansen 2010; Schlaufman & Winn 2013). We also conclude that HATS-54 and NGTS-23 are both towards the end of their lifetimes on the main-sequence and will evolve off the main-sequence within a few Gyrs. Again, this seems to be in good agreement with their ages and spectral type. These conclusions are further supported from the findings in Section 4.1, in which all three stars show significant evolution within relatively short periods of time.

3.2 Global modelling

We determined the physical parameters of the systems, including planet radius and mass, via a simultaneous fit to the photometric and spectroscopic data using the open-source astronomy software package ALLESFITTER (Günther & Daylan 2019, 2021). ALLESFITTER combines the packages ELLC (light curve and RV models; Maxted 2016), EMCEE (MCMC sampling; Foreman-Mackey et al. 2013),

³<https://github.com/joshbriegal/roto>

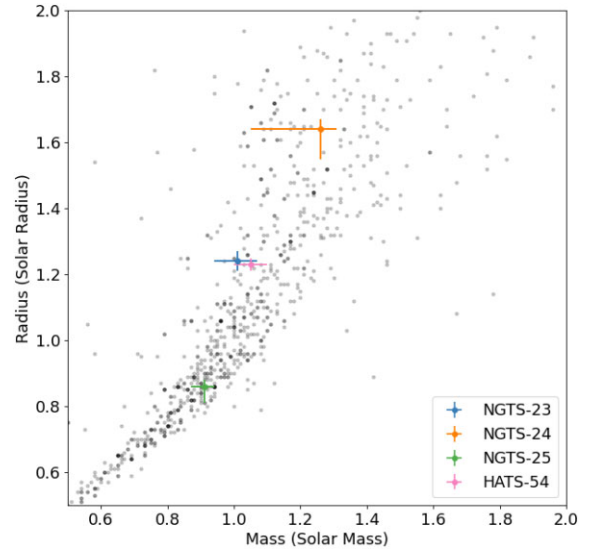


Figure 3. Stars that host exoplanets with defined radii and masses (grey points) with errorbars suppressed for clarity. The stars presented in this paper are plotted as coloured points with errorbars. It can be seen that three of the stars presented in this paper lie in relatively sparse regions of (stellar) mass-radius space while NGTS-25 lies in the general population. The stellar sample was obtained from the NASA Exoplanet Archive (<https://exoplanetarchive.ipac.caltech.edu/>).

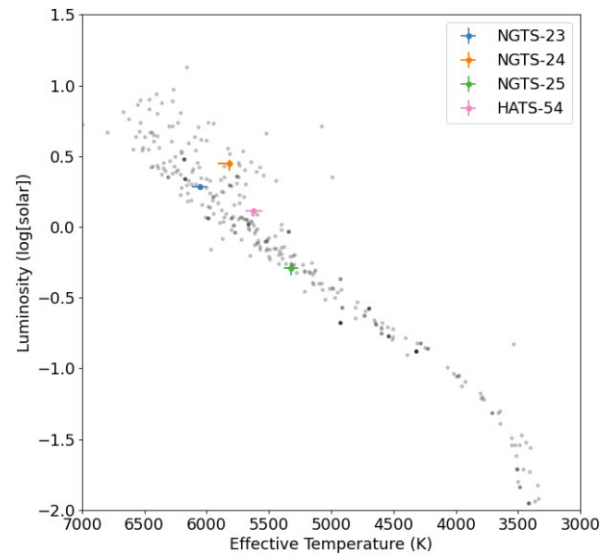


Figure 4. A Hertzsprung–Russell diagram for the same stellar sample described in Fig. 3. The errorbars are suppressed for clarity. The stars presented in this paper are plotted as coloured points with errorbars. Of particular note is NGTS-24, which appears to be beginning to transition off the main-sequence. HATS-54 also appears to be towards the end of its main-sequence lifetime.

DYNesty (nested sampling; Speagle 2020), and CELERITE (GP models; Foreman-Mackey et al. 2017). This combination allows ALLESFITTER to model a variety of signals, including stellar variability and transit-timing variations.

Given the length of time for the initial MCMC fits to converge, a nested sampling approach was taken to produce the global fits. However, quick MCMC fits (100 walkers and 5000 steps) to the

Table 5. Planetary parameters for the four systems presented in this paper. The final two columns display the updated planetary parameters for HATS-54b/NGTS-22b alongside those from the discovery paper (Espinoza et al. 2019).

Property	NGTS-23b	NGTS-24b	NGTS-25b	HATS-54b / NGTS-22b	
				Updated	Discovery
P (days)	4.0764326 ± 0.0000041	3.4678796 ± 0.0000070	2.8230930 ± 0.0000033	2.5441765 ± 0.0000024	2.5441828 ± 0.0000043
T_C (BJD)	$2458431.77134 \pm 0.00048$	2458100.6229 ± 0.0011	$2458204.26373 \pm 0.000502457780.01023 \pm 0.00061$	$2457780.01102 \pm 0.00089$	
T_{14} (hours)	3.919 ± 0.031	$2.929^{+0.095}_{-0.083}$	2.466 ± 0.025	$2.468^{+0.044}_{-0.041}$	2.5008 ± 0.0456
R_p/R_*	$0.1048^{+0.0011}_{-0.0011}$	$0.0768^{+0.0028}_{-0.0023}$	$0.1231^{+0.0011}_{-0.0013}$	0.0848 ± 0.0014	0.0832 ± 0.0025
a/R_*	$8.72^{+0.10}_{-0.13}$	$6.37^{+0.29}_{-0.31}$	$9.78^{+0.12}_{-0.16}$	6.48 ± 0.14	6.15 ± 0.16
b	$0.134^{+0.094}_{-0.091}$	0.821 ± 0.026	$0.113^{+0.097}_{-0.072}$	$0.715^{+0.024}_{-0.026}$	$0.740^{+0.022}_{-0.023}$
i ($^\circ$)	$89.12^{+0.53}_{-0.64}$	$82.61^{+0.52}_{-0.61}$	$89.34^{+0.42}_{-0.58}$	83.67 ± 0.34	83.08 ± 0.36
e	0 (fixed)	0 (fixed)	0 (fixed)	0 (fixed)	<0.126
K (km s $^{-1}$)	0.078 ± 0.011	$0.062^{+0.011}_{-0.010}$	0.100 ± 0.006	0.108 ± 0.006	0.105 ± 0.014
$R_p(R_J)$	1.267 ± 0.030	$1.214^{+0.058}_{-0.062}$	$1.023^{+0.035}_{-0.052}$	1.015 ± 0.024	1.067 ± 0.052
$M_p(M_J)$	0.613 ± 0.097	$0.520^{+0.120}_{-0.110}$	$0.639^{+0.058}_{-0.052}$	0.753 ± 0.057	0.76 ± 0.10
ρ_p (g cm $^{-3}$)	$0.373^{+0.070}_{-0.064}$	$0.359^{+0.120}_{-0.087}$	$0.892^{+0.099}_{-0.089}$	$0.940^{+0.120}_{-0.110}$	0.77 ± 0.16
a (au)	0.0504 ± 0.012	$0.0479^{+0.0028}_{-0.0030}$	$0.0388^{+0.0014}_{-0.0020}$	0.0370 ± 0.0010	0.03763 ± 0.00024
T_{eq} (K)	1327^{+17}_{-16}	1499^{+41}_{-36}	1101^{+15}_{-14}	1429 ± 24	1625 ± 22
Irradiation (W m $^{-2}$)	$(1.03 \pm 0.07) \times 10^6$	$(1.66^{+0.24}_{-0.26}) \times 10^6$	$(4.62^{+0.63}_{-0.57}) \times 10^5$	$(1.29^{+0.08}_{-0.09}) \times 10^6$	$(1.57^{+0.09}_{-0.08}) \times 10^6$

NGTS data were used to refine priors on parameters such as transit epoch and transit depth. Gaussian priors were placed on the stellar mass, radius, and temperature using the results from the SED fitting (see Table 4). We adopted a quadratic limb darkening law for the photometric data as parameterized in Kipping (2013). Gaussian processes with a Matern 3/2 kernel were used to fit any out-of-transit variability and for this reason we removed light curves from the final models that did not have a sufficient out-of-transit baseline. Therefore, for NGTS-23b we did not include the EulerCam data or the SAAO 1.0-m data in our final fits, while for HATS-54b we did not use the EulerCam data taken on the 2019/03/25 or either night of the photometry taken by the SAAO 1.0-m in our final fits. Transits that were not included in the final models can be found in appendix A1 for completeness. However, for HATS-54b we made use of both photometry and radial velocity measurements from Espinoza et al. (2019) in our global models. Furthermore, as multiple spectrographs were used to obtain radial velocity measurements for each system, we fit for a constant instrumental velocity offset for each spectrograph.

For each system, we ran multiple fits before performing our final fit. We first tested removing RV points that were measured during transit as these points could have been affected by the Rossiter-McLaughlin effect (McLaughlin 1924; Rossiter 1924). We found that removing these points did not significantly alter the results and therefore these in-transit RV points were included in the final fits.

Finally, we ran two additional fits for each planet – one where the planet was fixed on a circular orbit (i.e. $e = 0$), and another where the eccentricity was allowed to vary freely. Each of the latter fits resulted in non-zero eccentricity values. However, many small eccentricity values are spurious and Lucy & Sweeney (1971) define a probabilistic test to determine whether a small e is statistically significant from 0. By adopting a 5 percent level of significance, they find that if the condition $e > 2.45 \sigma_e$ is satisfied (where σ_e is the observational uncertainty on e), then one can be confident that the measured eccentricity is real. We find that all of our measurements fail to meet this criteria, and we therefore fixed the eccentricity to be zero in our final fits. This is in-line with both theoretical predictions,

and the observed population of hot Jupiters, where orbital eccentricity is often consistent with zero (e.g. Anderson et al. 2012).

We use ALLESFITTER to fit our data for the key physical parameters of each system, including the planet’s orbital period P , the radius ratio R_p/R_* , and the radial velocity semi-amplitude K . A full list of physical properties and their values for each planet can be found in Table 5. Fig. 1 shows the phase-folded light curves for each object. Similarly, the phase-folded radial velocity data for each star is shown in Fig. 2.

4 RESULTS AND DISCUSSION

All the planets presented in this work have radii between 1.02 and 1.27 R_J and masses between 0.52 and 0.75 M_J . This places them in the general population of hot Jupiters (Fig. 5), with planetary bulk densities ranging from 0.359 g cm $^{-3}$ to 0.940 g cm $^{-3}$.

Our stellar parameters for HATS-54 (NGTS-22) are generally in good agreement with those from Espinoza et al. (2019). T_{eff} , $\log g$, M_* , $[Fe/H]$, and v_{mic} all agree within the error bars. The radius values are within 2σ while the luminosity values are within 3σ . There are significant differences between the $v \sin i_*$ and v_{mac} values reported in the discovery paper and the values determined in this work. However, as Espinoza et al. (2019) assumed a v_{mac} value that is lower than the value in this paper, this will result in a larger $v \sin i_*$ due to the degeneracy between the two properties. Given the age and spectral type of HATS-54, our slowly rotating value seems more appropriate. There was also disagreement on the extinction values but the value from Espinoza et al. (2019) is high given the limit from both Schlegel et al. (1998) and Schlafly & Finkbeiner (2011).

The planetary parameters obtained for HATS-54b (NGTS-22b) are in good agreement with those reported in the discovery paper. Most of the values agree within error bars but a/R_* and the impact parameter, b , differ by less than 2σ . However, the main difference is that Espinoza et al. (2019) could only constrain $e < 0.126$ whereas we fixed $e = 0$ as non-zero eccentricity is likely spurious given that our eccentric models returned an eccentricity value of $e = 0.023^{+0.033}_{-0.017}$.

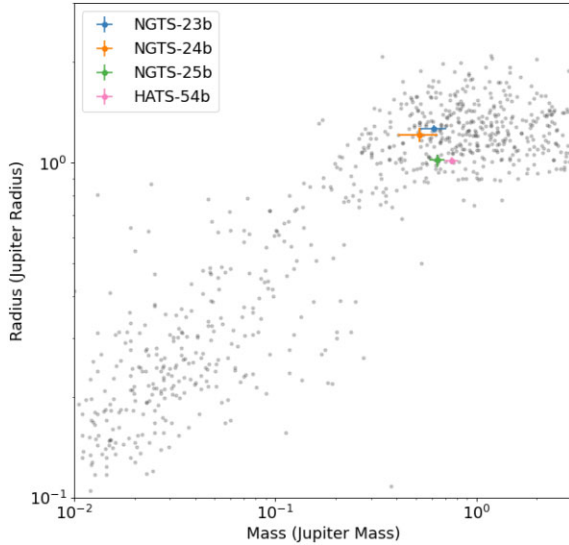


Figure 5. Exoplanets with defined masses and radii (grey points) with errorbars suppressed for clarity. The planets presented in this paper are plotted as coloured points with errorbars. All four planets presented are found to lie within the general population of hot Jupiters. The exoplanet sample was obtained using the NASA Exoplanet Archive (<https://exoplanetarchive.ipac.caltech.edu/>).

4.1 Orbital decay time-scales

Given the late age of two of our systems (NGTS-24 and HATS-54), we thought it useful to calculate the orbital decay time-scales for our planets. We use the following equation to calculate the rate of orbital decay (Goldreich & Soter 1966; Ogilvie 2014; Patra et al. 2020):

$$\frac{dP}{dt} = -\frac{27\pi}{2Q'_*} \left(\frac{M_p}{M_*} \right) \left(\frac{R_*}{a} \right)^5, \quad (1)$$

where $Q'_* = 3Q_*/2k_2$. Q_* and Q'_* are the tidal quality factor and the reduced tidal quality factor, respectively, while k_2 is the dimensionless quadrupolar Love number. We also only consider the equilibrium tide within the convection zone in this analysis as none of our systems satisfy the criterion for the dynamical tide to be considered. This is because the inertial waves (that constitute the dynamical tide) are only excited if the orbital period is twice the stellar rotation period (i.e. $P_{\text{orb}} > P_*/2$). Therefore, for our analysis, we can use Q'_* as defined in Strugarek et al. (2017).⁴

Using the orbital period and the rate of decay, we are able to make a prediction for the length of time for the planetary orbit to decay completely (i.e. where its orbital period is zero). This assumes a constant rate of decay and thus will likely over-estimate the time taken for the orbit to decay. We also have ignored the effects of stellar wind which would likely decrease the inspiral time as well. Finally, as the star evolves off the main-sequence the tidal quality factor can change by orders of magnitude, which dramatically impacts the decay time. This could in-part explain the lack of hot Jupiters observed around sub-giant stars (Villaver & Livio 2009; Hansen 2010; Schlaufman & Winn 2013).

A summary of all the estimated decay time-scales can be found in Table 6. Although we have estimated the rate of decay for these

systems, we are likely unable to measure these due to their long decay time-scales. For example, the observed orbital decay rate of WASP-12b is several orders of magnitude higher than for any of the stars presented here (Maciejewski et al. 2016; Patra et al. 2017; Yee et al. 2020). The long orbital decay time-scales, while an estimation, suggest that all four planets are in a relatively stable configuration while the respective host star remains on the main-sequence.

However, we see that when the star evolves, the orbital decay time-scale rapidly decreases, which agrees with the observations that there is a lack of hot Jupiters around sub-giant stars. Of particular interest is NGTS-24b as its host star will soon evolve off the main-sequence if it is not already doing so. There is some tentative evidence to suggest that NGTS-24b is already beginning to evolve off the main-sequence (see Table 4 and Fig. 4). Therefore, we expect that the star will change significantly within 500 million years causing the inspiral time to reduce significantly. This leads to the planet being engulfed within 600 million years. For comparison, HATS-54 will take 2 billion years before it evolves to a stage where it will have a similar inspiral time.

4.2 Inflation

Many hot Jupiters that are highly irradiated are observed to have low densities (Laughlin, Crismani & Adams 2011; Weiss et al. 2013; Thorngren & Fortney 2018). Previous studies have suggested that this is driven by inflation mechanisms that become effective above an incident flux of approximately $2 \times 10^5 \text{ W m}^{-2}$ (Demory & Seager 2011; Miller & Fortney 2011).

We calculated the irradiation received by each planet and found that all four are irradiated above this threshold (see Table 5). Therefore, we would expect that our planets might exhibit larger than expected radii due to inflation.

We follow the procedure laid out by Tilbrook et al. (2021) and Costes et al. (2020), who utilize the work of Baraffe et al. (2008; hereafter B08) and Sestovic et al. (2018; hereafter S18) to determine whether a planet is inflated. Using the work from B08 we obtain estimated radii for non-inflated planets that we can compare with expected inflated radii from S18 (see Fig. 6).

B08 produced planetary evolution models that account for previous limitations, such as quantifying the effect of metallicity. Notably, radius decreases with increasing heavy element mass so by assuming a low metallicity of 0.02–0.1 for each planet, we are able to provide an upper limit for the non-inflated radii. Differences between the predicted value and the observed radius could indicate that inflation mechanisms have affected the radius.

S18 implement hierarchical Bayesian modelling and a forward model to infer relationships between incident flux and radius. The increase in radius, ΔR , is also dependent on the mass of the planet. All four of the planets presented in this paper fall into the same range ($0.37\text{--}0.98 M_J$), and we therefore use the corresponding equation to derive the expected radius increase due to inflation for each,

$$\Delta R = 0.70 \times (\log_{10} F - 5.5). \quad (2)$$

The baseline radius in S18 is $R_{\text{base}} = (0.98 \pm 0.04) R_J$ for this mass regime. Therefore, the inflated radius is calculated as:

$$R_{\text{inflated}} = 0.98 + \Delta R. \quad (3)$$

These inflated radii can then be compared with the derived radii from our global models.

For NGTS-23b, the observed radius is consistent with the inflated radius as predicted by S18 within error-bars. It is also larger than the estimated non-inflated radius by more than 6σ . It is therefore

⁴This analysis requires the depth of the convection zone. Therefore the radius and mass of the core were determined using <http://www.astro.wisc.edu/~townsend/static.php?ref=e2-web>.

Table 6. Summary of the decay time-scales for the four planets presented in this paper.

Star	Age (Gyrs)	M_* (M_\odot)	R_* (R_\odot)	L_* (L_\odot)	M_c (M_\odot)	R_c (R_\odot)	Q'_*	$\frac{dP}{dt}$	t_a (Gyrs)
NGTS-23	6.39	1.01	1.24	1.920	1.00	0.95	6.49×10^6	-7.43×10^{-14}	150
	8.89	1.01	2.32	3.371	0.63	0.94	3.15×10^5	-3.51×10^{-11}	0.32
NGTS-24	4.77	1.26	1.64	2.792	1.23	1.18	2.43×10^6	-7.02×10^{-13}	14
	5.27	1.26	2.94	4.485	0.52	0.93	3.16×10^5	-1.00×10^{-10}	0.09
NGTS-25	3.49	0.91	0.86	0.511	0.88	0.57	1.06×10^7	-3.12×10^{-14}	248
	13.49	0.91	1.06	0.912	0.87	0.70	4.60×10^6	-2.05×10^{-13}	38
HATS-54	7.96	1.05	1.23	1.302	1.02	0.89	4.17×10^6	-6.05×10^{-13}	12
(NGTS-22)	9.96	1.05	2.11	2.360	0.66	0.88	4.98×10^5	-7.54×10^{-11}	0.09

Notes. In column 2, ‘Age’, we list two values for each star. The first is the stellar age estimation from Table 4. For comparison, the second value indicates the age at which the star is estimated to begin evolving off of the main-sequence. For NGTS-25b, this later age was decided to be 10 Gyrs older than its determined age. M_c and R_c are the mass and radius of the core, respectively. Q'_* is the reduced tidal quality factor, $\frac{dP}{dt}$ is the rate of orbital decay and t_a is the associated time for the planetary orbit to decay. The stellar rotation period was assumed to be 30 d in each scenario. We note that the assumption of an incorrect stellar rotation period does not affect the results much unless the true period is $<< 1$ d.

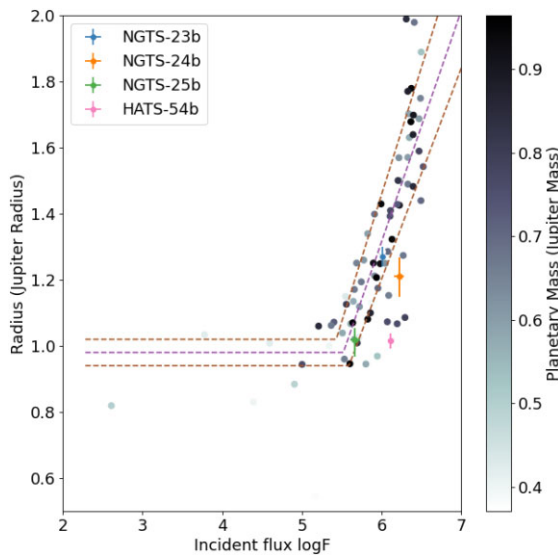


Figure 6. Exoplanets from the NASA Exoplanet Archive (<https://exoplanetarchive.ipac.caltech.edu/>) with masses between $0.37\text{--}0.97 M_J$. The purple dashed line shows the predicted ‘inflated’ radii (Sestovic et al. 2018) with the associated standard deviation shown as dashed brown lines. The hot Jupiters presented in this paper are plotted as coloured points with errorbars.

reasonable to conclude that NGTS-23b is inflated. NGTS-24b has a radius greater than the estimated non-inflated radius, but it is not as large as the expected inflated radius. Although there is less than a 3σ difference between the observed radius and the expected non-inflated radius, the model from S18 fails to predict the behaviour. However, due to the high irradiation and the low planetary density it is likely that this planet is inflated. Therefore, our results could be indicative that density is a more suitable parameter than planetary radius to determine whether a planet is inflated. HATS-54b has a radius that is consistent with the non-inflated model and disagrees with the inflated prediction by more than 6σ . Therefore, it appears that HATS-54b is not inflated. Finally, NGTS-25b is consistent with both inflated and non-inflated predictions so we are not able to draw any firm conclusions for this planet. However, this is not altogether surprising given that NGTS-25b has a lower level of irradiation compared to the other planets presented in this paper. A summary of the inflation results can be found in Table 7.

Similar to the results from Tilbrook et al. (2021), our results are in tension with the claim from S18 that there is not a population of non-inflated hot Jupiters in the mass range $0.37\text{--}0.98 M_J$. Although all of the incident fluxes are larger than the threshold for inflation ($2 \times 10^5 \text{ W m}^{-2}$; Demory & Seager 2011; Miller & Fortney 2011), it appears that HATS-54b is not inflated. It is also unclear whether NGTS-24b and NGTS-25b are inflated. Given that this is simply the minimum limit at which planets begin to appear inflated and that the incident flux for NGTS-25b is not much greater than this limit, it is not too surprising that NGTS-25b is not inflated. However, as new evidence suggests that gaseous planets may experience ‘re-inflation’ as the star evolves off the main-sequence (Hartman et al. 2016; Komacek et al. 2020; Thorngren et al. 2020, 2021), it is somewhat surprising that NGTS-24b cannot be confirmed to be inflated as it experiences relatively high levels of irradiation. Therefore, NGTS-24b may be an interesting candidate to study to better understand the impact of additional parameters on the mechanisms driving inflation.

In particular, metallicity might play an important role as, for the systems reported in this paper, the only planet orbiting a metal-poor star was also the only planet that agreed with the inflated radius predictions. This agrees with the expectation that radii decrease as the fraction of heavy elements in a planet increases (Thorngren et al. 2016). It is already known that hot Jupiter radii vary with age (Baraffe et al. 2008; Miller, Fortney & Jackson 2009; Thorngren & Fortney 2018), and additionally, differing system ages could indicate a difference in planetary migration times, and therefore variations in the level of stellar irradiation compared to what they currently experience. The detailed effects of additional parameters, such as metallicity, eccentricity and system age, are not yet fully understood. However, they could potentially explain why planets may exhibit different levels of radius inflation despite receiving similar levels of irradiation.

Furthermore, the number of known hot Jupiters has increased greatly since the publication of S18 therefore it would be beneficial to have an updated model with an aim of understanding the impact of various factors as this could describe the inflation relationships more accurately. Understanding the dependence on these parameters is crucial to determining the underlying inflation mechanisms.

5 CONCLUSIONS

We report the discovery of three new hot Jupiters and the independent detection of HATS-54b. Each planet was originally identified from photometry from the Next Generation Transit Survey (NGTS), and

Table 7. Quantification of the inflation of the four planets.

Planet	Irradiation (W m^{-2})	M_* (M_J)	$R_{\text{non-inflated}}$ (R_J)	ΔR (R_J)	R_{inflated} (R_J)	R_{observed} (R_J)
NGTS-23b	$(1.03 \pm 0.07) \times 10^6$	0.5	1.017–1.093	$0.36^{+0.01}_{-0.02}$	1.34 ± 0.04	1.27 ± 0.03
NGTS-24b	$(1.66^{+0.24}_{-0.26}) \times 10^6$	0.5	1.009–1.072	0.50 ± 0.03	1.48 ± 0.05	1.21 ± 0.06
NGTS-25b	$(4.62^{+0.63}_{-0.57}) \times 10^5$	0.5–1.0	1.009–1.186	0.12 ± 0.03	1.10 ± 0.05	$1.02^{+0.04}_{-0.05}$
HATS-54b	$(1.29^{+0.08}_{-0.09}) \times 10^6$	0.5–1.0	1.009–1.106	0.42 ± 0.01	1.40 ± 0.04	1.02 ± 0.02

Note the masses in this table are not the true planetary masses but are the masses used to calculate the non-inflated radius. ΔR is the expected change in radius caused by inflation mechanisms.

was confirmed through additional photometry with the Transiting Exoplanet Survey Satellite (TESS). Other ground-based follow-up observations were carried out at the South African Astronomical Observatory (SAAO) with the 1.0-m and 1.9-m telescopes, as well as with EulerCam on the 1.2-m *Euler Telescope* at La Silla Observatory. Radial velocity measurements made with the HARPS, CORALIE and FEROS spectrographs allowed for the masses of all four planets to be determined.

Spectral analysis via iSPEC and SED fitting via ARIADNE revealed the properties of the host stars in each system, which were found to be F-, G-, or K-type stars (see Table 4). We also attempted to constrain the rotation periods of the stars and from this determined that the four stars were likely to be slowly rotating. Global fits to the data were produced using the open-source astronomy software package ALLESFITTER. Our results for HATS-54b (NGTS-22b) are consistent with the original discovery paper (Espinoza et al. 2019). From the orbital decay time-scales of the four planets presented in this paper, they all appear to be on stable orbits while the host star remains on the main-sequence.

All four planets receive a level of irradiation above the minimum threshold for the onset of planetary inflation mechanisms (Demory & Seager 2011; Miller & Fortney 2011). Therefore, we sought to determine whether any of the planets experienced inflation by comparing our derived radii with predictions from evolutionary models (Baraffe et al. 2008). In addition, we examined the predicted additional radius change due to inflation, ΔR , through the flux-mass-radius relations outlined in Sestovic et al. (2018). We found that NGTS-23b is likely inflated while NGTS-24b is more likely to be inflated than not inflated. NGTS-25b is consistent with both inflated and non-inflated models so we are unable to determine whether it is inflated. HATS-54b is unlikely to be inflated as it is consistent with the non-inflated models.

Finally, similar to Tilbrook et al. (2021), we note that there is a disparity between some of our derived radii and those predicted by the inflationary forward model of Sestovic et al. (2018). Through the inclusion of both new hot Jupiter data and additional parameters, such as system age and heavy metal fraction, the accuracy of inflationary parametrization can be improved, which would help improve our understanding of the physical processes behind the inflation of hot Jupiters.

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DATA AVAILABILITY

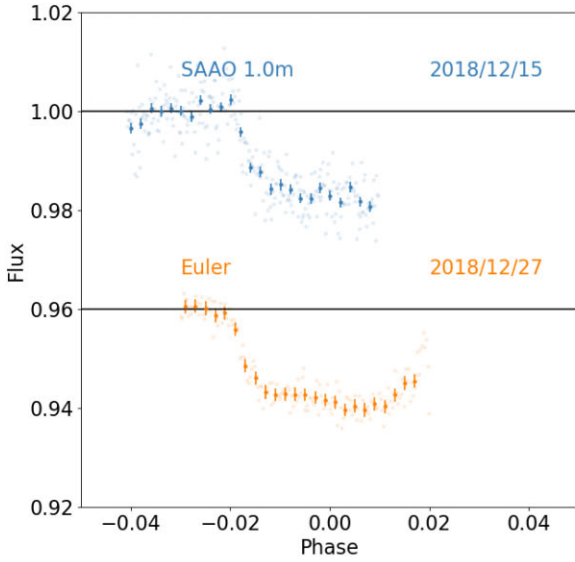
The data underlying this article will be shared on reasonable request to the corresponding author.

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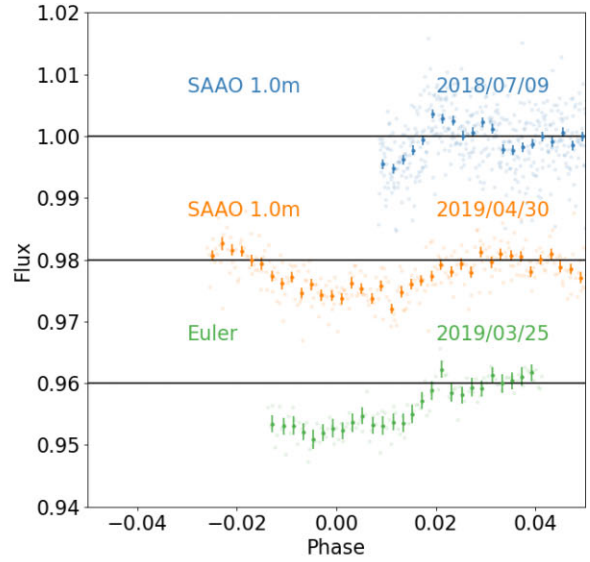
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APPENDIX A: ADDITIONAL PHOTOMETRY



(a) Additional Photometry for NGTS-23b.



(b) Additional Photometry for HATS-54b (NGTS-22b).

Figure A1. Phase-folded photometry for NGTS-23b and HATS-54b (NGTS-22b). These transits were not used in the fitting procedure as there was not sufficient baseline to characterize the out of transit behaviour.

APPENDIX B: SAMPLE NGTS PHOTOMETRY

Table B1. Sample NGTS photometry for each object. The full tables are available in a machine-readable format from the online journal.

(a) NGTS photometry for NGTS-23.		
Time (BJD-2450000)	Normalized flux	Flux error
7981.8851	1.016	0.022
7981.8852	1.002	0.022
7981.8854	0.978	0.022
7981.8855	0.999	0.022
7981.8857	1.022	0.022
...
8195.5415	1.020	0.020
8195.5416	1.007	0.020
8195.5417	1.056	0.020
8195.5419	1.012	0.020
8195.5420	0.956	0.020
(b) NGTS photometry for NGTS-24.		
8097.8107	0.993	0.010
8097.8108	0.985	0.010
8097.8110	0.977	0.010
8097.8111	1.011	0.010
8097.8113	1.013	0.010
...
8295.5989	0.969	0.014
8295.5991	1.001	0.014
8295.5992	0.975	0.014
8295.5993	1.013	0.014
8295.5995	1.007	0.014

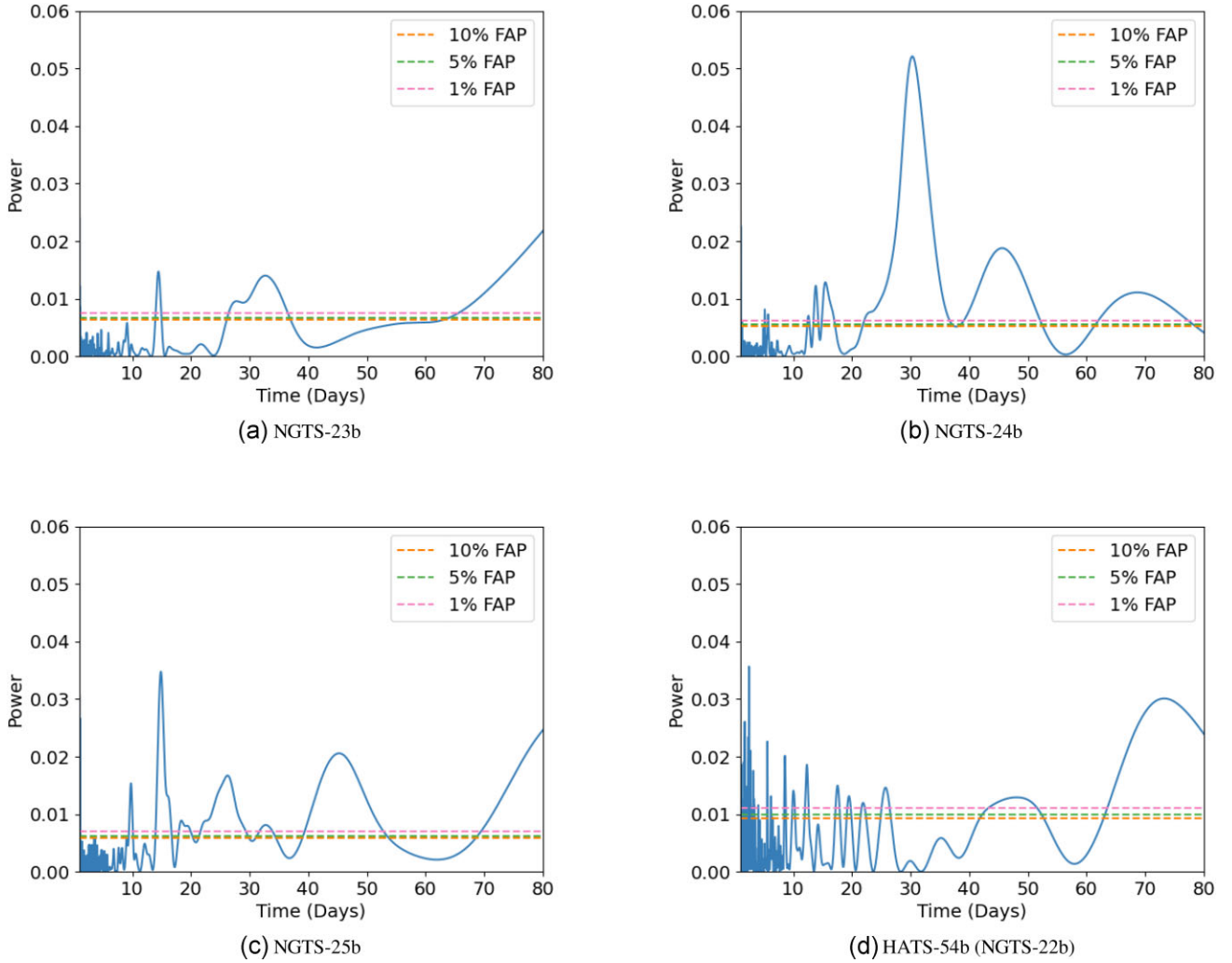
Table B1 – *continued*

(a) NGTS photometry for NGTS-23.		
Time (BJD-2450000)	Normalized flux	Flux error
8203.8697	1.009	0.022
8203.8698	0.984	0.022
8203.8700	1.051	0.022
8203.8701	0.983	0.022
8203.8702	0.959	0.022
...
8423.5825	1.001	0.022
8423.5826	1.020	0.022
8423.5828	0.968	0.022
8423.5829	0.954	0.022
8423.5831	0.984	0.022
(c) NGTS photometry for NGTS-25.		
8203.8697	1.009	0.022
8203.8698	0.984	0.022
8203.8700	1.051	0.022
8203.8701	0.983	0.022
8203.8702	0.959	0.022
...
8423.5825	1.001	0.022
8423.5826	1.020	0.022
8423.5828	0.968	0.022
8423.5829	0.954	0.022
8423.5831	0.984	0.022
(d) NGTS photometry for HATS-54 (NGTS-22).		
7374.8169	0.991	0.017
7374.8171	0.999	0.017
7375.8065	0.979	0.019
7375.8068	0.970	0.019
7375.8069	1.004	0.019
...
7632.4995	0.996	0.018
7632.4997	0.994	0.018
7632.4998	0.982	0.018
7632.5000	0.968	0.018
7632.5001	1.018	0.018

APPENDIX C: RADIAL VELOCITY DATA

Table C1. Radial velocity data for the four systems.

Target	BJD (−2450000)	RV (km s ^{−1})	RV error (km s ^{−1})	FWHM (km s ^{−1})	contrast	BIS (km s ^{−1})	Instrument
NGTS-23	8780.748	22.344	0.058	8.499	29.492	0.085	HARPS
	8781.748	22.325	0.033	8.376	29.162	0.016	HARPS
	8783.744	22.217	0.033	8.406	28.960	0.024	HARPS
	8784.756	22.315	0.040	8.445	29.733	0.162	HARPS
	8843.670	22.306	0.034	8.477	28.427	0.198	HARPS
	8844.657	22.180	0.023	8.376	28.400	0.104	HARPS
	8868.586	22.212	0.018	8.374	27.577	0.017	HARPS
	8418.650	22.415	0.080	8.946	31.387	−0.019	CORALIE
	8424.713	22.254	0.076	8.926	31.042	0.703	CORALIE
	8428.803	22.292	0.075	9.093	31.454	−0.112	CORALIE
	8736.888	22.281	0.028	–	–	0.082	FEROS
	8798.817	22.176	0.032	–	–	−0.101	FEROS
	8800.690	22.245	0.025	–	–	0.025	FEROS
	8847.709	22.288	0.023	–	–	0.106	FEROS
	8848.621	22.216	0.023	–	–	−0.009	FEROS
	8850.645	22.390	0.019	–	–	−0.014	FEROS
	8851.688	22.291	0.022	–	–	0.050	FEROS
	8852.642	22.199	0.019	–	–	0.111	FEROS
NGTS-24	8595.671	−15.272	0.084	9.264	42.373	−0.011	CORALIE
	8640.590	−15.180	0.084	9.176	42.392	−0.088	CORALIE
	8667.515	−15.296	0.084	9.022	43.598	0.109	CORALIE
	8829.810	−15.364	0.074	9.026	37.061	−0.004	CORALIE
	8843.768	−15.394	0.085	9.113	43.453	0.237	CORALIE
	8847.794	−15.335	0.011	–	–	0.014	FEROS
	8848.802	−15.247	0.012	–	–	0.026	FEROS
	8849.777	−15.305	0.014	–	–	−0.011	FEROS
	8850.770	−15.347	0.011	–	–	0.027	FEROS
	8851.764	−15.229	0.012	–	–	0.034	FEROS
NGTS-25	8852.769	−15.270	0.011	–	–	−0.007	FEROS
	9377.789	−22.926	0.009	6.801	41.105	0.057	HARPS
	9411.722	−22.913	0.011	6.878	40.947	0.054	HARPS
	9412.775	−23.060	0.014	6.819	41.233	0.073	HARPS
	9426.658	−23.009	0.021	6.841	40.176	0.007	HARPS
	9428.610	−22.931	0.015	6.897	40.651	0.011	HARPS
	9429.675	−23.068	0.014	6.742	41.160	−0.027	HARPS
	9430.706	−23.112	0.056	6.927	37.081	0.117	HARPS
	9432.638	−23.093	0.014	6.934	40.522	0.007	HARPS
	9433.636	−23.032	0.016	6.829	40.429	0.068	HARPS
	9461.668	−23.055	0.017	6.748	41.118	0.073	HARPS
	9462.671	−22.923	0.017	6.761	40.849	0.044	HARPS
	9470.521	−22.987	0.019	6.845	40.277	0.089	HARPS
	8704.642	−23.285	0.165	8.407	92.508	0.168	CORALIE
HATS-54 (NGTS-22)	8717.567	−22.566	0.227	8.508	193.087	0.105	CORALIE
	8613.692	46.189	0.011	7.858	50.059	−0.064	HARPS
	8615.664	46.066	0.007	7.837	50.219	−0.031	HARPS
	8616.656	46.191	0.008	7.837	49.736	−0.003	HARPS
	8617.649	45.981	0.014	7.847	48.371	−0.010	HARPS
	8642.655	46.064	0.009	7.798	50.454	0.033	HARPS
	8546.699	46.124	0.151	9.206	58.315	−4.184	CORALIE
	8547.824	46.269	0.099	8.644	55.326	−0.176	CORALIE
	8565.688	46.198	0.118	8.889	49.903	0.080	CORALIE
	8566.837	45.971	0.084	8.581	51.157	−0.218	CORALIE
	8593.533	46.200	0.132	8.649	49.514	0.141	CORALIE
	8601.602	46.208	0.094	8.737	53.746	−0.085	CORALIE
	8591.597	46.100	0.011	–	–	−0.022	FEROS
	8592.680	46.135	0.029	–	–	0.004	FEROS
	8593.569	46.227	0.013	–	–	0.005	FEROS
	8594.619	45.938	0.013	–	–	0.006	FEROS
	8597.574	45.965	0.014	–	–	−0.075	FEROS
	8599.601	46.040	0.013	–	–	−0.037	FEROS
	8599.648	45.968	0.014	–	–	0.006	FEROS
	8599.798	46.007	0.015	–	–	−0.066	FEROS

**Figure D1.** GLS periodograms for the NGTS photometry.**APPENDIX D: NGTS PERIODOGRAMS**

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