

# Signatures of resonance and phase mixing in the Galactic disc

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## ABSTRACT

*Gaia* DR2 has provided an unprecedented wealth of information about the kinematics of stars in the Solar neighbourhood, and has highlighted the degree of features in the Galactic disc. We confront the data with a range of bar and spiral models in both action-angle space, and the  $R_G-v_\phi$  plane. We find that the phase mixing induced by transient spiral structure creates ridges and arches in the local kinematics which are consistent with the *Gaia* data. We are able to produce a qualitatively good match to the data when combined with a bar with a variety of pattern speeds, and show that it is non-trivial to decouple the effects of the bar and the spiral structure.

**Key words:** Galaxy: bulge – Galaxy: disc – Galaxy: fundamental parameters – Galaxy: kinematics and dynamics – solar neighbourhood – Galaxy: structure.

## 1 INTRODUCTION

The European Space Agency (ESA)’s *Gaia* mission (Gaia Collaboration et al. 2016) is transforming our view of the Milky Way. The second data release (DR2; Gaia Collaboration et al. 2018a) contains 5 parameter astrometry for around  $1.3 \times 10^9$  stars, allowing us to trace the kinematics of the disc to unprecedented distances. In addition, DR2 also provides radial velocities for around  $7.2 \times 10^6$  stars, enabling a detailed analysis of the Solar neighbourhood (e.g. Gaia Collaboration et al. 2018b). This exciting new data revealed numerous disequilibria features in the Galactic disc, such as the ‘ridges’ in the  $R_G-v_\phi$  plane (e.g. Antoja et al. 2018; Kawata et al. 2018) and we now know that the classical moving groups in the  $v_R - v_\phi$  plane (e.g. Hyades, Pleiades, Coma Berenices, Sirius & Hercules) are a local manifestation of this radially extended structure. With DR2, Antoja et al. (2018) also discovered the ‘phase spiral’ in the  $z-v_z$  plane, highlighting the ongoing phase mixing in the vertical direction. This vertical asymmetry which was first discovered by Widrow et al. (2012) has now been well measured (Bennett & Bovy 2019), but the discussion of its origin is ongoing.

One potential explanation for the radial structure is resonant interaction with the Galactic bar, or a fixed pattern speed spiral, although a large number of resonances are required to create every feature. Alternately, Quillen et al. (2018) proposed that the ridges in the Solar neighbourhood could result from individual spiral arm

crossings, linking the divisions in the moving groups to specific spiral arms. Another potential explanation for the radial structure is phase wrapping (e.g. Minchev et al. 2009) either resulting from interaction with a satellite or dwarf galaxy (e.g. Laporte et al. 2018) or internal dynamical processes such as transient spiral structure (e.g. Hunt et al. 2018b). The vertical disequilibria are naturally explained by the external perturbation from a satellite or dwarf galaxy such as Sagittarius (e.g. Antoja et al. 2018; Bland-Hawthorn et al. 2019; Laporte et al. 2019). It has also been proposed that the vertical disequilibria features can arise from secular effects such as the buckling of the Galactic bar (Khoperskov et al. 2019) although the similarity of the spiral for stars with different ages (Laporte et al. 2019) and the lack of a spiral for stars on hotter orbits (Li & Shen 2019) argues against this scenario for its origin. There is likely a combination of both resonance and phase wrapping occurring in the Galactic disc, and it is not trivial to disentangle which feature arises by which specific mechanism, or which combination of mechanisms.

In Hunt et al. (2018b), we showed that transient spiral arms combined with a long slow bar reproduce the ridges in the  $R_G-v_\phi$  plane, as well as producing a qualitatively good match to the Solar neighbourhood kinematics in the  $v_R - v_\phi$  plane. Recently, Fragkoudi et al. (2019) reproduced the ridge structure with a short fast bar in an  $N$ -body simulation, and highlighted that these ridges also stand out well when colouring the  $R_G-v_\phi$  plane by Galactocentric radial velocity. Khanna et al. (2019) showed that both transient spiral structure and an external perturber are able to create these ridges via phase mixing. Martinez-Medina et al. (2018) present a model which

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reproduces the ridges through the resonances of a bar and density wave spiral arms, and link the ridges with observed ‘wiggles’ in the rotation curve of the model Galaxy, which are similar in amplitude to the wiggles in the rotation curve of the Milky Way, and external galaxies.

Another way of exploring kinematic structure is through the use of action-angle coordinates. For example, Trick, Coronado & Rix (2019a) use actions to show that the local moving groups which have long been observed in the Solar neighbourhood  $v_R$ - $v_\phi$  plane are local, low- $J_R$  manifestations of extended structure in orbit space. They show that while the Hercules, Hyades, and Sirius moving groups show diffusion to higher  $J_R$ , creating high- $J_R$  ridges at constant slope  $\Delta J_R/L_z$ , the Pleiades and Coma Berenices moving groups do not, which suggests a potentially different origin. They also point out that the ridges are related to asymmetries in the  $v_R$  distribution and low average vertical action  $J_z$ .

Sellwood et al. (2019) explore the effect of various models of spiral structure on the action-angle distribution, specifically a classical density wave model (e.g. Lin & Shu 1964), a model with overlapping transient modes (e.g. Sellwood & Carlberg 2014), a swing amplification like model (e.g. Toomre 1981) and a dressed mass clump model (e.g. Toomre & Kalnajs 1991). They compare the features apparent in the action-angle distributions with the data and determine that the data best resembles the transient spiral modes model, and strongly disfavours the dressed mass clump model. They state that the swing amplification model does resemble the data, but the features are broad, whereas the data also contains fine structure, closer resembling the transient mode model. They also argue against the quasi-stationary density wave model, but use arguments from previous works as opposed to directly confronting the data.

Monari et al. (2018) also examine the action-angle distribution for the long slow bar model of Pérez-Villegas et al. (2017), and find that various resonances fall in the right location to explain features in the Solar neighbourhood kinematics. In this model, the corotation resonance (CR) causes the lower velocity component of Hercules, the 2:1 outer Lindblad resonance (OLR) causes the hat feature, and the 3:1, 4:1, and 6:1 OLRs cause structure in the centre of the velocity distribution near the low-velocity moving groups. The importance of considering such higher order bar components was originally demonstrated in Hunt & Bovy (2018), where we showed that the 4:1 OLR of the long slow bar model created structure around the low-velocity moving groups, and also that for a slightly faster bar the 4:1 OLR would create a Hercules-like feature. The strength of Hercules and the hat are not an exact match to the data, and the model does not predict a multiple component Hercules, but other factors like non-axisymmetries such as the spiral arms or an external perturber may account for this. E.g. in Hunt et al. (2018b) we showed that the combination of a long bar potential similar to that explored by Monari et al. (2018) in combination with transient winding spiral arms could reproduce the Solar neighbourhood kinematics, including a double Hercules feature.

As we gain the ability to trace kinematic signatures further across the disc, it will become more important to think in actions and angles because velocity labels will no longer trace coherent structure, even when our imperfect knowledge of the potential and non-axisymmetries become an issue. As such, it is important not only to look at the data in action-angle space, but also the models which we use to explore the effects of non-axisymmetric structure.

In this paper, we explore the effects of the Galactic bar and the spiral structure on the kinematics of stars in the Solar neighbourhood in both the distribution of actions and angles, and  $R_G$ - $v_\phi$  plane. In

Section 2, we describe the selection and treatment of the *Gaia* data. In Section 3, we describe the construction of our models and then compare them with the data. Finally, in Section 5 we summarize the results.

## 2 THE GAIA DATA

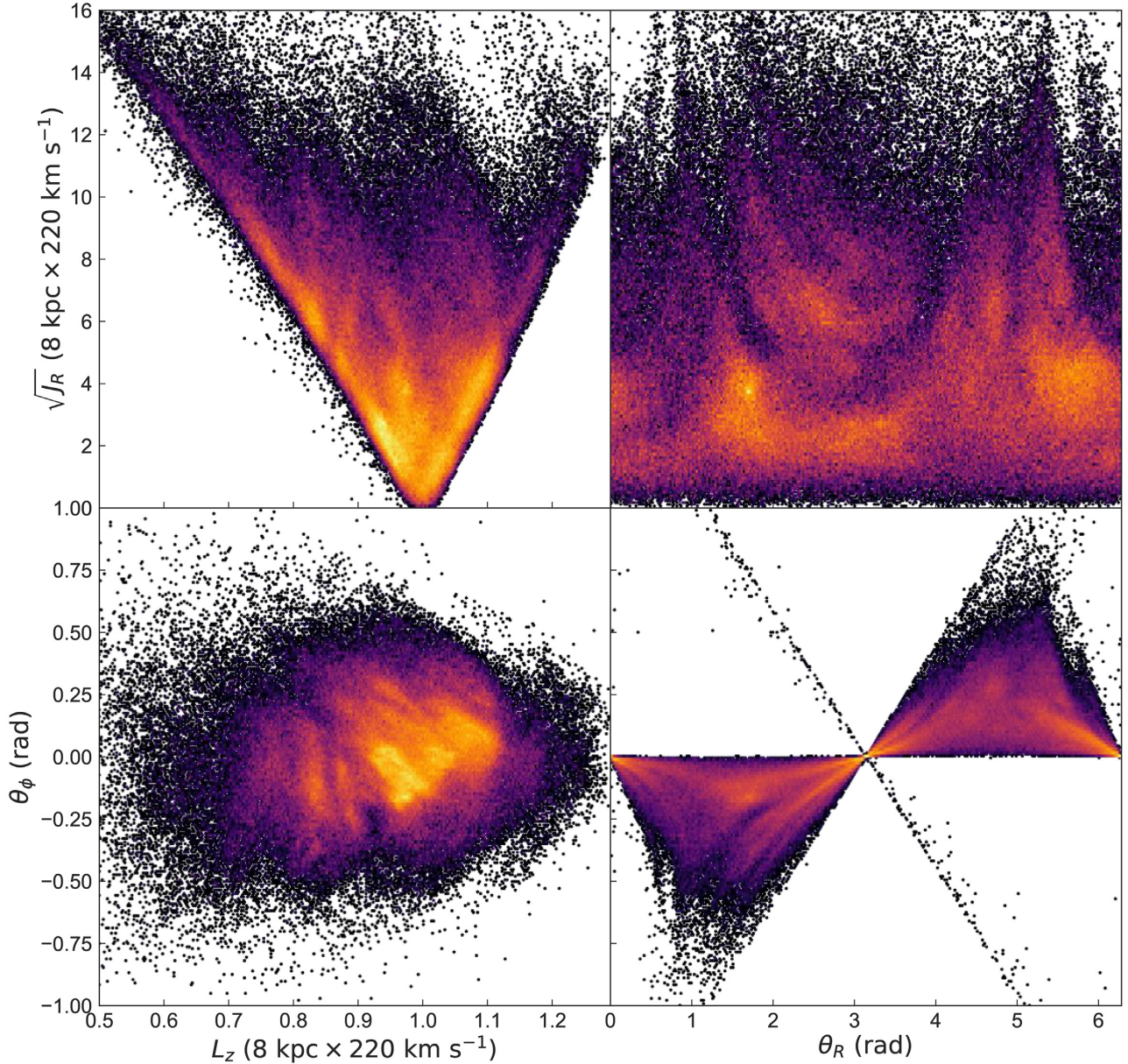
In this section, we describe the selection and treatment of the *Gaia* data, and the resulting action and angle distribution. We perform a simple selection of all stars with the full 6D phase space information which are within 200 pc, with fractional parallax errors of less than 10 per cent. We calculate distances naively as  $d = 1/\pi$ , which is reliable only for such low parallax errors. We transform the velocities from the *Gaia* frame to Galactocentric cylindrical coordinates using *galpy* (Bovy 2015), assuming a distance to the Galactic centre of  $R_0 = 8$  kpc, a circular velocity at the Solar radius of  $v_{\text{circ}} = 220$  km s<sup>-1</sup> (Bovy et al. 2012), and a Solar peculiar motion of  $v_\odot = (U, V, W) = (11.1, 12.24, 7.25)$  km s<sup>-1</sup> (Schönrich, Binney & Dehnen 2010). While these are fairly standard choices for the above parameters, the effect of these assumptions on the resulting action-angle distribution will be described later.

### 2.1 The Solar neighbourhood in action-angle space

We then calculate the actions, angles, and frequencies of the stars using the `actionAngleStaeckel` (Binney 2012) function in *galpy*, assuming the `MWPotential2014` potential, which is fit to various observational constraints (Bovy 2015). We calculate the delta parameter using the `estimateDeltaStaeckel` routine (e.g. as described in Sanders 2012).

Fig. 1 shows the resulting distribution of planar actions and angles for stars in our sample, similar to that explored in Trick et al. (2019a) and Sellwood et al. (2019). The left column is the azimuthal action  $J_\phi$ , more commonly called angular momentum (around the  $z$ -axis),  $L_z$ , and the top row  $y$ -axis is radial action (essentially a measure of radial eccentricity), presented as  $\sqrt{J_R}$  following Trick et al. (2019a). We set the units of the actions to be  $(8 \text{ kpc} \times 220 \text{ km s}^{-1})$  such that a circular orbit at the assumed Solar position has  $L_z = 1$ . The choice to square root  $J_R$  highlights structure at low radial action, which contains the majority of the substructure, in particular the classical moving groups. However, resonant features will lie along a curve for  $\sqrt{J_R}$ , as opposed to a straight line when visualizing  $J_R$ . The location and slope of such features were already analysed in Trick et al. (2019a), and are not a focus of this work. The lower row is the azimuthal angle  $\theta_\phi$ , which is the azimuth of the guiding centre with respect to the Sun. The right column is the radial angle  $\theta_R$  which is the phase of the star on its epicycle (with pericenter at  $\theta_R = 0$ , apocenter at  $\theta_R = \pi$ ). The lower left panel of Fig. 1 closely resembles the  $v_R$ - $v_\phi$  plane which is explored in numerous other works, with clear divisions between the classical moving groups. This is unsurprising as  $L_z = R \times v_\phi$ , and  $R_G$  is very similar for all stars in the volume. Similarly, over a small area,  $\theta_\phi$  should be an approximately linear map to  $v_R$  (for an illustration, see McMillan 2011). For example, stars with prograde orbits with  $0 < \theta_R < \pi$  are on the outwards moving part of their epicycle, and have guiding centre azimuths  $\theta_\phi < 0$ , and stars with  $\theta_R > \pi$  are on the inwards moving part of their epicycle, with guiding centre azimuths  $\theta_\phi > 0$ . For example, a star with a larger  $|\theta_\phi|$  has its guiding centre azimuth further from the Sun, and larger  $|v_R|$ . This divides the  $\theta_R$ - $\theta_\phi$  plane clearly into two sections, with the thin diagonal line of stars which does not follow this prescription being on counter rotating orbits.

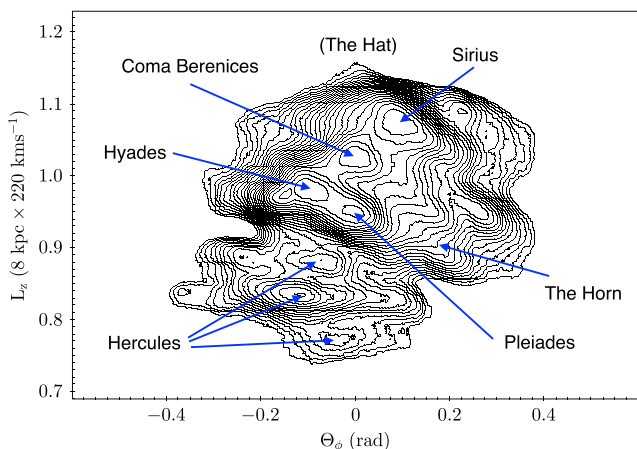




**Figure 1.** Action-angle distribution for *Gaia* stars with radial velocities, within 200 pc, with fractional parallax errors of less than 10 per cent. The left column shows angular momenta against radial action (upper row) and against azimuthal angle (lower row). The right column shows radial angle against radial action (upper) and azimuthal angle (lower).

Fig. 2 shows the density contour map of the  $\theta_\phi$ – $L_z$  plane in the standard orientation, with the moving groups labeled. Note that Hercules clearly decomposes into multiple peaks, as shown in other works (e.g. Antoja et al. 2018; Ramos, Antoja & Figueras 2018; Trick et al. 2019a). Also note the vertex deviation in the contours such that moving groups at higher  $L_z$  occur at higher  $\theta_\phi$ . The Hyades cluster is also visible, as the tiny peak directly to the left of the Hyades moving group. We also label the horn feature (e.g. Dehnen 2000; Fragkoudi et al. 2019) and the hat feature which is visible around  $L_z = 1.2$  in Fig. 1, but is below the minimum strength for the contour map. Fig. 3 shows contour maps for the other three projections, with the same moving groups labelled. The top panel of Fig. 3 shows the density contour map for the  $L_z$ – $\sqrt{J_R}$  plane.

The middle panel shows the density contour map for the  $\theta_R$ – $\sqrt{J_R}$  plane. Note that Coma Berenices is missing from this projection, because it is located around  $\theta_R = 0$  (or  $2\pi$ ) and is split across both edges of the plot. The lower panel of Fig. 3 shows the density contours of the  $\theta_R$ – $\theta_\phi$  plane. In this projection, Hercules does not form distinct peaks, but follows ridges (more clearly visible as three ridges in Fig. 1), and as with the centre panel Coma Berenices is split across either edge and is not labelled. The hat and the horn do not appear as distinct peaks in these projections. However, there are other obvious structures in the action-angle distributions such as the peaks to the ‘left’ of Sirius in the  $\theta_R$ – $\sqrt{J_R}$  and  $\theta_R$ – $\theta_\phi$  planes around  $\theta_R = 4.5$ , which correspond to the diagonal extension of the contour lines around  $(\theta_\phi, L_z) = (0.2, 0.95)$  in Fig. 2 (see Trick et al.



**Figure 2.** Logarithmic number density contours of the  $\theta_\phi$ – $L_z$  plane as seen by *Gaia* DR2, which as expected closely resembles the  $v_R$ – $v_\phi$  plane. The classical moving groups are labeled along with the horn and the location of the hat, which is below the minimum density for the contours.

2019a, for a more thorough examination of the additional structure in action space).

It is not surprising that structure in the Solar neighbourhood velocity distribution is clear in the action angle distribution. Over the small area of the sample, all stars essentially have the same  $R_G$ ,  $\phi$ ,  $z$  coordinates, and the velocities become effectively orbit labels. For spatially more extended survey volumes, actions are the preferred orbit labels. Over a highly localized volume any orbit structure should show up in velocity space, and any velocity structure should show up in orbit space, irrespective of the assumed Galactic potential. So the substructure in Fig. 1 is largely unaffected by the assumption of MWPotential2014. The exact values of the actions and angles will change with a different assumption of  $R_0$ ,  $v_{\text{circ}}$ , and  $v_\odot$ , but the features will stay qualitatively the same.

## 2.2 The resonance criteria

As discussed in the introduction, various models exist to explain the substructure in the kinematics of the Solar neighbourhood. A significant fraction propose that the substructure arises from various resonances of the Galactic bar or spiral structure. Assuming that this is the case, then we should be able to fit the overdensities and gaps in the action-angle distribution to specific resonances, arising from a rigidly rotating structure, such as is done in Monari et al. (2018).

The frequencies which are calculated alongside the actions and angles enable us to compute the ‘resonance criteria’, for which stars with a specific action and angle (and assuming a Milky Way potential) are affected by such a resonance. For example, the corotation resonance for a rigidly rotating structure with pattern speed  $\Omega_p$  occurs when  $\Omega_p - \Omega_\phi = 0$ . The inner (ILR) and outer (OLR) 2:1 Lindblad resonances occur when  $\Omega_p - \Omega_\phi \pm \Omega_R/2 = 0$ , and similarly for the 4:1 ILR and OLR, when  $\Omega_p - \Omega_\phi \pm \Omega_R/4 = 0$ , and the 1:1 ILR and OLR when  $\Omega_p - \Omega_\phi \pm \Omega_R = 0$ .

Fig. 4 shows the action angle distribution from Fig. 1 overlaid with the stars which are close to the resonance criteria for the CR (red), the 4:1 OLR (orange) and the 2:1 OLR (yellow), when assuming a rigidly rotating structure with a pattern speed of  $\Omega_p = 1.15 \times \Omega_0$ , where we assume a local circular frequency of  $\Omega_0 = 27.5 \text{ km s}^{-1} \text{ kpc}^{-1}$ . We are not claiming this is a best-fitting

model for the Milky Way, merely using it as an illustration of the axisymmetric resonance lines, and the similarity between the slope of the resonance lines and the gaps and overdensities in the data.

In this simple model, the 2:1 OLR (yellow) lies in the high angular momenta region of orbit space, and could be responsible for the faint arch feature found there. The 4:1 OLR (orange) corresponds roughly to the sharp transition around  $L_z \sim 1.1$  and the CR (red) falls roughly in line with Hercules. However, such a correlation does not imply causation, and a different choice of pattern speed will align different resonances with different features. In addition, the resonance criteria can only be calculated for the axisymmetric potential used to calculate the actions, angles, and frequencies, and thus while the resonance lines are symmetric in angle, this is not the case in the *Gaia* data. In addition, calculations of frequency are affected by the choice of potential which we will illustrate in a later section. Thus, because the assumed potential affects the location of the resonance criteria lines, and the choice of Solar peculiar motion affects the action-angle distribution directly, the combination of effects make it very challenging to directly measure the pattern speed of the structure causing gaps or overdensities from such an analysis alone, even if we assume that such features arise from a single rigidly rotating structure such as the Galactic bar, which may not be the case.

## 3 THE MODEL

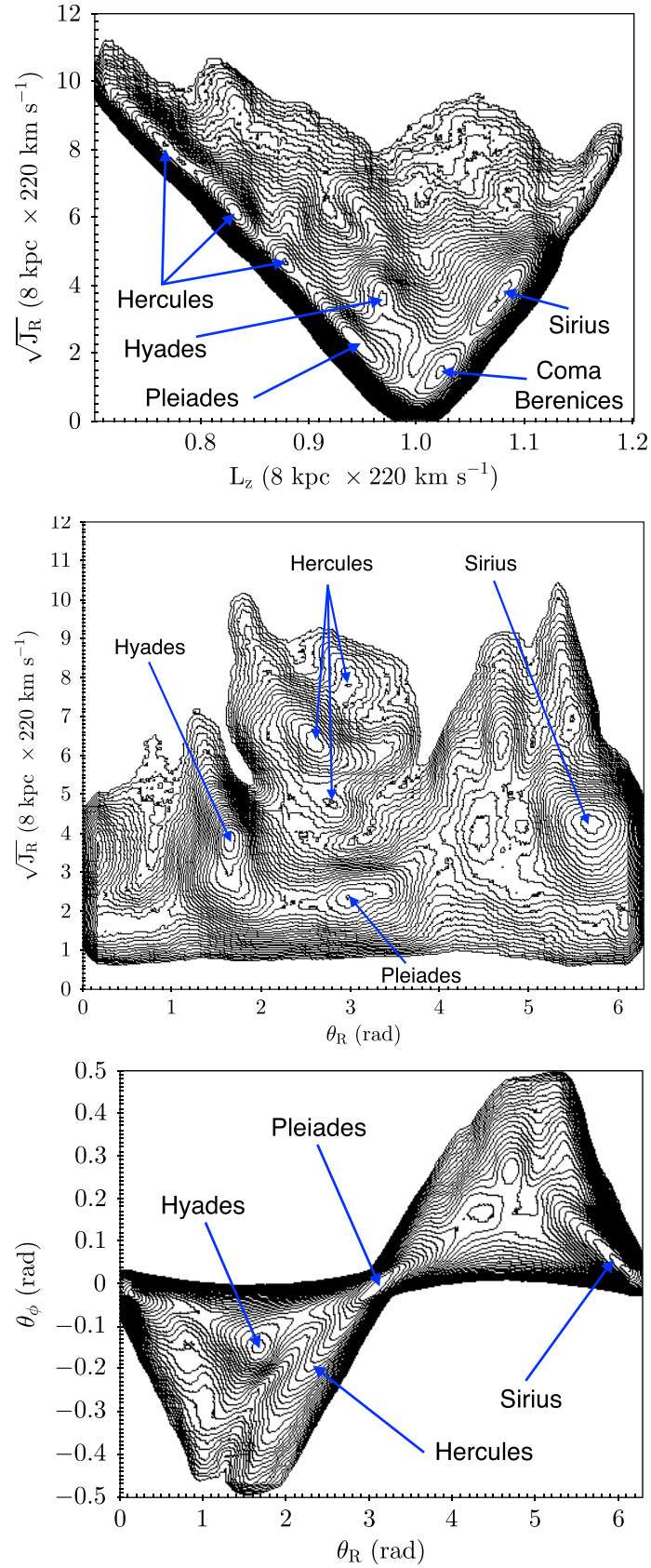
To explore the effects of calculating actions for a non-axisymmetric system, using an axisymmetric potential, and the assumption of a potential, we create 2D models using the backwards integration technique of Dehnen (2000) (e.g. following Hunt & Bovy 2018; Hunt et al. 2018b). We transform them into action-angle space by sampling  $10^7$  phase-space points from the resulting distribution, which was constructed over a grid of  $1001 \times 1001$  velocities at the location of the Sun with the `evolveddiskdf.meanVR` function, and converting these to action-angle space using the `actionAnglesSpherical` function with the `LogarithmicHaloPotential`.

One issue with the backwards integration technique is that it provides the fine-grained distribution function, which never smooths out on sufficiently small scales, whereas ideally we would compare the data with the course grained distribution function (Fux 2001). In addition, when integrating for a low number of bar periods some of the features in the velocity distribution are transient, resulting from incomplete phase mixing (e.g. see Fux 2001, for a thorough exploration of this effect in the context of the short fast bar model). Despite these caveats we used the backwards integration method to explore a large range of parameter space.

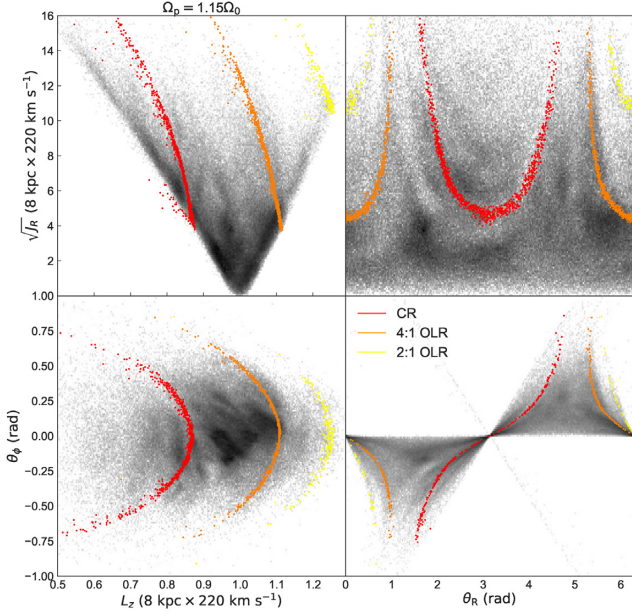
As discussed further below, we have performed tests of the backwards integration method that demonstrate that while the transient features due to the rapid growth of the bar are most apparent in the short fast bar model as explored in Fux (2001), they are only minor in the case of a long bar alone and are negligible in the presence of transient spiral structure, the main focus of this work. We also performed test particle simulations for two of the more interesting cases as shown in Section 4.5, both to explore the time evolution of the models, and as a consistency check for the backwards integration models.

As was done in Hunt et al. (2018a), Hunt & Bovy (2018), and Hunt et al. (2018b) we use a Dehnen distribution function (Dehnen 1999), which is a function of energy,  $E$ , and angular momentum,  $L$ , to model the stellar disc before bar and spiral





**Figure 3.** Upper: The  $L_z - \sqrt{J_R}$  plane. Middle: The  $\theta_R - \sqrt{J_R}$  plane. Lower: The  $\theta_R - \theta_\phi$  plane, all in logarithmic number density with the classical moving groups labelled when visible.



**Figure 4.** Same as Fig. 1, except overlaid with the CR (red), the 4:1 OLR (orange) and the 2:1 OLR (yellow) calculated with the axisymmetric frequencies for a rigidly rotating pattern of  $\Omega_p = 1.15 \times \Omega_0$ .

formation. We reproduce the equations here for convenience. The Dehnen distribution function represents the distribution of stellar orbits such that

$$f_{\text{dehnen}}(E, L) \propto \frac{\Sigma(R_c)}{\sigma_R^2(R_c)} \exp \left[ \frac{\Omega(R_c)[L - L_c(E)]}{\sigma_R^2(R_c)} \right], \quad (1)$$

where  $R_c$ ,  $\Omega(R_c)$ , and  $L_c$  are the radius, angular frequency, and angular momentum, respectively, of a circular orbit with energy  $E$ . We assume a simple power law for the gravitational potential such that the circular velocity is given by

$$v_c(R) = v_0(R/R_0)^\beta, \quad (2)$$

where  $v_0$  is the circular velocity at the solar circle at radius  $R_0$ .

For all subsequent bar models, we use the general form of the  $\cos(m\phi)$  potential shown in Hunt & Bovy (2018), adapted from the quadrupole potential from Dehnen (2000). In *galpy*, this is the *CosmphiDiskPotential* model. The bar potential is given by

$$\Phi_b(R, \phi) = A_b(t) \cos(m(\phi - \phi_b t)) \times \begin{cases} -(R/R_0)^p, & \text{for } R \geq R_b, \\ ([R_b/R]^p - 2) \times (R_b/R_0)^p, & \text{for } R \leq R_b, \end{cases} \quad (3)$$

where  $\phi_b$  is the angle of the bar with respect to the Sun–Galactic-center line, and the bar radius,  $R_b$ , is set to 80 per cent of the corotation radius. The potential is equivalent to the Dehnen (2000) quadrupole bar for  $m = 2$  and  $p = -3$ , where  $m$  is the integer multiple of the  $\cos$  term, and  $p$  is the power-law index.

The bar is grown smoothly using *galpy*’s *DehnenSmoothWrapperPotential* such that

$$A_b(t) = \begin{cases} 0, & t/t_b < t_1 \\ A_f \left[ \frac{3}{16}\xi^5 - \frac{5}{8}\xi^3 + \frac{15}{16}\xi + \frac{1}{2} \right], & t_1 \leq t/t_b \leq t_1 + t_2, \\ A_f, & t/t_b > t_1 + t_2, \end{cases} \quad (4)$$

where  $t_1$  is the start of bar growth, set to half the integration time,  $t_2$  is the duration of the bar growth and  $T_b = 2\pi/\Omega_b$  is the bar period such that

$$\xi = 2 \frac{t/T_b - t_1}{t_2} - 1, \quad (5)$$

and

$$A_f = \alpha_m \frac{v_0^2}{p} \left( \frac{R_0}{R_b} \right)^p, \quad (6)$$

where  $\alpha_m$  is the dimensionless ratio of forces owing to the  $\cos(m\phi)$  component of the bar potential and the axisymmetric background potential,  $\Phi_0$ , at Galactocentric radius  $R_0$  along the bar’s major axis. This growth mechanism ensures that the bar amplitude along with its first and second derivatives are continuous for all  $t$ , allowing a smooth transition from the non-barred to barred state (Dehnen 2000).

For our spiral arm potential, both density wave and corotating, we use the *SpiralArmsPotential* from *galpy*, which is an implementation of the sinusoidal potential from Cox & Gómez (2002) such that

$$\Phi(R, \phi, z) = -4\pi G H \rho_0 \exp \left( \frac{r_0 - R}{R_s} \right) \times \sum \frac{C_n}{K_n D_n} \cos(n\gamma) \left[ \text{sech} \left( \frac{K_n z}{\beta_n} \right) \right]^{B_n}, \quad (7)$$

where

$$K_n = \frac{nN}{R \sin(\theta_{\text{sp}})}, \quad (8)$$

$$B_n = K_n H (1 + 0.4 K_n H), \quad (9)$$

$$D_n = \frac{1 + K_n H + 0.3(K_n H)^2}{1 + 0.3 K_n H}, \quad (10)$$

$$\gamma = N \left[ \phi - \phi_{\text{ref}} - \frac{\ln(R/r_0)}{\tan(\theta_{\text{sp}})} \right], \quad (11)$$

$N$  is the number of spiral arms,  $\theta_{\text{sp}}$  is the pitch angle,  $\rho_0$  is the density at  $r_0$ ,  $\phi_{\text{ref}}$  is the reference angle,  $R_s$  is the radial scale length of the arm, and  $H$  is the scale height of the arm. Setting  $C_n$  to 1 gives a purely sinusoidal potential profile. Alternatively, setting  $C_n = [8/3\pi, 1/2, 8/15\pi]$  results in a potential which behaves approximately as a cosine squared in the arms, and is flat in the interarm region (Cox & Gómez 2002). Note that while equation (7) gives the full form available in *galpy*, we use the planar form  $\Phi(R, \phi, z = 0)$ , which sets the sech term to 1.

For the quasi-stationary density wave like spiral arms, we assign a fixed pattern speed  $\Omega_{\text{sp}}$ , and grow the potential in the same fashion as the bar. In this work, we also grow the spiral over the same time-scale.

For the corotating, winding spiral potential, we wrap the *SpiralArmsPotential* from equation (7) in *galpy*’s *CorotatingRotationWrapperPotential*, such that

$$\phi \rightarrow \phi + \frac{V_p(R)}{R} \times (t - t_0) + a_p \quad (12)$$

and

$$V_p(R) = V_{p,0} \left( \frac{R}{R_0} \right)^\beta, \quad (13)$$

**Table 1.** Model parameters for Models A–H, including the bar half length,  $R_b$  (kpc), the bar pattern speed,  $\Omega_b \times \Omega_0$ , the spiral pattern speed  $\Omega_{sp} \times \Omega_0$  and the origin of the model if applicable.

Label	$R_b$ (kpc)	$\Omega_b (\times \Omega_0)$	$\alpha_2$	$\alpha_4$	$\Omega_{sp} (\times \Omega_0)$	Model
A	3.5	1.85	0.01	0	None	Dehnen (2000)
B	5.0	1.3	0.024	0.001	None	Pérez-Villegas et al. (2017)
C	5.0	1.4	0.01	0.0005	None	Hunt & Bovy (2018)
D	3.5	1.85	0.01	0	0.56	
E	3.2	1.82	0.01	0	0.92	Martinez-Medina et al. (2018)
F	5.0	1.3	0.01	0.0005	Winding	Hunt et al. (2018b)
G	3.8	1.6	0.01	0.0005	Winding	
H	3.5	1.85	0.01	0.0005	Winding	

where  $V_p(R)$  is the circular velocity curve,  $t_0$  is the time when the potential is unchanged by the wrapper, and  $a_p$  is the position angle at time  $t_0$ . This causes the arm to wind up over time, as seen in  $N$ -body simulations. This model is designed to mimic the transient recurrent arms which corotate with the stars at all radii, e.g. as described in Grand, Kawata & Cropper (2012a) and Hunt et al. (2018b). It is not designed to reproduce the classic picture of swing amplification (e.g. Toomre 1981).

We then weight the amplitude with a Gaussian using the `GaussianAmplitudeWrapperPotential` to control the strength of the transient arm, where the amplitude gets multiplied with the function

$$A(t) = \exp\left(-\frac{[t - t_0]^2}{2\sigma^2}\right), \quad (14)$$

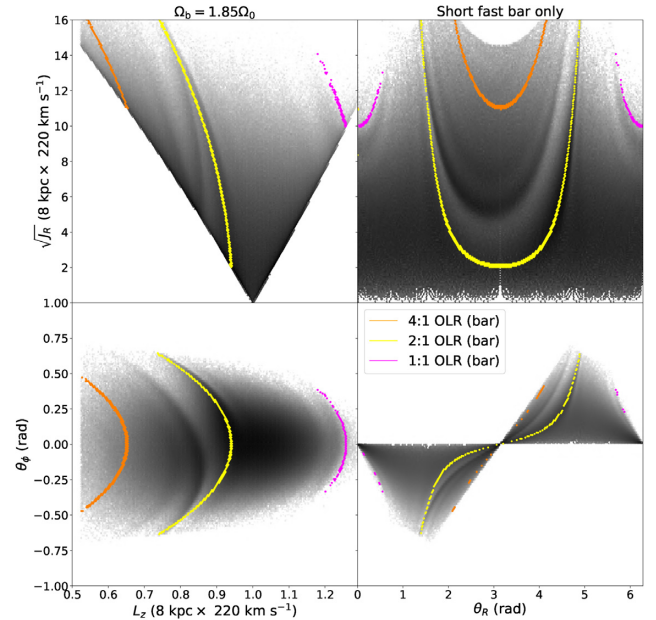
and  $\sigma$  is the standard deviation of the Gaussian, which controls the lifetime of the transient spiral potential. Although the wings of the Gaussian technically stretch to infinity, the density enhancements last approximately  $\mathcal{L} \approx 5.6 \times \sigma$  from formation to disruption. This is a simple model potential to approximate the corotating recurrent arms observed in  $N$ -body simulations.

## 4 MODEL RESULTS

### 4.1 Bar only

As an initial illustration, we present the action angle distribution for a few bar models from the literature which have been claimed to reproduce the Hercules stream, i.e. a short fast bar (from Dehnen 2000), a long slow bar (similar to Pérez-Villegas et al. 2017) and a slightly faster long bar (from Hunt & Bovy 2018). All models in this work are integrated for 10 bar periods, with their parameters given in Table 1. As mentioned above, note that 10 bar periods is not enough for the system to reach a quasi-stationary regime (e.g. Fux 2001). However, in this section we merely show existing models from the literature in action-angle space without providing a detailed critique of the method.

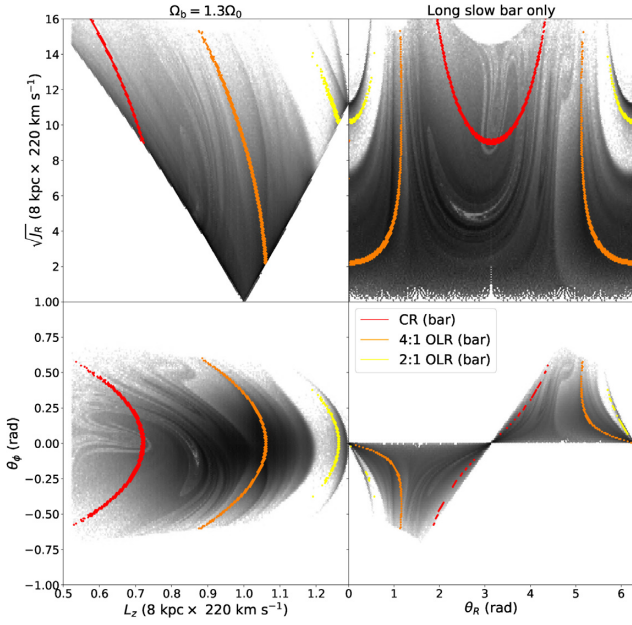
Fig. 5 shows the action-angle distribution for Model A, a short fast bar, which has been proposed to reproduce the Hercules stream through the OLR (e.g. Dehnen 2000; Hunt et al. 2018a). The model parameters used here are  $R_b = 3.5$  kpc,  $\Omega_b = 1.85 \times \Omega_0$ ,  $\alpha_2 = 0.01$ ,  $\alpha_4 = 0$ ,  $R_0 = 8$  kpc, and  $v_0 = 220$ . The gap and overdensity in the distribution corresponding to the OLR (yellow) is clear in each panel. Very little structure exists in the rest of the distribution, including the area around the 4:1 OLR (orange), which is unsurprising as this simple bar model consists of a purely  $m = 2$  bar. The exact shape of the Hercules-like feature does not perfectly match that of the *Gaia* data, although a better fit can likely be constructed by altering the bar angle, or other parameters

**Figure 5.** Action-angle distribution for Model A, the short-fast bar of Dehnen (2000), with  $R_b = 3.5$  kpc, and  $\Omega_b = 1.85 \times \Omega_0$ . The coloured lines show the location of the 4:1 OLR (orange), the 2:1 OLR (yellow) and the 1:1 OLR (magenta).

(e.g. see Dehnen 2000, for an exploration of the model parameter space). Interestingly, even this simple model creates three ridges in the  $L_z$ – $J_R$  plane, which are qualitatively consistent with the data. The rightmost ridge which lies exactly under the yellow OLR line corresponds to the horn-like feature, then slightly to its left is a ridge with very little slope corresponding to the strongest ridge of the Hercules-like feature, and then another ridge to its left with a higher slope which is a result of the incomplete phase mixing described above (see Fux 2001). The ability for this simple model to reproduce a multiple component Hercules feature is not surprising, smaller secondary peaks are visible in some of the models in Dehnen (2000). There is a slight hat feature resulting from the 1:1 OLR (magenta). In general, this model has been well explored in the literature (e.g. Bovy 2010; Antoja et al. 2014; Monari et al. 2017), and comfortably reproduces the Hercules stream, but none of the other moving groups, or features such as the arches. This is not new, merely shown for comparison.

Fig. 6 shows the action-angle distribution for Model B, a long slow bar, which has been proposed to reproduce the Hercules stream through the CR (e.g. Pérez-Villegas et al. 2017). Note that we are not exploring the model of Pérez-Villegas et al. (2017) specifically,

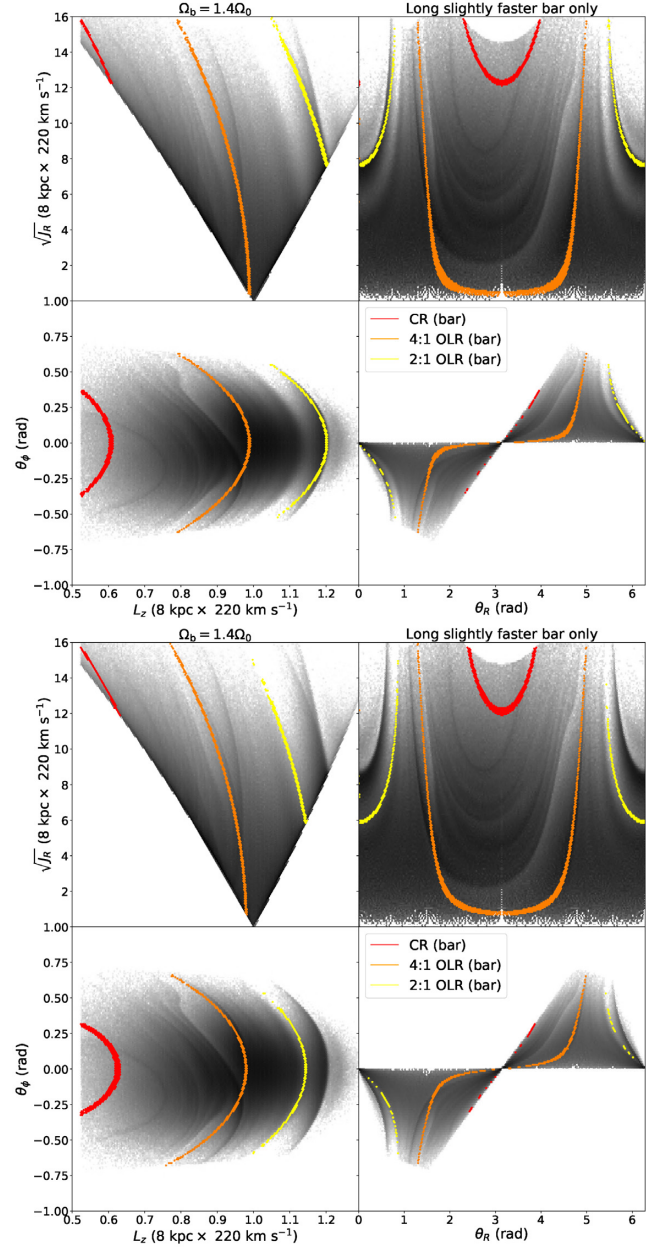




**Figure 6.** Action-angle distribution for Model B, the long slow bar resembling Pérez-Villegas et al. (2017), with  $R_b = 5$  kpc, and  $\Omega_b = 1.3 \times \Omega_0$ . The coloured lines show the location of the CR (red), the 4:1 OLR (orange), and the 2:1 OLR (yellow).

(which was recently done in Monari et al. 2018), but the major features of a bar with this pattern speed should be qualitatively the same. The model parameters used here are  $R_b = 5$  kpc,  $\Omega_b = 1.3 \times \Omega_0$ ,  $\alpha_2 = 0.024$ ,  $\alpha_4 = -0.001$ ,  $R_0 = 8.4$  kpc, and  $v_0 = 242$ . The model creates a hat-like feature at high  $L_z$  with the bar’s OLR (yellow). As with other works (e.g. Hunt & Bovy 2018), we find that the CR (red) has only a slight effect on local kinematics, even when including a higher bar mass. As shown in Monari et al. (2018), the resonance criteria lines fall in the correct locations to explain features in the velocity distribution, but as with other studies (e.g. Binney 2018; Hunt et al. 2018a) the Hercules-like feature is weaker than observed, and the hat-like feature is stronger than observed. However, the addition of the  $m = 4$  bar component to the model (as proposed in Hunt & Bovy 2018) produces an excellent qualitative reproduction of the division between the Sirius and Coma Berenices like streams in the  $L_z$ - $\theta_\phi$  plane, including its shape and angle, and qualitatively similar in the  $\theta_R$ - $J_R$  plane. Also, while the CR does not make a strong Hercules feature alone, it’s been shown that the addition of spiral structure in combination with the CR can create a strong distinct Hercules (Hunt et al. 2018b), and Monari et al. (2018) show that the  $m = 3$  and  $m = 6$  resonances also create structure for the model of Pérez-Villegas et al. (2017). Again, many of the fine features in between the resonance lines are the result of incomplete phase mixing following the growth of the bar, which are visible owing to the fine grained nature of the distribution function as discussed above. We do not expect them to be present in the data.

The top panel of Fig. 7 shows the action-angle distribution for Model C, a slightly faster long slow bar with an  $m = 4$  Fourier component which has been proposed to reproduce the Hercules stream through the 4:1 OLR (Hunt & Bovy 2018). The model parameters used here are  $R_b = 5$  kpc,  $\Omega_b = 1.4 \times \Omega_0$ ,  $\alpha_2 = 0.01$ ,  $\alpha_4 = -0.0005$ ,  $R_0 = 8$  kpc, and  $v_0 = 220$ . The 4:1 OLR (orange) creates a Hercules-like feature in the distribution, which also has a similar



**Figure 7.** Action-angle distribution for Model C, the long slow  $m = 4$  bar of Hunt & Bovy (2018), with  $R_b = 5$  kpc, and  $\Omega_b = 1.4 \times \Omega_0$ . The coloured lines show the location of the 2:1 OLR (yellow), the 4:1 OLR (orange), and the CR (red) for actions, angles and frequencies calculated in the correct axisymmetric potential (upper) and a different axisymmetric potential with the wrong rotation curve (lower).

angle to the gap in the *Gaia* data, but is weaker than observed. In addition, the 2:1 OLR (yellow) creates a distinct hat feature, which is qualitatively similar to the *Gaia* data, but is too strong and occurs at too low  $L_z$ . The lower panel of Fig. 7 shows the same model, C, where the actions, angles, and frequencies are calculated using the wrong axisymmetric potential, namely MWPotential2014 from galpy, instead of the LogarithmicHaloPotential which was used to construct the model. MWPotential2014 has a falling rotation curve, and while the action angle distribution barely changes, the resonance criteria lines which are based on the frequencies do shift. The effect on the CR (red) and 4:1 OLR

(orange) is relatively small, but the OLR (yellow) shows a significant shift between panels. This is as expected because the falling rotation curve will be more different from the flat rotation curve of the `LogarithmicHaloPotential` at larger Galactic radii.

All three bar models can qualitatively reproduce one or more of the features in the local velocity distribution, but none of them can reproduce all the features of the observed Solar neighbourhood kinematics. For example, the short fast bar model (e.g. Dehnen 2000) produces a good reproduction of Hercules but nothing else. The long slow bar model of Pérez-Villegas et al. (2017) produces good fits to the hat and low-velocity moving groups but does not produce a distinct Hercules. The slightly faster long bar model of Hunt & Bovy (2018) produces Hercules, but not the low-velocity moving groups, and the hat is too strong. It is worth noting that both of the long slow bar models are consistent within error with the recent measurements of the bar pattern speed (e.g. Clarke et al. 2019; Sanders, Smith & Evans 2019), depending on the choice of local circular frequency. For further exploration of local kinematics in bar only models, see Trick et al. (2019b). A brief exploration of integration time confirmed the findings of Fux (2001) for the short fast bar. Similar behaviour was observed in the long bar cases, but to a lesser degree.

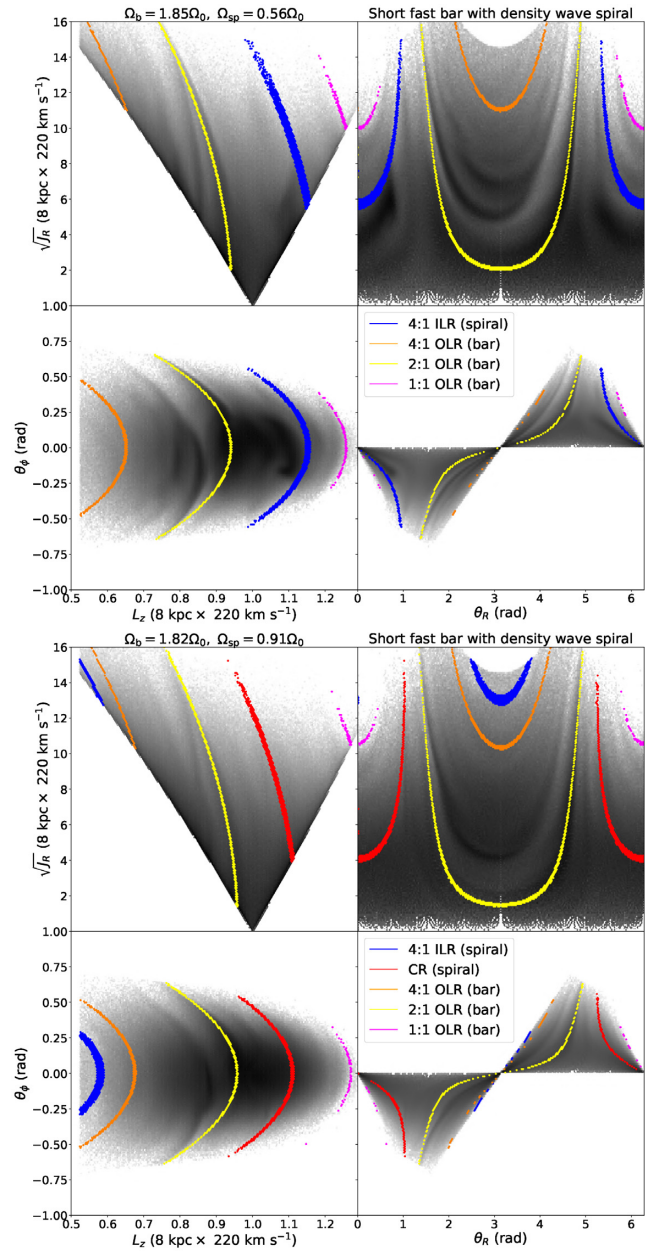
It is likely that the bar model used here is still too simple (e.g. see Monari et al. 2018, for the effects of higher order components on the velocity distribution). However, we do not expect the bar only models to produce a perfect representation of Solar neighbourhood kinematics because spiral structure will have some effect regardless of its nature, and so will the phase wrapping known to be occurring in the disc (e.g. Antoja et al. 2018) regardless of whether it originates from transient spiral arms (Hunt et al. 2018b), or an external perturbation (e.g. Laporte et al. 2018).

## 4.2 Density wave spirals

We do not perform a detailed comparison of the different theories of spiral structure alone (see e.g. Dobbs & Baba 2014, for a review of spiral structure), which was the focus of the recent work of Sellwood et al. (2019). Instead, we show the effect of the combination of both a bar and spiral structure, for both the classical quasi-stationary density wave fixed pattern speed spiral of Lin & Shu (1964) in this section, and the corotating winding arm as seen e.g. in Grand et al. (2012a) in Section 4.3 for bars of varying lengths and pattern speeds. We do not focus on making the best possible reproduction of the Solar neighbourhood kinematics, but merely show that the combination of bar and spirals lead to significant deviations from the velocity distribution for the bar alone.

A density wave like spiral, with a fixed pattern speed will have resonances at fixed radii in the same way as the bar described above. However, the spiral wave is expected to be slower than the bar, and thus the 4:1 ILR (blue) and the CR (red) of the spiral are the most likely to affect the velocity distribution in the Solar neighbourhood, whereas the OLRs will affect stars further from the Galactic centre.

We performed the action-angle modelling for an  $N = 2$  and  $N = 4$  density wave like spiral arm with pitch angle 12 deg (e.g. following the average measured angle of the Perseus arm; Vallée 2015), and with a range of pattern speeds from  $\Omega_{sp} = 0.2 \times \Omega_0$  to  $1.2 \times \Omega_0$  in combination with each of the three bar models described above. There are too many plots to display within the paper but we would be happy to provide any model upon request. Fig. 8 shows an example from the  $N = 4$  density wave spiral combined with a short fast bar for spiral pattern speeds  $\Omega_{sp} = 0.56 \times \Omega_0$  (upper), and  $\Omega_{sp} = 0.91 \times \Omega_0$  (lower).



**Figure 8.** Action-angle distribution for Model D, the short-fast bar of Dehnen (2000), with  $R_b = 3.5$  kpc, and  $\Omega_b = 1.85 \times \Omega_0$  combined with a  $N = 4$  density wave spiral with pattern speed  $\Omega_{sp} = 0.56 \times \Omega_0$  (upper) and Model E, a short-fast bar, with  $R_b = 3.2$  kpc, and  $\Omega_b = 1.82 \times \Omega_0$  combined with a  $N = 4$  density wave spiral with pattern speed  $\Omega_{sp} = 0.91 \times \Omega_0$  (lower). The coloured lines show the location of the bar 2:1 OLR (yellow), the bar 4:1 OLR (orange), the spiral CR (red), and the spiral 4:1 ILR (blue).

The upper panel of Fig. 8 shows Model D, the model with  $\Omega_{sp} = 0.56 \times \Omega_0$ , for which the 4:1 ILR (blue) lies around  $L_z = 1.1$ , corresponding to the ridge at the top of the main velocity distribution. The orange, yellow, and magenta lines are the bar 4:1, 2:1 and 1:1 OLR, respectively. The addition of the spiral density wave to the short bar model causes arches in the  $L_z - \theta_\phi$  plane, which are qualitatively similar to the split between the moving groups in the *Gaia* data, and the  $\theta_R - J_R$  plane shows similar features around  $\theta_R = 0$  and  $2\pi$ . More interestingly, the second component of the Hercules like feature is stronger than in the bar only case, becoming



more consistent with the data. The distribution as a whole does not have the correct vertex deviation, but otherwise it is a reasonably good match to the *Gaia* data considering the simplicity of the model. While there is a strong edge to the velocity distribution around  $L_z = 1.1$  in the  $L_z$ – $\theta_\phi$  plane, the ridge does not extend to high  $J_R$  as is seen in the data, and the hat is not as distinct. We selected this combination of pattern speed for the bar and spiral arms because it was qualitatively the best-fitting combination of a bar and density wave spiral from the sampled parameters when comparing the action-angle model to the data.

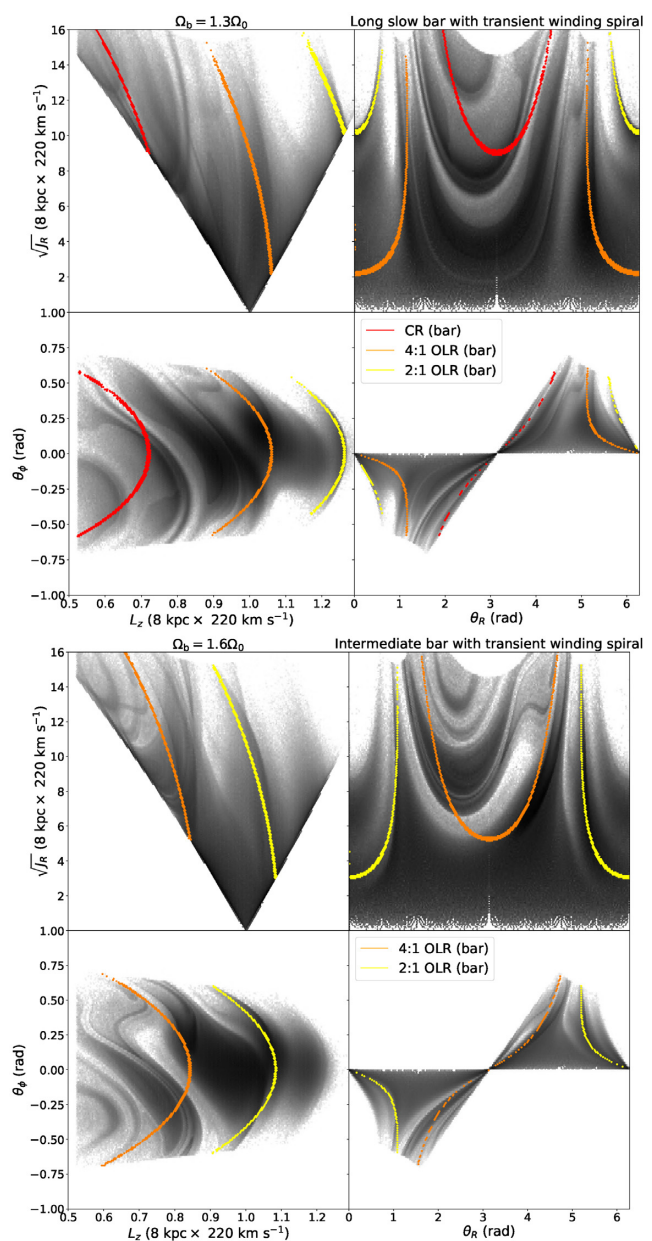
The lower panel of Fig. 8 shows Model E, a model chosen to mimic the model from Martinez-Medina et al. (2018), where the ridge structure in the  $R_G$ – $v_\phi$  plane was proposed to arise solely from a combination of bar and spiral resonances. The model is similar to the above panel except in this model the bar length is  $R_b = 3.2$  kpc, and has pattern speed  $\Omega_b = 1.82 \times \Omega_0$ . The spiral has a pattern speed of  $\Omega_{sp} = 0.91 \times \Omega_0$  and pitch angle 15.5 deg. Note that we do not have the same potential or the full parameters of their model. Hence, this will be only a qualitative reproduction based on the relevant spiral pattern speed, but the resonances should fall in the same area of the disc. Unlike the upper panel there is very little structure outside the bar OLR region. The spiral CR (red) creates a small amount of arching around  $L_z = 1$ , but it is very weak compared to the data.

### 4.3 Corotating recurrent arms

The origin and processes behind the transient winding arms commonly seen in  $N$ -body simulations are still under debate, i.e. whether they arise from a series of overlapping transient modes (e.g. Quillen et al. 2011; Comparetta & Quillen 2012; Sellwood & Carlberg 2014), or whether they are fully corotating recurrent arms which arise through a non-linear process similar to swing amplification (e.g. Wada, Baba & Saitoh 2011; Grand et al. 2012a; Grand, Kawata & Cropper 2012b) which is difficult to explain via the superposition of modes (e.g. Grand et al. 2012a; Kumamoto & Noguchi 2016). However, for this example we choose to reproduce the corotating recurrent arms (e.g. as shown in Grand et al. 2012a), where the arm corotates with the stars at all radii, and does not have a fixed pattern speed. Note that this is a similar but not identical process to the classic swing amplification model proposed by Toomre (1981), and our model is not designed to reproduce a swing amplified arm as suggested in Sellwood et al. (2019).

We again perform a series of action-angle models for a transient winding arm using the method described in Section 3, combined with a bar with varying length and pattern speed from  $3 < R_b < 5$  kpc, and  $1.0 \leq \Omega_b \leq 2.0 \times \Omega_0$ . The parameters of the spiral remain the same across models, with  $N = 2$ ,  $R_s = 0.3$ , and  $C_n = 1$ . We increase the pitch angle at the peak of the density enhancement from 12 deg as assumed in Hunt et al. (2018b) to 25 deg which better matches the value found in Grand, Kawata & Cropper (2013). This increases the strength of the perturbation on the Solar neighbourhood kinematics for an equivalent peak amplitude. The shape of the features are qualitatively the same, but more pronounced for the higher pitch angle at peak density.

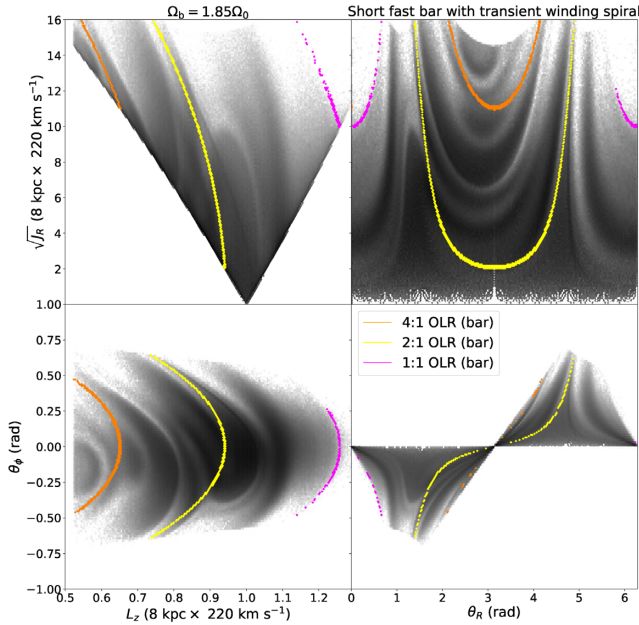
The top panel of Fig. 9 shows Model F, the combination of a long slow bar with pattern speed  $\Omega_b = 1.3 \times \Omega_0$ , and a series of three winding transient arms described above which peak at  $t = -460$ ,  $t = -230$  and  $t = 0$  Myr in the past (e.g. the model from Hunt et al. 2018b, but with a 25 deg pitch angle at the peak). This places the Sun in the interarm region similar to the Milky Way. The model clearly shows the a hat like feature at high  $L_z$  caused by the OLR



**Figure 9.** Action-angle distribution for Model F, a long slow bar with  $\Omega_b = 1.3 \times \Omega_0$  (upper) and Model G, a bar with  $\Omega_b = 1.6 \times \Omega_0$  (lower), both combined with transient winding spirals. The spiral arms in these models corotate at all radii, and thus do not have resonance criteria to plot.

which is similar to the observations from *Gaia*, and the central area of the velocity distribution contains three separate regions qualitatively similar to the Hyades + Pleiades, Coma Berenices, and Sirius moving groups. There is also a distinct Hercules like feature, which is divided into separate substructures as seen in the data. The combination of the long slow bar and winding transient spiral structure naturally leads to the vertex deviation observed in the  $L_z$ – $\theta_\phi$  plane, and the asymmetry around  $\theta_R = \pi$  in the  $\theta_R$ – $\sqrt{J_R}$  plane. It is worth noting that while we selected  $\Omega_b = 1.3 \times \Omega_0$  as the ‘best’ qualitative reproduction of the Solar neighbourhood kinematics for this model, the features are similar over an approximate range of  $\Omega_b \sim 1.26$ – $1.36 \times \Omega_0$ . The bar pattern speed found by Monari et al. (2018) falls within this range, and the features from the bar are consistent across both models. However, the addition of spiral





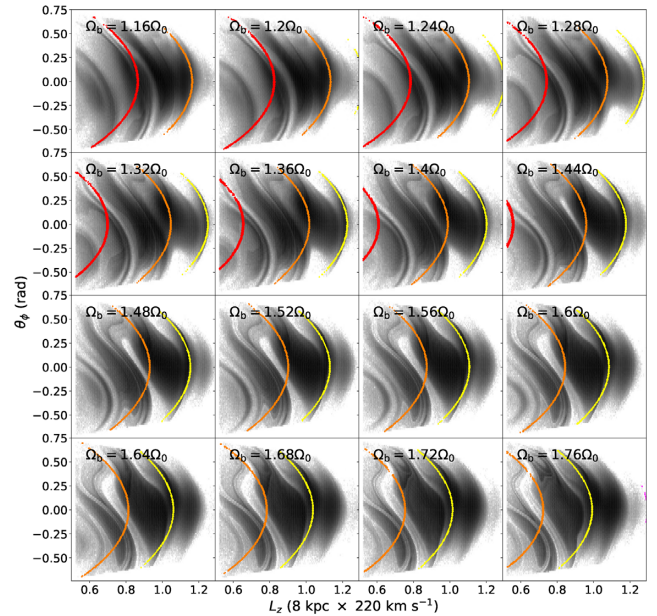
**Figure 10.** Action-angle distribution for Model H, the short fast bar combined with transient winding spirals.

structure adds additional features and strengthens the Hercules like feature.

The lower panel of Fig. 9 shows Model G, the same combination of potentials, but with a bar pattern speed of  $\Omega_b = 1.6 \times \Omega_0$ , which shows a surprisingly good reproduction of the Solar neighborhood kinematics. The Hercules like division is both the cleanest of any combination of parameters which were explored, and both matches the observed angle and location in  $L_z$ . The Hercules like feature is comprised of multiple components, and is asymmetric in  $\theta_R$ . It results from the 4:1 OLR of the bar, and the transient winding spiral arms. The 2:1 OLR occurs at an  $L_z$  of around 1–1.1 and makes a qualitatively similar feature to Sirius. There is then a small arch above this, which matches the hat in the data well. The vertex deviation of the distribution matches the observations, but there is no clear separation into low-velocity moving groups.

Fig. 10 shows Model H, the combination of a short fast  $m = 2$  bar potential with  $\Omega_b = 1.85 \times \Omega_0$  with a series of transient winding spiral arms. These transient arms have the same parameters, but their spacing has been altered such that the peaks occur slightly more frequently at 210 Myr years apart instead of 230 Myr apart. Fig. 10 shows a multiple component Hercules like feature, with a prominent main ridge, a small ridge between it and the Hyades/Pleiades like feature and smaller ridges at lower  $L_z$ . The model contains a small hat feature resulting from the 1:1 OLR of the bar, the correct vertex deviation, along with the strong feature resembling Sirius, although this extends further from  $\theta_\phi = 0$  than is seen in the data.

Again, the full range of pattern speeds are available upon request. We do not claim that any of the above solutions are the correct representation of the Milky Way’s non-axisymmetric structure, but merely that there are a range of solutions that are consistent with the kinematics of the Solar neighbourhood. When exploring a bar only model (e.g. Dehnen 2000; Hunt & Bovy 2018; Monari et al. 2018) the features are distinct, both in action-angle space and the  $v_R$ – $v_\phi$  plane, and there are a very limited number of pattern speeds which are able to account for Hercules, or other features such as the hat. Similarly, when exploring spiral only models such as was

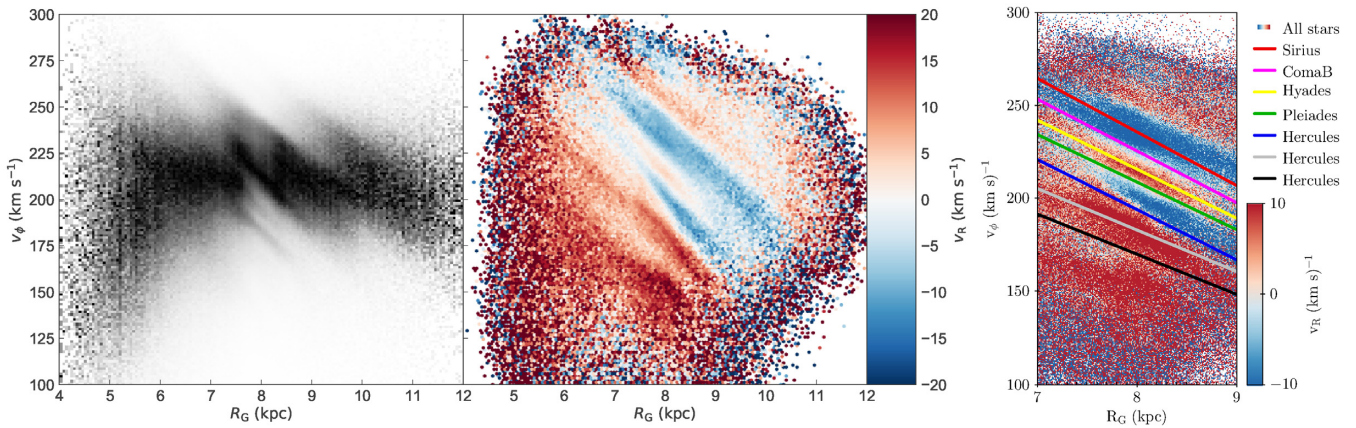


**Figure 11.**  $L_z$ – $\theta_\phi$  distribution for a transient winding spiral model combined with a bar model with a range of pattern speeds from  $\Omega_b = 1.16 \times \Omega_0$  to  $\Omega_b = 1.76 \times \Omega_0$ , and bar radius which decreases with increasing pattern speed. As with earlier figures, the coloured lines mark the location of the CR (red), 4:1 OLR (orange), and the 2:1 OLR (yellow).

done in Sellwood et al. (2019), the features in  $v_R$ – $v_\phi$  and action-angle space are clean and distinct from one another. However, when analysing a combination of bar and spiral structure, the overlapping of resonances and phase mixing make the picture less clear.

To illustrate this, Fig. 11 shows the  $L_z$ – $\theta_\phi$  plane for the same model as Fig. 9, but with a range of pattern speeds from  $\Omega_b = 1.16 \times \Omega_0$  to  $\Omega_b = 1.76 \times \Omega_0$ , and a bar length which decreases with pattern speed. As with earlier figures, the coloured lines show the CR (red), 4:1 OLR (orange), and the 2:1 OLR (yellow) for each bar pattern speed. The panels with  $\Omega_b = 1.28 \times \Omega_0$  and  $\Omega_b = 1.6 \times \Omega_0$  are the closest to the data, where both the CR and 4:1 OLR show divisions around the location of Hercules, and arch features at high  $L_z$ , with similar kinematics for similar pattern speeds. Note that the 4:1 OLR creates a more distinct gap between Hercules and the other moving groups compared to the CR, or the OLR in Fig. 10. Also note that while some features are broad, some features are very fine, in agreement with the data.

We also note that the inclusion of transient spiral arms to either the short fast bar or the long slow bar can be tailored to reproduce the structure which is otherwise missing. E.g. adding spirals to a long slow bar can reproduce a distinct Hercules, which it does not do alone, and adding spirals to a short fast bar can reproduce the structure at higher  $L_z$ , which it does not do alone. We also explored models with higher order bar components, e.g.  $m = 6$  and  $m = 8$ , similar to Monari et al. (2018), and found that the effects of the transient spiral arms washed out the higher order resonances. This is unsurprising considering the lower amplitude of the higher order components (see Buta et al. 2006, for a comparison of the strength of the Fourier components in observed bars), and thus we neglect them here for simplicity. Also note that the transient features induced by the comparatively recent spiral structure wash out the transient features resulting from the incomplete phase mixing owing to the integration over only 10 bar periods as discussed for the bar only models above. We repeated model F over 100 bar periods



**Figure 12.**  $v_\phi$  as a function of Galactocentric radius  $R_G$ , colored by normalized number density (left), and  $v_R$  (centre), for the *Gaia* DR2 data. The right-hand panel shows a linear fit to the approximate location of the local moving groups in this projection.

(with the same occurrences of spiral structure) and obtained almost identical results.

We do not expect any of these models to be a perfect reproduction of the Solar neighbourhood because this is a drastic simplification of the complex dynamical behaviour ongoing in the Milky Way disc, and we have not attempted to fit the time-scale of the phase mixing.

#### 4.4 The $R_G$ – $v_\phi$ plane

In addition to the kinematics of the Solar neighbourhood, we also compare the models further from the Solar circle. Thus, in this section, we show  $v_\phi$  as a function of Galactocentric radius  $R_G$  for the models described above, and compare them with the *Gaia* data, both in terms of stellar number density as we did previously in Hunt et al. (2018b), and also coloured by  $v_R$  as shown in Fragkoudi et al. (2019).

Fig. 12 shows the  $R_G$ – $v_\phi$  plane, coloured by normalized number density (left), and  $v_R$  (centre), for the *Gaia* DR2 data. Unfortunately, even with the unparalleled wealth of data from *Gaia* DR2, we are still only able to trace the ridges a few kpc from the Solar circle. The vertex deviation of the velocity distribution means the moving groups are not fully distinct in  $v_\phi$ . However, some moving groups do dominate some band of the  $v_\phi$ – $R$  plane. For example, the multiple Hercules streams are moving outwards (red) as expected with the uppermost of the two dark red feature being the main peak of Hercules, and the paler feature just above being the smaller peak at slightly higher  $v_\phi$ . The horn feature moving inwards (blue) is just ‘above’ Hercules (as explored in more detail in Fragkoudi et al. 2019). The wide inwards moving (blue) band of stars is mostly Sirius. The outwards moving band just ‘below’ Sirius is mostly the Hyades, but Coma Berenices lies on the transition from the Sirius band to the Hyades band, and the Pleiades also lies along the transition from the Hyades band to the horn. The Coma Berenices and Pleiades moving groups have approximately  $\bar{v}_R = 0$  and thus merely decrease the amplitude of the  $v_R$  signal of the adjacent bands. The right-hand panel of Fig. 12 shows the approximate location of the moving groups in this projection. The lines are a linear fit to the distribution of stars selected within  $2 \text{ km s}^{-1}$  of the density peaks shown in Fig. 2. Note that features arising from resonances will not follow a linear relation in this projection, and the lines are only an approximate representation of their location to aid the discussion.

The range of  $v_R$  has been reduced to accentuate the division between bands.

Fig. 13 shows the logarithmic number density (left column), and the mean radial velocity  $v_R$  (right column) for the previously shown models. Note that the mean  $v_R$  shown in the right-hand panels is not weighted by number density, so features may appear strong even when there are very few stars. Thus, this should not be taken to indicate that the models have an abnormally large velocity dispersion.

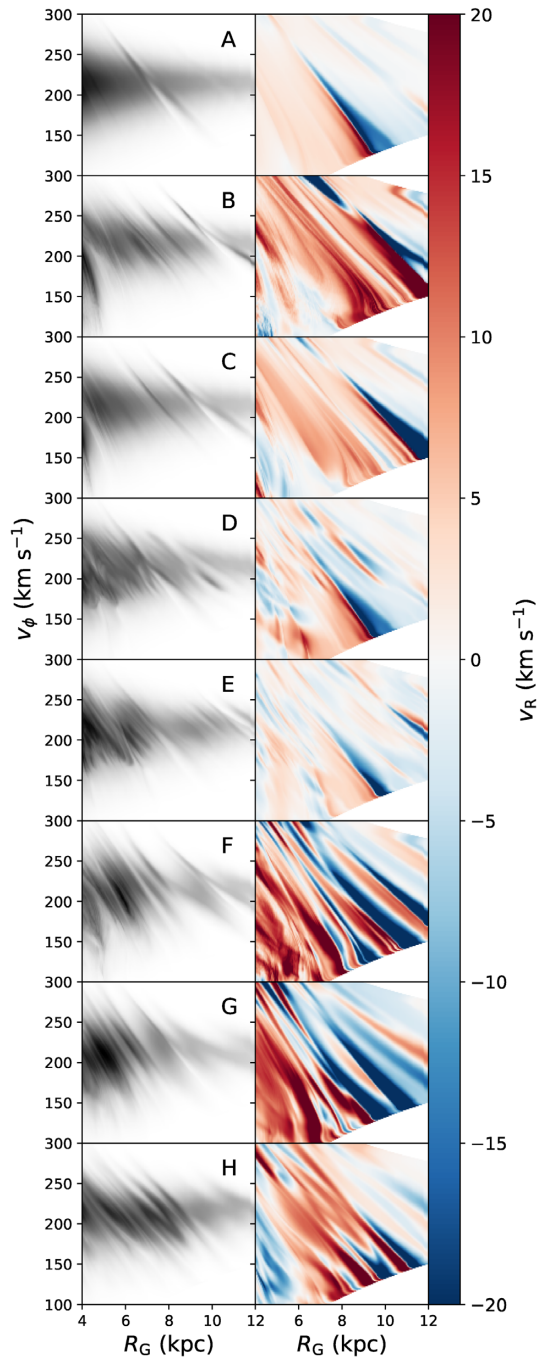
The short fast bar in model A (top row) shows the single OLR feature around the Solar neighbourhood as expected, with the red outwards moving stars corresponding to the Hercules stream, and the blue inwards moving stars corresponding to the horn (the orbital behaviour is explored in detail in Fragkoudi et al. 2019). There are minor effects present at other radii including a weak red-blue feature around the 1:1 OLR, and a small secondary Hercules-like ridge, but the distribution is mostly smooth away from the 2:1 OLR.

The long slow bar in model B (second row) shows a lot more structure away from the 2:1 OLR which occurs further out in the Galaxy and is identifiable from the sharp diagonal transition from red to blue. The 1:1 OLR makes another weaker transition in the top right corner of the panel. There are multiple red ‘stripes’ resembling Hercules, but the model lacks the alternating inwards and outwards moving groups of stars which are present in the data, with very little blue. In particular, there is no blue horn-like feature.

The slightly faster long bar model C (3rd row) shows very similar behaviour to model B, which is unsurprising considering they both contain an  $m = 4$  component. However, the slightly faster pattern speed puts the OLR nearer the Solar circle, and the lower bar strength reduces the relative strength of the inwards and outwards mean velocities. This lower strength allows a small blue horn-like feature to become apparent above the Hercules-like feature (for a comparison of the effect of the strength of the  $m = 4$  component see Hunt & Bovy 2018).

The short fast bar + four armed density wave model D (4th row) more closely resembles model A. E.g. the 2:1 OLR is the only strong feature in the Solar neighbourhood. The combination of the bar and spiral resonances create multiple ridges in the number density  $R_G$ – $v_\phi$  plane, which is qualitatively consistent with the data. However, when coloured by  $v_R$  there is significantly less variation than is seen in the data. There is a red blue feature around the 4:1 ILR of the density wave, but weaker than seen in the data.





**Figure 13.**  $v_\phi$  as a function of Galactocentric radius  $R_G$ , coloured by number density (left column), and  $v_R$  ( $\text{km s}^{-1}$ ) (right column), for models A, B, C, D, E, F, G, and H (top to bottom). Their parameters are given in Table 1.

The short fast bar + four armed density wave spiral with pattern speeds following Martinez-Medina et al. (2018), model E (fifth row) shows very similar behaviour to model D. There is ridge structure in the density panel, essentially following the same distribution as in Martinez-Medina et al. (2018), but the  $v_R$  panel shows only weak structure outside the OLR region, including at the spiral resonances.

The long slow bar + transient winding spirals model F (6th row) shows more strongly alternating inwards and outwards moving stars,

which is to be expected from a phase mixing model (e.g. see Khanna et al. 2019). That these overdensity ridges at approximately constant  $L_z = R \times v_\phi$  are related to regions of predominantly outwards or inwards motion was first shown by Trick et al. (2019a). The OLR again occurs further out in the disc, but the Solar neighbourhood contains multiple Hercules like features, a blue Sirius like feature, and then the alternating red and blue in the hat like feature caused by the 2:1 OLR, which is a good match to the data. There is also a very thin blue horn like feature, but the Hyades like feature is too strong compared to the data.

The bar with  $\Omega_b = 1.6 \times \Omega_0$  + transient winding spirals model G (seventh row) shows a similar mix of inwards and outwards moving stars consistent with phase mixing, but the amplitude of the variation is much higher. Again, there are multiple Hercules like features, and multiple blue horn like features are present, as are the red Hyades and the blue Sirius like features. However, the inwards and outwards moving components of the hat-like feature are opposite to what is observed, with the inwards moving stars occurring at lower  $v_\phi$  than the outwards moving stars, contrary to the data.

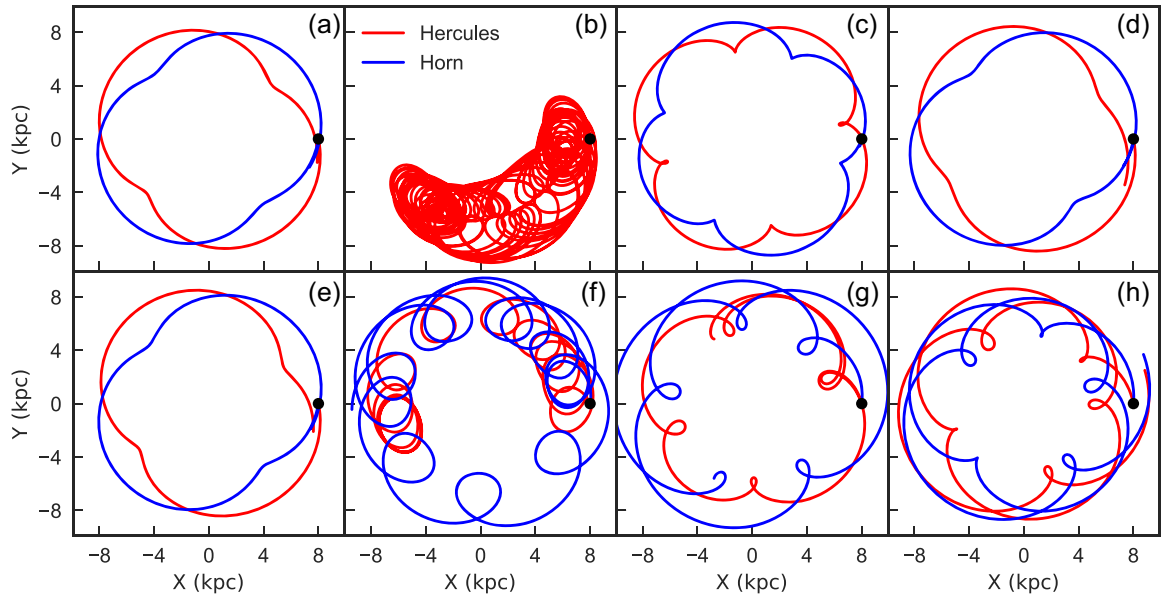
The short fast bar + transient spirals model H (bottom row) again shows some alternating ridges. However, interestingly the amplitude of the mean  $v_R$  is significantly lower outside the area affected by the bar OLR. This appears to be consistent with models F and G, but the effect is stronger here. While there are distinct multiple Hercules-like features, and a small horn-like feature, there are considerably too many red outwards moving features which are not consistent with the data.

In general, the models combining the bar with transient spiral structure are qualitatively a better fit to the *Gaia* data, with strong alternating inwards and outwards moving groups of stars. The bar only models affect too narrow a region of the phase space with various resonances, and the density wave spirals models do not leave a strong signal in  $v_R$  with the spiral induced features being approximately symmetrical around  $v_R = 0$  (e.g. see Fig. 8). The three transient spiral phase mixing models are qualitatively consistent with the *Gaia* data, but none are an excellent match. This is unsurprising because no attempt has been made to match the (currently unknown) time-scale of the phase mixing event, and the backwards integration models described above in Section 3 are evaluated only once at a single time-step.

We also include here a brief illustration of the orbits of stars in the Hercules-like and horn-like features for the above models in the reference frame of the bar. Fig. 14 shows an example orbit for a star currently at the position of the Sun (black dot) which is in the Hercules-like region of velocity space (red) and the horn-like region of velocity space (blue) for models A–H. Note that model B does not contain a clear horn. The orbits which (approximately) close are integrated for one orbit in the bar frame, and the orbits which do not close are integrated longer to show the behaviour.

For the bar only and spiral density wave cases, the orbital dynamics are already well explored in the literature (e.g. see Contopoulos & Grosbøl 1989, for a detailed exploration of orbit structure in a barred and density wave spiral potential). In the models where the Hercules-like and horn-like features originate from the 2:1 OLR (Models A, D, and E) the orbits belong to the  $x_1(1)$  and  $x_1(2)$  orbit families (see also Dehnen 2000; Fragkoudi et al. 2019), where they experience two radial epicycles per rotation in the bar frame. For the model where the Hercules-like feature originates from the CR (Model B) the orbits are banana-like orbits trapped around the  $\mathcal{L}_4$  and  $\mathcal{L}_5$  Lagrange points (see also Pérez-Villegas et al. 2017). In the model where the Hercules-like and horn-like features originate from the 4:1 OLR (Model C) the stars belong to





**Figure 14.** Example orbit for a star currently at the position of the Sun (black dot) which is in the Hercules-like region of velocity space (red) and the horn-like region of velocity space (blue) for models A–H. Note that model B does not contain a clear horn.

the  $x_1(2)$  orbit family and experience 4 radial epicycles per rotation in the bar frame as expected.

Adding the density wave potential (Models D and E) has only a slight impact on the 2:1 orbits in the Hercules-like and horn-like features. The density wave spiral of course has its own resonances, and resulting resonant orbits which are not a focus of this work (for more detail see e.g. Contopoulos & Grosbøl 1989). However, when adding the transient winding spirals (Models F, G, and H) the orbits no longer close in the bar frame. This is unsurprising considering the highly time varying nature of the potential. We defer a more detailed exploration of the orbital structure in such potentials to a future work.

#### 4.5 Test particle simulations

We also run high-resolution 3D test particle models for models F and H (selected as the two closest to the data) in order to explore how the phase mixing features change with time. This also serves as a useful check that the results from the simple models constructed with the backwards integration method are consistent with the test particle simulations.

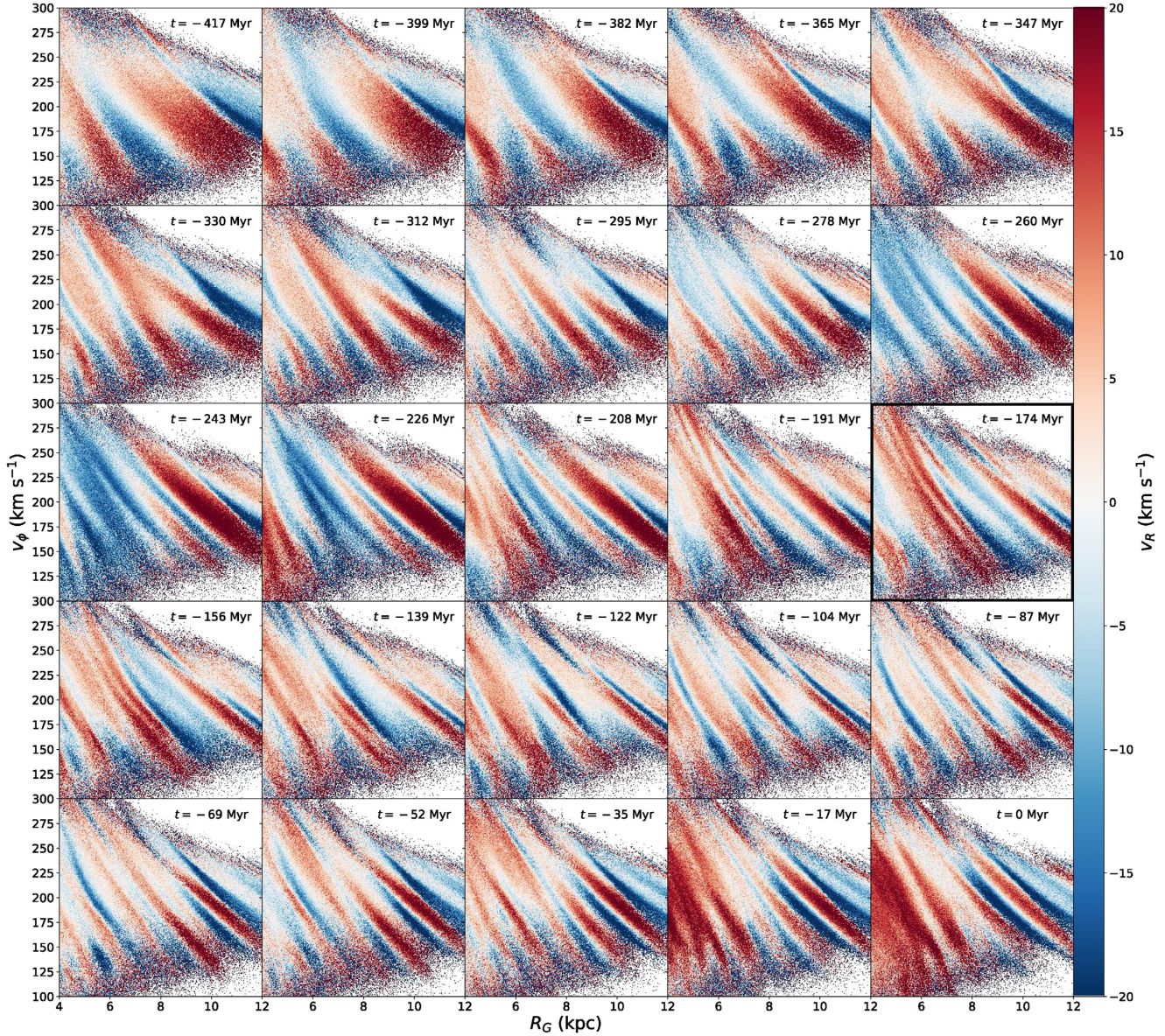
The initial conditions for the test particle models are sampled using `galpy`'s 3D `quasiisothermaldf`, a distribution function adapted from Binney (2010) which is expressed in terms of action-angle variables instead of the 2D `Dehndf` used earlier. The distribution function is initialized with a radial scale length of  $R_0/3$ , a local radial and vertical velocity dispersion of  $0.15 \times V_c(R_0)$  and  $0.075 \times V_c(R_0)$ , respectively, and a radial scale length of  $R_0$  for the velocity dispersions. Action-angles are calculated using the Staackel approximation via `galpy`'s `actionAngleStaackel`. The potential is evaluated with `MWPotential2014`.  $2 \times 10^8$  position and velocity samples are generated between radii of 1 and 13 kpc, and are integrated forwards in time for 7 Gyr using `MWPotential2014` to allow the disc to reach equilibrium prior to the model integration. The bar and spiral potentials used for models F and H are the 3D versions of those given above, and both

models are then integrated forwards for 10 bar periods using the `MWPotential2014` potential.

Fig. 15 shows the time evolution of the  $R_G-v_\phi$  plane for the test particle realization of model F, i.e. a long slow bar combined with transient winding spirals coloured by mean  $v_R$ . The time is given in the top corner of each panel from  $t = -417$  to  $t = 0$  Myr, with  $t = 0$  being the present day. The spirals peak at  $t = -460$ ,  $t = -230$ , and  $t = 0$  Myr. The simulation has the same overall time-scale as the backwards integration models, as we integrate forwards from  $t = -1.72$  Gyr to 0 (after the relaxation phase). All panels are taken at a Galactic azimuth such that the bar is at 25 deg from the line of sight from the Galactic centre to the observer, and the particles are selected in a wedge of  $\pm 15$  deg around the Galactic centre-Sun line. The panel at  $t = 0$  is equivalent to the right-hand panel in the sixth row (F) of Fig. 13, and they show excellent agreement, validating the backwards integration  $R_G-v_\phi$  planes. The test particle realization shows more blurring between features owing both to the finite resolution and the sampling of a wedge of particles, whereas the backwards integration model is evaluated along an exact line from the Galactic centre to anticentre. We remind the reader that as with Fig. 13, the mean  $v_R$  shown in the right-hand panels is not weighted by number density, so features may appear strong even when there are very few stars. Thus, this should not be taken to indicate that the models have an abnormally large velocity dispersion.

The bar OLR is visible in every panel, with the transition from outwards to inwards moving features occurring around 10–12 kpc. The strength of the feature varies between time-steps, but never gets fully obscured by the effects of the phase mixing from the transient spiral arms. The data in Fig. 12 is very noisy in the hat region, but the trend of a slight inwards moving component at high  $v_\phi$ , with the outwards moving component between it and Sirius is consistent with the tip of the OLR feature. While the  $t = 0$  panel is a reasonable match to the data, with a strong Hercules stream, and also Sirius and Hyades-like features, the panels for  $t = -174$  (outlined) and  $t = -156$  show a significantly better reproduction of the Solar





**Figure 15.**  $v_\phi$  as a function of Galactocentric radius  $R_G$ , coloured by mean  $v_R$  for a series of time-steps from  $t = -417$  Myr to  $t = 0$  (present day) for the test particle realization of Model F. The best-fitting panel at  $t = -174$  Myr is outlined.

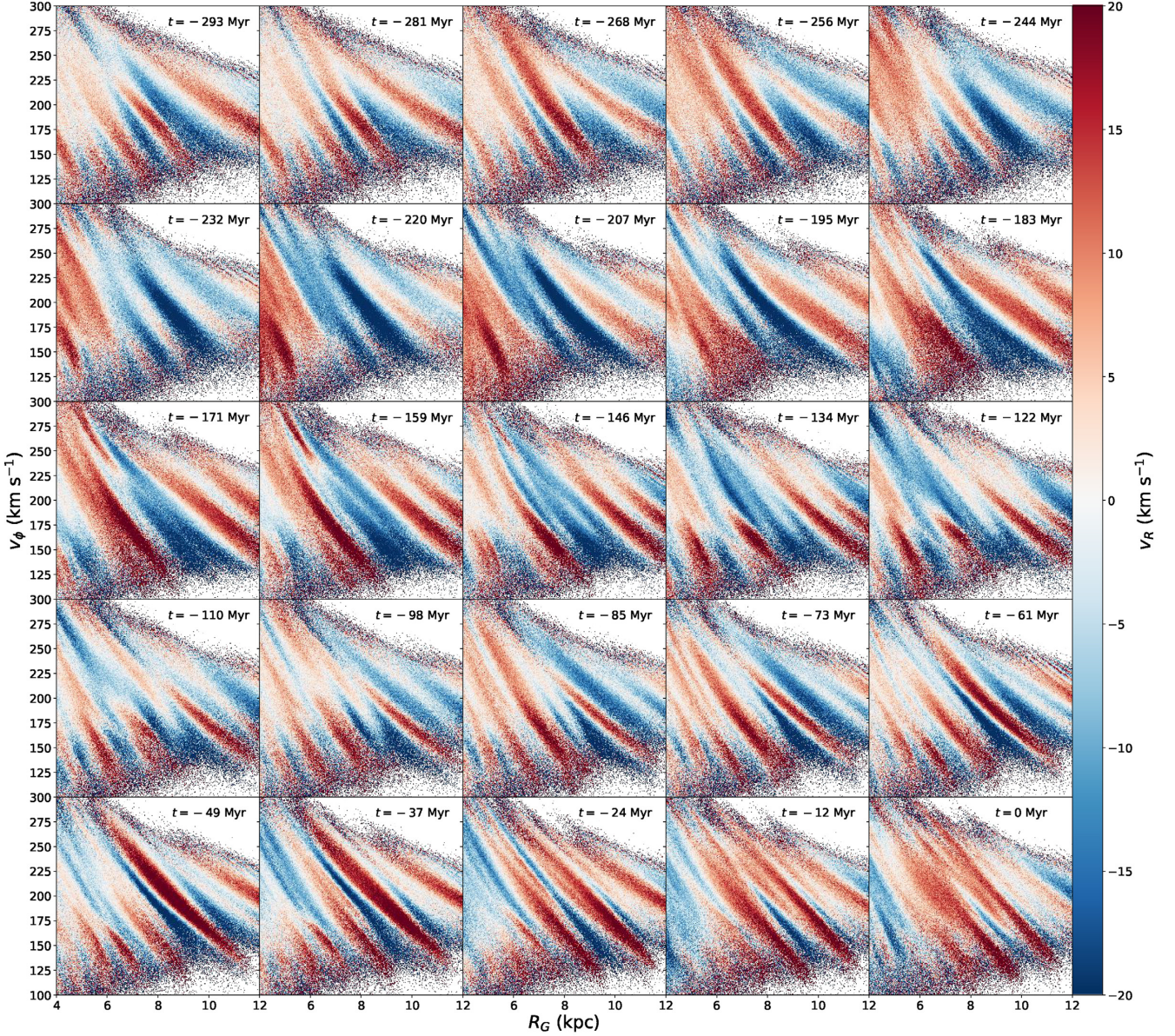
neighbourhood kinematics. For example, around  $R_G = 8$  kpc there is a double Hercules-like outwards moving feature, a tapering inwards moving feature resembling the horn, a thin Hyades-like feature and a broad Sirius-like feature. However, the red/blue transition in the hat is stronger than in the data.

In general, the inwards moving part (blue) of the OLR feature maintains a similar shape across most panels, where it is broad at large  $R_G$  and low  $v_\phi$  and becomes narrower towards lower  $R_G$  and high  $v_\phi$ . The strength and extent of the blue OLR feature does fluctuate between panels, but overall the feature remains qualitatively the same, even when it becomes very thin. However, the outwards moving (red) part of the OLR feature shows significant variation over the time-scale displayed in Fig. 15. In some panels, e.g.  $t = 0$ , the outwards moving part of the OLR feature closely follows the inwards moving feature, while maintaining an approximately constant width across the covered  $R_G$  and  $v_\phi$  range. However, in

other panels, the outwards moving part of the OLR feature splits significantly from the inwards moving part, either leaving a gap between them at higher  $v_\phi$  (e.g.  $t = -295$  Myr), or leaving a thin strip of outwards moving stars next to the inwards moving stars, while the stronger feature splits (e.g.  $t = -347$  Myr). This behaviour is qualitatively similar to the Hercules stream observed in the Solar neighbourhood albeit at a different  $R_G$  and  $v_\phi$ .

For example, Fig. 12 shows that the horn becomes broader at higher  $R_G$ , lower  $v_\phi$  and while the strongest part of the Hercules stream itself (the gray line in Fig. 12) splits from the horn, there is a smaller group of outwards moving stars which remain next to the horn (blue line in Fig. 12). These separate groups are also visible as two of the multiple components of Hercules in the  $v_R$ - $v_\phi$  or  $L_z$ - $\theta_\phi$  plane. While the OLR of the long slow bar cannot be responsible for this behaviour in the Solar neighbourhood, it is conceivable that the same effect can happen around a different





**Figure 16.**  $v_\phi$  as a function of Galactocentric radius  $R_G$ , colored by mean  $v_R$  for a series of time-steps from  $t = -293$  Gyr to  $t = 0$  Gyr (present day) for the test particle realization of Model H.

resonance, e.g. the CR or the 4:1 OLR. This interpretation, if correct, would mean that the smaller peak (or ridge) is the outwards moving component of the resonance feature, whereas the larger peak of Hercules originates from the coupling of this resonance and the phase mixing. We defer a more detailed exploration of this to future work.

Fig. 16 shows the time evolution of the  $R_G$ - $v_\phi$  plane for the test particle realization of model H, i.e. a short fast bar combined with transient winding spirals, where the arms peak at  $t = -420$ ,  $t = -210$ , and  $t = 0$  Myr. In this model, the OLR feature which occurs around 7–9 kpc is no longer clear at every time-step, despite the constant bar angle. However, multiple panels show a clear horn-like feature, and Hercules-like feature, with a tapering inwards moving group of stars, and a strong outwards moving group at slightly lower  $v_\phi$  resulting from the 2:1 OLR. Note that this is similar to the findings of Fujii et al. (2019), who observe a Hercules-like

feature slightly outside the OLR in approximately 50 per cent of the sampled time-steps in their high-resolution  $N$ -body simulations, depending on the complex interaction between the bar and transient spiral structure.

The panel at  $t = -73$  Myr contains a strong horn-like feature and a double Hercules-like feature, and also a broad Sirius-like feature. There is a thin strip of outwards moving stars qualitatively consistent with Hyades, but the amplitude of the  $v_R$  is much stronger than seen in the data. The panels where  $t = -159$ ,  $-146$  Myr contain similar behaviour as discussed above, where a small strip of outwards moving stars remain next to the inwards moving feature, and then a stronger feature splits away. These two panels also contain a thin slightly outwards moving feature just outside the horn-like feature which is qualitatively consistent with Hyades, and then a broad inwards moving feature just outside that which is consistent with Sirius. The hat-like feature then contains the



split outwards and inwards moving features, although the outwards moving group is stronger than seen in the data. The panel at  $t = 0$  is a significantly worse fit to the data, with the majority of the stars in the Solar neighbourhood moving outwards. There is a small horn-like feature, and a strong Hercules-like feature, but the rest of the distribution is a poor match to the data. Some of the phase mixing ridges survive across 100 Myr, whereas others only appear in one or two panels, i.e. only around 25 Myr, especially when interacting with the OLR. The trend observed in Fig. 13, where the phase mixing ridge features were weaker outside the bar 2:1 OLR does not appear to be consistent across panels, and thus cannot be reliably used to infer the location of the OLR.

Further work is necessary to fully explore the time-scale of the phase mixing, both regarding how long the features last from a single perturbation when combined with a barred potential, and also the effects of how frequently the perturbations occur, but we defer this to a future work.

## 5 SUMMARY

In this work, we have investigated the effects of the combination of bar and spiral potentials on the action-angle distribution of stars in the Solar neighbourhood.

In general, in both the action-angle planes and  $R_G-v_\phi$  planes we find it is relatively easy to qualitatively reproduce observations via a combination of a bar and transient winding spiral arms. In particular, we can explain a Hercules-like feature with the 2:1 OLR, 4:1 OLR or CR when combining the bar with transient spiral structure when comparing only to the Solar neighbourhood kinematics. With respect to the spiral structure, quasi-stationary density wave spirals make significantly less impact on the explored kinematics, even when resonances fall in the right area of phase space, and bar-only models simply cannot account for all the observed structure, which is not surprising.

Both the long slow bar and short fast bar combined with transient spiral structure make a good fit to both the action-angle distribution and  $R_G-v_\phi$  plane. The phase mixing effects from the transient spiral arms are highly time-dependent, and when selecting a particular time-step with respect to the phase mixing, we are able to reproduce the shape of the horn, and the multiple components of Hercules. However, we are not suggesting that this is the correct sequence of spiral structure, and it is likely that an even better fit to the data could be constructed with some other combination of transient spirals. Currently, the kinematic signatures in the Solar neighbourhood are more in favour of a shorter faster bar, whereas the more direct measurements of the bar length (e.g. Wegg & Gerhard 2013), or pattern speed (e.g. Sormani, Binney & Magorrian 2015; Portail et al. 2017; Clarke et al. 2019; Sanders et al. 2019) are consistently in favour of a longer slower bar. In this work, we have shown that the combination of a long slow bar and transient spiral structure can make a strong Hercules-like feature, with a distinct gap between Hercules and the Hyades/Pleiades moving groups removing one of the key objections to a long slow bar. It is also possible that the other moving groups originate from higher order resonances as shown in Monari et al. (2018), however we found that the transient spiral structure washed out the effects of components above  $m = 4$ .

Once we can trace the moving groups and ridges further across the disc, we may be able to distinguish the origin of individual features. For example, the behaviour around the OLR is distinctive, and if the bar is long and slow we should be able to observe an OLR-like feature further out in the disc, a radius at which we cannot currently sample reliably, but which will be accessible in the near future with

*Gaia* DR3 if not before. We may be seeing the edge of it in the hat, but the signal is weak. Currently there are too many models which can qualitatively reproduce Hercules and the other moving groups in the Solar neighbourhood to conclusively infer properties such as the bar pattern speed from Solar neighbourhood kinematics alone. However, the combination of a bar and transient winding spiral structure can reproduce the Solar neighbourhood moving groups, and the radially extended ridges seen in the *Gaia* data.

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