

# Multi-wavelength study of the Seyfert 1 galaxy NGC 3783 with XMM-Newton

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**Abstract.** We present the analysis of multi-wavelength XMM-Newton data from the Seyfert galaxy NGC 3783, including UV imaging, X-ray and UV lightcurves, the 0.2–10 keV X-ray continuum, the iron K $\alpha$  emission line, and high-resolution spectroscopy and modelling of the soft X-ray warm absorber. The 0.2–10 keV spectral continuum can be well reproduced by a power-law at higher energies; we detect a prominent Fe K $\alpha$  emission line, with both broad and narrow components, and a weaker emission line at 6.9 keV which is probably a combination of Fe K $\beta$  and Fe XXVI.

We interpret the significant deficit of counts in the soft X-ray region as being due to absorption by ionised gas in the line of sight. This is demonstrated by the large number of narrow absorption lines in the RGS spectrum from iron, oxygen, nitrogen, carbon, neon, argon, magnesium, silicon and sulphur. The wide range of iron states present in the spectrum enables us to deduce the ionisation structure of the absorbing medium. We find that our spectrum contains evidence of absorption by at least two phases of gas: a hotter phase containing plasma with a log ionisation parameter  $\xi$  (where  $\xi$  is in  $\text{erg cm s}^{-1}$ ) of 2.4 and greater, and a cooler phase with log  $\xi$  centred around 0.3. The gas in both phases is outflowing at speeds of around  $800 \text{ km s}^{-1}$ . The main spectral signature of the cold phase is the Unresolved Transition Array (UTA) of M-shell iron, which is the deepest yet observed; its depth requires either that the abundance of iron, in the cold phase, is several times that of oxygen, with respect to solar abundances, or that the absorption lines associated with this phase are highly saturated. The cold phase is associated with ionisation states that would also absorb in the UV.

**Key words.** galaxies: active – galaxies: Seyfert – galaxies: individual (NGC 3783) – X-rays: galaxies – ultraviolet: galaxies – techniques: spectroscopic

## 1. Introduction

NGC 3783 is a Seyfert 1 galaxy at redshift 0.00973. Many observers have found a deficit of counts under its power-law X-ray continuum in the soft X-ray range, which they interpret as photoelectric absorption by ionised gas in our line of sight to the active nucleus of the galaxy (e.g. George et al. 1998) - a warm absorber. It has also been claimed, alternatively (Ghosh et al. 1992), that there is an excess of counts over the power law in the soft X-ray range. De Rosa et al. (2002), using data from a  $\sim$  five day observation with BeppoSAX in 1998, detect both a soft excess and warm absorption in their soft X-ray spectrum. They also found a high-energy cut-off of the power-law continuum at  $340^{+560}_{-107} \text{ keV}$ .

With the advent of XMM-Newton and Chandra it has become possible, due to the high resolution of the X-ray spectrometers carried on these missions, to study warm absorbers in unprecedented detail, as was evident after the first high resolution spectrum of a Seyfert galaxy became available (Kaastra et al. 2000). The first high resolution X-ray spectrum of NGC 3783 (Kaspi et al. 2000) was taken in January 2000 using the Chandra HETGS, with an exposure time of 56 ks, and showed a large number of narrow absorption lines from the H-like and He-like ions of O, Ne, Mg, Si, S and Ar, as well as L-shell transitions of Fe XVII–Fe XXI, and a few weak emission lines mainly from O and Ne. The blueshifts of the absorption lines indicated that the warm absorber was outflowing at  $-440 \pm 200 \text{ km s}^{-1}$ . Further analysis of this dataset (Kaspi et al. 2001) also confirmed the presence of Fe XXII and Fe XXIII, and revised the estimate of the absorber’s blueshift to  $-610$

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$\pm 130 \text{ km s}^{-1}$ ; the emission lines were at the systemic velocity of the galaxy. Using regions of the continuum where line absorption is not present, a continuum model was fitted, which required deep O VII and O VIII absorption edges (implying an  $N_{\text{H}}$  of order  $10^{22} \text{ cm}^{-2}$ ). The absorption and emission in the gas were then modelled using photoionisation calculations. The model used involved two phases of gas at different levels of ionisation (with an order of magnitude difference between the ionisation parameters) and with different global covering factors.

NGC 3783 is known to show UV absorption lines intrinsic to the active nucleus (Maran et al. 1996), and there is much interest in attempting to connect the X-ray and UV warm absorbers in this object (e.g. Shields & Hamann 1997). Kaspi et al. (2001) concluded that the lower-ionisation component in their warm absorber model could give rise to the UV absorption observed in this source, but that this was highly sensitive to the unobservable UV-to-X-ray continuum.

Kaspi et al. (2001) also obtained the first high resolution spectrum of the region around the Fe K $\alpha$  emission line, as the HETGS range extends to these energies. They showed that the Fe K $\alpha$  line was narrow, unresolved even at the HETGS resolution, implying that it originated in the torus region of the AGN. Nandra et al. (1997) had previously fitted a relativistically broadened Fe K $\alpha$  line to ASCA spectra of NGC 3783, and the De Rosa et al. (2002) analysis of BeppoSAX data detects the presence of both broad and narrow components to the line.

The most recent Chandra paper on NGC 3783 (Kaspi et al. 2002) presents a 900 ks exposure obtained from several observations between February and June 2001, which produced an outstandingly high signal-to-noise HETGS spectrum, the best yet published from an AGN source. H-like and He-like Ca, H-like C, various more lowly ionised states of Si and S, iron species probably up to Fe XXV and cooler M-shell iron are added to the list of ions detected. The mean blueshift relative to the systemic velocity was found to be  $-590 \pm 150 \text{ km s}^{-1}$ , with the absorption lines of many ions being resolved to have a mean FWHM of  $820 \pm 280 \text{ km s}^{-1}$ . The profiles of the absorption lines show asymmetry which, at least in the case of O VII, originates from the presence of two absorbing systems whose velocity shift and FWHM are consistent with those identified in the UV. Evidence has also been found from this dataset by Behar & Netzer (2002), using the inner-shell absorption lines of silicon, that the warm absorber has a continuous distribution of ionisation parameters, and not just two discrete phases. The narrow Fe K $\alpha$  line is now resolved to have a FWHM of  $1720 \pm 360 \text{ km s}^{-1}$ , and it is accompanied by a Compton shoulder redwards of the line itself, though still by no relativistically (or otherwise) broadened component.

This paper describes the results of work on a 40 ks observation of NGC 3783 with the XMM-Newton Observatory. We obtained UV imaging and photometry from the Optical Monitor (OM; Mason et al. 2001) simultaneously with X-ray data from the EPIC-PN (Strüder et

al. 2001) and RGS (den Herder et al. 2001) instruments. The EPIC cameras, although they only have CCD spectral resolution, have a far larger effective area than HETGS around 6.4 keV where Fe K $\alpha$  emission is observed, requiring far shorter exposures to gain high-statistics spectra at these energies. Again, although the spectral resolution of the RGS is slightly less than that of the HETGS, the combined effective area of the RGS is several times that of the HETGS over much of the energy band where they are both sensitive. Also, the RGS bandpass extends up to about 38 Å (0.326 keV), whereas the HETGS effective area falls off quickly past 23 Å (0.539 keV), so we can make important additions to the large amount of knowledge already gained on NGC 3783 by Kaspi et al. (2000, 2001 and 2002).

Partly motivated by the lower statistical quality and resolution of our current RGS spectrum, which makes it harder to disentangle the many absorption features displayed in the 900 ks Chandra dataset, we have concentrated on self-consistent global fitting of the warm absorber rather than detailed measurements of the individual absorption lines. The RGS spectrum of NGC 3783 stands on the threshold between traditional X-ray astronomy - where spectroscopy relied on the  $\chi^2$  fitting of models to often low spectral resolution and poor statistics data - and traditional optical astronomical spectroscopy, where the instrumental resolution and collecting area are sufficiently good to allow precision measurements of individual spectral features. Analysing an RGS spectrum as complex as this requires far more careful usage of  $\chi^2$  fitting techniques than with spectra of lower resolution; there are so many datapoints, and so many interlinked parameters in the increasingly complex spectral models which are fitted to them, that the practical considerations in the global modelling of such spectra will become increasingly important to X-ray astronomers who work on high-resolution AGN spectroscopy. In this paper, within our analysis of the RGS spectrum of NGC 3783, we attempt to define a practical methodology for the modelling and fitting of such datasets.

This work is a prelude to the analysis of a 280 ks XMM-Newton observation of NGC 3783, which has now been carried out, that will provide an RGS spectrum with equivalent signal-to-noise to the 900 ks HETGS spectrum. This new dataset will extend to higher wavelengths the range which has already been observed by the HETGS, and, since it will be accompanied by high time resolution UV photometry by the OM, will be used to investigate further the multiwavelength variability of this AGN. The long observation will also enable time-resolved spectroscopy of the Fe K $\alpha$  line region.

## 2. Observations and data analysis

NGC 3783 was observed by XMM-Newton, in an RGS Guaranteed Time observation, on 28–29<sup>th</sup> December 2000 for a total of about 40 ks. The instrument modes and total exposure times are given in Table 1.

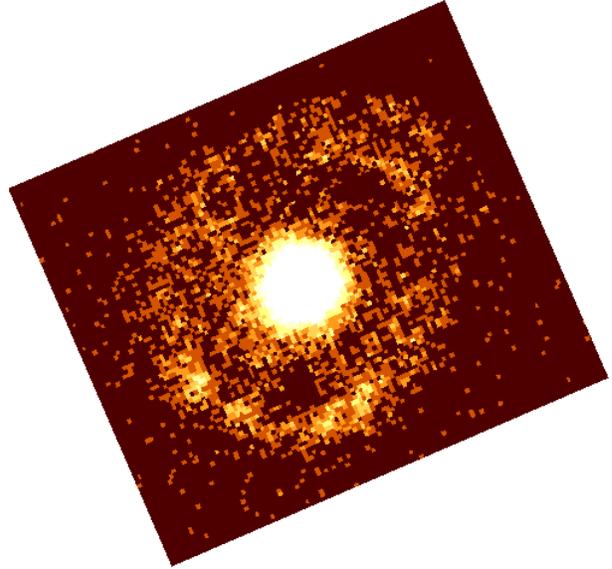
**Table 1.** Instrument modes and exposure times.

Instrument	Mode	Exposure time (s)
EPIC-MOS 1	Large Window (medium filter)	37724
EPIC-MOS 2	Large Window (medium filter)	37724
EPIC-PN	Small Window (medium filter)	37263
RGS 1	Spectroscopy	40456
RGS 2	Spectroscopy	40456
OM	Image (UVW2 filter)	40016

The MOS1 and MOS2 datasets were processed using `emchain` under SAS Version 5.2. Both were found to be piled-up, so the X-ray lightcurve, spectral continuum, and parameters of the Fe  $K\alpha$  line were obtained from the PN data, processed using `eprocc` under SAS Version 5.2. Since the PN was operated in Small Window Mode, the observed count rate ( $\sim 19$  counts  $s^{-1}$  before background subtraction) was well below the threshold (130 counts  $s^{-1}$ , XMM-Newton User’s Handbook (UHB)) where pile-up becomes a problem. The PN spectrum and lightcurve were extracted from within a circular region  $1'$  in radius, centred on the source. The background spectrum and lightcurve were extracted from a region of the same size  $3.1'$  away. The background-subtracted count rate of the Seyfert nucleus was  $18.73 \pm 0.03$  counts  $s^{-1}$  over the PN band of 0.1–12 keV. The spectrum was analysed in XSPEC and, since the data were taken in small window mode, a response matrix optimised for this mode, `epn_sw20_sY9_medium.rmf` (provided by F. Haberl) was used. The lightcurve was analysed using the XRONOS suite of tools.

The RGS1 and RGS2 spectra were extracted using `rgsproc` under SAS V5.2. Background-subtraction is performed with the SAS using regions adjacent to those containing the source in the spatial and spectral domains. The two spectra were then combined; the RGS1 and RGS2 response matrices are corrected for the differences in effective area between the two instruments using the residuals of a power law fit to a pure continuum source, the BL Lac Mrk 421. This also corrects for the instrumental absorption edge due to neutral oxygen. The two response matrices - and spectra - can then be added channel by channel. The resulting spectrum was analysed using SPEX 2.00 (Kaastra et al. 2002a).

The OM was operated in Rudi-5 imaging mode during our observation. This gave a series of 30 UVW2 filter (140–270 nm) images of the galaxy at  $0.5''$  spatial resolution, which were combined using FTOOLS to form the image in Fig. 1. This is the first image to be taken of this galaxy in this waveband and shows much UV emission from the spiral arms, about  $30''$  in diameter, as well as from the nucleus. The FWHM of the central source is about  $2.5''$  which is consistent with the OM Point Spread Function (PSF) with the UVW2 filter.

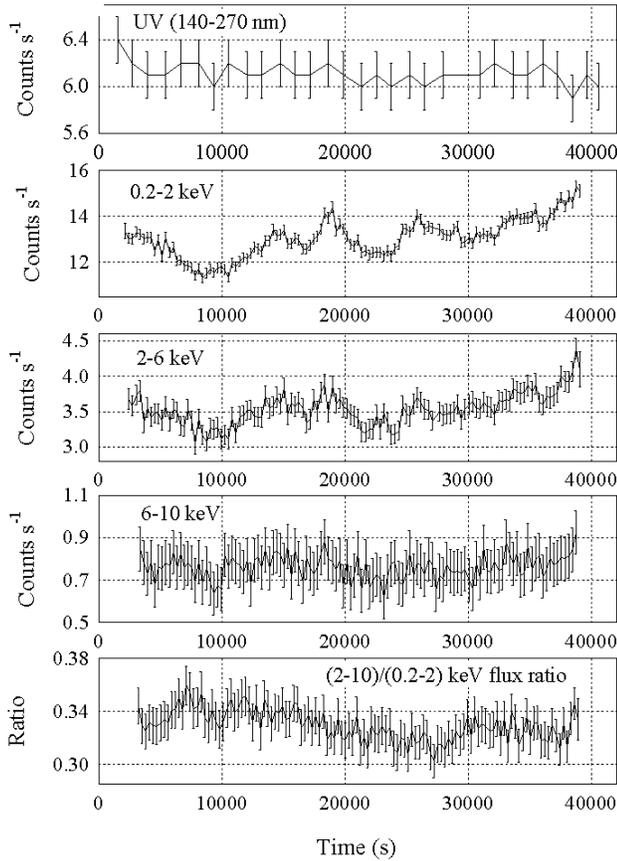
**Fig. 1.** UV image of NGC 3783; North is up, East to the left, and the image is about  $50 \times 40''$  in size. The nuclear image is saturated in order to show the spiral arms.

The long series of images also allowed us to produce a UV lightcurve for the AGN. They were processed using the `omichain` routine under a development version of SAS (`xmmsas_20020107_1901-no-aka`), and photometry of the active nucleus was performed with GAIA using an aperture of  $10''$  diameter. Background subtraction was performed using counts extracted from an annular region around the source, avoiding the spiral arms, and dead-time and coincidence loss corrections were applied to the count rates. The resulting lightcurve was analysed using the XRONOS tools.

### 3. Lightcurves

The background-subtracted lightcurves of NGC 3783 in different energy bands are shown in Fig. 2. The total 0.2–10 keV X-ray flux varied by about 30% between its highest and lowest points during the course of the observation, whilst the UV flux shows very little variation.

The variations in the soft (0.2–2 keV) and hard (2–10 keV) bands follow each other very closely. A plot of the hardness ratio against the total 0.2–10 keV intensity (Fig. 3), for 30 minute time bins, may indicate that the

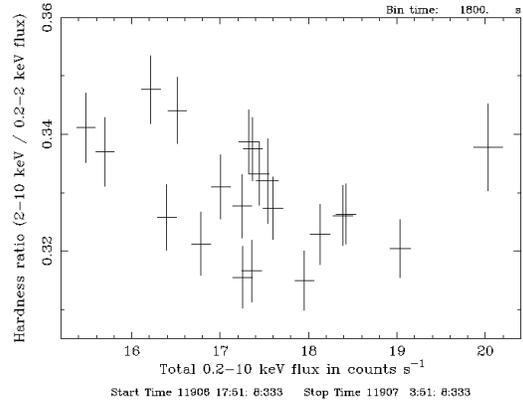


**Fig. 2.** Multiwavelength background-subtracted lightcurves: top, UV (140–270 nm) lightcurve from the OM, in 1000 s time bins; middle three, PN lightcurves in the bands 0.2–2, 2–6 and 6–10 keV, in 300 s time bins; bottom, the hardness ratio (2–10 keV / 0.2–2 keV flux) in 300 s time bins.

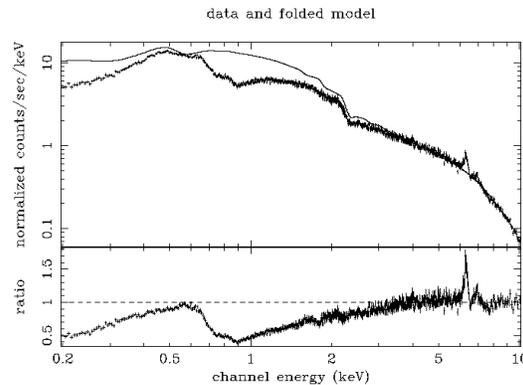
PN spectrum gets softer as the source intensity increases. The linear correlation coefficient for the data in this plot is 0.35, with a probability of about 3% of obtaining this value or greater from a random sample, implying a weak correlation significant at 2 sigma. The final point, at the highest intensity, does not fit in with this trend and corresponds to a sudden hardening of the spectrum at the end of our observation. If this point is disregarded, the softness-intensity correlation is significant at 3 sigma.

#### 4. Continuum and Fe K $\alpha$ line

To derive the X-ray spectral continuum of NGC 3783, the PN spectrum was fitted in XSPEC with a power-law model and neutral absorption (using the tbabs module), with the spectrum binned to 300 counts per bin. It was found that a simple power-law could not fit the data over the whole energy range, due to the significant deficit of counts below about 3 keV. However, it was possible to produce an acceptable power-law fit at higher energies, and in the ranges 3.5–5.5 and 7.5–10 keV (chosen to avoid



**Fig. 3.** Hardness – intensity plot in 30 minute time bins.



**Fig. 4.** Continuum power-law fit to the 3.5–5.5 and 7.5–10 keV ranges of the PN spectrum, extrapolated over 0.2–10 keV. The ratio of the data to the continuum model is shown in the bottom panel.

the prominent Fe K emission around 6.4 keV), the best-fit parameters given in Table 2 were obtained. The neutral absorption was held at the Galactic value of  $8.7 \times 10^{20} \text{ cm}^{-2}$  (as in Kaspi et al. 2001). This fit is shown, extrapolated over 0.2–10 keV, in Fig. 4, where the form and extent of the warm absorption is clearly visible (although there are still some calibration uncertainties in the PN small window mode, especially below about 0.7 keV). No soft excess is required.

With this model, the intrinsic 2–10 keV source flux is  $8.5 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ .

There is a prominent Fe K $\alpha$  emission line visible in our PN spectrum (see Fig. 4). The form of this line implies that it has both broad (line 1) and narrow (line 2) components, and we have fitted each of these as a gaussian, using the power-law of Table 2 as the underlying continuum. There is a second, much fainter emission line at a slightly higher energy (line 3), also parameterised as a gaussian, which could be identified as Fe K $\beta$  or Fe XXVI. The parameters of the resulting fit are given in Table 3;

**Table 3.** The parameters of three gaussians fitted to the emission lines seen in the PN spectrum.

Line	Identification	$E_{\text{rest}}^{\text{a}}$	$E_{\text{meas}}^{\text{b}}$	Flux <sup>c</sup>	FWHM <sup>d</sup>	$V_{\text{broad}}^{\text{e}}$	$cz^{\text{f}}$
1	broad Fe K $\alpha$	6.39	$6.29 \pm 0.03$	$7 \pm 1$	$240 \pm 80$	$11000 \pm 4000$	$5000 \pm 1000$
2	narrow Fe K $\alpha$	6.39	$6.40 \pm 0.01$	$7.1 \pm 0.9$	$70 \pm 50$	$3000 \pm 3000$	$-600 \pm 600$
3	Fe K $\beta$	7.06	$6.98 \pm 0.04$	$3.7 \pm 0.9$	$200 \pm 100$	$9000 \pm 6000$	$3000 \pm 2000$
	Fe XXVI	6.96					$-1000 \pm 2000^{\text{g}}$

<sup>a</sup> expected rest frame energy in keV<sup>b</sup> measured rest frame energy in keV<sup>c</sup> in  $10^{-5}$  photons  $\text{cm}^{-2} \text{s}^{-1}$ <sup>d</sup> Full Width Half Maximum in eV<sup>e</sup> velocity broadening in  $\text{km s}^{-1}$ <sup>f</sup> redshifted velocity in  $\text{km s}^{-1}$  (negative values indicate blueshifts)<sup>g</sup> if line 3 is identified with Fe XXVI**Table 2.** The best fit power-law continuum parameters, fitted to the ranges 3.5–5.5 and 7.5–10 keV.

Photon index ( $\Gamma$ )	Normalisation <sup>a</sup>	$N_{\text{HGal}}^{\text{b}}$	$\chi_{\text{reduced}}^2$
$1.60 \pm 0.02$	$0.0180 \pm 0.0007$	8.7	$1.4^{\text{c}}$

<sup>a</sup> in photons  $\text{keV}^{-1} \text{cm}^{-2} \text{s}^{-1}$  at 1 keV<sup>b</sup> in  $10^{20} \text{cm}^{-2}$ , fixed<sup>c</sup> for 168 degrees of freedom

over the range of 5.5–7.5 keV, this fit has a  $\chi_{\text{red}}^2$  of 1.1 for 70 degrees of freedom, and is shown in Fig. 5.

The FWHM of the narrow component of our Fe K $\alpha$  line,  $3000 \pm 3000 \text{ km s}^{-1}$ , is consistent with that of Kaspi et al. (2002) ( $1860 \pm 340 \text{ km s}^{-1}$ ), although it is poorly constrained. The flux in our line is slightly greater,  $(7.1 \pm 0.9) \times 10^{-5}$  photons  $\text{cm}^{-2} \text{s}^{-1}$  as opposed to  $(5.26 \pm 0.63) \times 10^{-5}$  photons  $\text{cm}^{-2} \text{s}^{-1}$ . The narrow Fe K $\alpha$  line of Kaspi et al. (2002) is not blueshifted; in our case, the blueshift is consistent with zero ( $-600 \pm 600 \text{ km s}^{-1}$ ), although again this is not well constrained. These authors detect a red wing to their Fe K $\alpha$  line which they identify as a Compton reflection hump, whereas our Fe K $\alpha$  line has a very strong broad component containing as much flux as the narrow component itself. De Rosa et al. (2002) also detect broad and narrow components in the Fe K $\alpha$  line. Their broad component, with  $\sigma = 0.72_{-0.27}^{+1.28}$  keV, is significantly broader than ours ( $\sigma = 0.10 \pm 0.03$  keV). It is also clear, from Fig. 5, that there is an excess of counts over the continuum redwards of Fe K $\alpha$  in our spectrum, and this could correspond to the Compton shoulder observed by Kaspi et al. (2002).

If the broad and narrow components of the Fe K $\alpha$  line are each associated with an Fe K $\beta$  line at the expected  $\sim 14\%$  of the flux of K $\alpha$ , then the combined flux in K $\beta$  is not sufficient to explain the measured flux in line 3. About

half the flux from this line probably originates, then, from Fe XXVI.

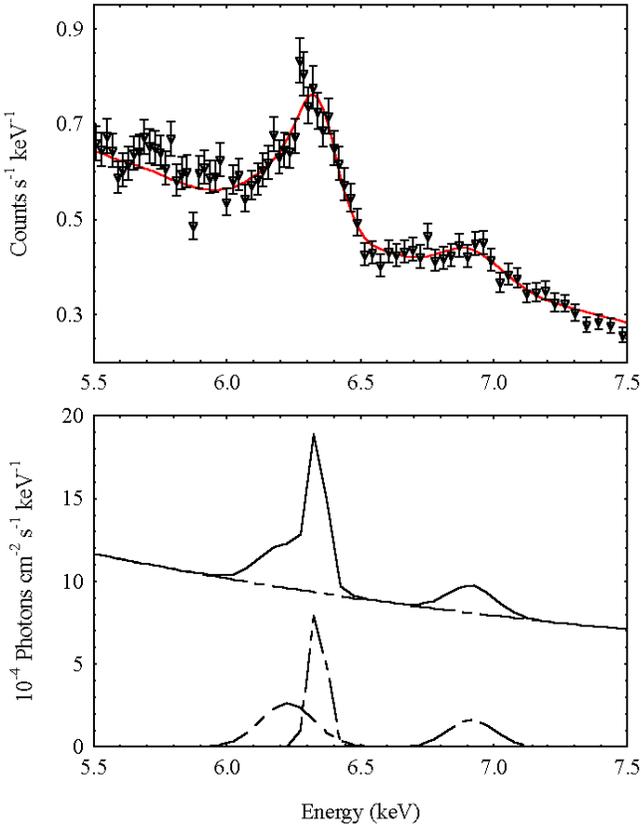
Noting the significant redshift of the broad component of Fe K $\alpha$ , it is possible to fit the entire line (or just the broad component) with relativistically broadened emission from an accretion disc. The broad component does not contrast sufficiently against the continuum for the shape of such a line to be determined unambiguously, although it can be fitted with the (almost gaussian) profile of a line from a face-on disc with an emissivity index below about 2.

## 5. RGS spectrum

### 5.1. Continuum

The continuum underlying the RGS spectrum was modelled with the power-law fitted to the PN data. The slope and normalisation of this power-law were corrected for the systematic calibration differences between the PN and RGS as described in den Herder et al. (2002). The photon index  $\Gamma$  became 1.54 and the new normalisation is 0.0175 photons  $\text{keV}^{-1} \text{cm}^{-2} \text{s}^{-1}$  at 1 keV, with the Galactic absorption again held constant at  $8.7 \times 10^{20} \text{cm}^{-2}$ .

With the chosen continuum, there was a significant deficit of flux below 17 Å (above 0.729 keV; see Fig. 6). There is no excess over the power law, and the overall form of the spectrum does not contain the extreme ‘saw-tooth’ features seen in Mrk 766 and MCG-6-30-15 (Branduardi-Raymont et al. 2001); however, there is a deep drop in the continuum at about 17 Å which has in the past been explained as an O VII absorption edge (e.g. George et al. 1998). There is also a small triangular feature centred at 15 Å (0.827 keV). Our task is then to work out how much of this spectral structure is due to emission, how much to absorption, and what physical model might explain the spectrum that we observe.



**Fig. 5.** Fit to the iron emission lines in the PN spectrum (left, lines 1 and 2 (Fe  $K\alpha$ ); right, line 3). Top panel: data (black) and model (red). Bottom panel: overall model (solid line) and model components (dotted line).

## 5.2. Modelling

We aim, in modelling this spectrum, to take into account both line and continuum absorption self-consistently. The narrow absorption lines from the photoionised absorber are our starting point; line absorption is a far more reliable indication of the physical state of the absorber than continuum absorption, since the contribution of the many different ions to the continuum absorption is very hard to determine (unless the columns are high enough to produce edges), whereas the lines, if they can be identified, are unambiguous.

This fact is at the heart of the difficulty with using standard  $\chi^2$  fitting techniques to model such spectra. Whilst the absorption lines are the best diagnostic of the warm absorption, they are represented by very few datapoints. Continuum absorption, however, is spread out over every datapoint in the spectrum, and this will bias a  $\chi^2$  minimisation process towards explaining the highly ambiguous continuum absorption rather than the lines themselves, thus providing misleading - and possibly even meaningless - results. This is why  $\chi^2$  fitting has to be used with great care with spectra such as these, and its

use must always accompany a thorough understanding of the physical properties of the model that is being fitted.

We used SPEX (Kaastra et al. 2002a) for the modelling, which provides two possibilities for a photoionised absorber. The *slab* model applies line and continuum absorption by individual ions to a spectral continuum, whilst *xabs* is a grid model of Xstar runs applying absorption by a column of photoionised gas at a range of different ionisation parameters, column densities, and elemental abundances. We define the ionisation parameter  $\xi$  here as

$$\xi = \frac{L}{nr^2}, \quad (1)$$

where  $L$  is the source luminosity (in  $\text{erg s}^{-1}$ ),  $n$  the gas density (in  $\text{cm}^{-3}$ ) and  $r$  the source distance in cm, so  $\xi$  has the units  $\text{erg cm s}^{-1}$  (Tarter et al. 1969) which are used throughout this paper. *Slab* and *xabs* incorporate new calculations, using HULLAC (Bar-Shalom et al. 2001), for absorption due to states of M-shell iron and lower ionisation states of oxygen (O II–O VI). *Slab* is the best tool for detailed measurements of individual ion columns in a spectrum, whereas *xabs* enables the physical parameters and composition of a warm absorber to be investigated directly. A more detailed description of these models is given in Kaastra et al. (2002a).

Our modelling strategy involves the complementary use of these two models. Firstly, *slab* was used to identify the narrow absorption lines, by matching their depths and positions with those corresponding to each ionisation state of each element that the model contains. *Slab* applies both narrow line and continuum absorption self-consistently to the spectrum, so it was possible to see, by looking at each separate ion in turn, the precise effect that the different ions had on the form of the spectral continuum.

For purposes of comparison with the published spectrum of Kaspi et al. (2002), and with future analyses of RGS spectra, Table 4 lists measured and rest wavelengths, blueshifts and equivalent widths for the deepest absorption lines of the most important ions of O, N, C and Fe, as well as these parameters plus the flux of the forbidden emission line of O VII. It also quotes values of flux and equivalent width measured by Kaspi et al. (2002) where applicable. Note that the deepest line of a particular ion is not always the one theoretically predicted; this is almost certainly because many of the deepest absorption lines are accompanied by re-emission (see Fig. 8 and Fig. 9 for an expanded view of the RGS spectrum and the warm absorber) as demonstrated first by Kaspi et al. (2002). Also, in the case of O VI, the theoretically predicted deepest line lies in the region of the O VII emission triplet, making it difficult to distinguish, so our table lists measurements for the second deepest line of O VI. C VI Ly- $\beta$  is listed as it occurs in a range where Galactic absorption has less of an effect than at the corresponding Ly- $\alpha$  transition, making it easier to obtain an accurate measurement.

We observe an O VII forbidden line in emission, but the signal-to-noise is not good enough to distinguish the accompanying intercombination and resonance lines. The

**Table 4.** The blueshifts and equivalent widths of some important spectral lines in the RGS data.

Line	$\lambda_{\text{meas}}^{\text{a}}$	$\lambda_{\text{rest}}^{\text{b}}$	$v_{\text{blue}}^{\text{c}}$	Flux <sup>d</sup>	Flux <sub>chan</sub> <sup>e</sup>	EW <sup>f</sup>	EW <sub>chan</sub> <sup>g</sup>
O IV K $\beta$	$22.67 \pm 0.02$	$22.777^{\text{h}}$	$-1400 \pm 300$			$80 \pm 10$	–
O V K $\alpha$	$22.27 \pm 0.02$	22.334	$-900 \pm 300$			$70 \pm 20$	–
O VI K $\alpha^{\text{i}}$	$19.293 \pm 0.006$	19.341	$-760 \pm 90$			$76 \pm 9$	–
O VII K $\alpha^{\text{j}}$	$21.5 \pm 0.1$	21.602	$-900 \pm 1000$			$20 \pm 20$	$40.1 \pm 34.6$
O VIII Ly- $\alpha^{\text{k}}$	$18.92 \pm 0.01$	18.969	$-900 \pm 200$			$50 \pm 10$	$53.6 \pm 15.9$
N VII Ly- $\alpha$	$24.67 \pm 0.06$	24.781	$-1300 \pm 700$			$90 \pm 10$	$72.6 \pm 34.9$
C VI Ly- $\beta^{\text{l}}$	$28.39 \pm 0.01$	28.466	$-800 \pm 200$			$100 \pm 10$	–
O VII (f)	$22.06 \pm 0.02$	22.101	$-600 \pm 200$	$7 \pm 2$	$10.13 \pm 2.74$	$-130 \pm 40$	$-270.4 \pm 73.2$
Fe XVII	$14.97 \pm 0.02$	15.014	$-800 \pm 400$			$92 \pm 6$	$26.2 \pm 5.3$
Fe XVIII	$14.17 \pm 0.01$	$14.210^{\text{m}}$	$-800 \pm 200$			$148 \pm 8$	$24.8 \pm 5.5$
Fe XIX	$13.50 \pm 0.01$	$13.523^{\text{n}}$	$-500 \pm 200$			$164 \pm 9$	$55.3 \pm 8.1$
Fe XX	$12.815 \pm 0.008$	$12.842^{\text{o}}$	$-600 \pm 200$			$167 \pm 8$	$56.7 \pm 5.6$

<sup>a</sup> measured wavelength in Å (in rest frame of NGC 3783)

<sup>b</sup> rest frame wavelength in Å

<sup>c</sup> blueshifted velocity in km s<sup>-1</sup>

<sup>d</sup> line flux in 10<sup>-5</sup> photons cm<sup>-2</sup> s<sup>-1</sup>

<sup>e</sup> line flux in Chandra spectrum as quoted in Kaspi et al. (2002) in 10<sup>-5</sup> photons cm<sup>-2</sup> s<sup>-1</sup>

<sup>f</sup> measured equivalent width in mÅ

<sup>g</sup> equivalent width in Chandra spectrum as quoted in Kaspi et al. (2002) in mÅ

<sup>h</sup> this value is yet to be benchmarked in the laboratory, and could be off by as much as 40 mÅ, so the blueshift of the line could probably be about 500 km s<sup>-1</sup> lower

<sup>i</sup> blended with a weaker absorption line of O V

<sup>j</sup> partially filled in by re-emission

<sup>k</sup> perhaps partially filled in by re-emission

<sup>l</sup> blended with a weaker absorption line of Si XI

<sup>m</sup> a blend of lines at 14.1548, 14.2042, 14.2120 and 14.2600 Å

<sup>n</sup> a blend of lines at 13.5070, 13.5146, 13.5210, 13.5540 and 13.5706 Å

<sup>o</sup> a blend of lines at 12.8130, 12.8270, 12.8319, 12.8470 and 12.9040 Å

fact that the forbidden line is more prominent than the other two indicates that they are formed in tenuous and photoionised gas (Porquet & Dubau 2000). The blueshift of the forbidden line,  $-600 \pm 200$  km s<sup>-1</sup>, is consistent with that of the warm absorber, which could support the idea that these lines originate from it. Kaspi et al. (2002), on the other hand, find that this line is not blueshifted in their spectrum.

From investigating the narrow absorption lines with *slab*, it became clear that we are dealing with two main ionisation ranges in the warm absorber. Kaspi et al. (2001) also found multiple phases in the warm absorber, although the two phases that they model both correspond to our high-ionisation range (see Table 7). Our high-ionisation phase gives rise to L-shell iron transitions, O VIII, O VII, N VII and C VI absorption, whilst the low-ionisation gas is associated with M-shell iron forming an Unresolved Transition Array (UTA) at about 16.5 Å (0.751 keV), as well as O IV, O V and O VI lines. Both phases are blueshifted by about  $-800 \pm 200$  km s<sup>-1</sup>, a figure obtained by fitting the blueshift of each phase as a whole. This is consistent with the value of Kaspi et al. (2002) ( $-590 \pm 150$  km s<sup>-1</sup>).

The many ions of iron apparent in the spectrum provide a tight constraint, independent of abundance, on the ionisation structure of the absorbing gas. Iron is used for this purpose as, in plasma at a given ionisation parameter, it will be present in several different ionisation states. Measuring the relative importances of these states in the spectrum gives a good measure of the ionisation parameter of the absorber. In general, as the plasma becomes less ionised, the main locus of iron absorption in the spectrum shifts to longer wavelengths.

From our identification of which iron states are present using *slab*, there seem to be two locations where iron absorption is particularly important in this spectrum. The first of these is deep absorption from L-shell iron at around 13–15 Å (0.954–0.827 keV), where we see Fe XVII, Fe XVIII, Fe XIX and Fe XX. At lower wavelengths, it becomes increasingly difficult to detect discrete absorption features from the more highly ionised states of iron.

If we consider the form of the spectrum around about 15 Å, where we might expect absorption from Fe XVI and Fe XV, we can see that there must be a far lower column of these ions as they would absorb away the low-wavelength side of the 15 Å feature otherwise. We claim then that this

represents the long-wavelength boundary of iron absorption from the high-ionisation phase of the warm absorber.

The other important locus of iron absorption is the Unresolved Transition Array of M-shell iron at around 16–17 Å (0.775–0.729 keV). This is the main spectral signature of the cold phase of the warm absorber. Once the blueshift of this phase has been established using the oxygen absorption lines which also originate from it, the position of the UTA in the spectrum can be used to determine the ionisation state of the gas. The high wavelength side of the 15 Å feature then forms a boundary to the iron absorption from the low-ionisation phase.

We claim, then, that the 15 Å feature provides us with a very clear cut-off between the iron absorption from two separate phases of gas. The height and shape of this feature places strict limits on exactly which iron states can be present in the absorbing medium.

Having established the existence of the two main ionisation regimes, we constructed a simple two phase warm absorber model using *xabs* in order to gain some insight into the overall physical structure and composition of the absorbing gas. We also constructed a detailed model using *slab* in order to obtain more precise measurements of the columns of individual ions within the absorber. These two models are described in the following sections.

### 5.3. Two-phase warm absorber model

The blueshift of the two phases was set to 800 km s<sup>-1</sup>. The narrow absorption lines are at the width of the RGS spectral resolution, so we cannot directly derive a turbulent velocity from them. For the purpose of these fits, we use a velocity width of 300 km s<sup>-1</sup> (where the velocity width is the rms width, i.e. gaussian sigma, for each individual velocity component of each line), similar to that used in the analysis of the Chandra HETGS spectrum of this source (Kaspi et al. 2001).

The first stage was to determine the ionisation parameters of the two phases using *xabs*. In the high-ionisation phase, if we take the most important iron states to be Fe XVII, Fe XVIII and Fe XIX, we can estimate an overall log  $\xi$  to be 2.4. In reality, of course, this is a simplification as the presence of the many higher iron states implies a range of ionisation parameters, from 2.4 upwards, in the high-ionisation phase. The lower-ionisation gas, on the other hand, is fairly well explained by gas at a single ionisation parameter. This is obtained from the position and form of the UTA, which is accurately reproduced with a log  $\xi$  of 0.3.

Once we had the ionisation parameters of the two phases, we were able to fit overall equivalent hydrogen columns, assuming Anders & Grevesse (1989) elemental abundances. The abundances of individual elements were then allowed to vary. *Xabs* includes the elements He, C, N, O, Ne, Na, Mg, Si, S, Ar and Fe; if it was not possible to establish the presence or abundance of a particular

**Table 5.** Properties of the two-phase warm absorber model

Parameter	High-ionisation phase	Low-ionisation phase
log $\xi^a$	2.4	0.3
$N_H^b$	$2.8^{+0.01}_{-0.3} \times 10^{22}$	$5.4^{+0.05}_{-0.5} \times 10^{20}$
$v_{\text{turb}}^c$	300	300
$v_{\text{blue}}^d$	800	800
C abundance <sup>e</sup>	0.8	1.0
N abundance	0.2	1.0
O abundance	0.7	1.0
Ne abundance	0.9	
Mg abundance	0.9	
Si abundance	2.5	
S abundance	1.0	
Ar abundance	0.7	
Fe abundance	0.9	10.7

<sup>a</sup> Log of the ionisation parameter, where  $\xi$  is in erg cm s<sup>-1</sup>

<sup>b</sup> Equivalent hydrogen column of phase in cm<sup>-2</sup>

<sup>c</sup> Turbulent velocity in km s<sup>-1</sup>

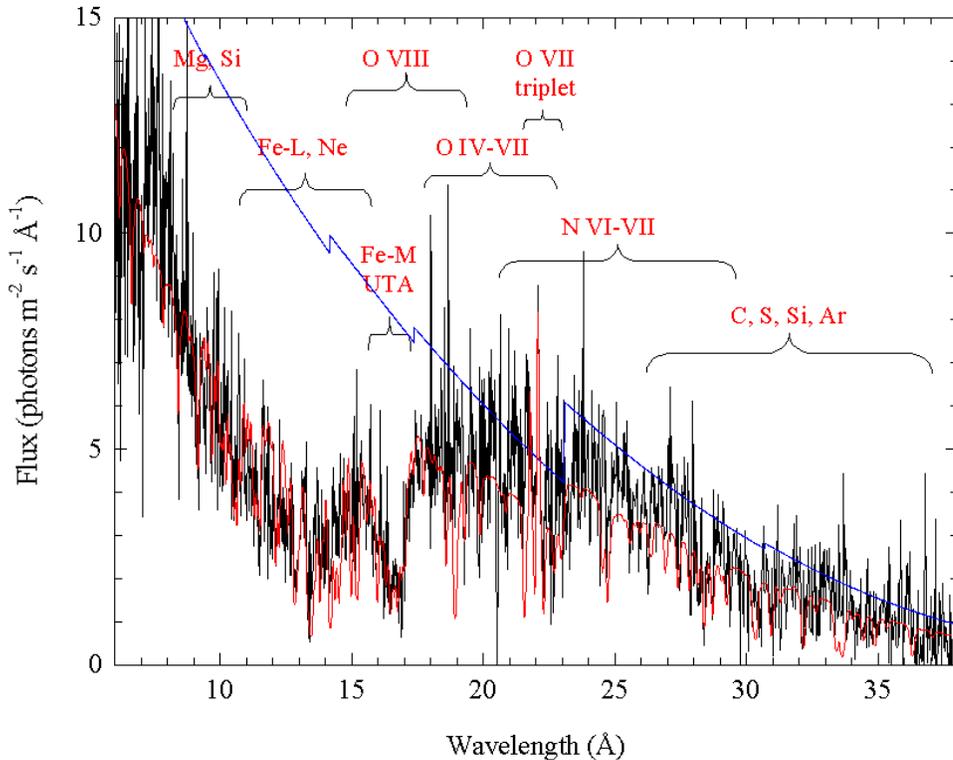
<sup>d</sup> Blueshift of phase in km s<sup>-1</sup>

<sup>e</sup> Abundances relative to solar; estimated errors on abundances are  $\pm$  about 0.5. Values are only given for elements whose abundances could be directly measured; the abundance of all other elements in the model was fixed at 1

element, it was kept in the model but fixed at its solar abundance.

This simple two-phase model, superimposed on the data, is shown in Fig. 6, and Fig. 7 shows the two phases of the absorber applied separately to the continuum. Table 5 gives the values of the parameters derived for the two phases. It can be seen that the overall model reproduces the global form of the spectrum very well although there are discrepancies. The assumption of a single ionisation parameter for the entire high-ionisation phase of the gas leads to an overestimation of the iron absorption at the low-wavelength side of the 15 Å feature and an underestimation of the iron absorption at wavelengths below this.

The ratios of the elemental abundances to each other in the high-ionisation phase are probably not too different from solar. Nitrogen appears to be underabundant, which is the opposite of what was found in IRAS 13349+2438 (Sako et al. 2001) and NGC 1068 (Kinkhabwala et al. 2002). The solar oxygen abundance used in SPEX (Anders and Grevesse 1989) has been shown by a new measurement (Allende Prieto et al. 2001) to be too high, and thus our oxygen abundances are probably underestimated by about 0.2. In the low-ionisation phase, iron needs to be about ten times more abundant relative to C, N and O in order to reproduce the form of the very deep UTA.



**Fig. 6.** Combined RGS1 and RGS2 rest frame spectrum of NGC 3783 (black), with the power-law continuum (blue) and the power-law plus two-phase warm absorber model (red) superimposed, plotted over the range 6–38 Å (0.326–2.066 keV). Species absorbing or emitting in various spectral ranges are indicated.

#### 5.4. Detailed modelling

The results of detailed modelling of the spectrum with *slab* are shown in Fig. 8 and Fig. 9. The individual ion columns and their estimated errors are listed in Table 6, alongside the log ionisation parameter at which the ion is most abundant. The L-shell iron absorption reproduces the form of the spectrum below 15 Å very well in some regions, but there are still some significant residuals in others. Fig. 10 shows the ionic column plotted against ionisation parameter of maximum abundance for the iron states (except Fe XV, for which only an upper limit could be obtained) listed in Table 6. It clearly shows the existence of the two main ionisation ranges, and demonstrates that the low-ionisation gas is more tightly focussed at a certain level of ionisation than the high-ionisation phase, which extends over a range of ionisation parameters (note the logarithmic scale) - as already inferred from the discrepancies of the two-phase model.

## 6. Discussion and conclusions

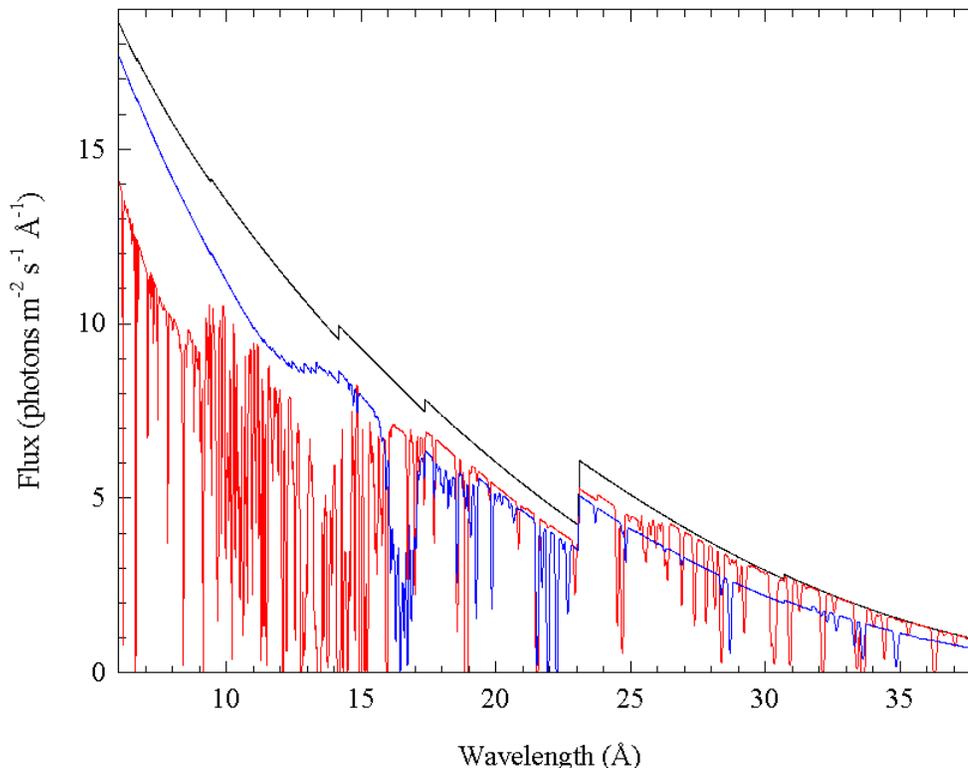
### 6.1. Ultraviolet image

The XMM OM provides us with the first ever UV image of the nuclear regions of NGC 3783. Although the active nucleus is by far the brightest source of ultraviolet light

in the galaxy, our image also shows UV emission from the spiral arms, which presumably traces the location of star formation activity. Although the spiral arms are visible in UV, the bar of the galaxy (oriented approximately North-to-South in our image) is not strongly present, indicating that star formation is much less important in the bar than it is in the arms, or alternatively that dust is obscuring the UV emission from the bar.

### 6.2. Lightcurves

It is clear from Fig. 2 that there is unlikely to be sufficient spectral variability in this source to have a serious effect on our modelling of the continuum underlying the RGS spectrum. However, there is clearly variability in the flux and spectral shape on many different timescales. Intensity variations on the scale of about one to three hours, which we see most prominently in the 0.2–2 keV band, seem to modulate both soft and hard bands in the same way. Also, the changes that we observe in the hardness ratio demonstrate variability in the relative intensities of the two bands over the course of a few hours. Fig. 3 indicates that the spectrum generally softens as the source intensity increases, although at the highest intensity point, that is at the very end of our observation, the source is actually in a harder state. The UV lightcurve is almost completely



**Fig. 7.** The two phases of the warm absorber model applied separately to the continuum (high-ionisation phase in red, low-ionisation phase in blue, and power-law continuum in black) plotted over the range 6–38 Å (0.326–2.066 keV; plotted in the rest frame).

flat, within the errors, implying less than 10% variability in the UV flux from the nucleus during our observation. There is therefore no evidence in our dataset for a direct correlation between UV and X-ray variability over the course of a few hours.

### 6.3. Continuum and iron $K\alpha$ line

We find that the 0.2–10 keV spectral continuum of NGC 3783 is dominated by a power-law at high energies, whilst there is significant absorption by ionised gas - the warm absorber - at lower energies. There is no evidence for a soft excess, although of course the presence of a soft excess could be masked by the warm absorber. There is also no requirement for a reflection component in our spectrum.

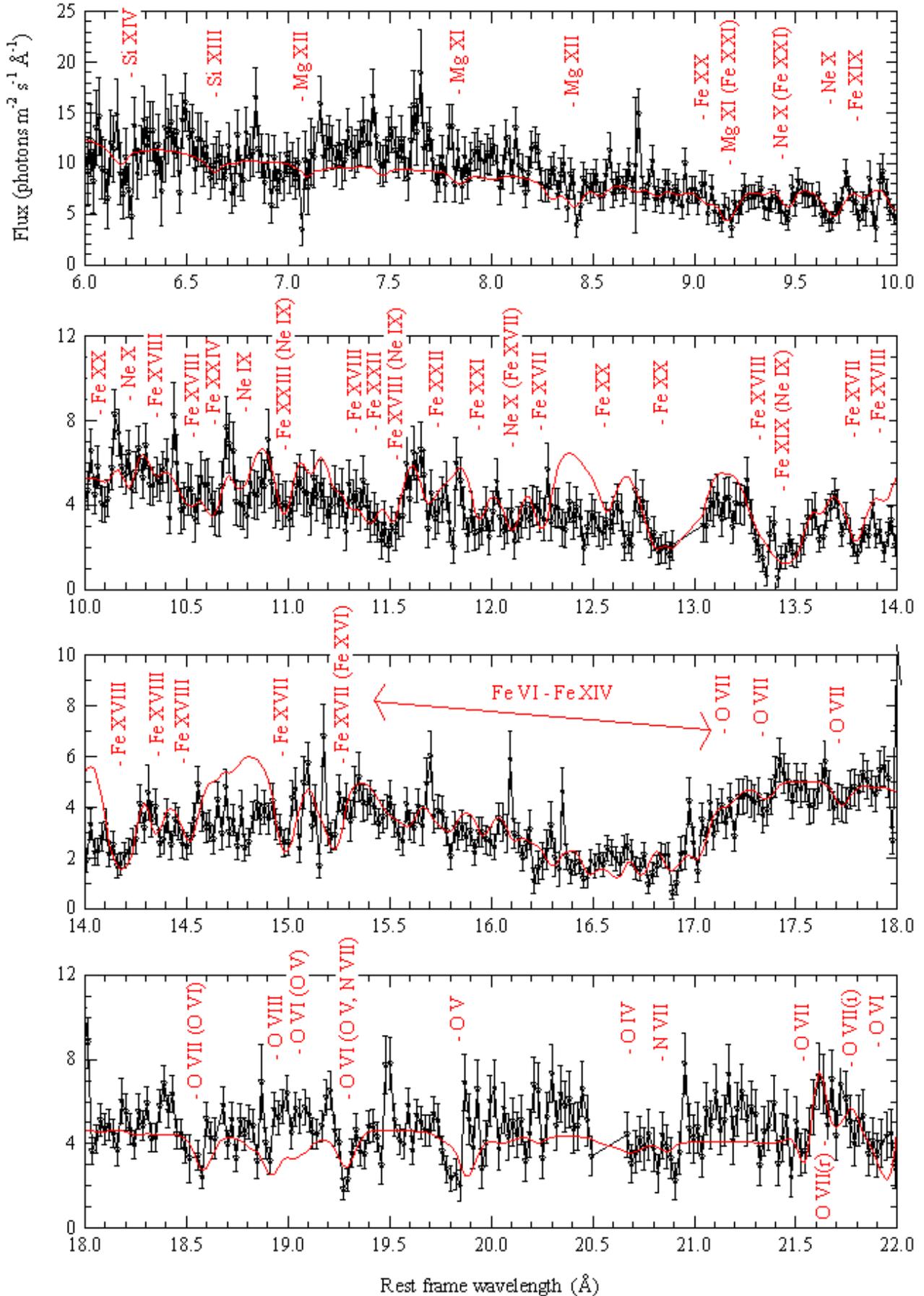
The iron  $K\alpha$  emission line has two components; one narrow component at 6.40 keV, and a broader component which is redshifted by about 5000 km s<sup>-1</sup>. The profile of this composite line looks superficially similar to that of the NGC 3783 spectrum (1) of Nandra et al. (1997), although our broad component is far narrower than the one they observed. We were able to fit our broad component with a gaussian; it does not contain enough flux to obtain good constraints on a relativistic disc line model. Our narrow Fe  $K\alpha$  line is apparently similar to that observed by Kaspi

et al. (2002), although our derived parameters are not well constrained.

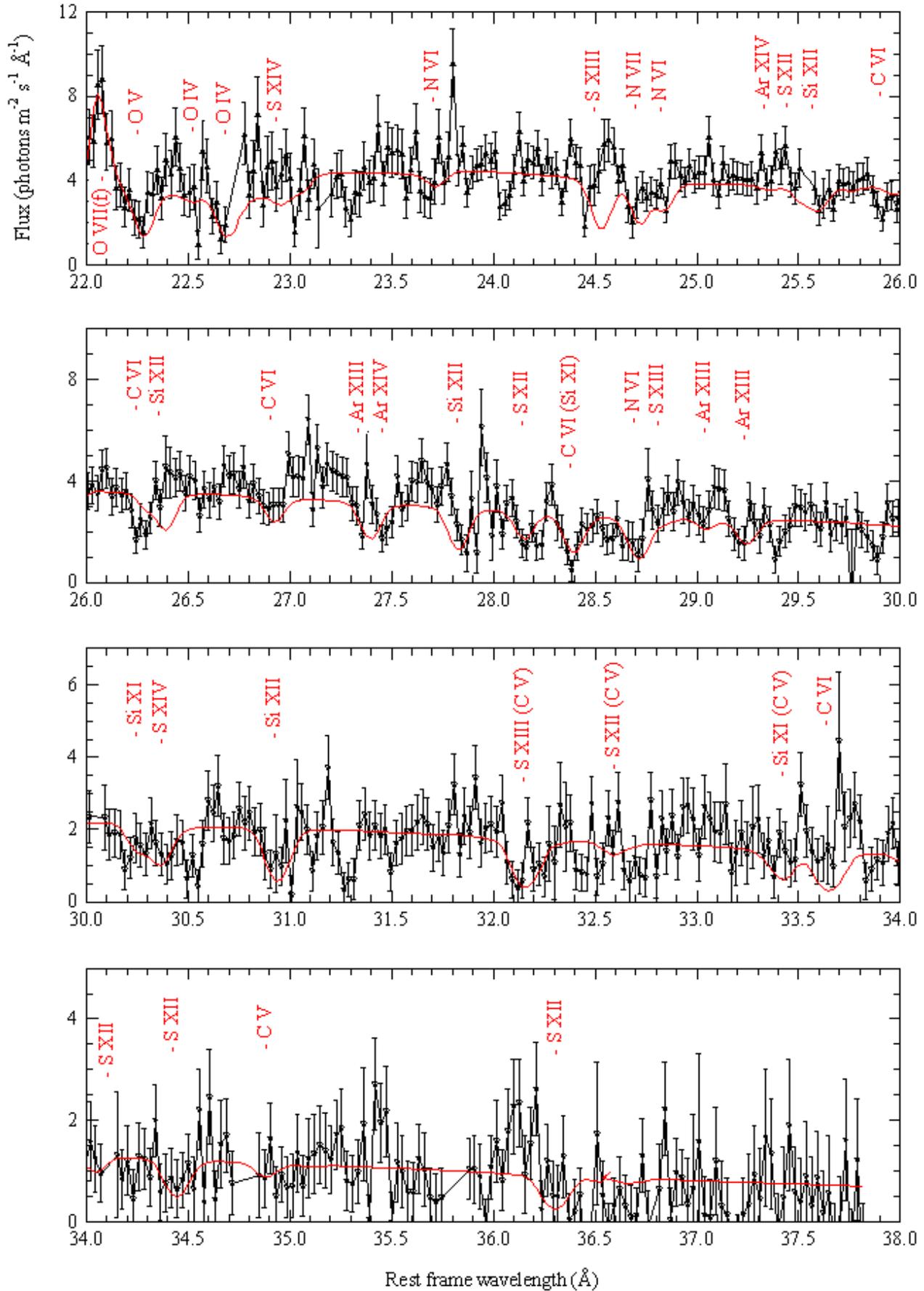
The broad component of the Fe  $K\alpha$  line implies emission from gas that is infalling close to the nucleus. The line broadening would presumably be dynamic rather than thermal, as Fe  $K\alpha$  arises from fluorescence of neutral iron, so the emitting matter is in a region of high speed motion under strong gravity. Certainly, the derived velocity broadening of this component, at  $11000 \pm 4000$  km s<sup>-1</sup>, would place it in the very broad line region of Winge et al. (1992). We were unable to ascertain whether the broadening is actually relativistic, although if we were dealing with emission from close enough to the black hole to be relativistically broadened, we might expect the iron to be rather more highly ionised than the energy of this line implies.

That the broad component of the Fe  $K\alpha$  line was not observed by Kaspi et al. (2002) in the 900 ks Chandra observation, where the signal-to-noise in this spectral region was very good, may imply that the spectrum has changed between the two observations. Infalling gas close to the central engine of the AGN, but not yet close enough to become too hot to emit Fe  $K\alpha$ , might be a transitory phenomenon which is not always observable.

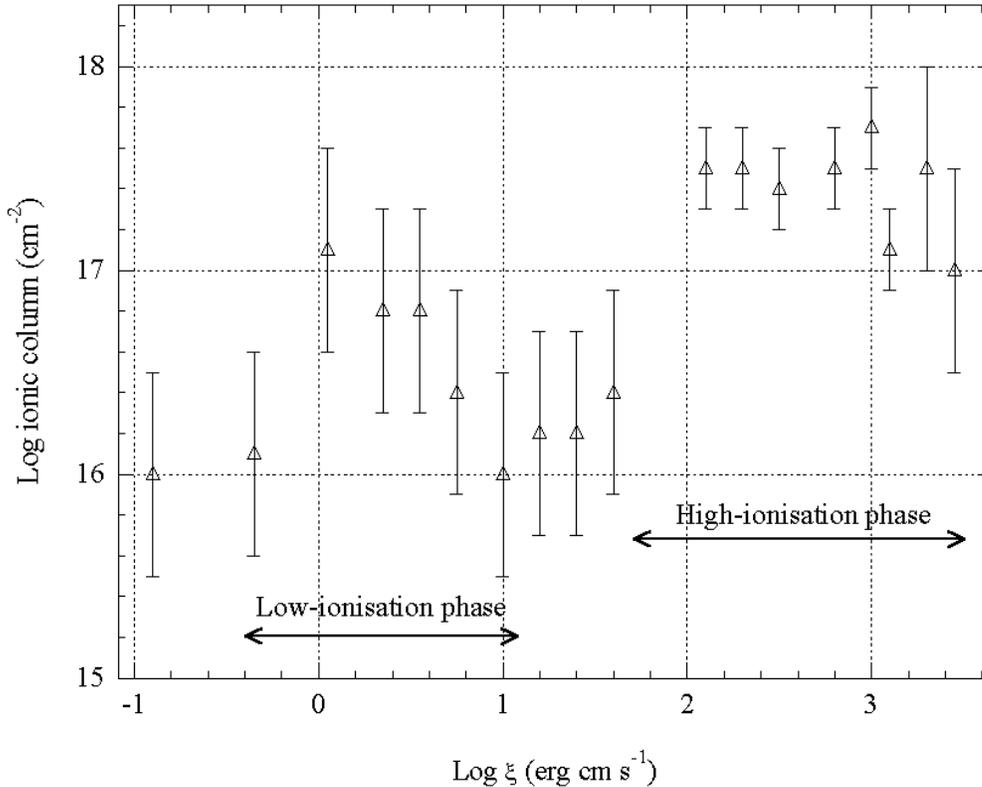
The Fe  $K\alpha$  emission is accompanied by another line at higher energy, which is probably a combination of Fe



**Fig. 8.** Combined RGS1 and RGS2 rest frame spectrum of NGC 3783 (black), binned into groups of two channels, with detailed warm absorber model superimposed (red), plotted over the range 6–22  $\text{\AA}$  (0.326–0.564 keV). The most



**Fig. 9.** Combined RGS1 and RGS2 rest frame spectrum of NGC 3783 (black), binned into groups of two channels, with detailed warm absorber model superimposed (red), plotted over the range 22–38  $\text{\AA}$  (0.564–2.066 keV). The most



**Fig. 10.** Ionic column plotted against ionisation parameter of maximum abundance for the iron states (except Fe XV, for which only an upper limit could be obtained) listed in Table 6.

$K\beta$  and Fe XXVI. The apparent broadening and redshift of this line are, then, a function of the flux and redshift of the components that it contains. Fe XXVI emission would not originate from the same material as Fe  $K\alpha$ ; instead, this must be emission from much hotter gas.

## 6.4. Warm absorber

### 6.4.1. Modelling strategy

We developed a strategy for modelling the complex RGS spectrum of this object using SPEX, with the aim of reproducing both line and continuum absorption self-consistently, and of obtaining some physical insight into the properties of the warm absorber.

This process has several steps. Firstly, *slab* - a model which applies absorption by individual ions in a photoionised medium to a given spectral continuum - is used to identify the patterns of line absorption by different ionic species in the warm absorber, and to measure the blueshift of the lines. This model is also used to investigate the contributions that each ion can make to continuum absorption. Once the lines have been identified, the relative strengths of line absorption by the different iron species are used to derive the ionisation structure of the absorber. Iron is used for this purpose due to its wide range of ionisation states, several of which can be observable in gas at a

given level of ionisation, thus providing a good constraint on the ionisation parameter of the medium. We contend that this gives a more reliable picture of the ionisation structure of the absorber than a simple measurement of O VII and O VIII edge opacities, which is dependent on continuum absorption (a more ambiguous diagnostic than line absorption) from only two ions.

At this stage, the *xabs* model, which applies line and continuum absorption by a given column of photoionised gas at a given ionisation parameter to the spectral continuum, is used to fit an overall column for the absorber at the ionisation parameter independently derived from the iron absorption. Multiple *xabs* models can be used to represent the different ionisation phases within the absorber. This is done with the assumption of solar abundances within the absorber, but, once an estimate of the overall column has been derived, the individual elemental abundances can be allowed to vary. To get columns for the individual ionic states, one can return to *slab* to fit values for these.

At every stage of the process outlined here, due to the likelihood of ‘blind’  $\chi^2$  fitting being misled by the enormous complexity of the spectrum and model, it is necessary to spend some time getting to know the model properties by trial and error. A probable range of realistic values needs to be established before using  $\chi^2$  minimisation to home in on the final value for a parameter, being careful

**Table 6.** The absorbing columns of individual ions in the detailed model of the warm absorber.

Ion	Log <sub>10</sub> column <sup>a</sup>	Log $\xi_{\text{ma}}$ <sup>b</sup>
C v	16.5 ± 0.5	0.15
C vi	17.5 ± 0.2	1.15
N vi	16.9 ± 0.2	0.70
N vii	16.8 ± 0.2	1.50
O iv	17.1 ± 0.2	-0.75
O v	17.5 ± 0.2	-0.05
O vi	17.0 ± 0.5	0.45
O vii	17.3 ± 0.2	1.15
O viii	16.8 ± 0.5	1.75
Ne ix	18.0 ± 0.5	1.70
Ne x	18.2 ± 0.5	2.25
Mg xi	17.5 ± 0.5	2.10
Mg xii	17.8 ± 0.5	2.60
Si xi	17.0 ± 0.2	1.70
Si xii	17.9 ± 0.2	2.05
Si xiii	17.5 ± 0.2	2.40
Si xiv	18.0 ± 0.5	2.90
S xii	17.0 ± 0.2	1.70
S xiii	17.5 ± 0.2	2.05
S xiv	16.7 ± 0.2	2.40
Ar xiii	16.1 ± 0.2	1.80
Ar xiv	16.4 ± 0.5	2.10
Fe vi	16.0 ± 0.5	-0.90
Fe vii	16.1 ± 0.5	-0.35
Fe viii	17.1 ± 0.5	0.05
Fe ix	16.8 ± 0.5	0.35
Fe x	16.8 ± 0.5	0.55
Fe xi	16.4 ± 0.5	0.75
Fe xii	16.0 ± 0.5	1.00
Fe xiii	16.2 ± 0.5	1.20
Fe xiv	16.2 ± 0.5	1.40
Fe xv	15.4 <sup>c</sup>	1.60
Fe xvi	16.4 ± 0.5	1.60
Fe xvii	17.5 ± 0.2	2.10
Fe xviii	17.5 ± 0.2	2.30
Fe xix	17.4 ± 0.2	2.50
Fe xx	17.5 ± 0.2	2.80
Fe xxi	17.7 ± 0.2	3.00
Fe xxii	17.1 ± 0.2	3.10
Fe xxiii	17.5 ± 0.5	3.30
Fe xxiv	17.0 ± 0.5	3.45

<sup>a</sup> where the column is in cm<sup>-2</sup><sup>b</sup> log  $\xi$  at which the ion is at its maximum abundance ( $\xi$  is in erg cm s<sup>-1</sup>)<sup>c</sup> upper limit

to ignore any spectral regions that will bias the fit towards unphysical results.

#### 6.4.2. Two-phase warm absorber model

The two-phase warm absorber model reproduces the global form of the spectrum very well. It indicates that the ratios of elemental abundances to each other in the high-

ionisation phase of the gas are probably not too different from solar. However, in order to fit the very deep UTA associated with the low-ionisation phase of the gas, iron would be ten times more abundant, relative to its solar value, than oxygen. This is sensitive to the turbulent velocity and overall absorbing column assumed for the low-ionisation phase, and the level of the spectral continuum itself. If the spectral continuum was higher (especially at the long wavelength end of the spectrum, where continuum absorption by C IV and C V becomes important), the equivalent hydrogen column of the low-ionisation phase would need to be higher and thus the iron abundance could be a lot lower. The value of  $\xi$  in the model is too high for some of the lower charge states of Fe, therefore requiring high Fe abundance to compensate for the low fractional ionic abundances that the model produces. The low-ionisation phase might also have a wider distribution of  $\xi$ . Moreover, the ionisation balance calculations of Fe-M are very uncertain; in particular, the dielectronic recombination rates for these ions have never been measured or calculated with modern codes.

The much shallower UTA in IRAS 13349+2438 (Sako et al. 2001) requires an iron abundance 2–3 times that of oxygen, and the UTA in the LETGS spectrum of NGC 5548 (Kaastra et al. 2002b) also suggests an overabundance of iron. It is possible that these apparent high iron abundances could be due to a covering factor effect (e.g. Arav et al. 2002). The iron abundance in the UTA is effectively measured against the abundances of C, N and O, and if these lines are highly saturated their equivalent widths will be determined by the covering factor and velocity width, and will be only weakly dependent on the column density.

The position of the UTA also coincides with the O VII absorption edge (16.78 Å (0.7389 keV) in the rest frame), and the deeper this edge, the lower the iron abundance required to fit the UTA. We would not expect significant O VII absorption to originate from the low-ionisation phase itself, as it is not highly ionised enough. The high-ionisation phase, on the other hand, could be a plausible source of O VII edge absorption. At the derived ionisation parameter of this phase, though, oxygen will be predominantly in the form of O VIII. Certainly, the fitted equivalent hydrogen column of the high-ionisation phase, at  $2.8 \times 10^{22}$  cm<sup>-2</sup>, is easily high enough to give rise to significant edge absorption. Examination of the detailed form of the spectrum where the edges are expected, however, shows that any such edges would be completely masked by L-shell and M-shell iron absorption. We propose, then, that although O VII and O VIII edge opacity is certainly part of the picture, absorption by L-shell and M-shell iron is likely to be more important in determining the detailed form of the spectrum in these regions.

The O VIII absorption is, in fact, seriously overpredicted in the two-phase model ( $N_{\text{O VIII}} \geq 10^{18}$  cm<sup>-2</sup>), implying a very deep O VIII Ly $\alpha$  absorption line (at 18.97 Å (0.6536 keV), see Fig. 6) that is not matched by the data. It is of course possible that this line is partially filled in by

O VIII line emission, but we cannot test this with the current dataset. It can also be seen from Fig. 6 that the iron absorption is overpredicted at the low-wavelength side of the 15 Å feature and increasingly underpredicted at lower wavelengths. Fig. 7 shows that the two-phase model compensates for this by generating neon edge absorption at  $\sim 9.5$  Å. This clearly demonstrates that it was not correct to approximate  $\log \xi$  of the high-ionisation phase with a single value of 2.4 - there is a lot of gas which is more highly ionised than this, so the high-ionisation phase should be more accurately modelled as containing a range of ionisation parameters from 2.4 upwards. The low-ionisation phase, too, is not perfectly explained by a single phase with  $\log \xi$  of 0.3, but the single ionisation parameter is a good approximation in this case.

A comparison between the parameters of our two-phase model, and that of Kaspi et al. (2001), is given in Table 7. Their low-ionisation phase corresponds most closely to our high-ionisation phase, whilst their high-ionisation phase is much more highly ionised than ours. It probably corresponds to the highly-ionised iron at the low-wavelength end of our spectrum which our two-phase model does not explain. Kaspi et al. (2001) do not model our low-ionisation (UTA) phase, although they report the presence of the UTA. The warm absorber modelled by De Rosa et al. (2002) is consistent with the low-ionisation component of Kaspi et al. (2001), and thus with our high-ionisation component.

The two models, although different, are not in fundamental disagreement with each other. The differences between them may be due to the methods used to model the spectrum. The Kaspi et al. (2001) modelling relies on the use of measured equivalent widths of the absorption lines as well as the fitting of a continuum model to ‘line-free zones’ where absorption lines are not expected. We, on the other hand, use the line absorption as the prime diagnostic and tie the continuum absorption directly to this. The elemental abundances in their model were kept at solar values, whilst ours were allowed to vary. These two phase models, though, are just approximations to a multi-phase absorber, as we discuss above, and as already pointed out in Behar & Netzer (2002).

### 6.4.3. Detailed modelling

Our detailed model, using *slab*, provides a list of the total columns of the main ions present in the multi-phase warm absorber, summed across all phases. This model (Fig. 8 and Fig. 9) provides a better reproduction of the detailed form of the spectrum than the simple *xabs* two-phase approximation. The agreement of our model with the data in the region of L-shell iron absorption, in particular, is excellent in some ranges. There are still some significant residuals in this region, though. The wavelengths and strengths of the L-shell iron absorption lines and blends are very well known, so it is unlikely that this is the source of the discrepancy. Some of the absorption may originate from

**Table 7.** Comparison of the properties of the two-phase warm absorber model presented here with that of Kaspi et al. (2001)

Phase	Property	Our model	Kaspi et al. 2001 model
High-ionisation	$\log \xi^a$	2.4	3.5
High-ionisation	$N_H^b$	$2.8^{+0.01}_{-0.3} \times 10^{22}$	$1.6 \times 10^{22}$
Low-ionisation	$\log \xi^a$	0.3	2.5
Low-ionisation	$N_H^b$	$5.4^{+0.05}_{-0.5} \times 10^{20}$	$1.6 \times 10^{22}$

<sup>a</sup> Log of the ionisation parameter, where  $\xi$  is in  $\text{erg cm s}^{-1}$

<sup>b</sup> Equivalent hydrogen column of phase in  $\text{cm}^{-2}$

elements, particularly calcium, which are not yet included in our model. It is possible to reduce some of the residuals by increasing the O VIII continuum absorption, although this then introduces very deep O VIII absorption lines which do not appear in our data. One interesting explanation for these residuals might be that our model lacks the higher level transitions associated with the UTA (Behar et al. 2001) which appear when there is a high column density in the UTA ions.

The derived columns of the L-shell iron states above Fe XX do indicate that there is a lot of gas in the hot phase of this warm absorber at  $\log$  ionisation parameters above 2.4, as shown in Fig. 10. This may be part of the explanation for the inability of the two-zone model to accurately reproduce the O VIII absorption; much of the gas in the hot phase may simply be too highly ionised for O VIII to form in significant quantities. The derived columns of O VII and O VIII ( $\sim 10^{17} \text{ cm}^{-2}$ ) in this detailed model are too low to produce significant edge absorption, but their values are fairly uncertain.

The effective area of the HETGS system falls off above around 23 Å (0.539 keV), so since the RGS spectrum extends up to 38 Å (0.326 keV), we can add quantitatively to the knowledge gained from the Chandra spectrum of NGC 3783. In this high wavelength range we see absorption lines from N VI, N VII, C VI and possibly C V, Si XI, Si XII, S XII, S XIII, S XIV, Ar XIII and Ar XIV. Direct superimposition of the model on the spectrum (Fig. 8 and Fig. 9) reveals clearly the inaccuracies in the model wavelengths of various lines, and shows that there are further absorption lines in the spectrum that are yet to be included in our model. This work will be carried forward using the 280 ks RGS exposure of NGC 3783.

What is the connection between the higher and lower ionisation phases of the warm absorber? They appear to be flowing towards us at a similar speed, around 800  $\text{km s}^{-1}$ . The high-ionisation phase has an equivalent hydrogen column about fifty times that of the low-ionisation phase in our line of sight, and the low-ionisation phase apparently has a very high iron abundance. The spread of values of the ionisation parameter in the two phases could give us

a clue as to what is going on. As indicated by the results of the *xabs* modelling, and also in Fig. 10, the low-ionisation phase appears to be concentrated around a single ionisation parameter, whilst the high-ionisation phase contains a range of levels of ionisation. When the central engine of NGC 3783 is radiating at a given luminosity, the ionisation parameter of the absorbing outflow is determined by the density and distance from the source of the absorbing gas.  $\xi$  is more sensitive to changes in distance than changes in density. One could, then, surmise that the wide range of  $\xi$  in the high-ionisation phase implies absorption by gas at a wide range of distances and (to a lesser extent perhaps) densities, whereas the low-ionisation, UTA phase absorption is quite localised to gas at a certain distance from the source, and at a particular density. This would imply the existence of a highly-ionised, multi-temperature outflow (as in the multi-temperature wind model posited by Krolik & Kriss 2001), spread over a range of distances from the source, and denser and perhaps iron-rich material at a fixed location within the circumnuclear environment.

Sako et al. (2001) claim that the low-ionisation phase of the warm absorber in IRAS 13349+2438 might originate in the dusty torus surrounding the active nucleus, which would fit in with the scenario described here. Perhaps the abnormally high iron content of the corresponding phase in NGC 3783 is due to the sublimation of iron-rich dust from the edge of a dusty torus. If this was the case, the high-ionisation phase would form the bulk of the warm absorber, and wherever it encountered the edge of the dusty torus it would sweep up denser material; this would then be photoionised and the dust it contained would be sublimated, giving rise to the high iron content of the low-ionisation phase. This places constraints, of course, on the composition of the dust, which would have to be justified. The higher signal-to-noise RGS spectrum which has now been obtained will allow us to investigate the relative abundance of other non-volatile elements which can be bound into dust grains (including silicon and carbon), to test this hypothesis further.

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