

The cold gas and dust properties of red star-forming galaxies

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ABSTRACT

We study the cold gas and dust properties for a sample of red star-forming galaxies called ‘red misfits.’ We collect single-dish CO observations and H I observations from representative samples of low-redshift galaxies, as well as our own James Clerk Maxwell Telescope CO observations of red misfits. We also obtain SCUBA-2 850 μm observations for a subset of these galaxies. With these data we compare the molecular gas, total cold gas, and dust properties of red misfits against those of their blue counterparts (‘blue actives’) taking non-detections into account using a survival analysis technique. We compare these properties at fixed position in the $\log \text{SFR} - \log M_\star$ plane, as well as versus offset from the star-forming main sequence. Compared to blue actives, red misfits have slightly longer molecular gas depletion times, similar total gas depletion times, significantly lower molecular- and total-gas mass fractions, lower dust-to-stellar mass ratios, similar dust-to-gas ratios, and a significantly flatter slope in the $\log M_{\text{mol}} - \log M_\star$ plane. Our results suggest that red misfits as a population are likely quenching due to a shortage in gas supply.

Key words: dust, extinction – ISM: molecules – galaxies: evolution – galaxies: ISM – galaxies: star formation – submillimetre: ISM.

1 INTRODUCTION

A key finding from large surveys of the local Universe such as the *Sloan Digital Sky Survey* (SDSS; York et al. 2000) is that the vast majority of galaxies in the nearby Universe tend to fall into one of two categories: a star-forming ‘main sequence’ (SFMS) where star formation rate (SFR) and stellar mass $M_\star$ are well-correlated, or the quiescent population where SFRs are low and not well-correlated with $M_\star$. In colour–magnitude space, star-forming galaxies are found in the diffusely populated region called the ‘blue cloud,’ while quiescent galaxies have red colours and form a tight correlation between colour and magnitude called the ‘red sequence.’ A small but significant number of galaxies lie in the so-called ‘green valley’ between the main sequence and red cloud (e.g. Salim 2014).

Studying the relationship between galaxy properties and galaxy position in the $\text{SFR} - M_\star$ plane has provided insight into which physi-

cal processes are responsible for evolution in this plane (Saintonge & Catinella 2022). One approach to tackle this question is to focus on populations with intermediate specific SFR ($\text{SSFR} \equiv \text{SFR}/M_\star$), which may be evolving away from or towards the SFMS (e.g. Salim 2014; Schawinski et al. 2014; Li et al. 2015; Smethurst et al. 2015; Belfiore et al. 2017; Lin et al. 2017; Coenda, Martínez & Muriel 2018; Eales et al. 2018; Salim, Boquien & Lee 2018; Mancini et al. 2019; Brownson et al. 2020; Lin et al. 2022). However, some works argue that the green valley exists due to observational biases rather than physical processes (Schawinski et al. 2014; Eales et al. 2018).

According to the gas-regulator model (Lilly et al. 2013), processes which affect inflows, outflows, and consumption of gas determine the star formation rate of a galaxy. Gas depletion time t_{gas}

$$t_{\text{gas}} [\text{yr}] \equiv \frac{M_{\text{gas}} [M_\odot]}{\text{SFR} [M_\odot \text{ yr}^{-1}]} \quad (1)$$

is the time it would take for a gas reservoir to turn into stars assuming none of this gas dissipates, no new gas is accreted, gas is not returned to the interstellar medium (ISM) via stellar evolution, and that the

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SFR is constant over time. Although these assumptions are not physically realistic, it is useful to think of t_{gas} as a proxy for the efficiency with which gas is converted into stars. In the literature, the reciprocal of gas depletion time is often referred to as ‘star formation efficiency’ (SFE, e.g. Leroy et al. 2008; Saintonge et al. 2017). To avoid confusion with the theoretical star formation efficiency ϵ_{SF} , namely the fraction of a gas reservoir that forms stars before it dissipates, or the more commonly used efficiency per free fall time, in this work we will write t_{gas}^{-1} instead of ‘SFE.’

Observations of the total cold atomic and molecular gas reservoirs in large samples of nearby galaxies such as the Galaxy Evolution Explorer (GALEX) Arecibo SDSS Survey (xGASS; Catinella et al. 2018), the CO Legacy Database for GASS (xCOLD GASS; Saintonge et al. 2011, 2017), and the James Clerk Maxwell Telescope (JCMT) dust and gas In Nearby Galaxies Legacy Exploration (JINGLE; Saintonge et al. 2018), have found that t_{gas} is correlated with offset from the SFMS such that t_{gas} decreases with increasing offset from the main sequence

$$\Delta\text{MS} [\text{dex}] \equiv \log \frac{\text{SFR}}{\text{SFR}_{\text{MS}}}, \quad (2)$$

where SFR_{MS} (SFR as a function of stellar mass) defines the SFMS. Tacconi et al. (2018) find that the trend between t_{gas} and ΔMS persists from $z = 4$ to 0. Colombo et al. (2020) find that declining molecular gas mass fractions drive galaxies off of the SFMS, and that once a galaxy is quenched, t_{gas} is more important than molecular gas mass in determining the SFR. With the ALMA-MaNGA QUENCHING and STAR FORMATION (ALMAQUEST) sample, Lin et al. (2020) and Ellison et al. (2020) find that local variations in t_{gas} cause regions to depart from the spatially resolved SFMS (that is, the SFMS based on SFR and M_* surface densities in sub-regions of galaxies rather than galaxy-integrated measurements). Brownson et al. (2020) study seven green valley galaxies, and find that t_{gas} and f_{gas} (gas mass divided by stellar mass) are equally important in driving departures from the SFMS. A recent analysis of xCOLD GASS and xGASS data (Feldmann 2020) found that after accounting for galaxy selection biases (e.g. stellar mass, SFR) and observational uncertainties the correlation between $\log t_{\text{mol}}$ (the depletion time of molecular gas only) and ΔMS flattens significantly, from $\log t_{\text{mol}} \propto -0.5 \Delta\text{MS}$ to $\log t_{\text{mol}} \propto -0.24 \Delta\text{MS}$. In other words, they find that t_{mol} has a small but significant dependence on offset from the SFMS after accounting for selection effects and observational uncertainties. This nearly flat relationship between $\log t_{\text{mol}}$ and ΔMS echoes the findings of Sargent et al. (2014).

A major motivation of this work is to improve our understanding of galaxy evolution by focusing on the gas and dust properties of a galaxy population (red misfits), that is selected differently from the green valley and shows differences from that population, but whose star formation, similar to the green valley, is possibly in the act of quenching. Another motivation is to use a large multiwavelength sample to compare the gas and dust properties of red misfits with the overall population of low-redshift star-forming galaxies.

We investigate a population of galaxies selected from SDSS to be optically red and actively forming stars (Evans, Parker & Roberts 2018). This population, called ‘red misfits,’ appears to have no preference for environment, has an elevated fraction of active galactic nuclei (AGNs), and accounts for about 10 per cent of low redshift galaxies across stellar masses from $\log M_* = 9.5$ to 11.5 (Evans et al. 2018). Evans et al. (2018) compared the properties of red misfits with green valley galaxies; they find that about 30 per cent of red misfits also lie in the green valley. Although there are similarities in these populations (e.g. both have significant AGN fractions, both

are dominated by intermediate morphologies, and star formation is likely slowing down in both populations), Evans et al. (2018) find several differences. Unlike green valley galaxies, red misfits are not simply in between blue star forming and red dead galaxies in the $\log \text{SFR} - \log M_*$ plane. Green valley galaxies with late-type morphologies are rarely found in haloes with masses larger than $10^{12} h^{-1} M_\odot$ (Schawinski et al. 2014); while red misfits are found in roughly the same proportions for all halo masses (Evans et al. 2018). Green valley galaxies lie between the blue star-forming, and red-and-dead populations, suggesting that they represent an intermediate stage of galaxy evolution (Salim et al. 2018), while red misfits, in contrast, lie below, on, and above the SFMS while being red in colour, making their average stage of evolution less obvious and making them an interesting population to explore further (Evans et al. 2018).

The primary goal of this work is to better understand the evolutionary state of red misfit galaxies by studying their cold gas and dust properties. We use a combination of two of the largest samples of CO in the local Universe (xCOLD GASS and JINGLE), and sub-millimeter observations of a large number of them (from JINGLE and our own observations), as well as H I measurements from xGASS and the Arecibo Legacy Fast ALFA (ALFALFA) $\alpha.100$ catalogue (Haynes et al. 2018) to compare the interstellar medium in red misfits with their blue counterparts (‘blue actives’) and to try to understand the nature of red misfits.

We assume a flat Λ CDM cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_{\text{m},0} = 0.27$, and $T_{\text{CMB},0} = 2.275 \text{ K}$.

2 DATA AND DATA PROCESSING

2.1 Star formation rates, stellar masses, and other basic properties

Optical colours and specific star formation rates are required to select red misfits. Star formation rates and stellar masses are also required to compute gas and dust-based quantities such as gas depletion times. These optical data are taken from the following sources:

(i) SDSS $g - r$ colours that have been extinction-corrected and inclination-corrected, taken from Evans et al. (2018). Some galaxies with CO measurements in our sample were not included in Evans et al. (2018). For this subset we computed $g - r$ colours using the same method.

(ii) SFR and stellar masses from UV + optical spectral energy distribution (SED) fitting taken from the GSWLC-M2 catalogue (Salim et al. 2018). The medium-depth (M2) measurements are ideal for star-forming galaxies, which are the focus of this work. Where available we use the M2 catalogue, and for a small subset of galaxies we use the A2 catalogue.

The distribution of all 118 769 galaxies in the intersection of the Evans et al. (2018) catalogue and GSWLC-M2, in the $\log \text{SSFR}$ versus $g - r$ colour plane ($0.01 \leq z \leq 0.12$, $9.0 \leq \log M_* \leq 11.9$) is shown in the left-hand panel of Fig. 1. We used the GSWLC-M2 star formation rates and stellar masses to compute the dividing line between star-forming and passive galaxies (the horizontal line at $\log \text{SSFR} [\text{yr}^{-1}] = -11.3$). This cut was determined by fitting the $\log \text{SSFR}$ histogram with a double-Gaussian and calculating where the two Gaussians intersect. In this work we focus on star-forming galaxies, namely red misfits (upper right quadrant of this figure) and ‘blue actives’ (upper left quadrant). Red misfits are defined as galaxies that are star forming (above the horizontal line) and red in colour ($g - r \geq 0.67$; Evans et al. 2018). The right-hand panel of Fig. 1 shows the relationship between ΔMS (equation 2) and $g -$

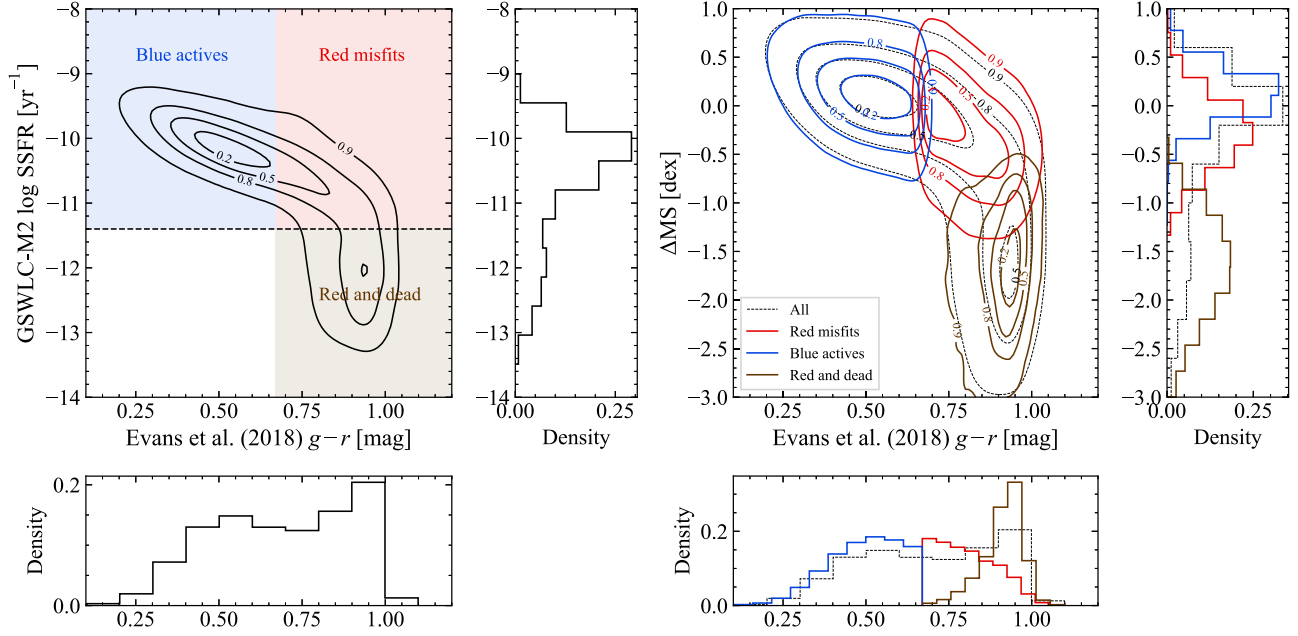


Figure 1. *Left:* SSFR versus $g-r$ colour of all 118 769 galaxies in the intersection of the Evans et al. (2018) and GSWLC-M2 catalogues. We used the GSWLC-M2 star formation rates and stellar masses to compute the dividing line between star-forming and passive galaxies (the horizontal line at $\log \text{SSFR} [\text{yr}^{-1}] = -11.3$). This cut was determined by fitting the log SSFR histogram with a double-Gaussian and calculating where the two Gaussians intersect. Red misfits are defined as galaxies that are star forming (above the horizontal line) and red in colour ($g-r \geq 0.67$; Evans et al. 2018). Red, passively evolving (‘red and dead’) galaxies are also indicated. *Right:* Offset from the SFMS versus $g-r$ colour, for all galaxies (black dashed), blue actives (blue), red misfits (red), and red and dead galaxies (brown). Red misfits clearly occupy a wide range in ΔMS that is systematically higher than red and dead galaxies.

r colour, indicating that red misfits occupy a wide range in ΔMS . Red misfits have a broader colour distribution and are systematically bluer than red and dead galaxies. Red misfits have a narrower ΔMS distribution than red and dead galaxies, and a median ΔMS that is 1.4 dex closer to the MS than the red and dead population (Table 1).

2.2 Single-dish CO observations

We use CO observations from the following three sources:

(i) **JCMT CO(2–1) measurements from the JINGLE survey** (Saintonge et al. 2018). JINGLE is a representative sample of galaxies ranging from just below the star-forming main sequence to the starburst regime. The entire JINGLE sample was observed with SCUBA-2, while a subset of about 75 galaxies were observed in CO(2–1). The JCMT beam at the frequency of CO(2–1) is 20 arcsec (Saintonge et al. 2018). Molecular gas mass is related to CO(2–1) luminosity $L_{\text{CO}(2-1)}$ by

$$M_{\text{mol}} [M_{\odot}] = r_{21}^{-1} \alpha_{\text{CO}} L_{\text{CO}(2-1)} [\text{K km s}^{-1} \text{pc}^2], \quad (3)$$

where α_{CO} is the CO-to-H₂ conversion factor (Bolatto, Wolfire & Leroy 2013) and r_{21} is the ratio of CO(2–1) to CO(1–0) intensities. Note that in this work we use the subscript ‘mol’ to indicate total molecular gas (hydrogen and helium). In normal star-forming regions α_{CO} is often assumed to be $4.35 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ (Bolatto et al. 2013), which includes the contribution from helium (a factor of 1.36). For CO(2–1) measurements, one must assume a value of r_{21} . Variations from $r_{21} \sim 0.6$ (Yajima et al. 2021) to ~ 0.8 (Saintonge et al. 2017) have been observed. We use the commonly used value of 0.7. The JINGLE analysis assumed a ratio of $r_{21} = 0.7$ and $\alpha_{\text{CO}} = 4.35$ (Saintonge et al. 2018).

(ii) **IRAM 30 m CO(2–1) and some CO(1–0) fluxes from the xCOLD GASS survey** (Saintonge et al. 2017). xCOLD GASS is

a representative sample of CO emission in nearby galaxies. These galaxies were primarily selected from the xGASS survey (Catinella et al. 2018). The IRAM 30-m beam sizes at the frequencies of the CO(2–1) and CO(1–0) lines are 11 and 22 arcsec, respectively (Saintonge et al. 2017). The molecular gas masses in the xCOLD GASS catalogue were computed using a metallicity-dependent α_{CO} . To be consistent with the JINGLE catalogue we recalculated these molecular gas masses using $\alpha_{\text{CO}} = 4.35$.

(iii) **Our own JCMT CO(2–1) measurements of red misfits.** These galaxies are from the JINGLE sample that were not scheduled to be observed in CO(2–1), but had already been observed with SCUBA-2. These data were reduced and converted into molecular gas masses using the same approach as for JINGLE galaxies [C. Wilson, private communication]. These measurements do not appear elsewhere in the literature and are provided in Table A1.

The number of galaxies with CO measurements, and the sources of these measurements are shown in the first row of Table 2.

2.3 H I observations

In addition to the molecular gas supply, we are interested in measuring the total gas mass

$$M_{\text{gas}} = 1.36(M_{\text{H}_2} + M_{\text{HI}}), \quad (4)$$

where M_{H_2} is the molecular hydrogen mass and M_{HI} is the neutral hydrogen mass. Note that in this work, the subscript ‘gas’ refers to the total molecular and atomic gas as shown in equation (4). All of our H I measurements were made using the Arecibo telescope, which has a beam size of ~ 3.5 arcmin (Catinella et al. 2018). We collected H I measurements from the following sources:

Table 1. Statistics of ΔMS and $g - r$ colour measured from galaxies in the parent sample.

Population (1)	Median (2)		Robust standard deviation (3)	
	ΔMS (dex)	$g - r$ (mag)	ΔMS (dex)	$\Delta g - r$ (mag)
All galaxies	-0.13	0.70	0.56	0.27
Red misfits	-0.27	0.80	0.37	0.11
Blue actives	0.13	0.51	0.25	0.12
Red and dead	-1.65	0.92	0.56	0.05

Note. (3) Defined as 1.4826 MAD (Astropy Collaboration 2018).

Table 2. Numbers of galaxies with CO, H I, and 850- μm measurements, and the sources of these measurements.

Measurement	Source(s)	# Galaxies	# Non-Detections
CO	- JINGLE - xCOLD GASS - Our own observations from JCMT	(f)427	(f)61
H I 21 cm ^a	- xGASS - JINGLE/Arecibo - ALFALFA α .100	(f)369	(f)65
Dust (850 μm)	- JINGLE - Our own observations from JCMT	209	106

Note. ^aOnly galaxies with both CO and H I observations are used.

(i) **The ALFALFA α .100 catalogue (Haynes et al. 2018).** We cross-matched the JINGLE sample with this catalogue, which provided H I measurements for 99 galaxies from the JINGLE sample.

(ii) **The xGASS representative sample (Catinella et al. 2018).** This sample provides H I measurements for most of the galaxies in the xCOLD GASS sample.

(iii) Observations of a subset of the JINGLE sample using the Arecibo telescope (obtained by private communication with M. Smith). This sample consists of 60 JINGLE galaxies which were not observed as part of the ALFALFA survey.

The number of galaxies with H I measurements, and the sources of these measurements are shown in the second row of Table 2.

2.4 Dust masses from sub-millimeter observations

We use SCUBA-2 850 μm flux densities $S_{850 \mu m}$ to estimate the cold dust mass of galaxies in our sample. The SCUBA-2 beam size at 850 μm is 13 arcsec. These measurements are from the following sources:

(i) **SCUBA-2 850 μm measurements from the JINGLE survey** Smith et al. (2019). These data are available at <http://www.star.ucl.ac.uk/JINGLE/data.html>. On that page is a catalog of far-infrared and sub-mm photometry, from which we obtained 850 μm flux measurements.

(ii) **Our own SCUBA-2 850 μm measurements of a sample of red misfits.** These galaxies were selected from the xCOLD GASS sample. xCOLD GASS does not overlap significantly with far-infrared surveys – this was the primary motivation for obtaining SCUBA-2 measurements of these galaxies. We present our 850- μm measurements in Table B1. These measurements were processed in the same way as in Smith et al. (2019) except we did not correct for CO(3–2) emission, which contributes a small amount to the observed 850- μm emission. Across the JINGLE sample, the mean

CO(3–2) correction is 10.1 per cent of the predicted 850- μm flux density (Smith et al. 2019).

The number of galaxies with SCUBA-2 850- μm measurements, and the sources of these measurements are shown in the third row of Table 2.

To convert 850- μm flux densities $S_{850 \mu m}$ into dust masses, we first consider the relationship between specific flux and dust mass at wavelength λ assuming it emits as a modified blackbody

$$F_{\lambda} [\text{W m}^{-2} \text{m}^{-1}] = M_{\text{dust}} d_L^{-2} \kappa_{\lambda}(\beta) B_{\lambda}(T), \quad (5)$$

where M_{dust} is the dust mass in kg, d_L is luminosity distance in m, κ_{λ} is the dust opacity in $\text{m}^2 \text{kg}^{-1}$, and $B_{\lambda}(T)$ is the Planck function

$$B_{\lambda}(T) = \frac{2hc^2}{\lambda^5} \frac{1}{\exp(hc/k_B T \lambda) - 1}. \quad (6)$$

Following Lamperti et al. (2019), dust opacity is given by

$$\kappa_{\lambda}(\beta) [\text{m}^2 \text{kg}^{-1}] = \kappa_0 \left(\frac{\lambda_0}{\lambda} \right)^{\beta}, \quad (7)$$

where $\kappa_0 = 5.1 \times 10^{-2} \text{ m}^2 \text{kg}^{-1}$ at 500 μm (Clark et al. 2016), $\lambda_0 = 500 \mu m$, and β is the spectral index.

The 850- μm flux density $S_{850 \mu m}$ in units of Jy can be converted into units of specific intensity via

$$F_{850 \mu m} = 10^{-26} c \lambda^{-2} S_{850 \mu m}. \quad (8)$$

Finally, we can rearrange equation (5) for M_{dust} which gives

$$M_{\text{dust}} = \frac{10^{-26} c \lambda^{-2} d_L^2 S_{850 \mu m}}{\kappa_{\lambda}(\beta) B_{\lambda}(T)}. \quad (9)$$

We use the scaling relation for β from equation (35) in Lamperti et al. (2019)

$$\beta = a_1 \log M_{\star} + a_2 \log A + a_3 \log[12 + \log(\text{O}/\text{H})] + a_4, \quad (10)$$

where $A = \pi r_{50}^2$ in kpc^2 is the surface area corresponding to the

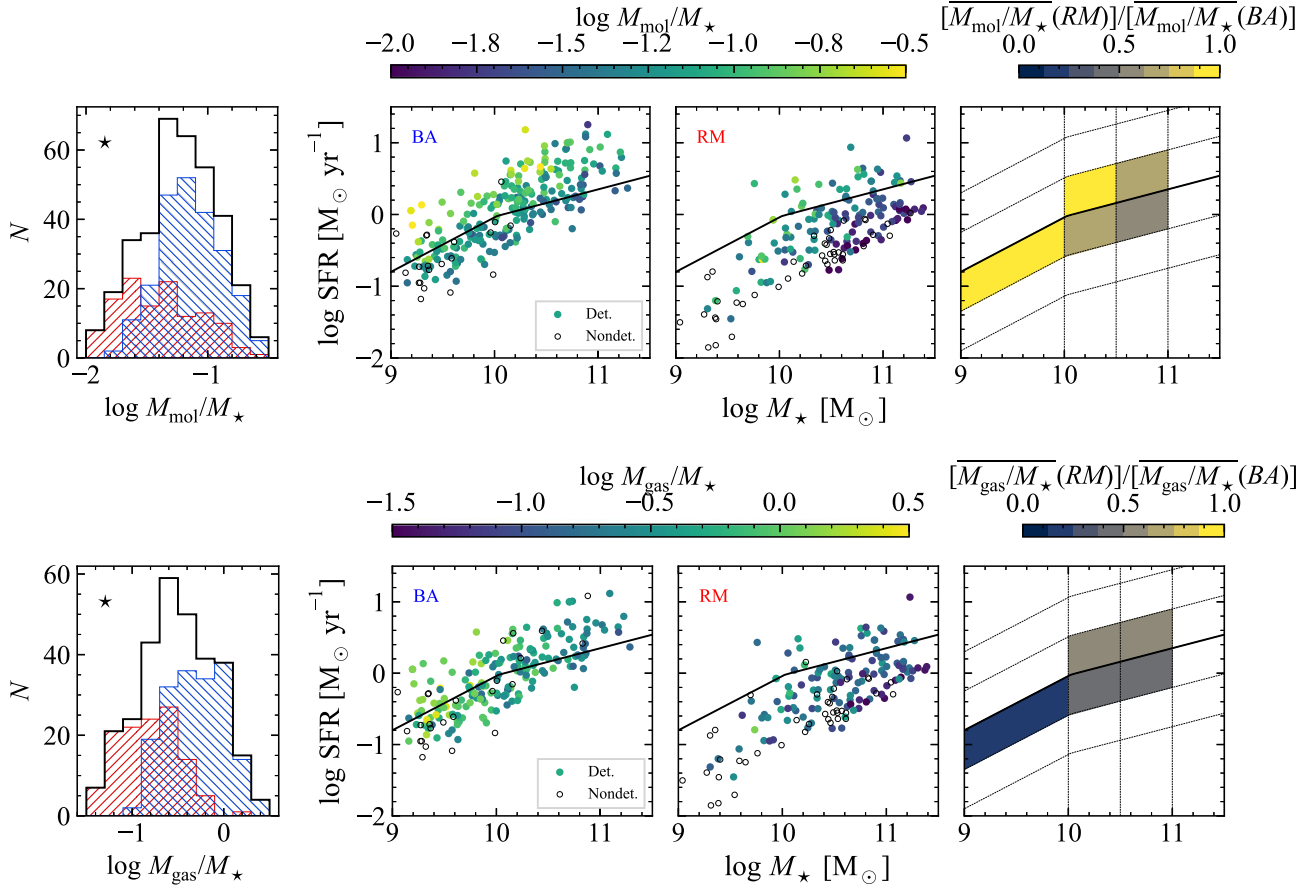


Figure 2. Left: Histograms of molecular (top) and total (bottom) gas mass fractions of all galaxies (black), red misfits (red histogram), and blue actives (blue histogram). Only detections are shown. The KS-test between the red misfit and blue active distributions are shown in Table 3; a ‘★’ symbol in the upper left of a histogram indicates that the distributions are likely different. Restricted means and the ratio of restricted means of red misfits to blue actives are shown in Table 3. Middle panels: gas mass fractions of blue actives (‘BA’) and red misfits (‘RM’) and their relationships to host galaxy position on the SFMS diagram. Coloured circles are detected in H₂ (and H I where those measurements are used). Open circles were observed but not detected. The solid line is the star-forming main sequence from table 1 of Popesso et al. (2019). The right-hand panels show the ratio of the average gas mass fractions of red misfits to blue actives.

SDSS *i*-band half-light radius r_{50} , $12 + \log(\text{O}/\text{H})$ is the gas-phase metallicity using the $[\text{O III}]/[\text{N II}]$ calibration of Pettini & Pagel (2004), and the fit parameters are $a_1 = 0.27$, $a_2 = -0.33$, $a_3 = 0.71$, and $a_4 = -6.62$. The r_{50} and $12 + \log(\text{O}/\text{H})$ measurements were taken from the NASA-Sloan Atlas.¹ The relation for dust temperature T in Kelvin is their equation (37)

$$T = b_1 \log \text{SFR} + b_2 \log M_* + b_3, \quad (11)$$

where $b_1 = 2.91$, $b_2 = -2.27$, and $b_3 = 45.42$. We use these scaling relations to estimate β and T for each galaxy, and then estimate dust mass using equation (9).

2.5 Note regarding beam sizes

Our measurements of CO, H I, and dust should be interpreted as galaxy-integrated totals rather than aperture-matched fluxes. As noted in section 6 of Catinella et al. (2018), although the IRAM/JCMT beams are significantly smaller than Arecibo, it is well known that H I emission extends much further than CO and dust, and so a larger beam is needed to capture all of the H I emission compared to CO and dust.

¹<http://nsatlas.org>

3 ANALYSIS AND RESULTS

3.1 Comparing the gas and dust properties of red misfit and blue active galaxies

To better understand the nature of red misfits and their role in galaxy evolution, it is critical to understand their gas and dust properties. In Figs 2, 3, and 4, we show their gas masses, gas depletion times, and dust mass fractions. We show each quantity from two perspectives in order to compare between red misfits and blue active galaxies. The first perspective is a comparison of distributions of detected measurements, shown in the left-hand panels, which allows us to compare the properties of the entire red misfit and blue active samples. We compare the two unbinned distributions using a two-sample Kolmogorov–Smirnov (KS) test implemented in `scipy.stats.ks_2samp`; if the resulting KS statistic is small or the p -value is large, then the distributions are consistent with each other. A ‘★’ in the upper left of these histograms indicates that the distributions are statistically different. The results of each KS test are shown in Table 3.

The second perspective shows gas mass properties, depletion times, and dust mass fractions in the $\text{SFR}-M_*$ plane (right-hand panels of Figs 2, 3, and 4). With this method we can explore how dust and gas properties for red misfits and blue actives depend on their

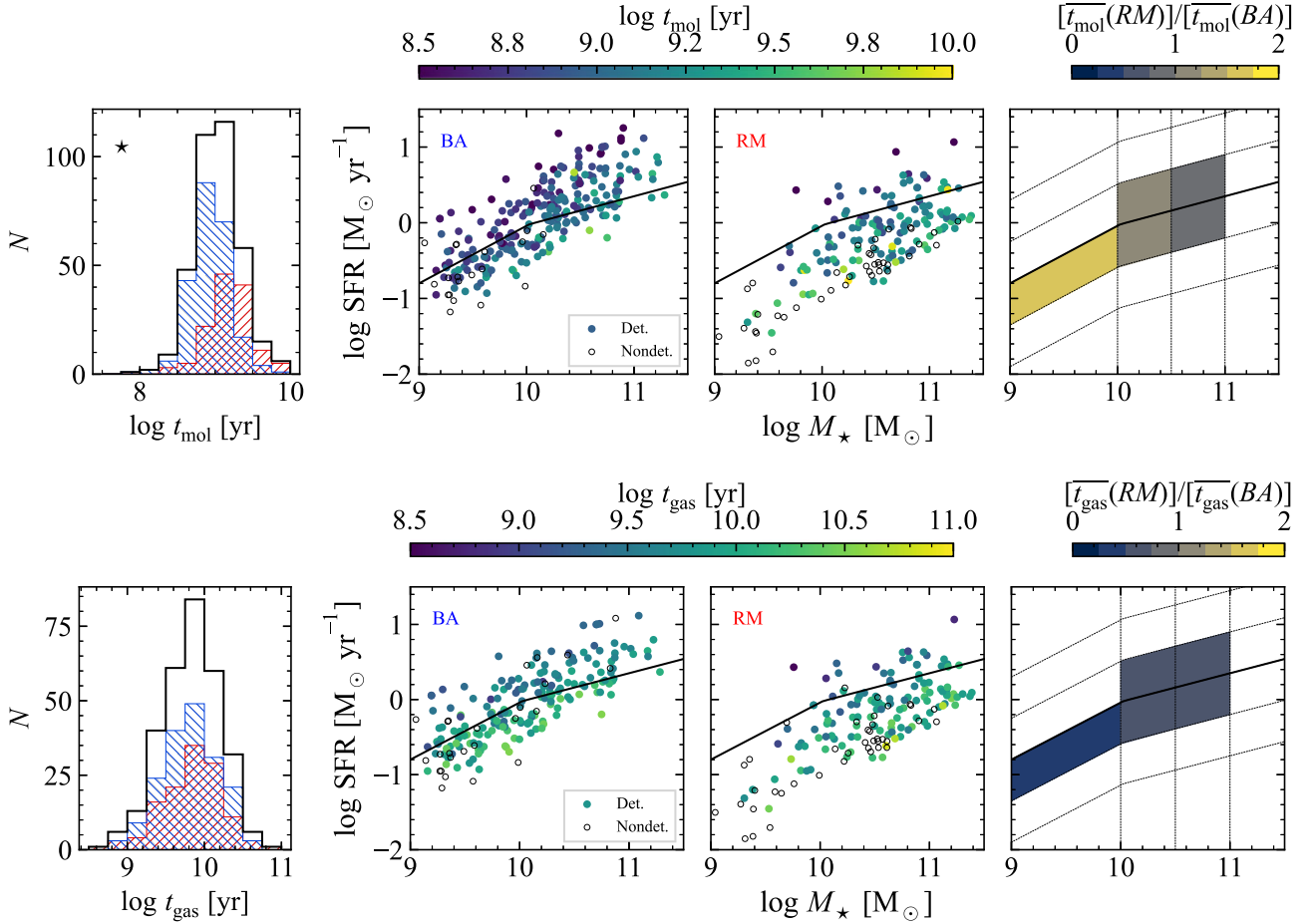


Figure 3. Same as Fig. 2 except with molecular gas (top) and total gas (bottom) depletion times.

position relative to the SFMS. For example, in the right-hand panel of the first row of Fig. 2, the colour of each bin shows the average M_{mol}/M_* for red misfits divided by that of blue actives in that bin. Viewing the sample this way allows us to examine the differences in gas and dust properties while controlling for the fact that red misfits and blue actives are distributed differently in the $\log \text{SFR}-\log M_*$ plane.

In Table 3, we also show the restricted mean and standard error of each quantity for red misfits and blue actives separately, taking non-detections into account. This was done using the Kaplan–Meier estimator (implemented in the `lifelines` PYTHON package), from which we extract a restricted mean and standard error. The Kaplan–Meier estimator is a survival analysis algorithm which estimates the probability distribution of a quantity when measurements of this quantity contain both detections and non-detections. The ‘restricted mean’ is an estimate of the mean of the true distribution. The restricted mean is defined as the integral of the estimated survival function up to the largest detected data point. A recent application of the Kaplan–Meier estimator to molecular gas measurements of galaxies can be found in Mok et al. (2016).

We compare the relative amount of cold gas in these two populations through two quantities: the molecular-to-stellar mass ratio M_{mol}/M_* , and the total gas to stellar mass ratio M_{gas}/M_* , shown in Fig. 2. The distributions on the left and the KS test results (Table 3) indicate that red misfits tend to have lower gas mass fractions than blue active galaxies. The middle two panels show that red misfits and

blue actives are distributed differently in the $\log \text{SFR}-\log M_*$ plane (red misfits tend to lie below the SFMS especially at high stellar masses), and that the gas fractions vary within this space. To compare the average properties as functions of position in the $\log \text{SFR}-\log M_*$ plane, we computed the average gas fractions (detections only) in 2D bins of $\log \text{SFR}$ and $\log M_*$ (right column). We require a minimum of three red misfits and three blue actives per bin. In Fig. 2, aside from the lowest- $\log M_*$ bin and the bin between $10 \leq \log M_* \leq 10.5$ which lies above the SFMS, red misfits have lower molecular gas and total gas mass fractions than blue actives. In the two exceptional bins, red misfits have higher molecular gas mass fractions, and lower total gas mass fractions than blue actives. These two exceptional bins contain few red misfits, and so may not adequately represent the whole population. The restricted mean gas fractions for red misfits and blue actives are shown in Table 3. The ratio of red misfit to blue active gas fractions are significantly less than unity, which supports our findings from detections alone (left-hand panel of Fig. 2). This indicates that red misfits have lower total gas content and molecular gas content relative to blue actives.

Next we compare the molecular and total gas depletion times (Fig. 3). Based on the KS test comparing the red misfit and blue active distributions (Table 3), the t_{mol} distributions are significantly different, while the t_{gas} distributions are not. This result is also supported by the ratios of the restricted means – compared to blue actives, the mean t_{mol} of red misfits is slightly larger and the difference is statistically significant. The mean t_{gas} of both populations are not

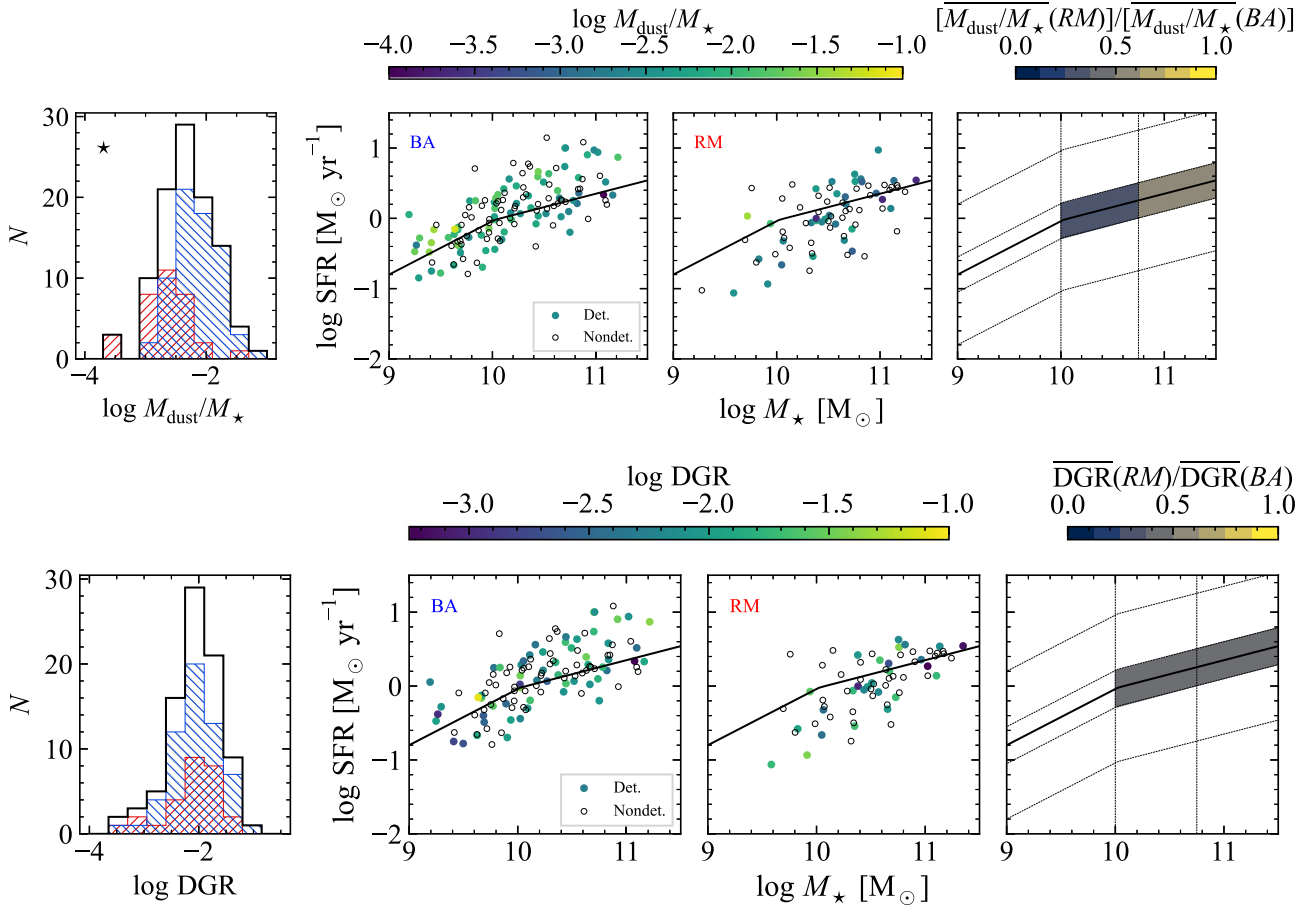


Figure 4. Same as Fig. 2 except with dust-to-stellar mass shown in the top row and the dust-to-gas ratio (DGR; equation 12) in the bottom row. Note that there are fewer galaxies in this figure compared to Figs 2 and 3 because fewer galaxies have both gas and dust measurements.

Table 3. Statistical comparisons of gas- and dust-based quantities between red misfits and blue active galaxies.

Quantity (1)	KS statistic (2)	p-value (3)	Different? (4)	Restricted mean (5)		$\log(\text{RM}/\text{BA})$ (6)	Figure (7)
				RM	BA		
$\log M_{\text{mol}}/M_{\star}$	0.428	1.40×10^{-14}	Y	-1.61 ± 0.03	-1.16 ± 0.02	-0.46 ± 0.04	2
$\log M_{\text{gas}}/M_{\star}$	0.592	3.33×10^{-16}	Y	-0.94 ± 0.03	-0.33 ± 0.02	-0.61 ± 0.04	2
$\log t_{\text{mol}} [\text{yr}]$	0.409	2.70×10^{-13}	Y	9.10 ± 0.03	8.91 ± 0.02	0.20 ± 0.03	3
$\log t_{\text{gas}} [\text{yr}]$	0.116	2.47×10^{-1}	N	9.76 ± 0.03	9.77 ± 0.03	-0.01 ± 0.05	3
$\log M_{\text{dust}}/M_{\star}$	0.536	1.79×10^{-6}	Y	-2.87 ± 0.07	-2.20 ± 0.05	-0.67 ± 0.09	4
$\log \text{DGR}$	0.112	9.44×10^{-1}	N	-2.32 ± 0.09	-2.07 ± 0.05	-0.2 ± 0.1	4

Notes. (2) Kolmogorov–Smirnov (KS) statistic comparing red misfits and blue actives. This includes detections only, by definition (Section 3.1).

(3) p-value corresponding to the KS statistic.

(4) Are the distributions statistically different based on the KS statistic (Y/N)? Y if $\text{KS} \geq 0.4$ and $p \ll 1$.

(5) Kaplan–Meier restricted mean and standard error (Section 3.1).

(6) Ratio of the restricted mean of red misfits to blue actives, in logarithmic units.

significantly different. This indicates that the molecular gas will be depleted more slowly in red misfits than blue actives, but the total gas reservoirs deplete at nearly the same rates. From the other panels of Fig. 3 there are no clear trends in the ratios of t_{mol} or t_{gas} .

Finally, we show the dust-to-stellar mass ratios $M_{\text{dust}}/M_{\star}$ and the dust-to-gas ratios

$$\text{DGR} \equiv \frac{M_{\text{dust}}}{M_{\text{gas}}}, \quad (12)$$

in Fig. 4. The $M_{\text{dust}}/M_{\star}$ distributions (top row of Fig. 4) are different based on a KS test (Table 3), with the blue actives having significantly higher values, while the DGR distributions do not show a significant difference. This finding is also supported by the fact that red misfits have significantly lower restricted mean $M_{\text{dust}}/M_{\star}$ value than blue actives. Red misfits have a smaller restricted mean DGR than blue actives, but this is not statistically significant (only 2σ). In the SFMS plane (middle panels of Fig. 4), aside from a few bins with small number of galaxies in them, red misfits tend to have lower $M_{\text{dust}}/M_{\star}$.

in all areas of this plane. Dust-to-gas ratios also do not show any strong differences between red misfits and blue actives in this plane. These results indicate that red misfits contain *less* dust than blue active galaxies, rather than more dust as one might initially expect based on their red optical colours.

3.2 Scaling relations

A key question in understanding the evolution of star-forming galaxies is what drives the scatter about the SFMS. Recently, assessing the relative importance of gas depletion time and gas mass fraction in driving the scatter about the SFMS has been a major focus (see e.g. Saintonge et al. 2016; Lin et al. 2019; Ellison et al. 2020; Feldmann 2020; Sánchez et al. 2021). Here, we explore whether there are differences in how depletion times and gas mass fractions of red misfits and blue actives correlate with offset from the SFMS. To answer these questions, we plot $\log t_{\text{mol}}$, $\log t_{\text{gas}}$, $\log M_{\text{mol}}/M_*$, and $\log M_{\text{gas}}/M_*$ versus offset from the SFMS. The offset from the star-forming main sequence is defined as

$$\Delta\text{MS} = \log \text{SFR}(M_*) - \log \text{SFR}_{\text{MS}}(M_*), \quad (13)$$

where $\log \text{SFR}(M_*)$ is the SFR of a galaxy with stellar mass M_* and $\log \text{SFR}_{\text{MS}}(M_*)$ is the star-forming main sequence (Popesso et al. 2019) at the same stellar mass

$$\log \text{SFR}_{\text{MS}}(M_*) = 0.38 \log M_* - 3.83. \quad (14)$$

We adopt this particular definition of the SFMS because it was derived from the same SFR and M_* measurements that we use here, namely those from the GSWLC-M2 catalog.

In Fig. 5, we show molecular gas t_{mol} (left column) and total gas t_{gas} depletion times (right column) versus ΔMS . The restricted mean and standard error (see Section 3.1) of each quantity is computed in bins of ΔMS . As a test of our method, in the top row we compare our relations with those from Feldmann (2020), which shows good agreement. There are some notable differences between their study and ours: in Feldmann (2020) the xCOLD GASS sample was used, whereas here we are using a larger sample and a slightly different definition of the SFMS; they took non-detections into account using a method that is different than ours (LeoPy; Feldmann 2019). We also compare our $\log t_{\text{mol}}-\Delta\text{MS}$ relationship using the average of detections only with the relationship found by Tacconi et al. (2018) with the IRAM Plateau de Bure high- z blue sequence CO(3–2) survey (PHIBSS; Tacconi et al. 2013), who did not incorporate non-detections. Although their sample is notably different than ours in terms of redshift ($z \sim 1$ to $z \sim 0$ in our work), this plot shows that our results are in good agreement with theirs. Overall, these relationships show that as galaxies move from above to below the main sequence their gas is used up more slowly (e.g. t_{gas} increases).

Having confirmed that our results for the sample as a whole agree with previous work, we move on to study these trends for red misfits and blue actives separately in the middle and bottom rows of Fig. 5. In the middle row, we see that both populations follow similar trends to the population as a whole; however, red misfits have higher t_{mol} values than blue actives below the SFMS and up to $\Delta\text{MS} \sim 0.2$ dex. This is in line with our previous result in Table 3, which showed that red misfits have longer t_{mol} and similar t_{gas} compared to blue actives. Here, however, we see that this difference is primarily coming from galaxies on and below the main sequence. In the bottom row, we show the same as the middle row except only galaxies with $10 \leq \log M_* \leq 11$, which is where both red misfits and blue actives are well-sampled. The red and blue points are closer together than in the middle row. This result indicates that the differences seen in the

left-hand panel of the second row are largely coming from galaxies outside of this stellar mass range.

In Fig. 6, we show molecular gas mass fractions M_{mol}/M_* and total gas M_{gas}/M_* mass fractions versus ΔMS . We use the same survival analysis approach as above to take non-detections into account in each ΔMS bin. Note that we do not have curves from the literature to show for comparison. The top row shows that M_{mol}/M_* increases as galaxies move from below to above the main sequence, while M_{gas}/M_* increases with increasing ΔMS up to $\Delta\text{MS} \sim -0.5$ dex and then remains constant at ~ 0.5 dex. In the middle row, we see that the trends for red misfits and blue actives are different, especially below the main sequence. Red misfits have significantly lower M_{mol}/M_* and M_{gas}/M_* relative to blue active galaxies, although the differences become less significant on and above the main sequence. This indicates that red misfits are quite gas-poor, despite their relatively similar gas depletion times compared to blue actives (Fig. 3). This result is echoed by the comparison of restricted means of these properties for red misfits and blue actives altogether (Table 3): red misfits have significantly lower gas mass fractions than blue actives. In the bottom row of Fig. 6, we show the same as the middle row but only for galaxies with $10 \leq \log M_* \leq 11$, which is where both red misfits and blue actives are well-sampled. Relative differences between red misfits and blue actives decrease slightly, indicating that controlling for stellar mass reduces differences between the populations. We explore this further in Section 3.4.

3.3 The log SFR–log M_{mol} relationship, and the molecular gas main sequence

Stellar mass, SFR, and molecular gas are correlated with each other, as shown by the SFMS and the Kennicutt–Schmidt relation (SFR surface density versus cold gas surface density; Kennicutt 1989; Kennicutt et al. 2007; Bigiel et al. 2008; Leroy et al. 2008, 2013). Recent work has introduced the ‘molecular gas main sequence’ (MGMS; Σ_{mol} versus Σ_*) as a companion to the aforementioned relationships (e.g. Lin et al. 2019). By simultaneously examining these three correlations one can gain insight into the physical mechanisms that lead to the star formation main sequence. Here, we compare the MGMS and log SFR–log M_{mol} relationships of red misfit and blue active galaxies (Fig. 7). Each of these relationships shows strong correlations (Pearson- r of detections ranging from 0.68 to 0.86). We used `linmix` (Kelly 2007) to fit lines to each of these plots, taking uncertainties in both variables and upper limits in M_{mol} into account. For the log SFR–log M_{mol} relation, the fits were done with log SFR on the x -axis in order to include upper limits, and the best-fitting equations were inverted to match how this relationship is usually shown with gas on the x -axis and SFR on the y -axis.

In the top row of Fig. 7, by comparing the points and fits with lines of constant M_{mol}/M_* (dotted lines), we see that red misfits have lower molecular fractions (between 1 and 10 per cent) than blue actives (mostly around 10 per cent). The fit to blue actives is nearly linear (similar to spatially resolved work e.g. Lin et al. 2019), and the intercept is close to the value of the restricted mean from Table 3 (dashed line). Red misfits, however, show a significantly shallower (sub-linear) slope than blue actives (a difference of $\sim 6\sigma$), causing the best-fitting intercept to differ more significantly from the restricted mean from Table 3 (dashed line). The difference between these populations is most striking at larger stellar masses. The shallower slope in the MGMS for red misfits suggests physical difference between these two populations – the molecular gas content of red misfits is lower than that of blue actives at fixed stellar mass, but only at high stellar masses. Additionally, the correlation between log M_{mol}

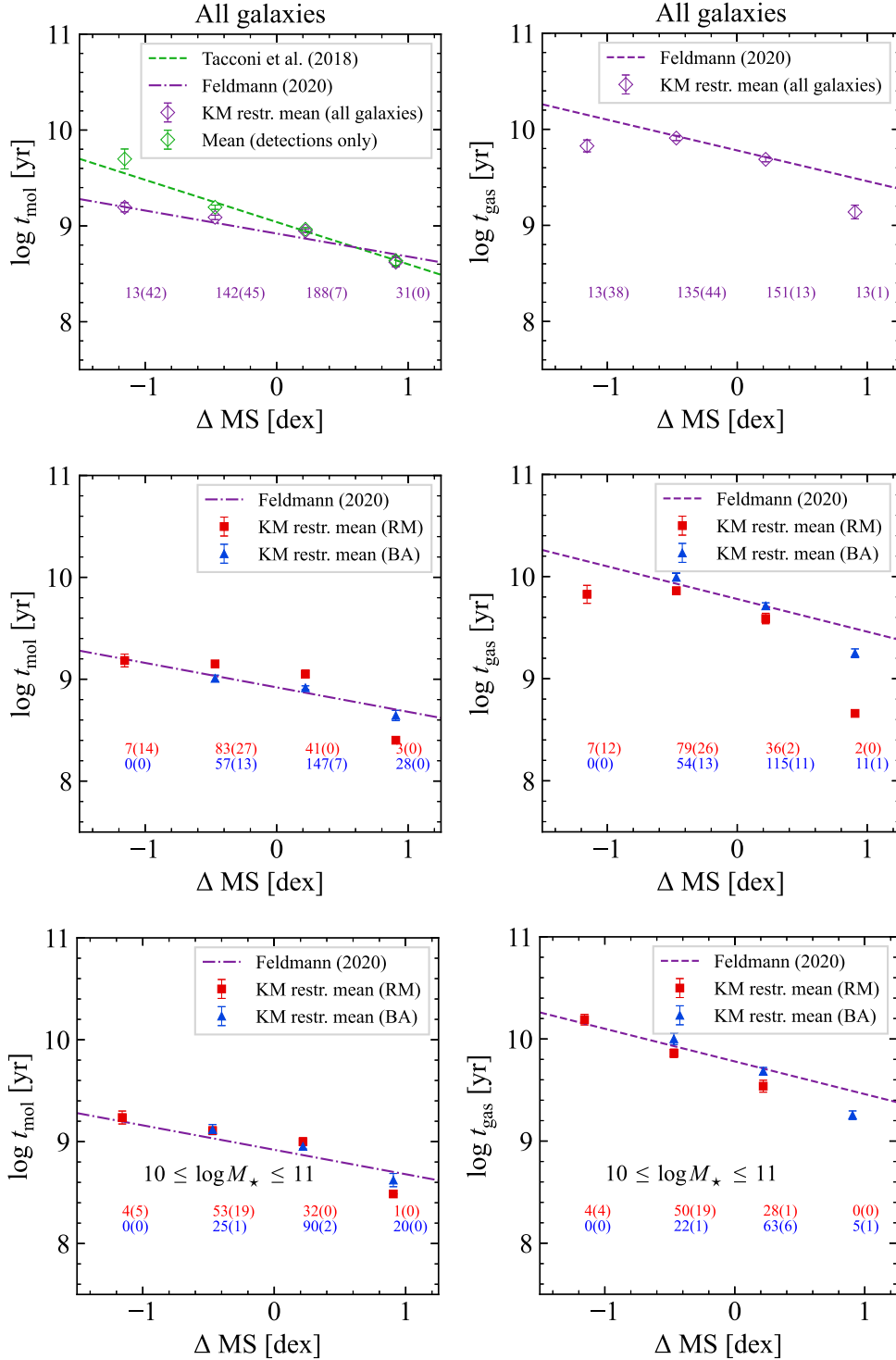


Figure 5. Molecular (left) and total (right) gas depletion times versus offset from the star-forming main sequence (equation 13). The purple dotted–dashed lines are from Feldmann (2020), which are fits to molecular (left) and total (right) gas depletion times from xCOLD GASS (+ xGASS) after taking non-detections into account. The green dashed line is the fit to $z = 0 - 4$ galaxies from the PHIBSS survey (Tacconi et al. 2018) for comparison. The Kaplan–Meier median is the value where the cumulative distribution reaches 0.5. The top row shows all galaxies in our sample. One can see that our results agree well with Tacconi et al. (2018) when we ignore non-detections and with Feldmann (2020) when we include non-detections. The middle row shows red misfits and blue actives separately. The number of detections are shown below each data point, with the number of non-detections shown in brackets. One can see that, when there is enough data to compute a median (≥ 5 points), the molecular and total gas depletion times of red misfits and blue actives are similar, although slightly (but statistically significantly) larger for red misfits. The bottom row is the same as the middle row except only showing galaxies with stellar masses between 10^{10} and $10^{11} M_{\odot}$, where both depletion times become more similar between red misfits and blue actives.

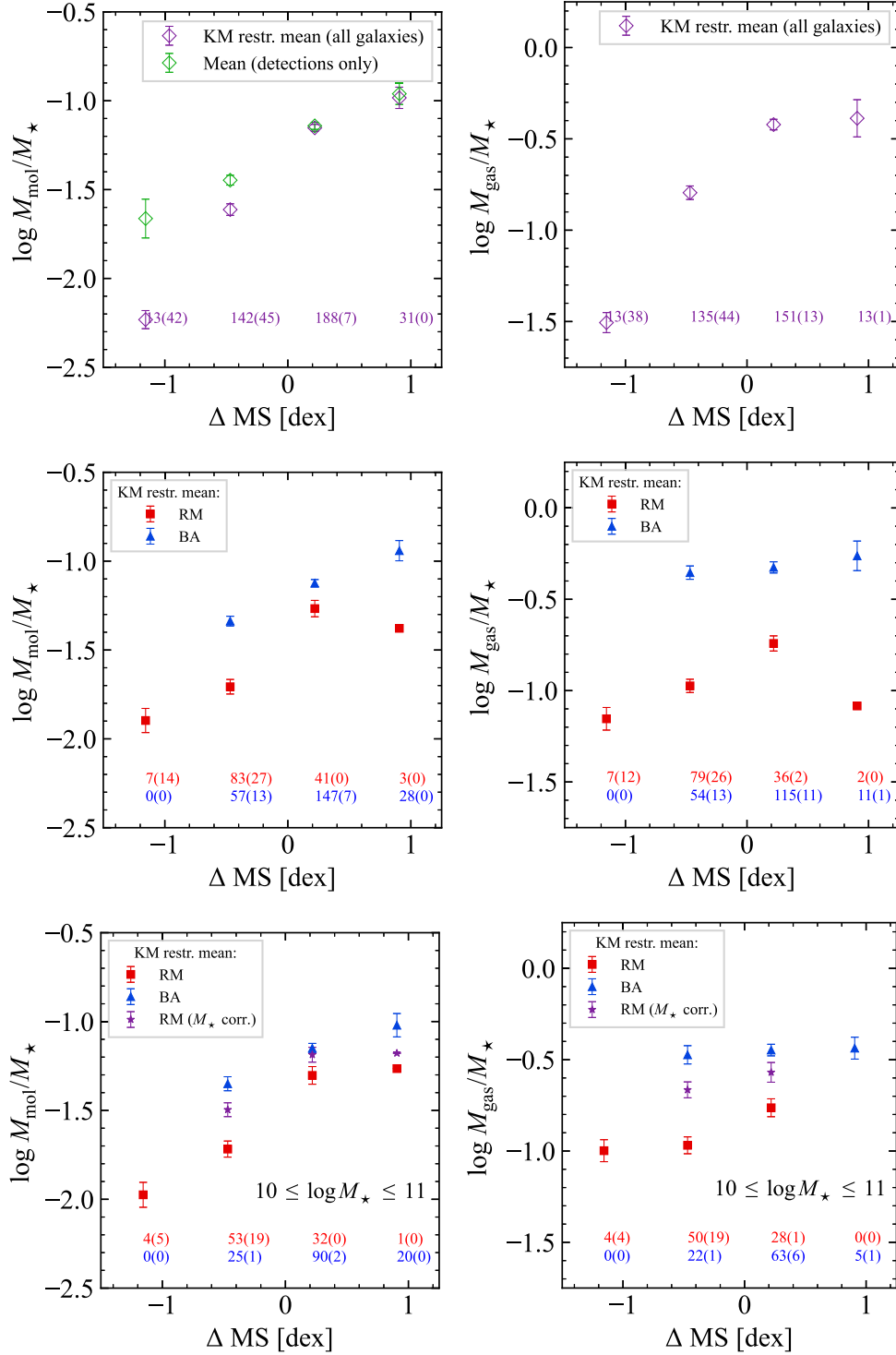


Figure 6. Molecular (left) and total (right) gas mass fractions versus offset from the star-forming main sequence (equation 13). The top row shows all galaxies in our sample. One can see that across the main sequence, molecular gas mass fraction rises, while the total gas mass fraction rises and then remains constant. The middle row shows red misfits and blue actives separately. The number of detections are shown below each data point, with the number of non-detections shown in brackets. One can see that, in contrast to Fig. 5, red misfits tend to have lower molecular and total gas fractions below and on the main sequence. The bottom row is the same as the middle row except only showing galaxies with stellar masses between 10^{10} and $10^{11} M_{\odot}$, where both gas mass fractions change slightly but remain significantly different between red misfits and blue actives. The ‘ $M_{\star}$ corr.’ points in the bottom row are gas fractions of red misfits, corrected for the difference in median stellar mass between red misfits and blue actives in each ΔMS bin (see Section 3.4).

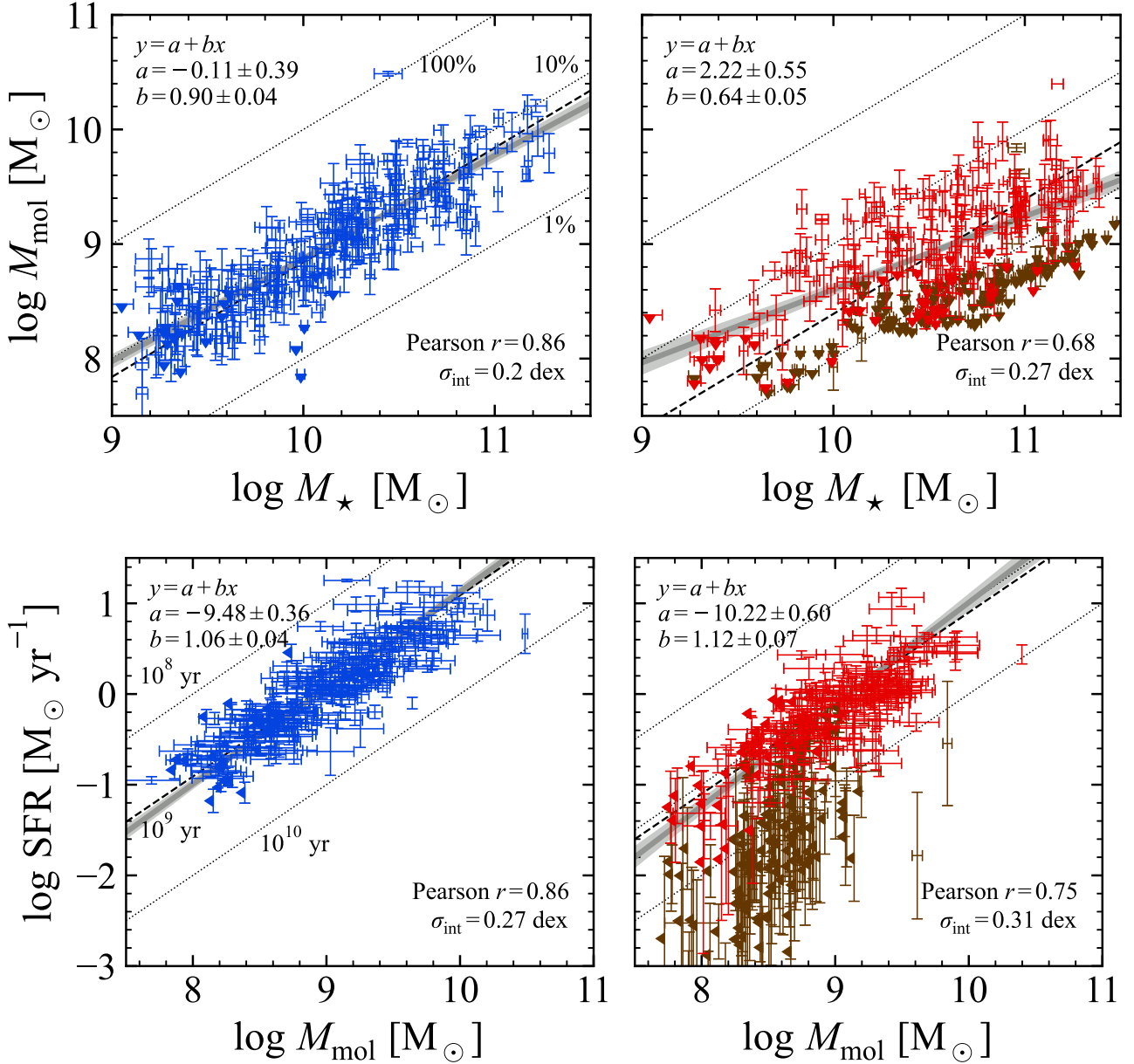


Figure 7. The molecular gas main sequence (top) and log SFR–log M_{mol} relationship with molecular gas (bottom) for blue active (left), red misfit galaxies (red points on the right), and red-and-dead galaxies (brown points on the right). The linear fits in all panels were performed using `linmix` (Kelly 2007) which incorporates uncertainties in x and y , and upper limits in y . In each panel, the intrinsic scatter σ_{int} (see appendix C of Chown et al. 2021) is indicated. In the bottom row the fits were done with log SFR on the x -axis, and the best-fitting equations were inverted to match how this relationship is usually shown with gas on the x -axis and SFR on the y -axis. In the top row, the dotted lines represent constant $M_{\text{mol}}/M_{\star}$ (1, 10, and 100 per cent), and the dashed lines correspond to the restricted mean $M_{\text{mol}}/M_{\star}$ for blue actives and red misfits from Table 3. The bottom rows show lines of constant t_{mol} , and the restricted means from Table 3. These plots show that red misfits tend to have lower $M_{\text{H}_2}/M_{\star}$ fractions, and slightly longer t_{mol} . The slope of the MGMS is significantly flatter for red misfits than for blue actives. Red and dead galaxies (brown points) are shown for comparison in the right-hand panels. Although most of their gas masses are upper limits, it is apparent that red and dead galaxies have longer gas depletion times and lower gas mass fractions than red misfits.

and $\log M_{\star}$ is significantly weaker for red misfits ($r = 0.68$) than for blue actives ($r = 0.86$).

In the bottom row of Fig. 7, by comparing the data points and linear fits with lines of constant t_{mol} , we see that red misfits have slightly longer t_{mol} than blue actives (echoing our earlier results). In contrast to the MGMS plots, the slope of the red misfit and blue active log SFR–log M_{mol} relations are not significantly different (agree within $\sim 1\sigma$). The log SFR–log M_{mol} slopes of red misfits

and blue actives are both slightly superlinear. This near-linearity results in the best-fitting intercepts being close to the restricted means from Table 3 (dashed lines). The similarity of the gas depletion time relationships, combined with the MGMS results, and the fact that red misfits tend to lie below the SFMS (e.g. Fig. 2) suggests that red misfits have lower than average star formation rates due to a lack of molecular gas rather than inefficient star formation.

3.4 Impact of M_* differences between red misfits and blue actives on cold gas scaling relations

The M_{mol}/M_* (and M_{gas}/M_*) scaling relations for red misfits and blue actives come closer together when we restrict the stellar mass range (Section 3.2). We investigate how much of the residual difference between the red and blue points in the bottom row of Fig. 6 can be explained by differences in the median stellar mass of blue actives and red misfits within each bin.

We correct the M_{mol}/M_* measurements of red misfits in a given ΔMS bin as follows. First, we assume that the two following empirical relationships for red misfits hold:

$$\log \text{SFR} = a_1 + b_1 \log M_*, \quad (15)$$

where $a_1 = -5.25$ and $b_1 = 0.48$ (middle panel of Fig. C1), and

$$\log \text{SFR} = a_2 + b_2 \log M_{\text{mol}}, \quad (16)$$

where $a_2 = -10.22$ and $b_2 = 1.12$ (bottom right panel of Fig. 7).

Consider a red misfit galaxy with stellar mass $\log M_{*, \text{RM}}$, molecular gas mass $\log M_{\text{mol, RM}}$ lying in ΔMS bin i . Let

$$\Delta \log M_* \equiv \log M_{*, \text{RM}} - \text{med}(\log M_*)_{\text{BA}, i}, \quad (17)$$

where $\text{med}(\log M_*)_{\text{BA}, i}$ is the median $\log M_*$ of blue actives in bin ΔMS bin i . We can estimate the molecular gas mass that this red misfit would have if its stellar mass was equal to $\text{med}(\log M_*)_{\text{BA}, i}$, by plugging the change in $\log \text{SFR}$ in equation (15) corresponding to $\Delta \log M_*$ into equation (16), and solving for the change in molecular gas mass

$$\Delta \log M_{\text{mol}} = \frac{b_1}{b_2} \Delta \log M_*. \quad (18)$$

The corrected $\log M_{\text{mol}}$ is then

$$\log M_{\text{mol, RM, corr.}} = \log M_{\text{mol, RM}} + \frac{b_1}{b_2} \Delta \log M_*. \quad (19)$$

We correct the $\log M_{\text{mol}}/M_*$ measurements of each red misfit in this way, and recompute the restricted mean $\log M_{\text{mol}}/M_*$, shown as the purple stars in Fig. 6. This correction for stellar mass differences brings red misfit and blue actives even closer together (bottom left panel of Fig. 6).

For the total gas scaling relations, we correct the molecular gas masses as outlined above, and we also correct the H I mass fractions such that the full sample relation between $\log M_{\text{HI}}$ and $\log M_*$ from Brown et al. (2015) holds (see their figs 4 and 5, where they find a slope of -0.85 for this relation). The corrected red misfit $\log M_{\text{gas}}/M_*$ points are shown in the bottom right panel of Fig. 6. This correction brings the red misfit and blue active $\log M_{\text{gas}}/M_*$ scaling relations even closer together, but red misfits still have lower fractions around $\Delta\text{MS} \sim 0$ dex.

4 DISCUSSION

Our findings show that red misfits and blue actives have different molecular and total gas mass fractions, different dust mass fractions, and slightly different molecular gas depletion times, but similar total gas depletion times, and similar dust-to-gas ratios. We showed that red misfits have lower M_{mol}/M_* and M_{gas}/M_* ratios than blue actives on average (Section 3.1) and as functions of offset from the main sequence (Section 3.2). We showed that red misfits have a significantly shallower slope than blue actives in the molecular gas main sequence (Section 3.3), and that red misfits and blue actives have consistent $\log \text{SFR}-\log M_{\text{mol}}$ relations (Section 3.3).

We found that the dust content of red misfits is similar (based on the DGR) or lower than (based on M_{dust}/M_*) that of blue actives, which supports the claims from Evans et al. (2018) that the red colours of red misfits are not due to dust reddening. Their red colours are therefore likely due to the presence of old stellar populations. However, $g-r$ colour is not as sensitive to young stellar populations as $u-r$ or $\text{NUV}-r$, and so a red $g-r$ colour does not necessarily indicate a red $u-r$ or $\text{NUV}-r$ colour. Indeed, by definition, red misfits are actively forming stars, and so they must host young stellar populations.

We find that red misfits have lower molecular gas fractions, and even lower total gas fractions, than blue actives, while t_{mol} and t_{gas} of red misfits and blue actives follow similar relationships. Red misfits tend to lie on or below the main sequence while blue actives tend to lie on or above the main sequence. After correcting for different median stellar masses between red misfits and blue actives, their M_{mol}/M_* and M_{gas}/M_* scaling relations become more similar. However, red misfits still tend to have lower total gas content particularly on the main sequence. Taken together, these results suggest that the lower star formation rates of red misfits lying on or near the main sequence are due to bottlenecks in the gas supply rather than reduced star formation efficiency. Our findings that the difference in total gas mass fraction is larger than that of molecular gas mass fraction suggests that the long-term fuel for star formation has been depleted. The fact that the molecular gas mass fraction of red misfits is lowest compared to blue actives at high stellar masses suggests that red misfits have depleted their gas supply by forming stars and are on their way towards the red sequence. Taking all of our findings together with those of Evans et al. (2018), we suggest that red misfits are not a single class of galaxies, but rather a mix of galaxies in different states whose behaviour depends on position relative to the SFMS. However, when we narrow in on galaxies on or slightly below the main sequence, and control for stellar mass biases, red misfits have lower total gas content than blue actives.

One limitation of this work is that we combined several data sets together, and so our sample has a complex selection function. Another limitation is that we only used 850- μm fluxes to estimate dust masses; a more optimal method would be to use infrared-to-submillimeter SED fitting. Unfortunately, our JCMT Semester 18B SCUBA-2 targets were selected from the xCOLD GASS sample and this sample does not overlap significantly with H-ATLAS and so the required infrared data do not exist like they do for JINGLE galaxies. In the interest of using the same method for all galaxies with SCUBA-2 data, we used the Lamperti et al. (2019) scaling relations to estimate a dust temperature and spectral index for each galaxy.

5 CONCLUSIONS

By analysing trends of molecular and total cold gas mass fractions and depletion times, we have found that red misfit and blue active galaxies do not show strong differences in depletion times, but their gas mass fractions are significantly different, and they exhibit significantly different scaling relations with offset from the main sequence and stellar mass. This suggests that red misfits are more limited than blue actives in both their near term and long-term gas supply rather than the rate with which they are turning the gas into stars. This is also likely due to the fact that red misfits below the main sequence tend to be more massive than blue actives. Thus red misfits have about the same amount of gas but are more massive. We also found that the dust-to-stellar ratios of red misfits are lower than that of blue actives, while their dust-to-gas ratios follow similar distributions.

Our results suggest that by selecting galaxies based on optical $g - r$ colour and specific star formation rate simultaneously, high-mass galaxies that are classified as red and star forming (red misfits) are actively quenching after depleting their gas supply through star formation, while red star-forming galaxies with low stellar masses either had limited gas supply to begin with or had their gas removed prematurely (e.g. due to environmental effects such as ram pressure stripping).

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DATA AVAILABILITY

The data underlying this article are available in the article (Table A1, Table B1), public domain sources, and a private catalogue (JINGLE).

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APPENDIX A: NEW JCMT CO(2–1) MEASUREMENTS OF RED MISFITS

Here we present the CO(2–1) measurements of red misfits using JCMT (Table A1).

Table A1. New JCMT CO(2–1) measurements of red misfits selected from the JINGLE sample.

ObjID	RA (J2000)	Dec. (J2000)	z	$\log M_*$	$\log \text{SFR}$	$\log L_{\text{CO}(1-0)}$	$\log M_{\text{mol}}$
(1)	deg	deg	(4)	$M_{\odot}$	$M_{\odot} \text{ yr}^{-1}$	$\text{K km s}^{-1} \text{ pc}^2$	$M_{\odot}$
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
1237665024374865968	203.305	33.110	0.0240	10.78	0.25	9.40 ± 0.05	10.04 ± 0.68
1237650762924621828	173.539	−1.595	0.0230	10.45	0.29	9.29 ± 0.06	9.93 ± 0.70
1237654604239274286	131.119	2.064	0.0250	10.32	0.51	9.22 ± 0.07	9.86 ± 0.71
1237650372092035132	178.878	−1.261	0.0190	10.08	0.34	8.73 ± 0.12	9.37 ± 0.75
1237650761854222501	181.212	−2.438	0.0200	10.21	−0.08	8.55 ± 0.12	9.19 ± 0.75
1237665126939295886	201.313	32.671	0.0400	10.82	0.52	<9.05	<9.69
1237671128051220863	175.912	−1.647	0.0430	10.46	0.60	<9.03	<9.67
1237648720695001218	182.575	−0.518	0.0350	11.02	0.27	<8.94	<9.57
1237648703516115092	212.612	−0.832	0.0250	10.71	0.59	<8.65	<9.29
1237648705663664205	212.740	1.036	0.0250	10.58	0.24	<8.64	<9.28
1237648722820661612	132.804	1.062	0.0270	9.94	−0.10	<8.55	<9.19
1237650762927308812	179.693	−1.466	0.0210	9.72	0.03	<8.49	<9.13

Notes. (1) *SDSS* photometric identification number.

(5) Stellar mass from the GSWLC-M2 or A2 catalogue (if unavailable in M2).

(6) SFR from the GSWLC-M2 or A2 catalogue (if unavailable in M2).

(7) Measured CO(1–0) luminosity (converted from 2–1 assuming $r_{21} = 0.7$).

(8) Measured molecular gas mass assuming $\alpha_{\text{CO}} = 4.35$.

APPENDIX B: NEW SCUBA-2 MEASUREMENTS OF RED MISFITS

Here, we present the SCUBA-2 850- μm measurements of red misfits using JCMT (Table B1).

Table B1. New SCUBA-2 measurements of red misfits selected from the xCOLD GASS sample.

Name	d_L Mpc	T K	β	Det?	r_{ap} arcsec	r_{90} arcsec	$S_{850\ \mu\text{m}}(r \leq r_{\text{ap}})$ mJy	$\log M_{\text{dust}}(r \leq r_{\text{ap}})$ $M_{\odot}$
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)
J142720.13 + 025018.1	115.53	22.64	1.94	Y	20.23	15.51	25.37 ± 3.95	8.18 ± 0.07
J104402.21 + 043946.8	116.29	22.25	1.65	Y	20.06	15.28	16.46 ± 2.90	7.95 ± 0.08
J101638.39 + 123438.5	138.88	22.14	1.86	Y	28.24	25.07	25.21 ± 4.36	8.34 ± 0.07
J100530.26 + 054019.4	196.55	20.16	1.94	Y	21.56	17.20	16.30 ± 2.51	8.53 ± 0.07
J095144.91 + 353719.6	117.99	22.93	1.93	Y	19.50	14.53	19.72 ± 3.98	8.08 ± 0.09
J235644.47 + 135435.4	159.92	22.81	1.82	Y	20.77	16.20	14.29 ± 2.98	8.19 ± 0.09
J105315.29 + 042003.1	184.23	22.54	1.91	Y	15.96	9.27	19.18 ± 3.68	8.46 ± 0.08
J100216.28 + 191256.3	71.90	21.88	1.92	Y	19.38	14.38	13.77 ± 3.51	7.52 ± 0.11
J080442.30 + 154632.6	128.22	21.32	1.86	Y	18.62	13.33	9.98 ± 2.94	7.89 ± 0.13
J112311.63 + 130703.7	208.32	22.26	1.90	Y	18.13	12.64	14.51 ± 4.23	8.45 ± 0.13
J094419.42 + 095905.1	44.09	20.38	2.03	N	16.58	10.28	<9.82	<7.02
J090923.67 + 223050.1	64.52	21.15	2.01	N	18.17	12.69	<4.15	<6.95
J135845.41 + 203942.7	69.95	22.29	2.00	N	16.56	10.26	<8.18	<7.28
J232326.53 + 152510.4	189.11	22.87	1.96	N	14.47	6.35	<10.56	<8.23
J151604.47 + 065051.4	162.09	22.99	1.93	N	21.02	16.52	<10.37	<8.08
J093953.62 + 034850.2	124.82	21.73	1.88	N	19.80	14.94	<3.89	<7.45
J104251.39 + 055135.5	146.88	20.90	1.94	N	23.48	19.55	<7.69	<7.92
J122006.47 + 100429.2	191.89	22.13	1.92	N	18.93	13.76	<6.80	<8.06
J152747.42 + 093729.6	136.77	21.57	2.02	N	17.33	11.46	<6.99	<7.82
J131934.30 + 102717.5	213.27	21.20	1.93	N	22.68	18.58	<11.19	<8.40
J102508.93 + 133605.1	80.82	21.56	1.72	N	19.07	13.96	<7.17	<7.30
J142846.66 + 271502.4	63.83	24.29	1.90	N	16.06	9.42	<7.62	<7.09
J021219.38 + 133645.6	182.82	22.01	1.89	N	16.66	10.42	<4.80	<7.87
J130035.67 + 273427.2	73.74	21.76	1.81	N	23.86	20.01	<6.54	<7.20
J001947.33 + 003526.7	76.84	22.52	1.70	N	17.88	12.28	<6.84	<7.21
J020359.14 + 141837.3	189.02	23.00	1.91	N	22.37	18.20	<12.11	<8.28
J130525.44 + 035929.7	193.31	22.13	1.87	N	15.77	8.93	<4.17	<7.85
J095439.45 + 092640.7	151.93	21.50	1.97	N	18.41	13.04	<6.71	<7.88
J111738.91 + 263506.0	210.84	22.38	2.07	N	13.61	4.03	<7.26	<8.20
J231816.95 + 133426.6	174.37	22.52	1.84	N	15.15	7.78	<5.19	<7.83
J011716.09 + 143720.5	167.65	19.62	1.98	N	15.68	8.76	<4.66	<7.87
J150926.10 + 101718.3	120.84	21.84	1.78	N	20.59	15.96	<9.99	<7.81
J150204.10 + 064922.9	204.97	21.47	1.75	N	16.90	10.79	<5.99	<8.04

Notes. (1) SDSS ID as shown in the xCOLD GASS catalogue.

(2) Luminosity distance.

(3) Dust temperature estimated using equation (11).

(4) Modified blackbody spectral index estimated using equation (10).

(5) Flag for whether this galaxy is classified as a detection or not.

(6) Aperture radius over which the 850 μm flux density was measured.

(7) SDSS r -band 90 per cent Petrosian radius.

(8) 850 μm flux density within r_{ap} .

(9) Dust mass within r_{ap} computed using equation (9).

APPENDIX C: LOG SFR VERSUS LOG M_* FOR RED MISFITS AND BLUE ACTIVES

Here, we present fits log SFR versus log M_* for red misfits and blue actives, and the whole population (Fig. C1). The solid black lines in

each panel are the best-fitting relation from Popesso et al. (2019). Blue actives and red misfits on their own do not follow the Popesso et al. (2019) relation, but the fit to both populations combined is consistent with Popesso et al. (2019). The slope of the fit to red misfits is used to correct gas mass fractions in Section 3.4.

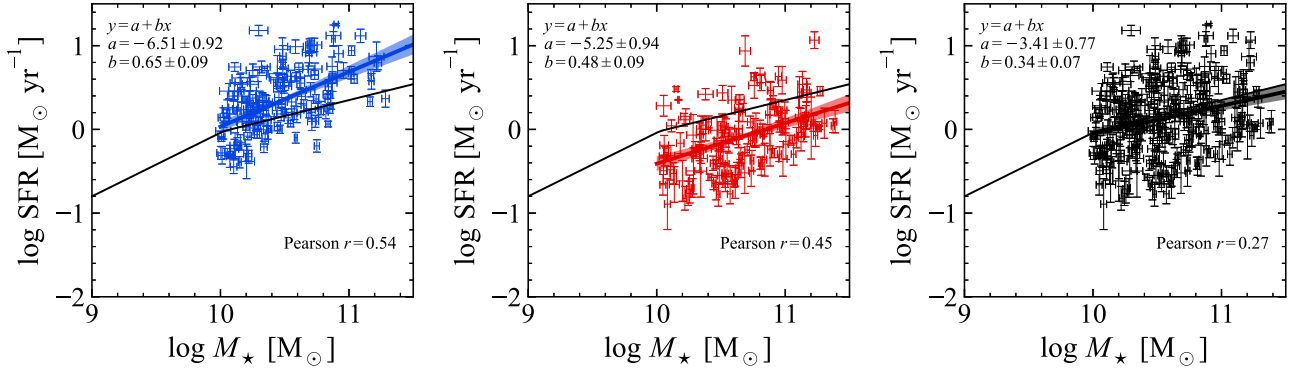


Figure C1. Linear fits to $\log \text{SFR}$ versus $\log M_{\star}$ for blue actives (left), red misfits (middle), and both populations combined (right). The fits were performed using LinMix, taking uncertainties into account. The solid black lines in each panel are the best-fitting relation from Popesso et al. (2019).

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