



## Euclid Preparation. XIV. The Complete Calibration of the Color–Redshift Relation (C3R2) Survey: Data Release 3

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## Abstract

The Complete Calibration of the Color–Redshift Relation (C3R2) survey is obtaining spectroscopic redshifts in order to map the relation between galaxy color and redshift to a depth of  $i \sim 24.5$  (AB). The primary goal is to enable sufficiently accurate photometric redshifts for Stage IV dark energy projects, particularly Euclid and the Nancy Grace Roman Space Telescope (Roman), which are designed to constrain cosmological parameters through weak lensing. We present 676 new high-confidence spectroscopic redshifts obtained by the C3R2 survey in the 2017B–2019B semesters using the DEIMOS, LRIS, and MOSFIRE multiobject spectrographs on the Keck telescopes. Combined with the 4454 redshifts previously published by this project, the C3R2 survey has now

obtained and published 5130 high-quality galaxy spectra and redshifts. If we restrict consideration to only the  $0.2 < z_p < 2.6$  range of interest for the Euclid cosmological goals, then with the current data release, C3R2 has increased the spectroscopic redshift coverage of the Euclid color space from 51% (as reported by Masters et al.) to the current 91%. Once completed and combined with extensive data collected by other spectroscopic surveys, C3R2 should provide the spectroscopic calibration set needed to enable photometric redshifts to meet the cosmology requirements for Euclid, and make significant headway toward solving the problem for Roman.

*Unified Astronomy Thesaurus concepts:* [Galaxy spectroscopy \(2171\)](#)

*Supporting material:* machine-readable table

## 1. Introduction

The Stage IV cosmology projects, such as the Vera C. Rubin Observatory Legacy Survey of Space and Time (hereafter LSST), Euclid, and the Nancy Grace Roman Space Telescope (hereafter Roman), will use deep imaging in multiple optical and near-infrared (near-IR) filters to calculate photometric redshifts ( $z_p$ ) of large numbers of galaxies for weak lensing cosmology. While high-quality photometric redshifts are crucial for many other investigations, the photometric redshift requirements for cosmology are the most exacting. Weak lensing cosmology requires *unbiased* redshift estimates (Huterer et al. 2006; Ma et al. 2006), which are very challenging for photometric redshift estimation techniques. The requirements for Euclid and Roman are that the mean redshift ( $\langle z \rangle$ ) of galaxies in  $\sim 10$  tomographic redshift bins for cosmic shear measurements must be known to better than 0.2% (that is,  $\Delta\langle z \rangle \leq 0.002(1 + \langle z \rangle)$ ) for each bin in the  $0.2 < z < 2.6$  range. Numerous tests have shown that this level of accuracy is not currently feasible for photometric redshift algorithms and realistic training samples (Desprez et al. 2020). The bias requirement, in particular, makes spectroscopic calibration samples necessary for the success of these missions (Newman et al. 2013).

Stage III cosmology surveys currently underway also face the photometric redshift estimation challenge, and constitute an important testbed for the photometric redshift techniques to be employed in the Stage IV projects. These include the Kilo-Degree Survey (de Jong et al. 2015, 2017; Hildebrandt et al. 2017), the Dark Energy Survey (DES; DES Collaboration et al. 2018; Troxel et al. 2018), and the Hyper-Suprime Cam Survey (Aihara et al. 2018). A variety of techniques have been employed in these surveys to constrain redshift distributions,  $N(z)$ , of galaxies in shear bins—such as clustering redshifts, which use the spatial distribution of overlapping spectroscopic redshift samples to infer the  $N(z)$  distribution of the photometric sample (e.g., Newman 2008; Ménard et al. 2013; McQuinn & White 2013; Schmidt et al. 2013; Morrison et al. 2017), template fitting (e.g., Benítez 2000; Brammer et al. 2008; Ilbert et al. 2009), and machine learning based on training samples (e.g., Collister & Lahav 2004; Carrasco Kind & Brunner 2013). Re-weighting of the spectroscopic sample to better match the photometric sample, as described in Lima et al. (2008), has also been used. It is generally agreed that a key limitation for the empirical photo- $z$  and re-weighting techniques is the need for a fully representative calibration sample of spectroscopic redshifts ( $z_s$ ) that explores the range of galaxy properties present in the surveys.

The Complete Calibration of the Color-Redshift Relation (C3R2) survey (Masters et al. 2017, hereafter M17 [DR1]) was initiated in response to this need, with the goal of mapping the empirical relation between galaxy redshift and color to the

Euclid depth of  $i \sim 24.5$  (AB). The survey strategy is based on the fact that the observed galaxy colors down to a given survey depth are both *limited* and *measurable*. Moreover, a well-defined, mostly nondegenerate relation between a galaxy’s position in multicolor space ( $C$ ) and its redshift ( $z$ ) can be discovered empirically. Elucidating this  $P(z|C)$  relation is the goal of the C3R2 survey.

The C3R2 survey strategy follows the method outlined in Masters et al. 2015 (hereafter M15). M15 illustrated the use of an unsupervised manifold learning algorithm, the *self-organizing map* (SOM; Kohonen 1982), to map the color distribution of galaxies in the high-dimensional color space ( $u-g, g-r, \dots, J-H$ ) anticipated for Euclid and Roman photometric redshift estimation. The SOM is a neural network model widely used to map and identify correlations in high-dimensional data. The algorithm uses unsupervised, competitive learning of “neurons” to project high-dimensional data onto a lower-dimensional grid. The SOM algorithm can be thought of as a type of nonlinear principal component analysis, and is also similar in some respects to the  $k$ -means clustering algorithm. In contrast to these and other methods, the SOM preserves the topology of the high-dimensional data in the low-dimension representation. Similar objects are thus grouped together on the self-organized map, and clusters that exist in the high-dimensional data space are reflected in the lower-dimensional representation. This feature makes the maps visually understandable and thus useful for identifying correlations that exist in high-dimensional data. This high-dimensional mapping allows us to directly determine which parts of galaxy color space are well sampled with existing spectroscopy and which are not, thus letting us focus spectroscopic calibration effort on those regions that are the least constrained. We have focused on obtaining only the necessary additional spectroscopy needed to build a representative calibration sample, such that direct inference of the  $P(z|C)$  relation can be made sufficiently accurate to meet the cosmology requirements.

In M17 and Masters et al. 2019 (hereafter M19, i.e., DR2), we presented the results of the C3R2 Keck observations taken in the 2016A, 2016B, and 2017A semesters. Here we present results obtained in the 2017B, 2018A, 2018B, and 2019B semesters, comprising 18 Keck nights (some of which were shared, or half-nights), of which  $\sim 11$  had good observing conditions. The new data come from nights allocated by NASA to PI D. Stern (two nights) and to PI D. Masters (16 nights). We refer the reader to M15 for background on the calibration approach, and to M17 and M19 for the observation and data reduction details of the C3R2 survey. Spectroscopic redshifts have also been obtained using the VLT as part of the C3R2 effort (Guglielmo et al. 2020).

This paper is structured as follows. In Section 2 we describe the observations and data reduction. The observations were all conducted in the fields listed in Table 1. In Section 3 we

**Table 1**  
Overview of the Deep Fields Targeted by C3R2 in DR3

Field	R.A. (J2000)	Decl. (J2000)	Area (deg <sup>2</sup> )	Optical Data (ugriz)	Near-IR Data (YJHK <sub>s</sub> )
VVDS-2h	02 <sup>h</sup> 26 <sup>m</sup>	−04°30′	1.0	CFHTLS	VISTA
COSMOS	10 <sup>h</sup> 00 <sup>m</sup>	+02°12′	2.0	CFHTLS	VISTA
EGS	14 <sup>h</sup> 19 <sup>m</sup>	+52°41′	1.0	CFHTLS	CFHTLS- WIRDS <sup>a</sup>

**Note.**

<sup>a</sup> Y-band obtained from CFHT-WIRCAM observations separate from the WIRDS survey.

present our spectroscopic results from 19 nights of Keck time in semesters 2017B–2020A, discuss the performance of the method, and investigate the status of the calibration effort and the issues still to be addressed.

## 2. Observations and Data Reductions

The observations were carried out at the Keck Observatory located on Maunakea in Hawaii. Both of the two 10 m telescopes were used, as the Low Resolution Imaging Spectrometer (LRIS) and Multi-Object Spectrometer For Infra-Red Exploration (MOSFIRE) are located on Keck I and the DEep Imaging Multi-Object Spectrograph (DEIMOS) on Keck II. The observing nights and weather conditions are summarized in Table 2. Note that approximately six of the 19 nights were lost to bad weather. Table 3 summarizes the observed slit masks.

### 2.1. Description of Observations

Observations for DR3 were carried out essentially in the same manner as described in M17 and M19 for the previous data releases. As discussed in more detail in M17 and M19, we estimated the exposure times needed to obtain redshifts for each target with the three main instruments available at Keck, i.e., DEIMOS, LRIS, and MOSFIRE, and assigned targets to the instrument with the lowest expected exposure time. Here we give a brief overview of the instruments and methods used for the observations.

#### 2.1.1. Deep Imaging Multi-object Spectrograph

DEIMOS (Faber et al. 2003) is a wide-field optical spectrograph that covers a roughly rectangular field of view measuring  $16.5 \times 5$  arcmin<sup>2</sup>. Custom slit masks typically target  $\sim 100$  galaxies at the same time, providing spectra from 5000 Å up to 1 μm at spectral resolution  $R \equiv \lambda/\Delta\lambda \sim 3000$ . We used the 600 groove mm<sup>−1</sup> grating blazed at 7200 Å and the GG400 blocking filter, with dithering performed to improve sky subtraction. We use a minimum slit length of 8'' as a balance between maximizing the number of targets on the mask and getting good sky measurements. Data were reduced using a modified version of the DEEP2 pipeline designed to deal with dithered data.

#### 2.1.2. Multi-object Spectrograph for Infrared Exploration

MOSFIRE (McLean et al. 2012) is a near-IR spectrograph that can observe up to 46 objects at the same time over a  $6.1 \times 6.1$  arcmin<sup>2</sup> field. Data may be obtained in one of the YJHK bandpasses at typical resolutions of  $R \sim 3000$ . We used MOSFIRE in its default configuration. For instrumental details, we refer the

reader to Steidel et al. (2014). For our *H*-band observations, we used integration times of 120 s with ABAB dithering to improve sky subtraction. Reductions were performed with the MOSFIRE Data Reduction Pipeline (DRP) made available by the instrument team.<sup>102</sup> We chose to observe only in the *H* band because the density of high-priority targets for which we could expect to get a secure redshift was notably higher in this band. This is largely due to the limited wavelength range of a single bandpass in MOSFIRE multislit spectroscopy, the  $z \sim 1.5$  redshifts of high-priority targets, and the resulting observed-frame wavelengths of prominent emission lines.

#### 2.1.3. Low-resolution Imaging Spectrograph

LRIS (Oke et al. 1995) is a dual-beam optical spectrograph that can simultaneously cover from  $\sim 3200$  Å to 1 μm over a  $6 \times 7.8$  arcmin<sup>2</sup> field. Spectra with resolution in the range  $300 < R < 5000$  may be obtained using custom masks, which typically allow one to observe  $\sim 25$  galaxies at the same time. We used LRIS with the 400 groove mm<sup>−1</sup> blue grism blazed at 3400 Å and the 400 groove mm<sup>−1</sup> red grating blazed at 8500 Å, with the D560 dichroic. Our choice of blue grism gives high sensitivity at bluer wavelengths where spectral features are likely to be found for objects with photometric redshifts of  $z \sim 1.5$ –3, while the red coverage allows for the detection of [O II]λ3727 for sources out to  $z \sim 1.6$ . Note that the blue camera was inoperative on the night of UT 2019 January 2. The LRIS spectra were reduced using the IRAF-based BOGUS software developed by D. Stern, S. A. Stanford, and A. Bunker, and flux calibrated using observations of standard stars from Massey & Gronwall (1990) observed on the same night using the same instrument configuration.

### 2.2. Redshift Determination

We refer the reader to M17 for a detailed description of the redshift determination procedure, as well as the quality flags and failure codes we adopt. To summarize, each observed source was assessed independently by two coauthors to determine the redshift and associated quality flag ( $Q = 0$ –4) defined as follows:

1.  $Q = 4$ : A quality flag of 4 indicates an unambiguous redshift identified with multiple features or the presence of the split [O II]λ3727 doublet.
2.  $Q = 3.5$ : A quality flag of 3.5 indicates a high-confidence redshift based on a single line, with a remote possibility of an incorrect identification. An example might be a strong, isolated emission line identified as  $H\alpha$ , where other identifications of the line are highly improbable due to the lack of associated lines or continuum breaks. This flag is typically only adopted for LRIS and MOSFIRE spectra; single line redshifts in DEIMOS spectra are usually the OII doublet, which is split by the DEIMOS spectral resolution.
3.  $Q = 3$ : A quality flag of 3 indicates a high-confidence redshift with a low probability of an incorrect identification. An example might be the low signal-to-noise ratio (S/N) detection of an emission line, possibly corrupted by telluric emission or absorption, identified as [O II] λ3727, but where the data quality is insufficient to clearly resolve the doublet.

<sup>102</sup> <https://keck-datareductionpipelines.github.io/MosfireDRP/>

**Table 2**  
List of Observing Nights

UT Date	Code	Instrument	# Masks	Notes
2017 December 11	N30-D	DEIMOS	2.3	light cirrus; 0 <sup>h</sup> 7–1 <sup>m</sup> 3 seeing
2017 December 12	N31-D	DEIMOS	2.3	0 <sup>h</sup> 8–2 <sup>m</sup> 3 seeing
2017 December 13	N32-D	DEIMOS	3.3	intermittent clouds
2018 January 11	N33-D	DEIMOS	2	thin cirrus to thick clouds; seeing 1 <sup>m</sup> 1–2 <sup>m</sup> 1
2018 January 12	N34-D	DEIMOS	3	thin cirrus; 1 <sup>m</sup> 3–1 <sup>m</sup> 7 seeing
2018 February 9	N35-M	MOSFIRE	4	1 <sup>m</sup> 0–1 <sup>m</sup> 3 seeing
2018 February 28	N36-M	MOSFIRE	2	thick clouds 1st half; clear and 0 <sup>h</sup> 8 seeing 2nd half
2018 April 3	N37-M	MOSFIRE	...	lost to weather: fog/humidity
2018 April 5	N38-M	MOSFIRE	...	lost to weather: fog/humidity
2018 April 6	N39-M	MOSFIRE	...	lost to weather: fog/humidity
2018 December 6 (first half)	N40-L	LRIS	2	clear, windy; 0 <sup>h</sup> 8–1 <sup>m</sup> 0 seeing
2018 December 7	N41-L	LRIS	4	clear; 0 <sup>h</sup> 8–1 <sup>m</sup> 2 seeing
2018 December 31 (first half)	N42-L	LRIS	2	lost to weather: thick clouds, 2 <sup>m</sup> 0 seeing
2019 January 1 (first half)	N43-L	LRIS	3.5	clear; 1 <sup>m</sup> 0–1 <sup>m</sup> 5 seeing
2019 January 2	N44-L	LRIS	5	clear; 1 <sup>m</sup> 0–1 <sup>m</sup> 4 seeing; blue camera inoperative
2019 January 29	N45-L	LRIS	1	high wind; poor seeing; closed early
2019 January 30 (first half)	N46-L	LRIS	1	lost to weather: ice/wind/moisture
2019 December 11	N47-M	MOSFIRE	4	clear; 0 <sup>h</sup> 6 seeing
2019 December 31	N48-M	MOSFIRE	4	clear; 0 <sup>h</sup> 6–1 <sup>m</sup> 0 seeing
2020 January 19	N49-M	MOSFIRE	4	clear; 0 <sup>h</sup> 5 seeing
2020 October 18 (2nd half)	N50-L	LRIS	2	partly cloudy; 1 <sup>m</sup> 0 seeing
2020 October 19 (2nd half)	N51-L	LRIS	2	partly cloudy; 1 <sup>m</sup> 0 seeing

- $Q = 2/1$ : A quality flag of 2 indicates a reasonable guess, while a quality flag of 1 indicates a highly uncertain guess. Sources with these low-confidence redshifts are not included in the data release.
- $Q = 0$ : A quality flag of 0 indicates that no redshift could be identified. As described above, a code indicating the cause of the redshift failure is assigned in place of the redshift.

These flags are meant to be qualitative indicators. Sample spectra associated with each flag are given in Figure 2 of DR1. Any conflicting assessments were reconciled through a joint review of the spectra with the help of a third, independent reviewer. Sources for which we failed to identify a redshift were assigned  $Q = 0$  and a failure code to indicate the most likely reason:

- Code = –91: Insufficient S/N;
- Code = –92: Well-detected but no discernible features;
- Code = –93: Problem with the reduction;
- Code = –94: Missing slit (an extreme case of –93).

As in M17, we also investigated all  $Q = 4$  sources for which the spectroscopic redshift ( $z_s$ ) was highly discrepant from the expected photometric redshift ( $z_p$ ), where the latter is an SOM-based value calculated by combining individual photometric redshifts for all galaxies in an SOM cell, as determined with the template-fitting *Le Phare* code (Arnouts et al. 1999; Ilbert et al. 2006) to deep 30 band photometry. The results of this analysis are presented in Section 3.3.

### 3. Redshift Results

In this data release, we targeted 1959 objects in 50 masks; 18, 19, and 13 were observed with MOSFIRE, LRIS, and DEIMOS, respectively. We were able to obtain 435 redshifts with  $Q = 4$  and 174 redshifts with either  $Q = 3$  or 3.5 for the targeted objects; these numbers do not include serendipitous sources, which are discussed below. Thus the success rate for all high-confidence redshifts ( $Q \geq 3$ ) is 31% of targeted

objects. If we restrict the results to the data obtained in good conditions (32 masks), then we obtained high-confidence redshifts for 511 out of 1346 targeted objects for a success rate of 38%. We return to a discussion of the relatively low success rate in Section 3.4.

Combined with the previous DR1 and DR2 results, C3R2 now has obtained 5131 high-confidence spectroscopic redshifts. The updated redshift and magnitude distributions of the entire spectroscopic sample are presented in Figure 1. The newly reported DR3 redshifts are available in a machine-readable table, of which the first few lines are presented in Table 4 in the Appendix.

#### 3.1. Redshift Performance of the SOM Technique

Figures 2–4 show comparisons of the highest-quality ( $Q = 4$ ) spectroscopic redshifts with the predicted SOM photometric redshifts (i.e., prior to obtaining the new spectroscopic redshifts) for the new DR3 redshifts (Figure 2), for the combination of all three C3R2 data releases (Figure 3), and for a combination of all of our C3R2 redshifts with high-quality literature redshifts for only the COSMOS field (Figure 4). For the latter plot, we use the median of the *Le Phare* photometric redshifts (Laigle et al. 2016) of the galaxies in each cell as the SOM photometric redshift for all galaxies in a cell.

For the case of all of our high-quality C3R2 redshifts, as seen in Figure 3, we calculate the scatter, defined as the normalized median absolute deviation

$$\sigma_{\text{NMAD}} = 1.48 \times \text{median} \left( \frac{|z_p - z_s|}{1 + z_s} \right),$$

to be 0.023, the true outlier fraction (defined as the sources with  $|z_p - z_s|/(1 + z_s) > 0.15$ ) to be 0.041, and the bias,

$$\text{mean} \left( \frac{z_p - z_s}{1 + z_s} \right),$$

**Table 3**  
List of Observed Slit Masks

Mask ID / Name	Night(s)	R.A. (J2000)	Decl. (J2000)	P.A. (deg)	Exposure (s)	# Spectra (targets / $Q = 4$ / $Q \geq 3$ serendips)
17B-D094 / VVDS32	N30-D	02 <sup>h</sup> 27 <sup>m</sup> 22 <sup>s</sup> .01	-04° 48' 02"	+90.0	15 × 1200	82 / 17 / 2
17B-D095 / COSMOS-1hr1	N30,N31,N32-D	09 <sup>h</sup> 59 <sup>m</sup> 44 <sup>s</sup> .20	+02° 35' 17"	+90.0	3 × 1200	80 / 29 / 2
17B-D096 / COSMOS-3hr1	N30-D	09 <sup>h</sup> 58 <sup>m</sup> 25 <sup>s</sup> .20	+01° 43' 16"	+90.0	11 × 1200	81 / 53 / 7
17B-D097 / VVDS33	N31-D	02 <sup>h</sup> 25 <sup>m</sup> 40 <sup>s</sup> .40	-04° 03' 11"	+90.0	15 × 1200	77 / 16 / 3
17B-D098 / COSMOS-3hr2	N31-D	09 <sup>h</sup> 59 <sup>m</sup> 57 <sup>s</sup> .50	+01° 55' 45"	+90.0	11 × 1200	88 / 6 / 0
17B-D099 / VVDS34	N32-D	02 <sup>h</sup> 25 <sup>m</sup> 30 <sup>s</sup> .10	-04° 31' 58"	+90.0	9 × 1200	79 / 16 / 3
17B-D100 / VVDS27	N32-D	02 <sup>h</sup> 25 <sup>m</sup> 19 <sup>s</sup> .50	-04° 42' 41"	+90.0	3 × 1200	80 / 24 / 2
17B-D101 / COSMOS-3hr3	N32-D	10 <sup>h</sup> 00 <sup>m</sup> 36 <sup>s</sup> .70	+01° 39' 47"	+90.0	12 × 1200	84 / 34 / 5
17B-D102 / VVDS35	N33-D	02 <sup>h</sup> 24 <sup>m</sup> 32 <sup>s</sup> .30	-04° 30' 23"	+90.0	9 × 1200	81 / 31 / 1
17B-D103 / COSMOS-3hr4	N33-D	09 <sup>h</sup> 58 <sup>m</sup> 44 <sup>s</sup> .10	+01° 51' 40"	+90.0	4 × 1200	81 / 2 <sup>a</sup> / 0
17B-D104 / VVDS36	N34-D	02 <sup>h</sup> 24 <sup>m</sup> 33 <sup>s</sup> .30	-04° 37' 23"	+90.0	10 × 1200	78 / 25 / 4
17B-D105 / COSMOS-3hr6	N34-D	10 <sup>h</sup> 00 <sup>m</sup> 28 <sup>s</sup> .80	+02° 42' 00"	+90.0	9 × 1200	85 / 33 / 0
17B-D106 / COSMOS-3hr7	N34-D	10 <sup>h</sup> 01 <sup>m</sup> 12 <sup>s</sup> .00	+01° 39' 00"	+90.0	9 × 1200	89 / 35 / 2
18A-M107 / VVDS-H-m1	N35-M	02 <sup>h</sup> 24 <sup>m</sup> 26 <sup>s</sup> .50	-04° 05' 27"	+55.0	30 × 120	23 / 0 / 0
18A-M108 / COSMOS-H-2hr5	N35-M	10 <sup>h</sup> 01 <sup>m</sup> 28 <sup>s</sup> .50	+01° 52' 20"	+45.0	60 × 120	15 / 0 / 4
18A-M109 / COSMOS-H-2hr1	N35-M	09 <sup>h</sup> 59 <sup>m</sup> 15 <sup>s</sup> .70	+02° 43' 05"	+5.0	60 × 120	18 / 9 / 3
18A-M110 / COSMOS-H-2hr3	N35-M	10 <sup>h</sup> 01 <sup>m</sup> 11 <sup>s</sup> .80	+02° 19' 00"	-10.0	60 × 120	19 / 1 / 0
18A-M111 / EGS_H_1hr1	N36-M	14 <sup>h</sup> 16 <sup>m</sup> 30 <sup>s</sup> .00	+52° 54' 00"	+15.0	44 × 120	20 / 3 / 0
18A-M112 / EGS_H_1hr2	N36-M	14 <sup>h</sup> 18 <sup>m</sup> 10 <sup>s</sup> .00	+52° 38' 30"	-25.0	46 × 120	15 / 0 / 0
18B-L113 / VVDS_1h1	N40-L	02 <sup>h</sup> 25 <sup>m</sup> 19 <sup>s</sup> .20	-04° 02' 38"	0.0	3 × 1200	29 / 2 / 0
18B-L114 / VVDS_3h1	N40-L	02 <sup>h</sup> 24 <sup>m</sup> 59 <sup>s</sup> .90	-04° 06' 47"	0.0	4 × 1200	28 / 2 <sup>a</sup> / 3
18B-L115 / VVDS_3h3	N41-L	02 <sup>h</sup> 25 <sup>m</sup> 57 <sup>s</sup> .12	-04° 49' 19"	0.0	9 × 1200	33 / 2 / 0
18B-L116 / VVDS_1h2	N41-L	02 <sup>h</sup> 24 <sup>m</sup> 26 <sup>s</sup> .32	-04° 10' 19"	0.0	4(5)×1200B(R)	32 / 7 / 0
18B-L117 / VVDS_1h3	N41-L	02 <sup>h</sup> 24 <sup>m</sup> 13 <sup>s</sup> .20	-04° 44' 13"	0.0	2 × 1200	29 / 0 / 0
18B-L118 / COSMOS_3h1	N41-L	10 <sup>h</sup> 01 <sup>m</sup> 28 <sup>s</sup> .32	+02° 21' 04"	-29.0	9 × 1200	31 / 0 / 0
18B-L119 / VVDS_1h4	N43-L	02 <sup>h</sup> 26 <sup>m</sup> 43 <sup>s</sup> .47	-04° 43' 41"	0.0	1 × 3600	32 / 0 / 0
18B-L120 / VVDS_1h5	N43-L	02 <sup>h</sup> 27 <sup>m</sup> 19 <sup>s</sup> .90	-04° 51' 40"	0.0	1 × 3600	31 / 3 / 1
18B-L121 / VVDS_1h6	N43-L	02 <sup>h</sup> 26 <sup>m</sup> 14 <sup>s</sup> .44	-04° 39' 38"	0.0	1 × 3600	31 / 0 / 1
18B-L122 / VVDS_3h2	N43-L	02 <sup>h</sup> 24 <sup>m</sup> 39 <sup>s</sup> .98	-04° 20' 20"	0.0	1 × 3600	29 / 0 <sup>a</sup> / 0
18B-L123 / VVDS_3h5	N44-L	02 <sup>h</sup> 27 <sup>m</sup> 31 <sup>s</sup> .33	-04° 43' 57"	0.0	8 × 1260	32 / 2 <sup>b</sup> / 0
18B-L124 / VVDS_1h7	N44-L	02 <sup>h</sup> 25 <sup>m</sup> 29 <sup>s</sup> .52	-04° 32' 10"	+9.0	3 × 1200	26 / 0 <sup>b</sup> / 0
18B-L125 / COSMOS_1h3	N44-L	10 <sup>h</sup> 01 <sup>m</sup> 36 <sup>s</sup> .00	+02° 02' 24"	0.0	3 × 1200	24 / 1 <sup>b</sup> / 0
18B-L126 / COSMOS_1h4	N44-L	10 <sup>h</sup> 00 <sup>m</sup> 19 <sup>s</sup> .20	+02° 42' 00"	0.0	3 × 1200	26 / 0 <sup>b</sup> / 0
18B-L127 / COSMOS_3h2	N44-L	10 <sup>h</sup> 00 <sup>m</sup> 27 <sup>s</sup> .33	+02° 23' 20"	0.0	5 × 1200	29 / 2 <sup>b</sup> / 0
19B-M128 / VVDS-H-3h1	N47-M	02 <sup>h</sup> 27 <sup>m</sup> 25 <sup>s</sup> .98	-04° 47' 17"	-60.0	78 × 120	23 / 6 / 1
19B-M129 / VVDS-H-2h2	N47-M	02 <sup>h</sup> 27 <sup>m</sup> 05 <sup>s</sup> .82	-04° 17' 31"	-70.0	63 × 120	21 / 2 / 2
19B-M130 / COSMOS-H-1h1	N47-M	10 <sup>h</sup> 01 <sup>m</sup> 22 <sup>s</sup> .26	+01° 51' 45"	-65.0	31 × 120	17 / 9 / 0
19B-M131 / COSMOS-H-3h3	N47-M	10 <sup>h</sup> 02 <sup>m</sup> 33 <sup>s</sup> .68	+02° 17' 56"	+10.0	51 × 120	16 / 7 / 4
19B-M132 / VVDS-H-3h3	N48-M	02 <sup>h</sup> 25 <sup>m</sup> 15 <sup>s</sup> .00	-04° 52' 35"	+55.0	80 × 120	24 / 6 / 3
19B-M133 / VVDS-H-2h3	N48-M	02 <sup>h</sup> 24 <sup>m</sup> 13 <sup>s</sup> .53	-04° 54' 40"	+350.0	40 × 120	21 / 1 / 0
19B-M134 / COSMOS-H-2h4	N48-M	10 <sup>h</sup> 01 <sup>m</sup> 29 <sup>s</sup> .88	+01° 48' 56"	+350.0	56 × 120	19 / 4 / 0
19B-M135 / COSMOS-H-2h2	N48-M	09 <sup>h</sup> 58 <sup>m</sup> 33 <sup>s</sup> .90	+02° 11' 07"	+350.0	55 × 120	17 / 8 / 2
19B-M136 / VVDS-H-1h3	N49-M	02 <sup>h</sup> 27 <sup>m</sup> 28 <sup>s</sup> .53	-04° 51' 52"	-10.0	20 × 120	22 / 1 / 3
19B-M137 / COSMOS-H-2h3	N49-M	10 <sup>h</sup> 00 <sup>m</sup> 16 <sup>s</sup> .79	+02° 41' 20"	+50.0	48 × 120	20 / 9 / 0
19B-M138 / COSMOS-H-3h1	N49-M	10 <sup>h</sup> 00 <sup>m</sup> 41 <sup>s</sup> .61	+01° 45' 50"	+119.0	80 × 120	18 / 6 / 2
19B-M139 / COSMOS-H-1h3	N49-M	09 <sup>h</sup> 58 <sup>m</sup> 05 <sup>s</sup> .91	+02° 27' 12"	+20.0	28 × 120	16 / 3 <sup>a</sup> / 0
20B-L128 / VVDS_1hr3	N50-L	02 <sup>h</sup> 25 <sup>m</sup> 23 <sup>s</sup> .92	-04° 14' 53"	-11.5	6 × 1200	25 / 6 / 0
20B-L129 / VVDS_1hr5	N50-L	02 <sup>h</sup> 26 <sup>m</sup> 11 <sup>s</sup> .88	-04° 06' 07"	+13.0	7 × 1200	26 / 6 / 3
20B-L130 / VVDS_1hr2	N51-L	02 <sup>h</sup> 24 <sup>m</sup> 38 <sup>s</sup> .10	-04° 56' 10"	+69.0	6 × 1200	29 / 5 / 3
20B-L131 / VVDS_1hr1	N51-L	02 <sup>h</sup> 24 <sup>m</sup> 37 <sup>s</sup> .28	-04° 14' 59"	-7.0	6 × 1200	28 / 2 / 2

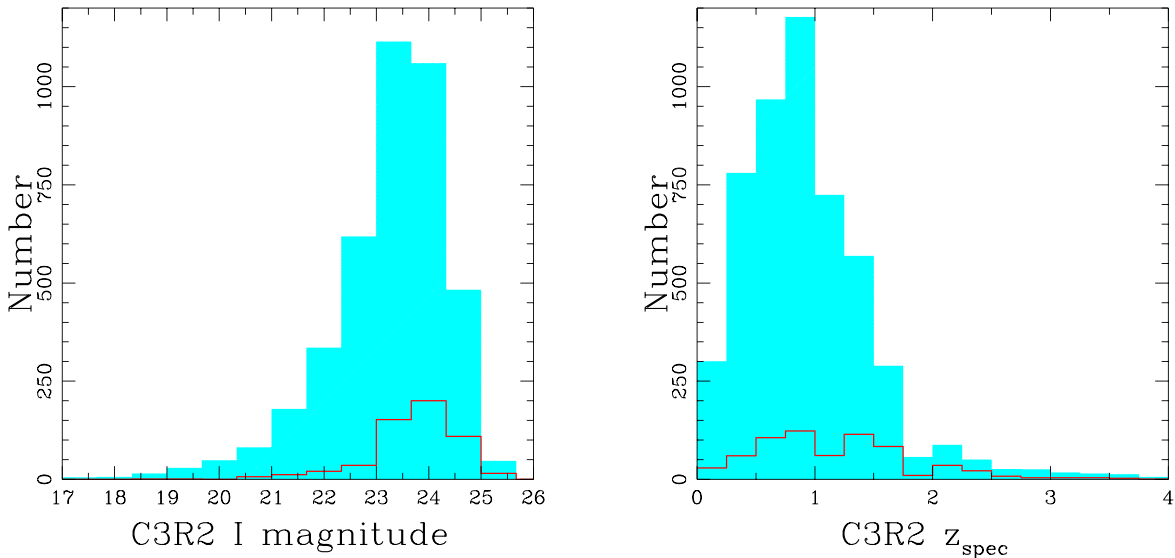
**Notes.** The “Night” column refers to the observing code in the second column of Table 2: night number, followed by letter indicating instrument used (D—DEIMOS, L—LRIS, M—MOSFIRE). R.A. and decl. refer to the mask center. The final column gives the total number of slitlets in mask, the total number of highest-quality ( $Q = 4$ ) redshifts measured, and the number of serendipitous sources with the highest-quality redshifts ( $Q = 4$ ).

<sup>a</sup> Exposure time significantly less than planned.

<sup>b</sup> No blue side data obtained.

to be  $-0.0025$ , where the scatter and bias are calculated after removing outliers. If we consider redshifts only in the range  $0.2 < z_p < 2.6$ , corresponding to the epoch that will be used by

Euclid for its cosmological weak-lensing survey, the bias is  $-0.0020$  (after removing outliers). If we include *good*-quality spectroscopic redshifts ( $Q = 3$  or  $3.5$ ) along with the  $Q = 4$



**Figure 1.** Magnitude and redshift histograms of the objects with  $Q \geq 3$  redshifts. The shaded cyan histograms represent objects in the first three C3R2 data releases, and the red histograms represent only objects in DR3.

redshifts from all three C3R2 data releases, then the scatter is 0.024, the true outlier fraction is 0.056, and the bias is  $-0.0018$  (after removing outliers) over the entire redshift range.

In Figure 4, which shows the results for the COSMOS field, we see that there is an increase in the fraction of outliers when all available high-quality spectroscopic redshifts are compared with the SOM photometric redshifts. We discuss the outliers in the DR3 sample in more detail in Section 3.3. More detailed discussion of the other samples from the literature is beyond the scope of this work; we plan to investigate the causes of the photometric and spectroscopic redshift outliers in the future. Once all of the outliers are removed, the bias is calculated to be  $-0.0020$  over the entire redshift range. This improvement relative to the case of only the C3R2 redshifts is likely due to the inclusion of brighter galaxies from the literature samples.

In M19, we compared C3R2 spectroscopic redshifts to photometric redshifts in the COSMOS field determined with the template-fitting code *Le Phare* (Arnouts et al. 1999; Ilbert et al. 2006) using the same photometry as used here to generate the SOM and SOM-based photometric redshifts. Considering the highest-quality ( $Q = 4$ ) redshifts from the three merged C3R2 data releases, these *Le Phare* photometric redshifts show a scatter  $\sigma_{\text{NMAD}}$  of 0.029, an outlier fraction of 0.038, and a bias (after removing outliers) of  $-0.013$ . While the scatter and outlier rates are similar to those from the SOM-based photometric redshifts (Figure 3), the *Le Phare* bias is substantially worse. It may be possible to combine an SOM-determined calibration with the *Le Phare* method in a two-step process to produce photometric redshifts for Euclid; although, it is important to remember the photometry that will be available for calculating Euclid photometric redshifts will not be as good as the existing 30 band photometry in the COSMOS field.

Figure 5 presents the SOMs based on all available high-confidence spectroscopic redshifts in C3R2. The fraction of cells with at least one spectroscopic redshift has increased from 49% in M15 to 84%, and these cells account for 91% of the galaxies. If we restrict consideration to only the  $0.2 < z_p < 2.6$  range of interest for the Euclid cosmological goals, then the fraction of

cells with at least one spectroscopic redshift is now 88%, and these cells account for more than 92% of the galaxies.

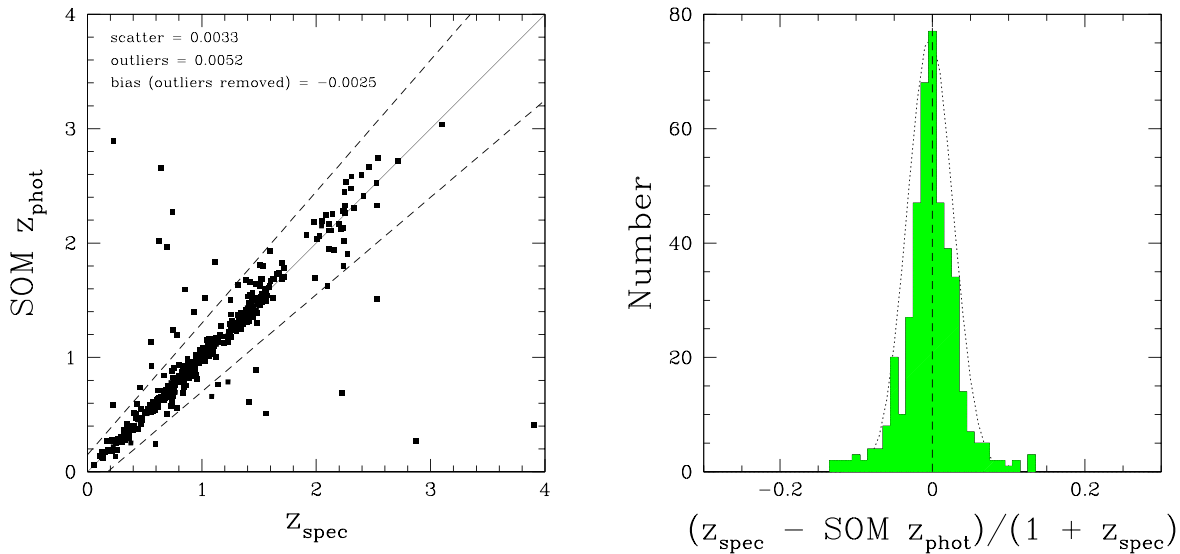
Figure 6 presents the distributions of the calibrated and uncalibrated cells in the SOM map as a function of magnitude and of redshift. We consider galaxy color cells to be “calibrated” if at least one galaxy in the cell has a high-confidence spectroscopic redshift, while “uncalibrated” cells are those for which no galaxy occupying that cell has a secure redshift measurement. These plots show that most of the uncalibrated cells are faint and located at  $z_p < 0.6$ . The next two sets of plots in Figure 7 further elucidate the nature of these populations. In particular we see that most of the uncalibrated galaxies at faint magnitudes and low photometric redshifts are likely to be passive rather than active, where the rest-frame UVJ colors, estimated with the COSMOS2015 (Laigle et al. 2016) data, are used to determine if a galaxy is active or passive. It is worth noting that the photometric redshifts based on Euclid and LSST data for faint, passive galaxies at low redshifts are expected to be robust.

### 3.2. Serendipitous Sources

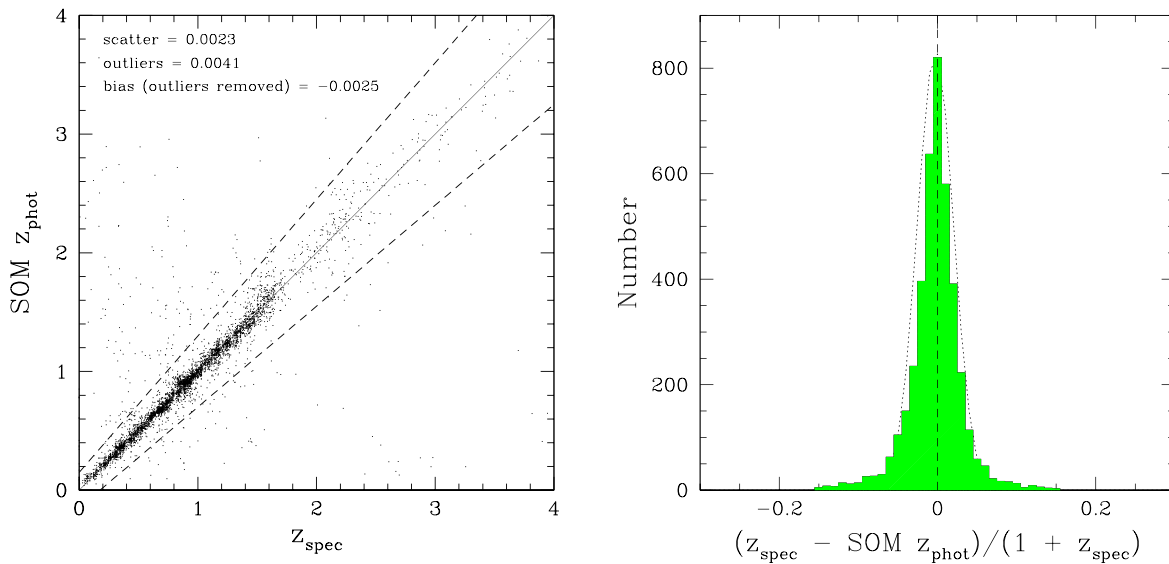
In addition to the targeted objects, we also determined  $Q = 4$  redshifts for 48 serendipitous sources in DR3, and an additional 19  $Q = 3$  or 3.5 redshifts. Some of these serendipitous sources turned out to be in the target catalogs; these are listed in the redshift table with the ID of the matched object and a prefix of “serendip.” A few of the serendipitous sources are located very close (e.g., within  $\sim 1''$ ) to the targeted galaxy. In such situations, the catalog photometry is probably affected by blending of the sources within the aperture used on the images to measure brightnesses. This might not be the case for Euclid VIS images, where the detection+segmentation will be used to perform template-fitting photometry. The extent to which apparently secure redshifts could affect photometric redshift calibration still needs further investigation.

### 3.3. Assessing the Photometric Redshift Outliers

As a final step in the analysis, we identified all sources in DR3 with secure redshifts (i.e., quality flag  $Q \geq 3$ ) where the



**Figure 2.** The performance of the SOM method on the DR3 redshifts. The plots show all of the objects for which high-confidence ( $Q = 4$ ) spectroscopic redshifts were obtained. The true outlier fraction is 0.052, the scatter is 0.033, and after removing outliers, the bias is  $-0.0025$  over the entire redshift range. In the left panel, those objects that are outliers are located outside of the dashed lines. The dotted line in the right panel is a Gaussian centered at zero with  $\sigma = 0.033$ . Note that the photometry used to place galaxies on the SOM here is generally deeper than that in the Euclid wide survey, so these results are not fully representative of what Euclid will achieve.



**Figure 3.** Same as in Figure 2, except here we show the performance of the SOM method on the redshifts published by all three C3R2 data releases. The true outlier fraction for this case is 0.041, the scatter is 0.023, and after removing outliers, the bias is  $-0.0025$ . The dotted line in the right panel is a Gaussian centered at zero with  $\sigma = 0.023$ .

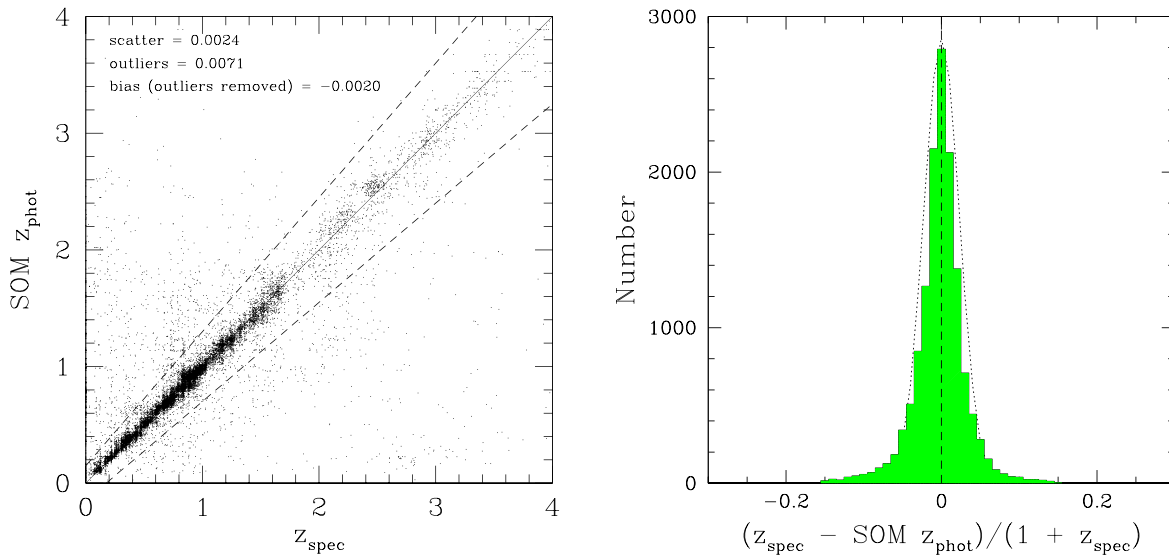
spectroscopic redshift was discrepant from the SOM-based redshift by at least 15% (i.e.,  $|z_s - z_{\text{SOM}}|/(1 + z_s) > 0.15$ ). At least two coauthors independently assessed each of the 80 such cases identified.

For three galaxies, this reanalysis lowered the spectroscopic quality flag to  $Q < 3$ . For the majority of the remaining cases (41/70; 59%), we revised the assigned spectroscopic redshift based on a reconsideration of the data and the photometric redshift. Most of these cases were one-line redshifts initially assigned a quality flag  $Q = 3$  observed with either MOSFIRE (with its limited spectral range) or LRIS. Typically, a target had a single faint emission line detected, initially assumed to be, for example,  $H\alpha$ . However, if the line had a different identification, such as [O III], the spectroscopic redshift matched the SOM-based redshift. In all such cases, corroborating lines were

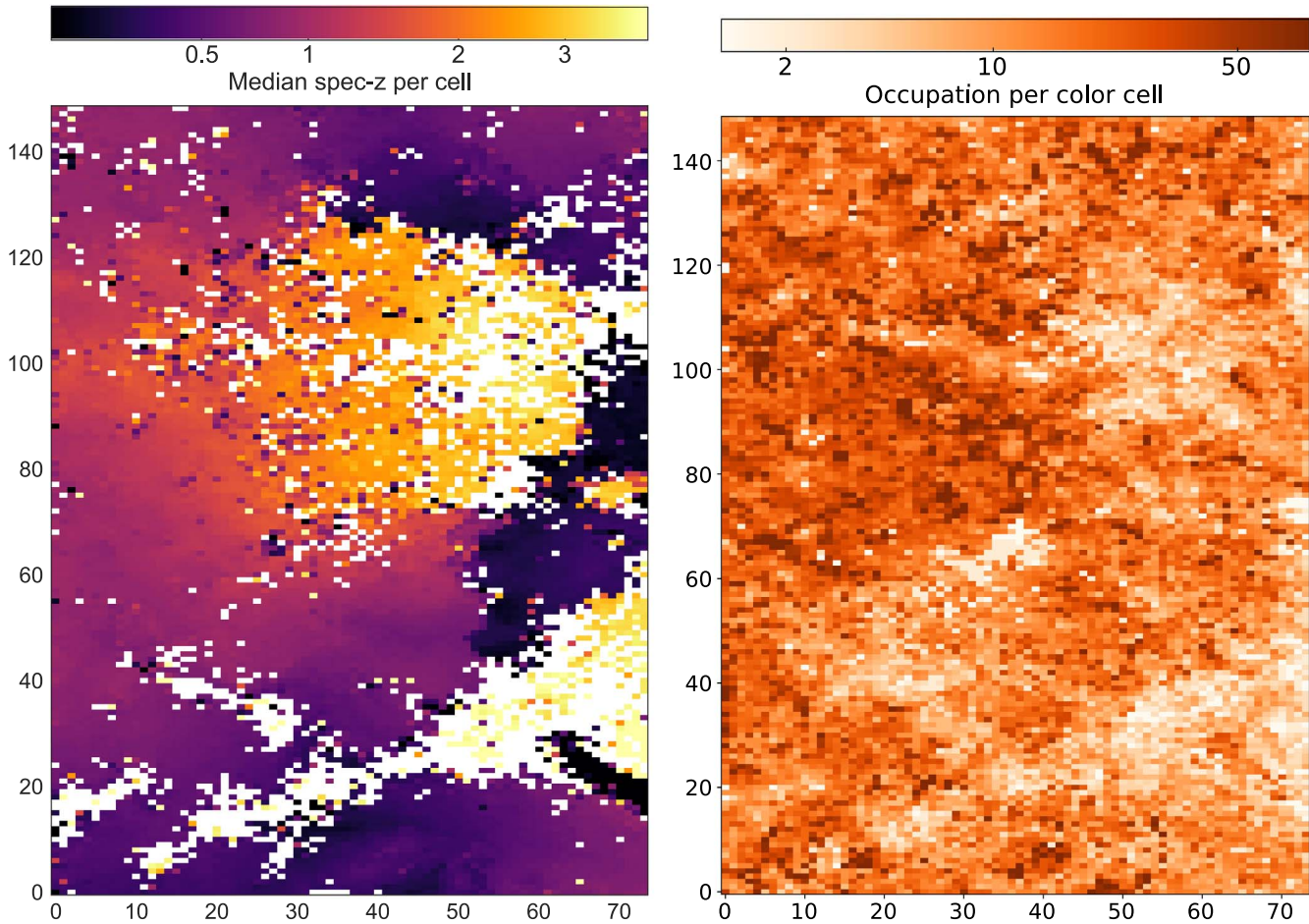
searched for, but generally were not found, either due to the limited spectral coverage or strong sky lines. In another two cases, this exercise made us realize that we had misidentified a targeted galaxy with a serendipitous galaxy. In only one case (out of the entire DR3 sample of 635 high-quality redshifts) was a  $Q = 4$  redshift corrected by this exercise. In that case, both initial spectroscopic redshift assessments matched, though a reanalysis clearly showed that redshift to be incorrect.

The remaining outliers were approximately evenly split between three categories. For approximately one-third of the outliers, the spectroscopic redshift was discrepant from the SOM-based photometric redshift but matched the *Le Phare* photometric redshift, suggesting an issue with the SOM-based redshift (e.g., potentially an SOM cell with a bimodal photometric redshift distribution). In fact, for two out of 11





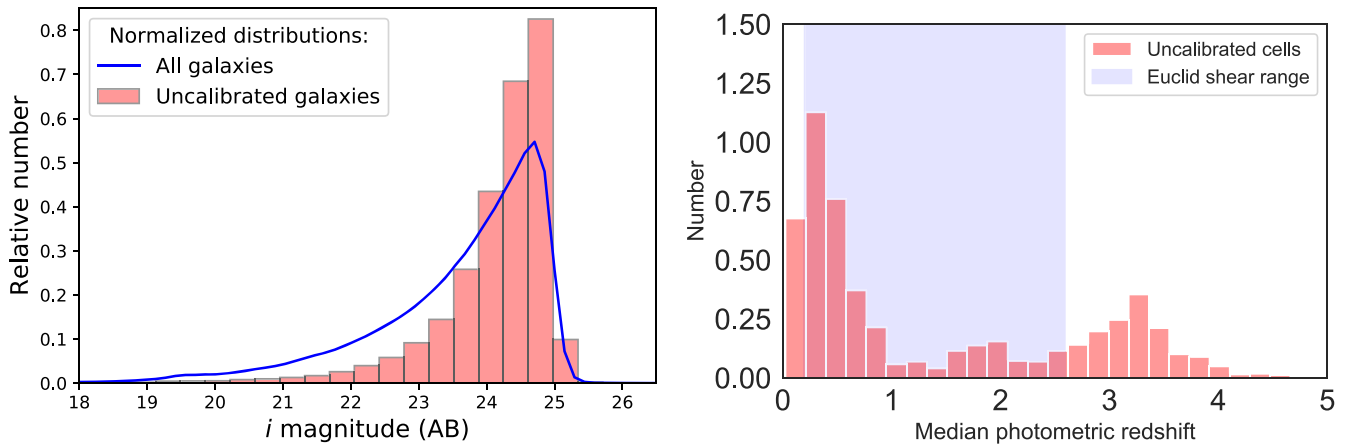
**Figure 4.** Same as in Figure 2, except here we show the performance of the SOM method only in the COSMOS field using all high-quality ( $Q = 4$ ) redshifts from C3R2 combined with all high-quality spectroscopic redshifts available from the literature. For this case, the true outlier fraction is 0.071, the scatter is 0.024, and after removing outliers, the bias is  $-0.0020$ . The dotted line in the right panel is a Gaussian centered at zero with  $\sigma = 0.024$ .



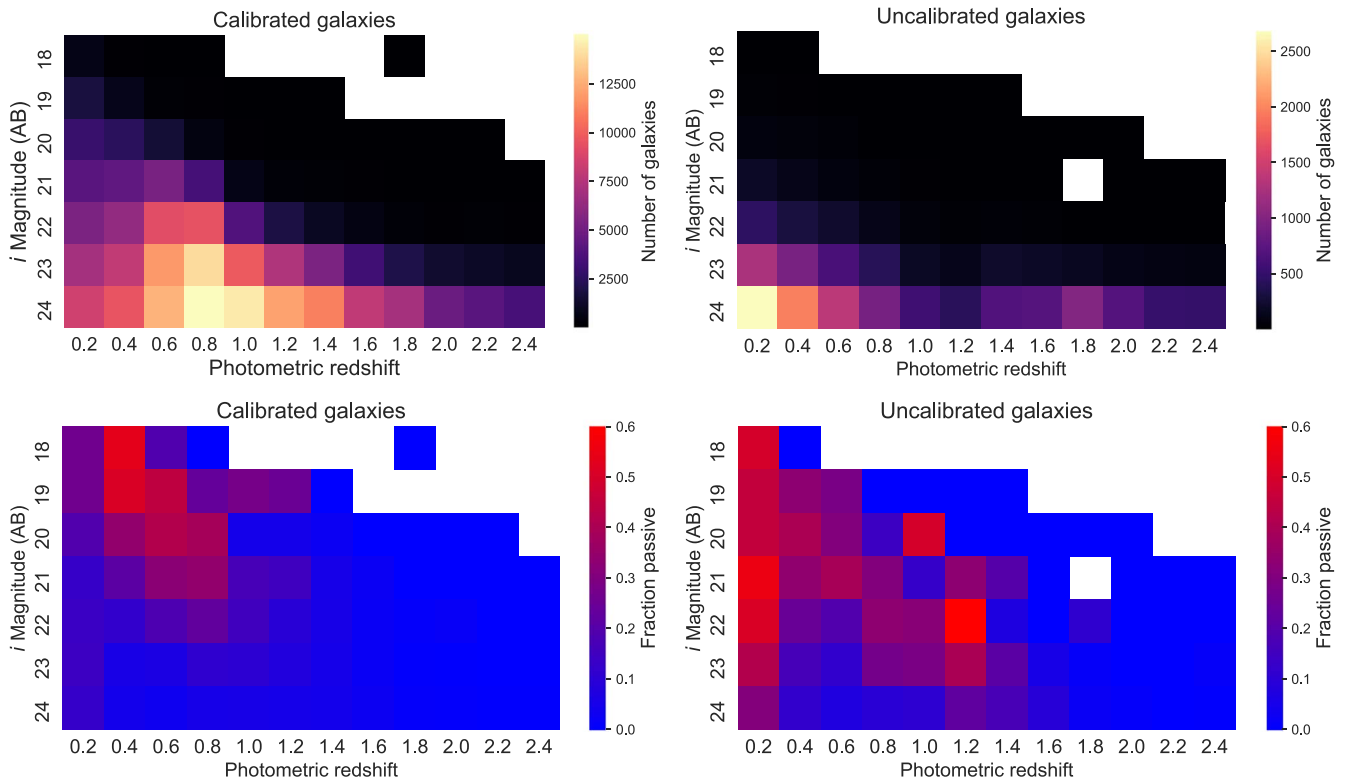
**Figure 5.** Left: the redshift coverage of the SOM color space based on all of the  $Q \geq 3$  redshifts in C3R2. Right: the current occupation density of the SOM, including the COSMOS and VVDS fields. Approximately 88% of color cells are covered in the range  $0.2 < z < 2.6$ , and 92% of galaxies live within calibrated cells. The axes are  $(x, y)$  indices into the SOM and have no (direct) physical meaning, though a given index corresponds to a particular galaxy spectral energy distribution.

of these cases, we find that the SOM cell of the galaxies is on a clear “caustic” separating low- and high-redshift galaxies with similar *Euclid* colors, while the others tend to be in cells with

intrinsically high redshift dispersion compared with most parts of the color space. It is plausible that the 30 band COSMOS data is capable of resolving these degeneracies, hence, the



**Figure 6.** Left: the histogram of the  $i$ -band median magnitudes of cells in the SOM, with all cells shown by the blue curve and the uncalibrated cells (those that do not yet have spectroscopic redshifts) in red. Right: the photometric redshift histogram of the uncalibrated cells in the SOM. The blue shading delineates the redshift range of the objects that will be used for Euclid shear measurements.



**Figure 7.** Top left: the distribution of “calibrated” galaxies (i.e., those in SOM cells with at least one spectroscopic redshift) in the plane of redshift and  $i$ -band magnitude. Top right: the same plot as in the top left panel, but for currently “uncalibrated” galaxies. Bottom left: the relative fractions of passive galaxies in the calibrated set as a function of magnitude and redshift, as estimated from rest-frame UVJ colors. Bottom right: the same plot but for the uncalibrated galaxies. It is clear from these figures that the currently uncalibrated galaxies are fainter and tend to be more quiescent than the calibrated galaxies. In all panels, the analysis uses galaxies in the COSMOS2015 sample, and the white areas represent no galaxies in a bin.

better *Le Phare* photometric redshift estimates. On the other hand, a detailed examination of the Hubble Advanced Camera for Surveys imaging of these sources showed that a majority seem to be affected by blending to some degree, which may also contribute to the incorrect SOM-based photometric redshifts.

Another third of the outliers are due to blends. While the input source catalog is based on ground-based, multiband photometry, in several cases Hubble imaging clearly resolved the source into two close galaxies, thereby compromising the photometric redshift. While this was easily assessed in the

COSMOS field, such higher-resolution imaging was not available for most of the other target fields. Finally, the remaining third of targets were deemed true outliers, where the spectroscopic redshift was deemed solid (e.g., generally  $Q=4$ ), but did not match either the SOM-based or the *Le Phare* photometric redshift. We expect that several of these cases are also unidentified blends where Hubble imaging was not available, highlighting a likely concern for future cosmology projects, both in developing the “gold sample” of calibration photometric redshifts, as well as in assigning photometric redshifts to the large sample of observed objects.

It is worth noting that near-Hubble quality imaging will in fact be available for deblending algorithms to use in creating photometric catalogs for Euclid.

### 3.4. Spectroscopic Failures

While we do not expect to spectroscopically confirm all targeted objects, it is useful to examine the possible failure modes that led to not being able to determine a redshift. Neglecting the obvious explanations having to do with poor seeing, clouds, shortened exposures, and/or instrument failure (one LRIS night, primarily targeting Lyman-break galaxies, was conducted without the blue side), there were 32 masks out of the 50 observed that were observed in at least good conditions. These 32 good masks had 1346 targeted objects, of which no secure redshift was obtained for 835 sources. Because the COSMOS catalog has the best photometry and so should have the best predicted exposure times, we restrict our investigation of the failures to only the 597 objects observed in the COSMOS field. The distributions of the good COSMOS masks by instrument are 10, 5, and 1 for MOSFIRE, DEIMOS, and LRIS, respectively. There were 70 objects for which a low-confidence ( $Q = 1$  or  $2$ ) redshift was determined, 162 with a failure code of  $-91$  (too faint), 39 with a failure code of  $-92$  (reasonable detection but no redshift), and 25 with a code of  $-93$  or  $-94$ , meaning there were issues with the data or reductions. The success rates by instrument are fairly similar for DEIMOS (43%) and MOSFIRE (37%) for the good COSMOS masks; there was only one mask done by LRIS in good conditions in the COSMOS field.

The success rate for the good LRIS observations in all three fields was lower compared to that obtained with DEIMOS and MOSFIRE. High-confidence redshifts were obtained for only 61 of the 413 targeted objects in the 14 good LRIS masks; no secure redshifts were obtained for 347 objects, and the reductions were problematic for five objects. The high-confidence redshift success rate for the good LRIS masks was 15%. The vast majority of the targeted objects in the LRIS masks are  $1.6 < z_p < 3$  Lyman-break galaxies, which often do not have any emission lines. In these cases, redshifts usually need to be determined from rest-frame UV absorption features, which are much more difficult to detect. This analysis suggests that for the  $z \sim 2$  Lyman-break galaxies, we should revise the way that we predict exposure times. Possible solutions to the low success rate with the LRIS observations of such galaxies would be to obtain deeper exposures, or to try observing such targets with MOSFIRE, which offers the possibility of detecting emission lines in the rest-frame optical.

## 4. Summary

We have presented 676 new high-confidence spectroscopic redshifts obtained by the C3R2 survey. Taken together with the previous two data releases, we have now collected and published 5130 high-confidence spectroscopic redshifts (and associated publicly released spectra) covering a wide range in galaxy type and redshift.<sup>103</sup>

While good progress has been made in covering the Euclid color space with spectroscopic redshifts, and the prospects for using these redshifts to calibrate photo- $z$  calculations for Euclid are promising, challenges remain. More work remains to be done to validate published samples of spectroscopic redshifts before they can also be used for this purpose. Also, the relatively low success rate of the observations (particularly with LRIS) of galaxies without emission lines emphasizes the technical challenges faced by redshift surveys such as C3R2 that are attempting to calibrate photometric redshifts with great precision down to the faint magnitudes of interest to upcoming space missions such as Euclid.

A significant portion of the research and manuscript preparation took place during the COVID-19 global pandemic. The authors would like to thank all of those who risked their lives as essential workers in order for us to safely continue our work from home. The authors also thank the referee whose comments helped improve the clarity and presentation of our results.

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## Appendix

Table 4 is published in its entirety in machine-readable format. A portion is shown here for guidance regarding its form and content. More detailed information is available in the header of the machine-readable table.









































<sup>103</sup> Redshifts and spectra released by C3R2 to date can be found at <https://sites.google.com/view/c3r2-survey/home> and <https://koa.ipac.caltech.edu/Datasets/C3R2>.

**Table 4**  
C3R2 DR3 Spectroscopic Redshifts

Object ID	R.A. (J2000)	Decl. (J2000)	Mask	Slit	<i>i</i> mag	<i>z</i>	<i>Q</i>	Instr.	Filename
VVDS-103460	02 27 32.10	-04 47 41.6	D094-VVDS32	15	24.09	0.5047	4.0	DEIMOS	spec1d.vvds32.015.VVDS-103460.fits
VVDS-98275	02 27 06.64	-04 48 14.5	D094-VVDS32	16	24.38	1.0881	4.0	DEIMOS	spec1d.vvds32.016.VVDS-98275.fits
VVDS-107542	02 27 35.45	-04 47 14.3	D094-VVDS32	25	24.34	0.8819	4.0	DEIMOS	spec1d.vvds32.025.VVDS-107542.fits
VVDS-121353	02 26 59.16	-04 45 39.2	D094-VVDS32	38	22.24	0.6153	4.0	DEIMOS	spec1d.vvds32.038.VVDS-121353.fits
VVDS-89949	02 27 29.90	-04 49 09.7	D094-VVDS32	61	24.02	0.8398	4.0	DEIMOS	spec1d.vvds32.061.VVDS-89949.fits
VVDS-125603	02 27 08.15	-04 45 04.7	D094-VVDS32	62	24.86	0.8691	4.0	DEIMOS	spec1d.vvds32.062.VVDS-125603.fits
VVDS-120554	02 26 59.91	-04 45 43.0	D094-VVDS32	63	24.78	0.9038	4.0	DEIMOS	spec1d.vvds32.063.VVDS-120554.fits
VVDS-118835	02 27 14.55	-04 45 53.9	D094-VVDS32	64	24.59	0.9076	4.0	DEIMOS	spec1d.vvds32.064.VVDS-118835.fits
VVDS-112287	02 27 46.85	-04 46 39.6	D094-VVDS32	67	24.86	1.2580	4.0	DEIMOS	spec1d.vvds32.067.VVDS-112287.fits
VVDS-97940	02 27 01.46	-04 48 17.7	D094-VVDS32	69	24.64	0.9837	4.0	DEIMOS	spec1d.vvds32.069.VVDS-97940.fits
VVDS-110496	02 26 53.92	-04 46 53.3	D094-VVDS32	71	24.82	1.1917	4.0	DEIMOS	spec1d.vvds32.071.VVDS-110496.fits
VVDS-111705	02 27 52.70	-04 46 44.1	D094-VVDS32	72	24.46	0.4468	4.0	DEIMOS	spec1d.vvds32.072.VVDS-111705.fits

(This table is available in its entirety in machine-readable form.)

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### References

- Aihara, H., Arimoto, N., Armstrong, R., et al. 2018, *PASJ*, 70, S4  
Arnouts, S., Cristiani, S., Moscardini, L., et al. 1999, *MNRAS*, 310, 540  
Benítez, N. 2000, *ApJ*, 536, 571  
Brammer, G. B., van Dokkum, P. G., & Coppi, P. 2008, *ApJ*, 686, 1503  
Carrasco Kind, M., & Brunner, R. J. 2013, *MNRAS*, 432, 1483  
Collister, A. A., & Lahav, O. 2004, *PASP*, 116, 345  
de Jong, J. T. A., Verdoes Kleijn, G. A., Boxhoorn, D. R., et al. 2015, *A&A*, 582, A62  
de Jong, J. T. A., Verdoes Kleijn, G. A., Erben, T., et al. 2017, *A&A*, 604, A134  
DES Collaboration, Abbott, T. M. C., Abdalla, F. B., et al. 2018, *PhRvD*, 98, 043526  
Desprez, G., Paltani, S., Coupon, J., et al. 2020, *A&A*, 644, A31  
Faber, S. M., Phillips, A. C., Kibrick, R. I., et al. 2003, *Proc. SPIE*, 4841, 1657  
Guglielmo, V., Saglia, R., Castander, F. J., et al. 2020, *A&A*, 642, 192  
Hildebrandt, H., Viola, M., Heymans, C., et al. 2017, *MNRAS*, 465, 1454  
Huterer, D., Takada, M., Bernstein, G., & Jain, B. 2006, *MNRAS*, 366, 101  
Ilbert, O., Arnouts, S., McCracken, H. J., et al. 2006, *A&A*, 457, 841  
Ilbert, O., Capak, P., Salvato, M., et al. 2009, *ApJ*, 690, 1236  
Kohonen, T. 1982, *Biol. Cybern.*, 43, 59  
Laigle, C., McCracken, H. J., Ilbert, O., et al. 2016, *ApJS*, 224, 24  
Lima, M., Cunha, C. E., Oyaizu, H., et al. 2008, *MNRAS*, 390, 118  
Ma, Z., Hu, W., & Huterer, D. 2006, *ApJ*, 636, 21  
Massey, P., & Gronwall, C. 1990, *ApJ*, 358, 344  
Masters, D., Capak, P., Stern, D., et al. 2015, *ApJ*, 813, 53  
Masters, D. C., Stern, D. K., Cohen, J. G., et al. 2017, *ApJ*, 841, 111  
Masters, D. C., Stern, D. K., Cohen, J. G., et al. 2019, *ApJ*, 877, 81  
McLean, I. S., Steidel, C. C., Epps, H. W., et al. 2012, *Proc. SPIE*, 8446, 84460J  
McQuinn, M., & White, M. 2013, *MNRAS*, 433, 2857  
Ménard, B., Scranton, R., Schmidt, S., et al. 2013, arXiv:1303.4722  
Morrison, C. B., Hildebrandt, H., Schmidt, S. J., et al. 2017, *MNRAS*, 467, 3576  
Newman, J. A. 2008, *ApJ*, 684, 88  
Newman, J. A., Cooper, M. C., Davis, M., et al. 2013, *ApJS*, 208, 5  
Oke, J. B., Cohen, J. G., Carr, M., et al. 1995, *PASP*, 107, 375  
Schmidt, S. J., Ménard, B., Scranton, R., Morrison, C., & McBride, C. K. 2013, *MNRAS*, 431, 3307  
Steidel, C. C., Rudie, G. C., Strom, A. L., et al. 2014, *ApJ*, 795, 165  
Troxel, M. A., MacCrann, N., Zuntz, J., et al. 2018, *PhRvD*, 98, 043528