

ATOMIC AND MOLECULAR INTERSTELLAR ABSORPTION LINES TOWARD THE HIGH GALACTIC LATITUDE STARS HD 141569 AND HD 157841 AT ULTRA-HIGH RESOLUTION

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ABSTRACT

We present ultra-high-resolution (0.32 km s^{-1}) spectra obtained with the 3.9 m Anglo-Australian Telescope (AAT) and Ultra-High-Resolution Facility (UHRF) of interstellar Na I D1, Na I D2, Ca II K, K I, and CH absorption toward two high Galactic latitude stars HD 141569 and HD 157841. We have compared our data with 21 cm observations obtained from the Leiden/Dwingeloo H I survey. We derive the velocity structure and column densities of the clouds represented by the various components and identify the clouds with ISM structures seen in the region at other wavelengths. We further derive abundances, linear depletions, and H_2 fractional abundances for these clouds wherever possible. Both stars are located in regions of IRAS $100 \mu\text{m}$ emission associated with high Galactic latitude molecular clouds (HLCs): HD 141569 lies, in projection, close to MBM 37 and the Lynds dark cloud L134N, whereas HD 157841 is in the vicinity of the MBM 151. Toward HD 141569 we detect two components in our UHRF spectra: a weak, broad $b = 4.5 \text{ km s}^{-1}$ component at -15 km s^{-1} , seen only in Ca II K absorption, and another component at 0 km s^{-1} , seen in Na I D1, Na I D2, Ca II K, K I, and CH absorption. The cloud represented by the -15 km s^{-1} component is warm and may be located in a region close to the star. The cloud represented by the 0 km s^{-1} component has a Ca linear depletion $\delta(\text{Ca}) = 1.4 \times 10^{-4}$ and shows evidence for the presence of dust, consistent with strong $100 \mu\text{m}$ emission seen in this region. The H_2 fractional abundance $f(\text{H}_2)$ derived for this cloud is 0.4, which is typically what is observed toward HLCs. We conclude that this 0 km s^{-1} cloud is associated with MBM 37 and L134N based on the presence of dust and molecular gas (CH) and good velocity agreement with CO emission from these two clouds. This places HD 141569 beyond MBM 37 and L134N, which are estimated to be at $\approx 110 \text{ pc}$. In the case of the HD 157841 sight line, a total of six components are seen on our UHRF spectra in Na I D1, Na I D2, Ca II K, K I, and CH absorption. Two of these six components are seen only in a single species. The cloud represented by the components at 1.85 km s^{-1} has a Ca linear depletion $\delta(\text{Ca}) = 2.8 \times 10^{-4}$, indicating the presence of dust. The $f(\text{H}_2)$ derived for this cloud is 0.45, and there is good velocity agreement with CO emission from MBM 151. To the best of our knowledge, this 1.85 km s^{-1} component toward HD 157841 is the first one found to have relative line widths that are consistent with pure thermal broadening only. We associate the 1.85 km s^{-1} cloud seen in our UHRF spectra with MBM 151 and conclude that HD 157841 must lie beyond $\sim 200 \text{ pc}$, the estimated distance to MBM 151.

Subject headings: atomic data — ISM: abundances — ISM: kinematics and dynamics — stars: individual (HD 141569 and HD 157841)

1. INTRODUCTION

In this paper, we report the results of an ultra-high-resolution (0.32 km s^{-1}) absorption-line study of two stars, HD 141569 and HD 157841. Both stars were chosen as background sources for detailed spectroscopic studies of high Galactic latitude molecular clouds (HLCs), as well as gas in the nearby halo.

HLCs show structure, with dense clumps as small as 0.03 pc , which appear to be virially unstable (Pound, Wilson, & Bania 1990). Some of these clouds appear to be embedded in the local hot interstellar medium as suggested by the

probable detection of X-ray shadowing from two MBM clouds (Burrows & Mendenhall 1991). If this is indeed the case, it is not clear what physical process or combination of processes is causing such enhanced molecular abundances. This has led some authors (Blitz 1990) to conclude that HLCs may represent a completely distinct population of clouds. More recent work by Gir, Blitz, & Magnani (1994) indicates that HLCs are almost always located along filamentary or looplike H I structures and appear to have condensed from atomic gas in situ rather than having been entrained in the H I.

These clouds are an important factor both in the composition and the energy balance of the nearby interstellar medium. They have molecular masses of about $50 M_\odot$, and the molecular gas contribution of these high-latitude clouds

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in the solar neighborhood up to 100 pc is about $5000 M_{\odot}$ (Magnani 1993). The relation and exchange of energy between the molecular gas in these clouds and the ambient highly ionized, moderately ionized, and atomic gas is unclear, although it has been suggested by Blitz (1990) that compression of these clouds could lead to a phase transition of atomic hydrogen into molecular hydrogen.

Absorption-line studies toward stars whose lines of sight intercept regions close to HLCs probe the diffuse component of the ISM associated with these high-latitude clouds. Typical values of the kinetic temperatures for the diffuse component of the ISM range between 45 and 130 K (Savage et al. 1977). To resolve absorption lines arising from clouds with cloud kinetic temperatures of $T_k = 50$ K, a velocity resolution of 0.30 km s^{-1} corresponding to a resolving power $R = 10^6$ is necessary. The Ultra-High-Resolution Facility (UHRF) at the Anglo-Australian Telescope (AAT) is currently the world's highest resolution optical astronomical spectrograph (Diego et al. 1995; Barlow et al. 1995) and is ideal for studies of such cool, interstellar clouds.

We have begun a study of the spatial distribution and properties of atomic and molecular interstellar species toward HLCs at ultra-high resolution (velocity resolution = 0.32 km s^{-1}) and have compiled a list of background stars with reliable photometric and spectroscopic data and lying in the line of sight to HLCs. We have already observed a small fraction of this sample, utilizing the UHRF (Blades et al. 1997), and in the following sections we present UHRF observations toward HD 141569 and HD 157841.

1.1. Overview of the HD 141569 Sight Line

HD 141569 has been variously classified as an isolated Herbig Ae/Be star (Thé, De Winter, & Perez 1994) and a B9.5 V star by Penprase (1992) and an AO Ve star by Dunkin, Barlow, & Ryan (1997). It is an emission-line star associated with an *IRAS* source, IRAS 15473–0346, which is not exactly coincident with it. The heliocentric radial velocity of the star is -6.4 km s^{-1} (Frisch 1987). In the direction of HD 141569, which is located in Galactic coordinates at $(l, b) = (4^{\circ}2, +36^{\circ}9)$, this translates into a LSR radial velocity of -20.1 km s^{-1} . The star exhibits excess far-IR emission (Oudmaijer et al. 1992) believed to arise due to emission from dust grains present in a circumstellar disk, and Sylvester et al. (1996) include it in their list of Vega-like systems.

A high Galactic latitude dark cloud complex at $(l, b) \sim (4^{\circ}, 36^{\circ})$ lies close to the star in projection. This cloud complex includes the Lynds dark clouds L134, L183 (which is also often referred to as L134N, hereafter we will refer to this cloud as L134N), and L1780. The dense molecular CO core MBM 36 is located within L134, whereas the CO core in L134N is referred to as MBM 37 (Magnani, Blitz, & Mundy 1985). Of these two dark clouds, L134N, centered at $(l, b) = (6^{\circ}0, 36^{\circ}8)$ and 2° north of L134, lies closest to HD 141569. HD 141569 lies well within the $100 \mu\text{m}$ emission contours of the L134N dark cloud (Laureijs, Clark, & Prusti 1991).

Penprase (1992) estimated a distance of 190 ± 110 pc for HD 141569 based on spectroscopic and UBV and Ström-gren photometric data. There are various distance estimates to the L134, L134N, and L1780 complex: 100 ± 50 pc, based on optical extinction and surface brightness obser-

vations by Mattila (1979), 110 ± 10 pc based on photometric data by Franco (1989), and 160 pc based on reddenings of a large number of stars in this direction by Snell (1981). The parallax of HD 141569 from the *HIPPARCOS* Catalogue is 10.10 ± 0.83 mas, corresponding to a distance of 99 ± 8 pc. This implies that the distance to the star is very similar to that of the L134, L134N, and L1780 complex. Penprase (1992) estimated an $E(B-V)$ of 0.12 ± 0.09 for HD 141569. The *HIPPARCOS* Catalogue lists Tycho magnitudes that transform to give $(B-V) = 0.078 \pm 0.007$, which for an A0 V star corresponds to an $E(B-V)$ of 0.10 ± 0.03 [the uncertainty of 0.03 is due to the uncertainty in the $(B-V)_0$ calibration for dwarf stars]. The star falls just outside the outermost CO contours that trace MBM 37, the CO core in L134N (Caillault, Magnani, & Fryer 1995). Based on available information, although HD 141569 appears to lie approximately at the distance of L134N and clearly does not lie within the dark cloud itself, it is not certain whether it lies in front of or behind this dark cloud. This point is discussed further in § 3.1.3, and our UHRF data indicate that HD 141569 lies beyond L134N and MBM 37.

1.2. Overview of the HD 157841 Sight Line

HD 157841, a B9 star (SIMBAD database) at $(l, b) = (16^{\circ}8, 15^{\circ}7)$, lies in projection near the HLC, MBM 151, which is located at $(l, b) = (21^{\circ}5, 20^{\circ}9)$. We have searched the SIMBAD database and related literature and have been unable to find any radial velocity measurement for this star.

Based on CO velocity agreement, Penprase (1993) suggested that MBM 57 located at $(l, b) = (5^{\circ}1, 30^{\circ}8)$ is related to MBM 151, which is $\sim 19^{\circ}$ away. Photometric data of stars in the region around MBM 57 (Franco 1989) show an increase in extinction beyond 100 pc. Penprase (1992) also observed this increase in the reddening and estimated a distance of 150–210 pc for the near edge of MBM 57. Assuming that MBM 57 and MBM 151 are related, the distance to MBM 151 is ~ 200 pc. The CO ($J = 1 \rightarrow 0$) LSR velocity of MBM 151 is -0.8 km s^{-1} (Magnani et al. 1985). The region near HD 157841 has not been mapped in CO, and it is not known whether there is any CO emission present at the position of the star or in its vicinity. The *HIPPARCOS* Catalogue lists a parallax for HD 157841 of 5.70 ± 1.16 mas corresponding to a distance of 175^{+45}_{-30} pc. It also lists Tycho magnitudes that transform to yield $(B-V) = 0.213 \pm 0.008$. For a B9 V star, this corresponds to an $E(B-V)$ of 0.28 ± 0.03 . In § 3.2.5, we identify the various components seen in the UHRF spectra toward HD 157841.

2. THE DATA AND THEIR REDUCTION

2.1. The UHRF Data

The spectra of HD 141569 and HD 157841 were obtained during 1994 April 25–28 and 1995 July 14 and 15 by means of the UHRF at the Coudé focus of the 3.9 m Anglo-Australian Telescope. For details about the UHRF refer to Diego et al. (1995). The configuration of the spectrograph was set to give close to the highest achievable resolution, which is $R = 938,000 \pm 15,000$. We used the 35 slice image slicer (Diego 1993) at the entrance aperture. We masked the image slicer down to give 18 slices on the detector, which allowed for the simultaneous detection of the interorder background on either side of the stellar spectrum.

The detector was a Tektronics CCD (1024×1024) with a pixel size of $24 \times 24 \mu\text{m}$. The width of each individual slice produced by the image slicer is about 30 pixels, and the 18 slices filled over half the detector format while the order background filled the remainder. The entire CCD was read out with 8 times binning in the direction perpendicular to the dispersion.

The instrumental setup gave a wavelength coverage of 2.5 Å, requiring separate exposures for each of the Na I D lines. Individual exposures were normally 1800 s, and the total integration times were 30 minutes each for the D1 and D2 spectra, 60 minutes for Ca II K, and 180 minutes each for the K I and CH spectra. Comparison exposures made by means of a Th-Ar lamp were used to establish the wavelength scale and showed that the plate scale is linear with wavelength over the 2.5 Å coverage (Barlow et al. 1995). Measurements of the 6328.160 Å stabilized He-Ne laser line taken during the observing run confirmed that we achieved a resolution of 0.32 km s^{-1} (FWHM).

Standard IRAF reduction packages were used to combine individual integrations after cosmic-ray events, and the interorder background had been removed. A scattered-light correction of 3.5% was applied to all the data, derived from observations of a highly saturated D line in a different star (HD 24263) recorded during the same period. The scattered-light fraction for the Ca II, CH, and K I lines is not the same as for the Na I D1 and D2 lines. In particular for the Ca II K line, in spite of background subtraction (by taking averages of the interorder regions above and below the stellar spectrum), there may still be a residual of as much as 10% (Crawford et al. 1994). We do not have observations of saturated Ca II K lines, as we do the Na I D lines, and therefore we do not have exact values for the scattered-light fraction in the Ca II line. However, we have compared our total equivalent widths with the values published by Penprase (1993), assuming that the Penprase data have no scattered-light contribution. The equivalent widths obtained for Na I and CH for both HD 141569 and HD 157841 are consistent with zero contribution from scattered light. For the case of the Ca II line, we estimate the contribution of scattered light to the measured equivalent widths to be $\sim 15\%$ for HD 141569 and $\sim 8\%$ for HD 157841. For the case of the K I lines, we are not able to make such scattered-light estimates since there are no published K I equivalent widths for either star. We have not subtracted this correction for the Ca II lines and estimate that the uncertainty due to scattered light has a small effect on the derived column densities ($\leq 10\%$) and has almost no effect on the measured line widths, which is our main concern in this paper.

The spectra were normalized by fits of straight lines to sections without telluric or interstellar absorption. Telluric absorptions around the D line region were removed by dividing by a UHRF telluric template spectrum of α Eri, which was rescaled so that the depth of individual telluric features matched those of unblended telluric lines in the spectra. Rest wavelengths in air of 5895.9234 Å for D1, 5889.9512 Å for D2, 3933.663 Å for Ca II K, 7698.974 Å for K I, and 4300.3133 Å for CH were adopted from Morton (1991). The signal-to-noise ratios derived from the rms fluctuation in the stellar continuum for HD 141569 are 20 for D1, 28 for D2, and 10 for the Ca II K, CH, and K I spectra. In the case of HD 157841, the S/N values are 35 for D1, 16 for D2, 10 for K I, 20 for Ca II K, and 42 for CH.

2.2. Profile Fitting

We used the method of profile fitting developed by Welty, Hobbs, & Kulkarni (1994) to derive the properties of the individual interstellar clouds. The essential assumptions are that each cloud is homogeneous and can be described by a column density N , radial velocity v , and a Gaussian line-of-sight velocity dispersion characterized by the line width parameter $b = (2kT_k/m + 2v_t^2)^{1/2}$, where v_t is the rms (Gaussian) turbulent velocity along the line of sight. The intrinsic line profiles were calculated assuming a minimum number of velocity components and were then convolved with the instrumental response. The three parameters (N , v , b) for each cloud were then determined by an iterative, non-linear, least-square fit to the observed profile. The parameters were fine-tuned until a satisfactory fit was achieved. The main criterion of a satisfactory fit is that the rms deviation (σ) from the observed fit is comparable to the rms fluctuation in the stellar continuum. The uncertainties were determined by varying one parameter while keeping others unchanged until the residual between the fit and observed profile goes beyond 3σ of the stellar continuum.

For both HD 141569 and HD 157841, we started with the K I profile. To provide a good fit, one component was required for HD 141569 and two components for HD 157841. We started with the components determined from the K I fitting and added more components to generate good fits for the other lines. For the case of the Ca II K profile, the prominent features coincide with the components seen in K I but we needed to include additional components for both HD 141569 and HD 157841 to generate good fits. In the case of HD 157841, for example, only one of the K I components coincides with the Ca II K components and we needed a minimum of five more components to achieve a good fit. For the Na D1 line the central part is saturated for both stars, but in the case of HD 157841, for example, the four components derived from the Ca II K fitting were sufficient to obtain a satisfactory fit. The theoretical profile was calculated by means of the same parameters for the Na D2 line, and a satisfactory fit was obtained, which confirms the choice of v , N , and b parameters used in the Na D1 fit.

The CH $R_2(1)$ line is split due to Λ -doubling in the $^2\Pi$ ground state of the molecule, with the two components being separated by 1.43 km s^{-1} (Black & van Dishoeck 1988). The two Λ -doublet components have equal strength, which results in a symmetric absorption-line profile in the absence of additional unresolved velocity structure. The CH line is weak toward both stars, and the S/N ratios of the spectra are probably not enough to clearly see this splitting. The slight asymmetry of the profiles suggests the possible presence of unresolved cloud structure. We have, however, used only one component to fit the CH profile for each of the stars.

Penprase (1993) obtained CH spectra for HD 141569 and HD 157841 at a much lower resolution of $\lambda/\delta\lambda \sim 75,000$. Although the values of the velocity centroids for the CH profiles were not published in this paper, a comparison of our UHRF spectra with Penprase's spectra shows consistency in the CH velocities.

Figures 1 and 2 show the normalized profiles plotted against LSR velocity in the various absorption lines for HD 141569 and HD 157841, respectively. The observed profiles are shown as histograms, whereas the best-fit model

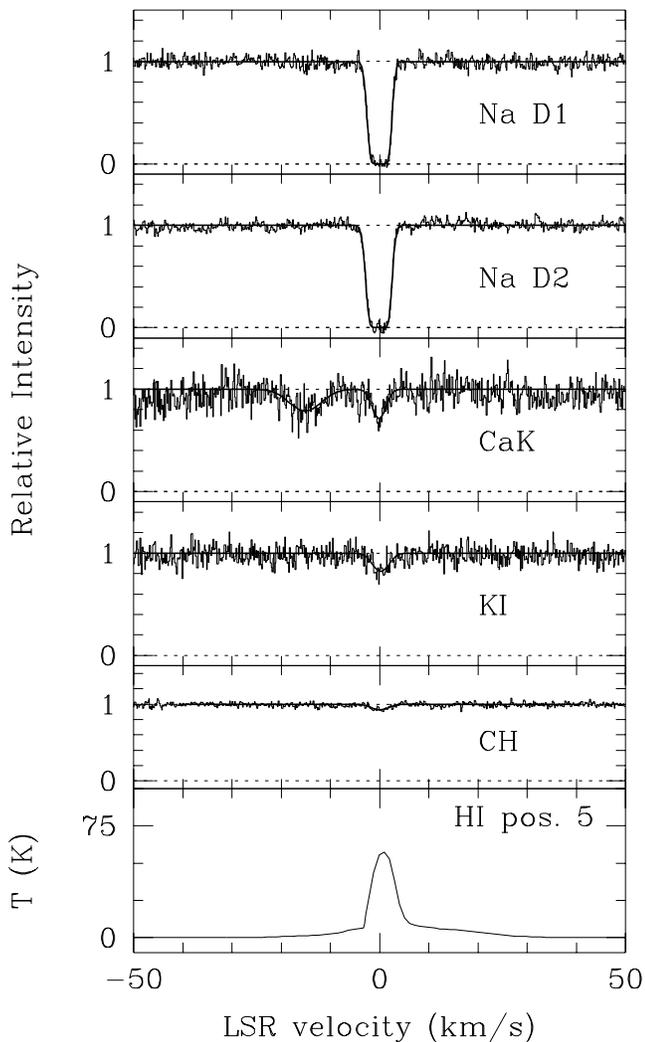


FIG. 1.—Normalized profiles (*histograms*) plotted against LSR velocity for HD 141569. The top five panels show UHRF absorption-line data in the Na D1, Na D2, Ca II K, K I, and CH lines, respectively. The best-fit model absorption line's profiles are plotted as continuous thin lines. The lowest panel shows the H I profile at position 5 (see Fig. 3 and § 2.3), the observation that is closest to the location of HD 141569. Note the presence of the strong absorption feature at 0 km s^{-1} , which coincides with the H I peak. This component shows the presence of both molecular as well as atomic hydrogen. For details see § 3.1.3.

absorption profiles are shown as continuous lines. The parameters resulting from the profile fitting are listed in Tables 1 and 2 for HD 141569 and HD 157841, respectively.

2.3. H I 21 cm Data

H I 21 cm data from the Galactic survey made with the Dwingeloo 25 m telescope (Hartmann 1994; Hartmann & Burton 1997) for the regions around HD 141569 and HD 157841 in the velocity range -100 to $+100 \text{ km s}^{-1}$ are presented in Figures 3 and 4, respectively. The H I observations were made by use of a beamwidth of $36'$. The mean sensitivity of the data is about 0.07 K , and the velocity resolution is 1.03 km s^{-1} . Details of the observations and

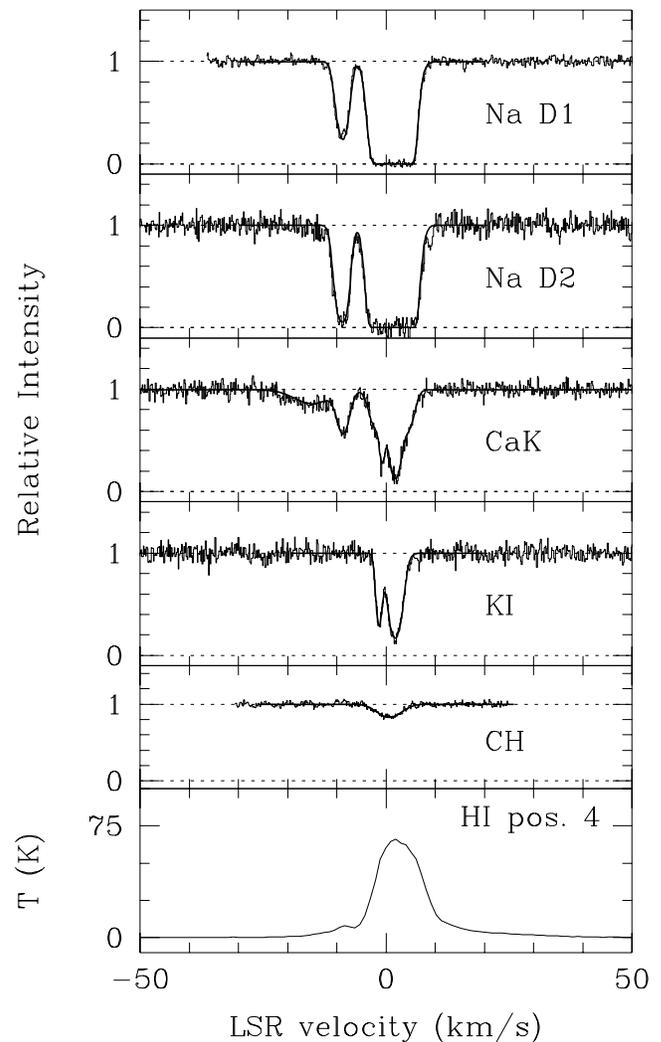


FIG. 2.—Normalized profiles (*histograms*) plotted against LSR velocity for HD 157841. The top five panels show UHRF absorption-line data in the Na D1, Na D2, Ca II K, K I, and CH lines, respectively. The best-fit model absorption line's profiles are plotted as continuous thin lines. The lowest panel shows the H I profile at position 4 (see Fig. 4 and § 2.3), the observation that is closest to the location of HD 157841. Note the presence of the strong absorption feature in the Na I, D1, D2, Ca II K, and K I lines at 1.85 km s^{-1} .

data reduction procedures can be obtained from the two references mentioned above.

The top portions of Figures 3 and 4 show the positions of the HD 141569 and HD 157841, respectively, in Galactic coordinates, with respect to the positions of the H I observations. The lower portions of Figures 3 and 4 show H I line profiles at the nine positions spaced 0.5° apart and centered approximately around the two stars. We have fitted Gaussian profiles to the H I data and estimated the velocities and column densities for the individual line components for the regions around HD 141569 and HD 157841, which are listed in Tables 3 and 4, respectively. Column densities have been estimated by means of equation (8-57) from Mihalas & Binney (1981), assuming that the material is optically thin and has a spin temperature of 100 K . As can be seen from

TABLE 1
MODEL FIT PARAMETERS FOR HD 141569

Parameter	Na I D1, D2	Ca II K	K I	CH
Component 1				
v_{LSR} (km s ⁻¹)	...	-15.1 ^a
N (10 ¹¹ cm ⁻²)	...	2.83 ^b
b (km s ⁻¹)	...	4.5 ^b
Component 2				
v_{LSR} (km s ⁻¹)	0.07 ^a	-0.11 ^a	0.23 ^a	0.09
N (10 ¹¹ cm ⁻²)	45.4 ^a	1.68 ^a	1.16 ^b	58.0 ^a
b (km s ⁻¹)	1.45 ^a	4.45 ^a	2.28 ^b	2.23 ^a

^a Error: <10%.

^b Error: 10%–30%.

Figures 3 and 4, there are large variations in line strength from profile to profile. The UHRF absorption-line profiles are compared to the H I profile in Figures 1 and 2. The top five panels in both figures show the Na D1, Na D2, Ca II K, K I, and CH absorption profiles, respectively, whereas the H I profile closest to the star is shown in the lowest panel.

3. DISCUSSION

3.1. The HD 141569 Sight Line

As mentioned in § 1.1, the main ISM structures in this direction are L134N and MBM 37. L134N is a prototypical cold, dark cloud with no internal energy source and has been extensively mapped at millimetric wavelengths. The kinetic temperature in the dense region of L134N has been derived to be $T_k \sim 12$ K by Snell (1981). Swade (1989) has made the most detailed mapping of this cloud in the molecular emission lines of C¹⁸O, CS, H¹³CO⁺, SO, NH₃, and C₃H₂. The beam sizes used for these observations range from 46'' to 99'', with a grid spacing of $\sim 1'$ and a velocity resolution of 0.16 km s⁻¹. The emission lines in all the molecular spectra are relatively narrow (FWHM ~ 0.5 –

TABLE 2
MODEL FIT PARAMETERS FOR HD 157841

Parameter	Na I D1, D2	Ca II K	K I	CH
Component 1				
v_{LSR} (km s ⁻¹)	...	-14.97 ^d
N (10 ¹¹ cm ⁻²)	...	1.94 ^d
b (km s ⁻¹)	...	4.97 ^d
Component 2				
v_{LSR} (km s ⁻¹)	-8.89 ^c	-8.7 ^c
N (10 ¹¹ cm ⁻²)	7.81 ^c	2.40 ^c
b (km s ⁻¹)	1.30 ^c	1.65 ^c
Component 3				
v_{LSR} (km s ⁻¹)	...	-2.17 ^c	-1.95 ^c	...
N (10 ¹¹ cm ⁻²)	...	2.09 ^c	2.39 ^c	...
b (km s ⁻¹)	...	1.68 ^c	0.71 ^c	...
Component 4				
v_{LSR} (km s ⁻¹)	0.66 ^c
N (10 ¹¹ cm ⁻²)	193.0 ^c
b (km s ⁻¹)	3.19 ^c
Component 5				
v_{LSR} (km s ⁻¹)	1.85 ^{a,b,c}	1.85 ^c	1.85 ^c	...
N (10 ¹¹ cm ⁻²)	605.0 ^c	8.11 ^c	7.53 ^c	...
b (km s ⁻¹)	2.01 ^c	1.47 ^c	1.64 ^c	...
Component 6				
v_{LSR} (km s ⁻¹)	4.65 ^{a,c}	4.65 ^c
N (10 ¹¹ cm ⁻²)	5.78 ^c	1.88 ^c
b (km s ⁻¹)	1.91 ^c	1.52 ^c

^a Component not well resolved in the spectra.

^b Possible weak counterpart to the Ca II and K I at -1.95 km s⁻¹ is hidden in this saturated Na I component.

^c Error: <10%.

^d Error: 10%–30%.

km s⁻¹) and are centered around v_{LSR} ranging between 2.16 and 2.65 km s⁻¹ with a typical uncertainty of ~ 0.35 km s⁻¹. These values are consistent with the H₂CO velocities of $\sim 2.4 \pm 0.7$ km s⁻¹ at 2 mm and $\sim 2.8 \pm 0.9$ km s⁻¹ at 2 cm obtained by Snell (1981) and the [C I] and CO velocities of ~ 2.3 km s⁻¹ obtained by Stark et al. (1996). MBM 37, the

TABLE 3
MODEL FIT PARAMETERS TO THE H I LINE PROFILES FOR HD 141569

Position Number	l, b (deg)	Velocity (km s ⁻¹)	FWHM (km s ⁻¹)	$N(\text{H I})$ (10 ¹⁹ cm ⁻²)
1	3.5, 36.5	-3.14 ± 0.53	15.03 ± 0.74	1.06 ± 0.06
		0.69 ± 0.01	5.48 ± 0.03	9.39 ± 0.07
		13.52 ± 0.31	14.27 ± 0.31	1.14 ± 0.02
2	3.5, 37.0	-0.26 ± 0.96	19.58 ± 1.21	1.09 ± 0.08
		0.59 ± 0.01	5.23 ± 0.03	9.62 ± 0.04
		15.18 ± 0.86	16.21 ± 0.93	0.86 ± 0.09
3	3.5, 37.5	0.46 ± 0.01	4.77 ± 0.03	9.74 ± 0.05
		4.65 ± 0.53	24.05 ± 1.07	1.42 ± 0.03
		20.25 ± 0.61	9.98 ± 2.78	0.22 ± 0.07
4	4.0, 36.5	-3.08 ± 0.47	14.09 ± 0.80	17.30 ± 1.40
		0.81 ± 0.01	5.43 ± 0.04	54.36 ± 0.69
		13.19 ± 0.40	15.45 ± 0.76	17.84 ± 1.05
5	4.0, 37.0	0.78 ± 0.02	4.75 ± 0.04	50.3 ± 0.6
		6.92 ± 0.39	25.63 ± 0.80	35.3 ± 1.13
		3.80 ± 0.38	17.69 ± 0.75	11.84 ± 1.35
6	4.0, 37.5	0.54 ± 0.01	4.99 ± 0.02	47.98 ± 0.24
		9.66 ± 0.48	31.52 ± 0.75	23.20 ± 1.37
		3.80 ± 0.38	17.69 ± 0.75	11.84 ± 1.35
7	4.5, 36.5	-3.14 ± 0.65	14.52 ± 0.92	17.56 ± 1.94
		0.92 ± 0.01	5.35 ± 0.04	53.42 ± 0.75
		12.80 ± 0.60	16.45 ± 0.95	17.70 ± 1.48
8	4.5, 37.0	0.73 ± 0.01	5.20 ± 0.03	51.37 ± 0.50
		3.65 ± 1.25	23.33 ± 1.80	30.62 ± 3.77
		18.13 ± 3.44	19.84 ± 4.58	5.25 ± 3.80
9	4.5, 37.5	0.60 ± 0.01	4.86 ± 0.03	45.07 ± 0.43
		2.88 ± 0.44	22.19 ± 1.11	27.24 ± 0.88
		20.10 ± 0.50	12.17 ± 1.27	4.71 ± 0.95

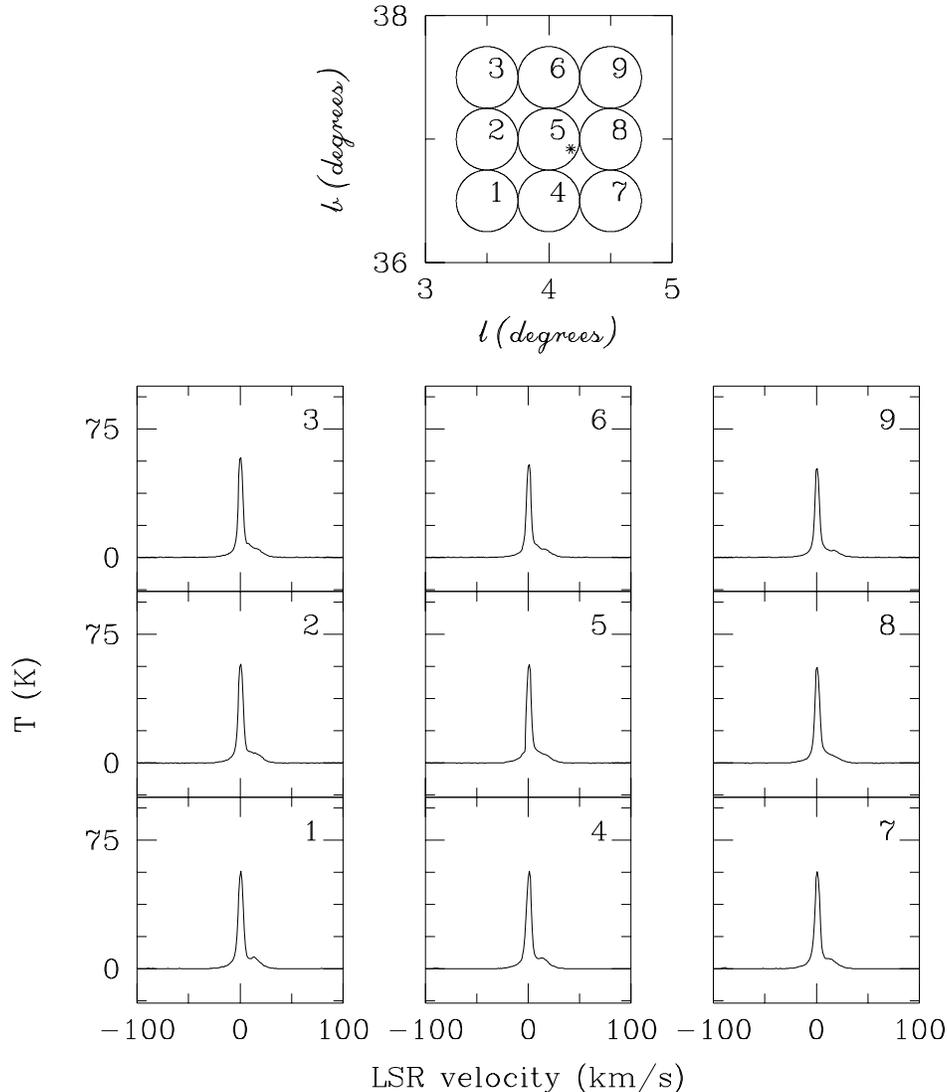


FIG. 3.—*Top*: Position of HD 141569 with respect to the positions of the beams used for the nine H I observations presented in § 2.3, in Galactic coordinates. The asterisk indicates the position of the star that falls just within the beam of position 5. *Bottom*: H I profiles plotted against LSR velocity at the nine positions taken from the Hartmann & Burton (1997) survey. Parameters of Gaussian fits to the profiles are listed in Table 3.

molecular core associated with L134N, has a similar velocity of $\sim 2.3 \text{ km s}^{-1}$ derived from the CO(1 \rightarrow 0) observations by Magnani et al. (1985), implying close association with L134N.

The dark clouds L134 and L134N appear to be connected to several clouds and have been described as resembling “beads on a string” by Magnani, Hartmann, & Speck (1996). These authors report that projected on these clouds is an H I filament with significantly higher LSR velocity. We checked the 21 cm survey of Heiles & Habing (1974), and there is a large-scale H I feature extending between $l \sim 10^\circ$ and $\sim 340^\circ$ and $b \sim 36^\circ$ and 38° and mainly at v_{LSR} between of the order -3 and $+3 \text{ km s}^{-1}$ but extending between of the order -10 and $+10 \text{ km s}^{-1}$. There is no clear evidence for any H I at significantly higher LSR velocity on the Heiles & Habing (1974) plots. Our H I profiles at position 5 do show a broad component at 6.9 km s^{-1} , but we are not certain that this is the H I feature referred to by Magnani et al. (1996). The column density for this H I feature is $N(\text{H I}) = (3.53 \pm 1.13) \times 10^{20} \text{ cm}^{-2}$, and it has no counter-

part in the UHRF spectra. We conclude that this H I component at 6.9 km s^{-1} must therefore lie beyond HD 141569.

3.1.1. The UHRF Spectra

We base our interpretation of the components seen in the UHRF spectra, within the context of the McKee & Ostriker (1977) three-phase model. According to this model, the neutral ISM (the cold neutral medium [CNM]) is contained in cold ($T \sim 80 \text{ K}$), dense ($n \sim 10\text{--}1000 \text{ cm}^{-3}$) clouds, which contain most of the mass of the ISM but have small filling factors. These cold, neutral clouds are surrounded by both a warm neutral medium (WNM) and a warm ionized medium (WIM). All these phases are embedded in a hot ($T \sim 10^5 \text{ K}$), highly ionized medium.

Toward HD 141569, the UHRF spectra show two components: one at $v_{\text{LSR}} = -15.1 \text{ km s}^{-1}$, which is seen only in Ca II K absorption, and another feature at $\sim 0 \text{ km s}^{-1}$, which shows Na I D1, D2, Ca II K, K I, and CH absorption. These components and their associations are discussed in detail in the following sections.

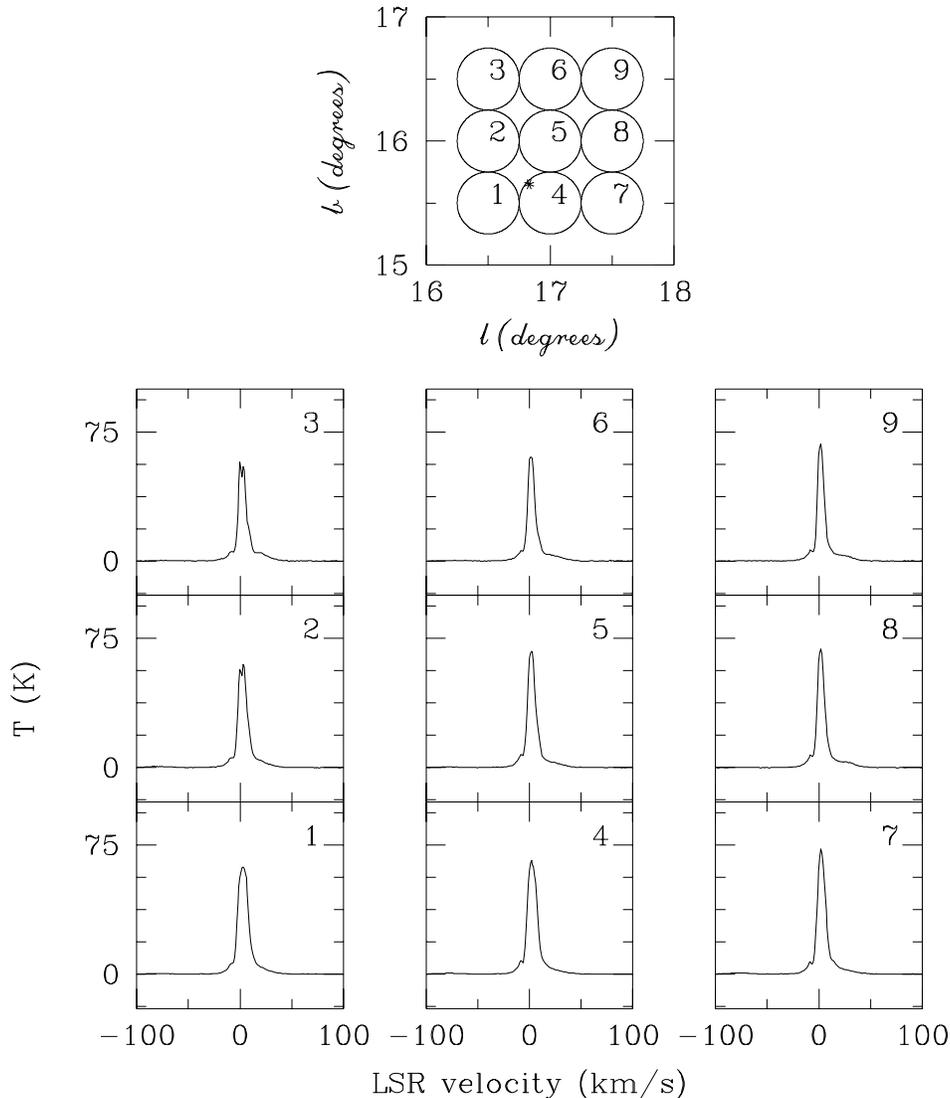


FIG. 4.—*Top*: Position of HD 157841 with respect to the positions of the beams used for the nine H I observations presented in § 2.3, in Galactic coordinates. The asterisk indicates the position of the star that falls just within the beam of position 4. *Bottom*: H I profiles in the LSR velocity range -100 to $+100$ km s^{-1} at the nine positions taken from the Hartmann & Burton (1997) survey. Parameters of Gaussian fits to the profiles are listed in Table 4.

Despite the ultra-high resolution of 0.32 km s^{-1} of our observations, the b values derived for the five components (listed in Table 1) are greater than 1.45 km s^{-1} . Welty et al. (1994) made a high-resolution survey of interstellar Na I D1 lines toward 38 stars. Their spectra had a resolution of 0.5 km s^{-1} and an average S/N ~ 100 . They detected a total of 276 clouds with a median b value of 0.73 km s^{-1} and a median $\log N$ value of 11.09 cm^{-2} . Hence there is the possibility that the b values reported in our paper are overestimates of the true b values of individual clouds, caused by our inability to resolve the component structure due to the relatively low S/N of our spectra. However, our earlier UHRF observations on the HD 28497 sight line (Blades et al. 1997) have S/N comparable to the S/N of the spectra presented here. In the case of HD 28497, four of the 15 components detected in Na I had b values ranging between 0.31 and 0.40 km s^{-1} , whereas we do not detect any component toward HD 141569 (or HD 157841) with b values less than 1.30 km s^{-1} . The minimum detectable column density for Na I in the spectra presented in this paper is

$\sim 2 \times 10^9$ cm^{-2} , which is about a factor of 100 less than the column densities of the typical interstellar clouds studied by Welty et al. (1994). One possibility is that the large b value components seen toward HD 141569 (and HD 157841) are a superposition of several low column density ($< 2 \times 10^9$ cm^{-2}) clouds with overlapping velocities. These are not the typical interstellar clouds described by Welty et al. (1994), and it is unlikely that all the components seen in our spectra are due to such overlapping low column density clouds. For this reason, we conclude that the b values we estimate in this paper probably reflect the true b values of the individual clouds.

3.1.2. The -15 km s^{-1} Component

Seen in absorption only in Ca II K, this weak component is relatively broad with $b = 4.5$ km s^{-1} . No H I counterpart to this component is seen at position 5. The H I beam ($36'$) is much larger than the absorption beam (which is infinitesimally small), and it could be that the H I emission associated with this weak component is lost in the general

TABLE 4
MODEL FIT PARAMETERS TO THE H I LINE PROFILES FOR HD 157841

Position Number	l, b (deg)	Velocity (km s ⁻¹)	FWHM (km s ⁻¹)	$N(\text{H I})$ (10 ¹⁹ cm ⁻²)
1	16.5, 15.5	-0.42 ± 0.12	5.81 ± 0.14	43.70 ± 2.81
		5.04 ± 0.13	6.99 ± 0.18	63.53 ± 3.05
		5.89 ± 0.19	27.84 ± 0.56	51.45 ± 0.98
2	16.5, 16.0	-0.99 ± 0.05	3.30 ± 0.15	10.87 ± 0.98
		3.00 ± 0.06	8.54 ± 0.14	81.44 ± 2.04
		5.85 ± 0.36	29.79 ± 1.03	44.77 ± 1.53
3	16.5, 16.5	-9.75 ± 0.39	6.65 ± 1.06	6.32 ± 0.78
		0.81 ± 0.24	6.79 ± 0.36	56.60 ± 11.78
		5.03 ± 0.99	9.14 ± 0.74	40.40 ± 10.93
4	17.0, 15.5	18.78 ± 1.81	20.02 ± 3.44	17.33 ± 3.29
		-8.74 ± 0.4	4.86 ± 0.63	4.82 ± 0.54
		2.49 ± 0.10	9.89 ± 0.19	125.0 ± 1.99
5	17.0, 16.0	14.51 ± 1.28	22.92 ± 2.20	21.7 ± 2.54
		-8.89 ± 0.18	4.96 ± 0.47	4.37 ± 0.50
		0.91 ± 0.10	6.59 ± 0.15	62.11 ± 7.18
6	17.0, 16.5	4.50 ± 0.45	8.57 ± 0.32	44.74 ± 6.94
		9.01 ± 0.81	36.78 ± 1.35	32.44 ± 1.71
		-8.91 ± 0.26	4.79 ± 0.63	33.18 ± 0.53
7	175, 15.5	1.28 ± 0.08	6.59 ± 0.12	69.41 ± 4.73
		5.51 ± 0.75	8.84 ± 0.58	20.62 ± 4.49
		10.42 ± 1.00	38.42 ± 1.67	30.98 ± 1.87
8	17.5, 16.0	-9.27 ± 0.12	3.85 ± 0.30	2.90 ± 0.27
		0.22 ± 0.10	6.02 ± 0.15	35.14 ± 5.18
		3.66 ± 0.20	7.88 ± 0.16	72.82 ± 5.24
9	17.5, 16.5	7.99 ± 0.35	31.54 ± 0.63	40.64 ± 1.04
		-8.92 ± 0.13	4.61 ± 0.33	4.09 ± 0.34
		1.13 ± 0.06	6.56 ± 0.08	70.44 ± 4.04
9	17.5, 16.5	4.88 ± 0.39	8.55 ± 0.30	30.47 ± 3.91
		11.38 ± 0.64	39.84 ± 0.99	34.26 ± 1.20
		-0.78 ± 0.04	2.76 ± 0.12	10.05 ± 0.88
9	17.5, 16.5	2.17 ± 0.04	5.98 ± 0.07	64.46 ± 1.30
		3.61 ± 0.18	22.61 ± 0.68	44.10 ± 0.81
		28.03 ± 0.52	13.62 ± 0.99	6.48 ± 0.55

background emission. The Ca II column density derived for this component is $8.11 \times 10^{11} \text{ cm}^{-2}$. Unfortunately, interstellar Ca II column densities are very poorly correlated with quantities like $N(\text{H I})$, $N(\text{Na I})$, $N(\text{K I})$, and $E(B-V)$ (Hobbs 1974), and it is not a useful exercise to derive $N(\text{H I})$ by means of the observed $N(\text{Ca II})$ for this component. We can put an upper limit to the $N(\text{H I})$ of this component, corresponding to the sensitivity of our H I data, which is $\sim 10^{18} \text{ cm}^{-2}$. Assuming that Ca II is the dominant ionization stage of this element in the gas, this implies the linear depletion $\delta(\text{Ca}) = [N(\text{Ca})/N(\text{H})]/[N(\text{Ca})/N(\text{H})]_{\odot} > 0.12$, which implies relatively low depletion for this component.

In the absence of turbulence, $b = 4.5 \text{ km s}^{-1}$ for this component, which corresponds to a kinetic temperature $T_k \leq 47,900 \text{ K}$, and such a warm component may be located in a region close to the star. HD 141569 has a v_{LSR} radial velocity of -20.1 km s^{-1} , and the association of the -15.1 km s^{-1} component with a region close to the star is not clear.

3.1.3. The 0 km s⁻¹ Component

The Na I D1, D2 velocities for this component match the Ca II K, K I, and CH component velocities to better than 0.32 km s^{-1} . In the absence of turbulent motions, the kinetic temperature of an absorbing component for Na I is given by $T_k = 1393b^2$, for Ca II by $T_k = 2422b^2$, for K I by $T_k = 2368b^2$, and for CH by $T_k = 787b^2$. If all these absorption components arise from coextensive gas where turbulent motions are absent, the b values for the various species

should differ systematically: $b(\text{Ca})$ should be $0.75b(\text{Na})$, $b(\text{K})$ should be $0.77b(\text{Na})$, and $b(\text{CH})$ should be $1.33b(\text{Na})$. From Table 1, it is clear that this is not the case since the b values for Ca II and K I exceed the $b(\text{Na})$ value, and the $b(\text{CH})$ value is more than $1.33b(\text{Na})$. On the other hand, if turbulent motions were dominant in the region where these absorption components arise, the b values for all the species should be exactly the same, and this is also not what is observed. Therefore the various components must arise from different regions of a cloud, which represent different phases of the ISM. This is similar to the interpretations by Barlow et al. (1995) for clouds toward ζ Oph and Blades et al. (1997) for clouds toward HD 28497.

In the absence of turbulent motions, $T_k(\text{Na}) = 2940 \text{ K}$ and $T_k(\text{Ca}) = 8570 \text{ K}$ and we interpret the Na I component to arise from gas associated with the cold neutral medium (CNM), enveloped by the warm gas delineated by the Ca component. The K I component has a $T_k = 12,310 \text{ K}$. It is possible that the K I absorption feature has velocity structure but the S/N ratio of 10 for the spectra has led us to estimate a large $b(\text{K})$ for this weak absorption feature. This issue can be resolved by higher S/N UHRF spectra in K I. If these future higher S/N UHRF spectra still show no velocity structure for this weak K I feature, it would imply that a warm K I component envelopes the cooler Ca II K and Na I components.

The kinetic temperature obtained for the CH component is $T_k = 3910 \text{ K}$, assuming no turbulent motions. The CH feature is weak, and the S/N ratio of 42 for the CH spectra is

probably not sufficient to resolve the two components of the CH molecule, and the b value that we have obtained may include a blend of both components, resulting in a large T_k value. CH spectra at higher S/N obtained with the UHRF would be required to clarify this issue of additional velocity components. If the higher S/N spectra still show no velocity structure, then the CH component toward HD 141569 must be associated with the warm, outer halo surrounding the cold molecular core of MBM 37. A similar interpretation was made for the broad $b = 1.8 \text{ km s}^{-1}$ CH component toward ζ Oph by Crawford et al. (1994). The CO gas associated with both L134N and MBM 37 (see § 3.1) is at $v_{\text{LSR}} \sim 2.4 \text{ km s}^{-1}$ and association of the 0 km s^{-1} component with MBM 37 would imply that HD 141569 lies beyond $\sim 110 \text{ pc}$.

We are interested in determining abundances and depletions for the entire cloud represented by the components at $\sim 0 \text{ km s}^{-1}$ seen in the various species. Although the various absorption components arise in different regions of the cloud, it is valid to use the column density values derived from the UHRF spectra for the various species for the purposes of studying abundances of the $\sim 0 \text{ km s}^{-1}$ cloud. Based on velocity agreement, we identify the H I component at position 5 with $v_{\text{LSR}} = 0.78 \text{ km s}^{-1}$ to be associated with this 0 km s^{-1} UHRF component. This component has $N(\text{H I}) = 5.03 \times 10^{20} \pm 5.57 \times 10^{18} \text{ cm}^{-2}$ (Table 4). The linear depletions derived are $\delta(\text{Na}) = [N(\text{Na})/N(\text{H})]/[N(\text{Na})/N(\text{H})]_{\odot} = 0.004$, $\delta(\text{Ca}) = [N(\text{Ca})/N(\text{H})]/[N(\text{Ca})/N(\text{H})]_{\odot} = 1.4 \times 10^{-4}$, and $\delta(\text{K}) = [N(\text{K})/N(\text{H})]/[N(\text{K})/N(\text{H})]_{\odot} = 0.0017$, which imply extremely high depletions for this cloud. High depletions imply low gas phase abundances and therefore high dust phase abundances since the missing elements are believed to be incorporated into dust, particularly in the case of calcium (Barlow & Silk 1977; Barlow 1978). Evidence for the presence of dust in this component from the UHRF data is consistent with the strong $100 \mu\text{m}$ emission from L134N (Laureijs et al. 1991).

The H_2 column densities in diffuse interstellar clouds are well correlated with the CH column densities through the relation $N(\text{H}_2) = 2.6 \times 10^7 N(\text{CH})$ with an rms scatter of 9.0×10^{19} (Somerville & Smith 1989). Using this relation and $N(\text{CH})$ derived from the UHRF data, we obtain $N(\text{H}_2) = 1.51 \times 10^{20} \pm 9.0 \times 10^{19} \text{ cm}^{-2}$. The H_2 fractional abundance, given by $f(\text{H}_2) = 2N(\text{H}_2)/[N(\text{H I}) + 2N(\text{H}_2)]$, for this 0 km s^{-1} cloud is therefore $\sim 0.4 \pm 0.2$. This is the $f(\text{H}_2)$ value that is typically observed for the diffuse component associated with HLCs (Penprase 1993). Both the evidence for the presence of dust and molecular gas and the H_2 fractional abundance value for this 0 km s^{-1} cloud strengthen the case for its association with MBM 37 and L134N.

3.2. The HD 157841 Sight Line

A total of six components are seen toward HD 157841 in Na I, Ca II K, K I, and CH (Fig. 2 and Table 2) in the UHRF absorption spectra. The absorption features in the various species broadly resemble each other, although they differ in details. Two of these six components are seen only in a single species. A comparison of the UHRF spectra with the H I profile at position 4, located closest to HD 157841 (lowest panel of Fig. 2), shows that the -8.8 km s^{-1} and 1.85 km s^{-1} UHRF components have H I counterparts. There is an additional H I feature at 14.5 km s^{-1} , which has

no corresponding absorption feature on the UHRF spectra, and this H I emission must be associated with gas beyond HD 157841. In the following sections, we discuss the individual components and their nature.

3.2.1. The -15 km s^{-1} Component

This weak, broad $b = 4.97 \text{ km s}^{-1}$ component is seen only in Ca II K absorption. If we assume no turbulent motions, the kinetic temperature derived for this component is $T_k = 59,900 \text{ K}$, and this component is likely to be associated with a region close to the star. As mentioned in § 1.2, we have been unsuccessful in finding any stellar radial velocity determinations for HD 157841, and therefore we cannot make any further comments on the association of this cloud with a region close to the star. This Ca II component has no associated H I emission that could be due to the larger H I beam size causing weak components with small angular extent to disappear in the general emission.

3.2.2. The -8.8 km s^{-1} Component

This component is seen in Na I and Ca II K absorption with $b(\text{Na}) \neq 0.75b(\text{Ca})$ or $b(\text{Na}) = b(\text{Ca})$ indicating that they do not arise from coextensive gas. If we assume no turbulent motions, $T_k(\text{Na}) = 2370 \text{ K}$ and $T_k(\text{Ca}) = 6595 \text{ K}$. We interpret the observed warm Ca II component to envelope the colder Na I cloud core, similar to what is observed in the case of ζ Oph (Barlow et al. 1995), HD 28497 (Blades et al. 1997), and the 0 km s^{-1} component toward HD 141569 (§ 3.1.3). The H I emission associated with this component has $N(\text{H I}) = 4.89 \times 10^{19} \pm 5.53 \times 10^{18} \text{ cm}^{-2}$. This results in $\delta(\text{Na}) = [N(\text{Na})/N(\text{H})]/[N(\text{Na})/N(\text{H})]_{\odot} = 0.0078$ and $\delta(\text{Ca}) = [N(\text{Ca})/N(\text{H})]/[N(\text{Ca})/N(\text{H})]_{\odot} = 0.002$, which indicate relatively high depletions and imply the presence of dust.

3.2.3. The -1.95 km s^{-1} Component

This component is seen in Ca II K and K I absorption. The Ca II feature is relatively narrow, with $b = 1.68 \text{ km s}^{-1}$, and assuming no turbulent motions the kinetic temperature $T_k(\text{Ca}) = 6820 \text{ K}$. The K I feature is narrow, with $b = 0.71 \text{ km s}^{-1}$, which corresponds to a $T_k(\text{K}) = 1210 \text{ K}$ in the absence of turbulence. The warm Ca II K component is interpreted as the warm envelope surrounding a cooler K I component.

3.2.4. The $0.66/1.85 \text{ km s}^{-1}$ Component

The 1.85 km s^{-1} component is a strong feature seen in the UHRF spectra in Na I, Ca II, and K I absorption. As noted in § 2.2, the CH absorption line toward HD 157841 has a slight asymmetry and suggests possible unresolved cloud structure. We believe that because of this asymmetry, this 0.66 km s^{-1} component seen only in CH is associated with the 1.85 km s^{-1} component seen in Na I, Ca II, and K I.

For the Na I, Ca II, and K I absorptions, $b(\text{Na}) \approx 0.75b(\text{Ca})$ (see Table 2) $\approx 0.77b(\text{K})$. The systematic difference in the line widths in these species implies that they arise from coextensive material where turbulent motions are not important. The kinetic temperatures corresponding to the derived b values of the 1.85 km s^{-1} component are 5630 K for Na I, 5230 K for Ca II, and 6370 K for K I, in good agreement with each other. This component therefore appears to correspond quite well to the warm neutral medium (WNM) phase posited by the McKee & Ostriker

(1977) model of the ISM. The CH absorption has $b(\text{CH}) = 3.19 \text{ km s}^{-1}$, which is more than $1.33b(\text{Na})$ (see § 3.1.3). However, as in the case of the CH spectra of HD 141569, the low S/N of the spectra may have led us to overestimate the b value for this component.

The velocity of this component closely matches the CO LSR velocity of MBM 151, which is $\sim 0.8 \text{ km s}^{-1}$. This points to an association between this component and MBM 151 and places HD 157841 beyond MBM 151, which is at $\sim 200 \text{ pc}$ (see § 1.2). In the absence of turbulence, the kinetic temperature T_k (CH) derived for this component is $\sim 8000 \text{ K}$, and this molecular gas may be a warm envelope surrounding the cold core of MBM 151.

There is a strong H I 21 cm emission at 2.49 km s^{-1} , which is most likely associated with the UHRF feature. The small velocity difference between this H I feature and the Na I, Ca II, and K I absorptions is not significant if one takes into consideration the differences in beam sizes for the H I and UHRF observations. The column density of this feature is $N(\text{H I}) = 1.27 \times 10^{21} \pm 2.15 \times 10^{19} \text{ cm}^{-2}$. Using the Somerville & Smith (1989) relation (§ 3.1.3) and the UHRF derived CH column densities, we estimate $N(\text{H}_2) = 5.02 \times 10^{20} \pm 9 \times 10^{19} \text{ cm}^{-2}$. The H_2 fractional abundance, $f(\text{H}_2)$, for this cloud is 0.45 ± 0.08 , which is typically what is found in HLCs (Penprase 1993) and which is similar to the value determined earlier in § 3.1.3 for the 0 km s^{-1} component seen toward HD 141569.

The linear depletions derived for this cloud are $\delta(\text{Na}) = [N(\text{Na})/N(\text{H})]/[N(\text{Na})/N(\text{H})]_{\odot} = 0.023$, $\delta(\text{Ca}) = [N(\text{Ca})/N(\text{H})]/[N(\text{Ca})/N(\text{H})]_{\odot} = 2.8 \times 10^{-4}$, and $\delta(\text{K}) = [N(\text{K})/N(\text{H})]/[N(\text{K})/N(\text{H})]_{\odot} = 0.0044$ for this component. This cloud is therefore highly depleted and has high dust phase abundances. Both the evidence for the presence of dust and the H_2 fractional abundance values make the association of the 0.66 km s^{-1} CH component and the 1.85 km s^{-1} Na I, Ca II, and K I components with the same cloud and with MBM 151 more secure.

3.2.5. The 4.65 km s^{-1} Component

This component is seen in both Na I and Ca II absorption. $b(\text{Na}) \approx 0.75b(\text{Ca})$ for this component, and the two features must arise from coextensive material.

4. CONCLUSIONS

We have made an ultra-high-resolution study of the Na I, Ca II K, K I, and CH interstellar absorption lines toward two stars, HD 141569 and HD 157841. These absorption spectra have been compared to 21 cm data obtained from the Leiden/Dwingeloo H I survey. Both stars probe the gas in regions close to high Galactic latitude molecular clouds: the HD 141569 sight line intercepts a region close to MBM

37 and L134N, whereas the HD 157841 sight line probes a region close to MBM 151.

The results of our investigation are as follows.

1. Toward HD 141569, two components are seen in our UHRF spectra: one at -15 km s^{-1} and another at 0 km s^{-1} . The -15 km s^{-1} component, seen only in Ca II K, is weak, broad, and shows low linear depletion, $\delta(\text{Ca}) > 0.12$. In the absence of turbulence, the kinetic temperature derived is $T_k \leq 47,900 \text{ K}$ and the cloud represented by this component may be located in a region close to HD 141569, whose stellar radial LSR velocity is -20.1 km s^{-1} .

The $\sim 0 \text{ km s}^{-1}$ component is seen in Na I, Ca II K, K I, and CH absorption. The cloud represented by these components has a linear Ca depletion of $\delta(\text{Ca}) = 1.4 \times 10^{-4}$, which implies the presence of dust, consistent with strong $100 \mu\text{m}$ emission from this region and H_2 fractional abundance $f(\text{H}_2) = 0.4$. We conclude that this 0 km s^{-1} cloud is associated with MBM 37 and L134N because of the presence of dust and molecular (CH) gas at velocities close to CO velocities derived for the two HLCs. The *HIPPARCOS* distance of $99 \pm 8 \text{ pc}$ for HD 141569 is comparable to the distance of $100\text{--}110 \text{ pc}$ that has been estimated for the MBM 37 and L134N clouds. However, our absorption-line data imply that these clouds must lie a little closer than HD 141569.

2. A total of six components are seen in our UHRF spectra in Na I, Ca II K, K I, and CH absorption toward HD 157841. The cloud represented by the 1.85 km s^{-1} absorption components has a Ca linear depletion $\delta(\text{Ca}) = 2.8 \times 10^{-4}$, and the derived H_2 fractional abundance is $f(\text{H}_2) = 0.45$. The relative b values determined for the 1.85 km s^{-1} Na I, Ca II, and K I absorption lines are found to be consistent with thermal broadening (with no significant turbulent component), corresponding to a kinetic temperature of $T_k = 5700 \pm 500 \text{ K}$. To our knowledge, this is the first velocity component found whose line widths can unambiguously be attributed to pure thermal broadening. We conclude that the 1.85 km s^{-1} cloud is associated with MBM 151 because of the presence of dust and molecular (CH) gas at velocities close to CO velocities derived for MBM 151. HD 141569 must therefore lie beyond MBM 151, estimated to be at $\sim 200 \text{ pc}$, and its *HIPPARCOS* distance of $175_{-30}^{+45} \text{ pc}$ is consistent with this conclusion.

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